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Time calibration with atmospheric muon tracks in the ANTARES neutrino telescope

S. Adrián-Martínez¹, A. Albert², M. André³, G. Anton⁴, M. Ardid¹, J.-J. Aubert⁵,
 B. Baret⁶, J. Barrios-Martí^{7,b}, S. Basa⁸, V. Bertin⁵, S. Biagi⁹, C. Bogazzi¹⁰,
 R. Bormuth^{10,11}, M. Bou-Cabo¹, M.C. Bouwhuis¹⁰, R. Bruijn^{10,12}, J. Brunner⁵,
 J. Busto⁵, A. Capone^{13,14}, L. Caramete¹⁵, J. Carr⁵, T. Chiarusi¹⁶, M. Circella¹⁷,
 R. Coniglione⁹, H. Costantini⁵, P. Coyle⁵, A. Creusot⁶, I. Dekeyser¹⁸,
 A. Deschamps¹⁹, G. De Bonis^{13,14}, C. Distefano⁹, C. Donzaud^{6,21}, D. Dornic⁵,
 D. Drouhin², A. Dumas²⁰, T. Eberl⁴, D. Elsässer²², A. Enzenhöfer⁴, K. Fehn⁴,
 I. Felis¹, P. Fermani^{13,14}, V. Flaminio^{23,24}, F. Folger⁴, L.A. Fusco^{16,26}, S. Galatà⁶,
 P. Gay²⁰, S. Geißelsöder⁴, K. Geyer⁴, V. Giordano²⁵, A. Gleixner⁴, R. Gracia-Ruiz⁶,
 J.P. Gómez-González^{7,aa}, K. Graf⁴, H. van Haren²⁷, A.J. Heijboer¹⁰, Y. Hello¹⁹, J.J.
 Hernández-Rey⁷, A. Herrero¹, J. Höfl⁴, J. Hofestädt⁴, C. Hugon^{28,29}, C.W James⁴,
 M. de Jong^{10,11}, M. Kadler²², O. Kalekin⁴, U. Katz⁴, D. Kießling⁴,
 P. Kooijman^{10,30,12}, A. Kouchner⁶, I. Kreykenbohm³¹, V. Kulikovskiy^{9,32},
 R. Lahmann⁴, G. Lambard⁷, D. Lattuada⁹, D. Lefèvre¹⁸, E. Leonora²⁵,
 S. Loucatos³³, S. Mangano⁷, M. Marcelin⁸, A. Margiotta^{16,26}, A. Marinelli^{23,24},
 J.A. Martínez-Mora¹, S. Martini¹⁸, A. Mathieu⁵, T. Michael¹⁰, P. Migliozzi³⁴,
 A. Moussa³⁵, C. Mueller²², M. Neff⁴, E. Nezri⁸, G.E. Pávālas¹⁵, C. Pellegrino^{16,26},
 C. Perrina^{13,14}, P. Piattelli⁹, V. Popa¹⁵, T. Pradier³⁶, C. Racca², G. Riccobene⁹,
 R. Richter⁴, K. Roensch⁴, A. Rostovtsev³⁷, M. Saldaña¹, D. F. E. Samtleben^{10,11},
 A. Sánchez-Losa⁷, M. Sanguineti^{28,29}, P. Sapienza⁹, J. Schmid⁴, J. Schnabel⁴,
 S. Schulte¹⁰, F. Schüssler³³, T. Seitz⁴, C. Sieger⁴, M. Spurio^{16,26}, J.J.M. Steijger¹⁰,
 Th. Stolarczyk³³, M. Taiuti^{28,29}, C. Tamburini¹⁸, A. Trovato⁹, M. Tselengidou⁴,
 C. Tönnes⁷, D. Turpin⁵, B. Vallage³³, C. Vallée⁵, V. Van Elewyck⁶, E. Visser¹⁰,
 D. Vivolo^{34,38}, S. Wagner⁴, J. Wilms³¹, J.D. Zornoza⁷, and J. Zúñiga⁷

^aMain corresponding author

^bCorresponding author

¹Institut d'Investigació per a la Gestió Integrada de les Zones Costaneres (IGIC) - Universitat Politècnica de València. C/ Paranimf 1, 46730 Gandia, Spain.

²GRPHE - Université de Haute Alsace - Institut universitaire de technologie de Colmar, 34 rue du Grillenbreit BP 50568 - 68008 Colmar, France

³Technical University of Catalonia, Laboratory of Applied Bioacoustics, Rambla Exposició,08800 Vilanova i la Geltrú,Barcelona, Spain

⁴Friedrich-Alexander-Universität Erlangen-Nürnberg, Erlangen Centre for Astroparticle Physics, Erwin-Rommel-Str. 1, 91058 Erlangen, Germany

⁵CPPM, Aix-Marseille Université, CNRS/IN2P3, CPPM UMR 7346, 13288, Marseille, France

⁶APC, Université Paris Diderot, CNRS/IN2P3, CEA/IRFU, Observatoire de Paris, Sorbonne Paris Cité, 75205 Paris, France

⁷IFIC - Instituto de Física Corpuscular, Parque Científico, C/Catedrático Jos Beltrán 2, E-46980 Paterna, Spain

⁸LAM - Laboratoire d'Astrophysique de Marseille, Pôle de l'Étoile Site de Château-Gombert, rue Frédéric Joliot-Curie 38, 13388 Marseille Cedex 13, France

- ¹⁶INFN - Sezione di Bologna, Viale Berti-Pichat 6/2, 40127 Bologna, Italy
- ²⁶Dipartimento di Fisica e Astronomia dell'Università, Viale Berti Pichat 6/2, 40127 Bologna, Italy
- ¹⁰Nikhef, Science Park, Amsterdam, The Netherlands
- ¹¹Huygens-Kamerlingh Onnes Laboratorium, Universiteit Leiden, The Netherlands
- ¹²Universiteit van Amsterdam, Instituut voor Hoge-Energie Fysica, Science Park 105, 1098 XG Amsterdam, The Netherlands
- ¹³INFN -Sezione di Roma, P.le Aldo Moro 2, 00185 Roma, Italy
- ¹⁴Dipartimento di Fisica dell'Università La Sapienza, P.le Aldo Moro 2, 00185 Roma, Italy
- ¹⁵Institute for Space Science, RO-077125 Bucharest, Măgurele, Romania
- ¹⁷INFN - Sezione di Bari, Via E. Orabona 4, 70126 Bari, Italy
- ⁹INFN - Laboratori Nazionali del Sud (LNS), Via S. Sofia 62, 95123 Catania, Italy
- ¹⁸Mediterranean Institute of Oceanography (MIO), Aix-Marseille University, 13288, Marseille, Cedex 9, France; Universit du Sud Toulon-Var, 83957, La Garde Cedex, France CNRS-INSU/IRD UM 110
- ¹⁹Géoazur, Université Nice Sophia-Antipolis, CNRS, IRD, Observatoire de la Côte d'Azur, Sophia Antipolis, France
- ²¹Univ. Paris-Sud , 91405 Orsay Cedex, France
- ²⁰Laboratoire de Physique Corpusculaire, Clermont Université, Université Blaise Pascal, CNRS/IN2P3, BP 10448, F-63000 Clermont-Ferrand, France
- ²²Institut für Theoretische Physik und Astrophysik, Universität Würzburg, Emil-Fischer Str. 31, 97074 Würzburg, Germany
- ²³NFN - Sezione di Pisa, Largo B. Pontecorvo 3, 56127 Pisa, Italy
- ²⁴Dipartimento di Fisica dell'Università, Largo B. Pontecorvo 3, 56127 Pisa, Italy
- ²⁵INFN - Sezione di Catania, Viale Andrea Doria 6, 95125 Catania, Italy
- ²⁷Royal Netherlands Institute for Sea Research (NIOZ), Landsdiep 4,1797 SZ 't Horntje (Texel), The Netherlands
- ²⁸INFN - Sezione di Genova, Via Dodecaneso 33, 16146 Genova, Italy
- ²⁹Dipartimento di Fisica dell'Università, Via Dodecaneso 33, 16146 Genova, Italy
- ³⁰Universiteit Utrecht, Faculteit Betawetenschappen, Princetonplein 5, 3584 CC Utrecht, The Netherlands
- ³¹Dr. Remeis-Sternwarte and ECAP, Universität Erlangen-Nürnberg, Sternwartstr. 7, 96049 Bamberg, Germany
- ³²Moscow State University,Skobeltsyn Institute of Nuclear Physics,Leninskie gory, 119991 Moscow, Russia
- ³³Direction des Sciences de la Matière - Institut de recherche sur les lois fondamentales de l'Univers - Service de Physique des Particules, CEA Saclay, 91191 Gif-sur-Yvette Cedex, France
- ³⁴INFN -Sezione di Napoli, Via Cintia 80126 Napoli, Italy
- ³⁵University Mohammed I, Laboratory of Physics of Matter and Radiations, B.P.717, Oujda 6000, Morocco
- ³⁶IPHC-Institut Pluridisciplinaire Hubert Curien - Université de Strasbourg et CNRS/IN2P3 23 rue du Loess, BP 28, 67037 Strasbourg Cedex 2, France
- ³⁷ITEP - Institute for Theoretical and Experimental Physics, B. Chermushkinskaya 25, 117218 Moscow, Russia
- ³⁸Dipartimento di Fisica dell'Università Federico II di Napoli, Via Cintia 80126, Napoli, Italy

Abstract

The ANTARES experiment consists of an array of photomultipliers distributed along 12 lines and located deep underwater in the Mediterranean Sea. It searches for astrophysical neutrinos collecting the Cherenkov light induced by the charged particles, mainly muons, produced in neutrino interactions around the detector. Since at energies of ~ 10 TeV the muon and the incident neutrino are almost collinear, it is possible to use the ANTARES detector as a neutrino telescope and identify a source of neutrinos in the sky starting from a precise reconstruction of the muon trajectory. To get this result, the arrival times of the Cherenkov photons must be accurately measured. A to perform time calibrations with the precision required to have optimal performances of the instrument is described. The reconstructed tracks of the atmospheric muons in the ANTARES detector are used to determine the relative time offsets between photomultipliers. Currently, this method is used to obtain the time calibration constants for photomultipliers on different lines at a precision level of 0.5 ns. It has also been validated for calibrating photomultipliers on the same line, using a system of LEDs and laser light devices.

1 Introduction

The ANTARES neutrino telescope [1] aims at the exploration of the high-energy Universe by using neutrinos as cosmic probes. One of the main goals of the experiment is the discovery of point-like sources of cosmic neutrinos. Typically, searches for sources of neutrinos are performed using the reconstructed directions of selected events. In this sense, a good angular resolution, i.e., the angle

between the direction of the reconstructed muon and of the neutrino, is of great importance to better discriminate between the signal and the background. At energies ~ 10 TeV the median value of the angular resolution in the ANTARES detector, as determined with simulations, is $\sim 0.4^\circ$ [2].

In the energy range of interest, the angular resolution is driven by the reconstruction accuracy. Since the reconstruction algorithm depends on the precise measurement of the photon arrival times on the photomultipliers (PMTs), an accurate detector time calibration is crucial to guarantee the best performance of the telescope. In particular, the precision in the measurement of the relative PMT timings is required to be ~ 1 ns [3].

In this paper, the methods to perform the time calibration of the detector at this level of accuracy and to determine the calibration constants used in the physics analyses of the ANTARES collaboration are presented. The method is based on the reconstructed trajectories of the down going atmospheric muons, which are the main source of physical triggers for a neutrino telescope. This calibration method does not require the physics data acquisition to be stopped and does not rely on additional electronic devices. It was initially implemented for the first ANTARES point-source analysis [4].

The paper is structured as follows. In section 2 a brief description of the ANTARES detector is presented. The calibration method is introduced in Section 3. In the fourth Section the reconstruction method used in the point-like source search, and in the majority of the physics analyses in ANTARES, is described. In Sections 5 and 6 the derivation of the calibration constants is described in detail. In Section 7 the results are discussed and compared to the results obtained with an independent system of calibration that uses an array of optical beacons [5]. A summary and the conclusions are given in Section 8.

2 The ANTARES detector

ANTARES is the first deep-sea neutrino telescope in operation in the world. Anchored at a depth of 2475 m in the Mediterranean Sea, 40 km off the coast of Toulon (France), the detector comprises 885 PMTs distributed on a three-dimensional array made up by 12 flexible lines, each of them 480 m long, which are arranged on an octagonal layout with an interline separation of 60-75 m. Each line holds triplets of PMTs and the control electronics boards needed for the power supply and data control at this level, at 25 vertical positions (storeys). The first storey of each line is placed 100 m above the seabed and the distance between adjacent storeys is 14.5 m. Each PMT (10" photocathode, model R7081-20 of Hamamatsu [6]), hosted in a high-pressure-resistant glass sphere together with high-voltage power supplies and internal calibration tools, constitutes an optical module (OM) [7].

The electronics boards are enclosed in a titanium container, the local control module. The main electronic components are the analogue ring sampler (ARS) circuits [8], which perform the digitisation of the electronic signals produced in the PMTs, providing information on amplitude, arrival time and signal shape. Two ARSs per OM work in a token ring configuration to minimise the acquisition dead time. Inside each ARS, an amplitude-to-voltage converter is used for the signal charge integration, and a time-to-voltage converter allows the measurement of hit times with a sub-nanosecond precision.

The 12 ANTARES detection lines are connected to a junction box that, through an electro-optical cable about 40 km long, links the detector to the shore station and provides the power and control signals needed for the full apparatus. All the signals passing a threshold condition, referred to as L0 trigger and typically set to 0.3 photoelectrons, activate the digitisation and the sending of the hit time and the integrated charge of the signal to the shore station [9]. They are then processed with a cluster of computers by applying a data-filter algorithm looking for space-time correlated hits

consistent with a specific physics signal to mitigate the effect of optical background. Finally, the detector includes a set of additional instruments for calibration purposes, such as hydrophones for the positioning system and optical beacons for timing calibration. In particular, the optical beacons system consists of a series of LED and laser light devices which are distributed along the detector to illuminate the PMTs.

3 Time calibration in ANTARES

An accurate muon track reconstruction requires a precise determination of the arrival times of the photons on the PMTs. In particular, all the PMTs have to be synchronised to within ~ 1 ns. The main uncertainties on the measured hit times come from the transit time spread of the PMTs and the optical properties of sea water. In the ANTARES experiment, the first time calibration [3] of the PMTs was performed at the integration laboratories before the deployment of each line. A laser device sending light to the PMTs through an optical fiber was used. After correcting for the propagation delay within the line, the time offsets between each ARS and an ARS chosen as a reference were calculated. These offsets defined a first set of time calibration parameters, usually referred to as the `ARS_T0` or the *intra line* calibration parameters. These values were stored in the ANTARES database and used for the prompt data analysis after deployment.

The value of `ARS_T0` measured in the laboratory may change after the line deployment due to, e.g., variations in the environmental conditions. Therefore, a system allowing for the *in situ* time calibration of the detector and the monitoring of the calibration constants is needed. The optical beacon system can be used to perform the relative time calibration of the optical sensors. There are two kinds of optical beacons: the LED beacons and the laser beacons. The LED optical beacons consist of 36 individual blue LED light sources arranged in groups of six on electronic boards placed side by side, forming a hexagonal prism, enclosed in a glass container. Four LED optical beacons are positioned regularly (in the 2nd, 9th, 15th and 21st storeys from the bottom) on each detector line. The non-consecutive LED optical beacon pairs within a line are fired simultaneously at maximum intensity during a calibration run. In total, 24 runs of 5 minutes duration are taken once per month. The runs are analysed later to check and update (if needed) the `ARS_T0` parameters.

In order to measure the relative offsets between the detector lines (*inter line* calibration), one laser beacon is placed at the bottom of a central line. This device consists of a Nd-YAG laser, which emits green light in high-intensity short pulses (FWHM ~ 1 ns) illuminating up to the 10th floor of every detector line. Currently, one laser beacon run of 5-10 minutes duration is taken every month. However, it is important to complement this system with the method proposed here, as mentioned in Sec. 6.

4 Track reconstruction

The muon track reconstruction algorithm [10] used by ANTARES both for physics analyses and for timing calibration consists of several fitting procedures that are performed consecutively. In the final step, estimates of all the parameters needed to describe the event trajectory are obtained by maximising the likelihood function

$$\mathcal{L}(\vec{d}, \vec{p}) = \prod P(t_i | t_i^{th}, \vec{d}, \vec{p}). \quad (1)$$

\mathcal{L} is the product of the likelihood, P , of the individual hits accounting for the probability of the time of the hits, where \vec{p} indicates a point in the muon track direction \vec{d} at a time t_0 and t_i indicates the measured hit time. t_i^{th} is the expected hit time given by

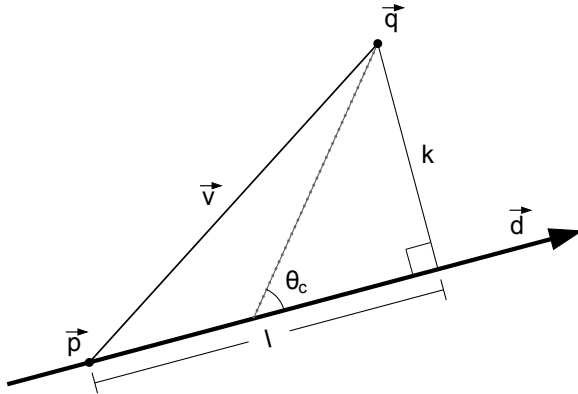


Figure 1: A muon passing near the detector along a direction \vec{d} and a position \vec{p} at a given time t_0 induces in the medium the production of Cherenkov photons. These photons are emitted at an angle θ_c with respect to the muon trajectory, \vec{d} . They are detected by the PMT located at position \vec{q} after travelling the light-grey line path. l and k indicate the components of the vector $\vec{v} = \vec{q} - \vec{p}$ in the parallel and perpendicular directions with respect to \vec{d} , respectively.

$$t_i^{th} = t_0 + \frac{1}{c} \left(l - \frac{k}{\tan \theta_c} \right) + \frac{1}{v_g} \left(\frac{k}{\sin \theta_c} \right), \quad (2)$$

derived from a given set of parameters (see Fig. 1), namely the position, \vec{q} , of the PMT which has detected the Cherenkov photon; the components of the vector $\vec{v} = \vec{q} - \vec{p}$ in the parallel, l , and perpendicular, k , directions to the muon track; the group velocity of light, v_g ; the angle of the Cherenkov light, θ_c ; and the speed of light in vacuum, c .

The likelihood in Eq. 1 can be expressed in terms of the probability density of the time residuals $r_i = t_i - t_i^{th}$. The probability density function used for the likelihood fit is obtained from simulations, and takes into account the contribution from hits arriving late due to Cherenkov emission by secondary particles and/or light scattering, and the effect of the transit time spread of the PMT. The probability of a hit being due to background is also accounted for.

The quality of the track fit is quantified by the parameter

$$\Lambda = \frac{\log(\mathcal{L})}{N_{DOF}} + 0.1(N_{comp} - 1), \quad (3)$$

which incorporates the maximum likelihood value \mathcal{L} , the number of degrees of freedom in the fit $N_{DOF} = N_{hits} - 5$, equal to the number of hits used minus the number of free parameters. N_{comp} is the number of compatible solutions found by the maximization algorithm, which in general takes large values for well-reconstructed tracks and small ones for badly reconstructed events. The Λ parameter is a useful tool to reject badly reconstructed events, mainly down going atmospheric muons misreconstructed as up going. Moreover, the uncertainty on the reconstructed track direction is used for further event selection. Assuming that the likelihood function near the fitted maximum follows a multivariate Gaussian distribution, the error on the zenith (β_θ) and azimuth (β_ϕ) angles can be derived from the likelihood fit covariance error matrix. From these errors, the parameter

$$\beta = \sqrt{\sin^2(\theta_{rec})\beta_\phi^2 + \beta_\theta^2}, \quad (4)$$

referred to as the angular error estimate, is obtained, where θ_{rec} is the reconstructed zenith angle. A cut on the angular error estimate is highly efficient for signal events and removes a large fraction of misreconstructed atmospheric muons [10].

5 Time calibration with muon tracks: method

The reconstructed trajectories of the down going muons produced in cosmic-ray interactions in the atmosphere are used for the time calibration of the ANTARES detector. These are collected at a rate of ~ 5 Hz. The method is an iterative procedure that uses the time residual distributions obtained for a subset of hits which are not used in the track fit. The complete procedure consists of the following steps:

1. A subset of hits (the *probe hits*), which were registered by the ARS or were within the line under investigation, is selected. These hits are not used in the track fit.
2. The muon trajectory is reconstructed using only the remaining hits (the *reco hits*).
3. Time residuals for the probe hits are calculated with respect to the fitted track.
4. A Gaussian function is fitted to the peak of the probe hit time residuals distribution (an example is given in Fig. 2), the mean value of which is taken as the time correction to be applied to the hit times for a new iteration of the method.

The exclusion of the probe hits guarantees that the resulting time residual distributions are not biased, since the muon track has not been fitted to minimise them. If the subset of probe hits includes only those collected on one single ARS, the so-called intra line calibration is performed. If it includes those collected on one detection line, the inter line calibration is performed. This procedure is repeated until the time corrections are small (~ 0.5 ns). The reconstructed muon tracks used for the time residuals are affected by the time corrections applied; at each new iteration, the number of high quality reconstructed events increases with the number of iterations until the method converges to an optimal solution. Fig. 3 shows how fast this solution converges to its noise limit of 0.5 ns when applying the method to determine the time offsets between different detection lines. In the same figure it is also shown how the median value of the fit quality distribution increases with the number of iterations of the method.

6 Results on the inter line timing

The inter line calibration corrects the time offsets with respect to a common reference for the full detector. Since at the integration laboratories the time calibration is only performed at the line level, this is the most important calibration performed off-shore. Indeed, the first hint of the existence of time offsets between the lines of the ANTARES detector was found when studying the distributions of the quality reconstruction parameter Λ . A large discrepancy between data and Monte Carlo simulations¹ was found for tracks crossing the detector diagonally and inducing hits on several lines, whereas nearly vertical tracks, the reconstruction of which is single-line dominated, showed better agreement.

¹In the ANTARES collaboration, the generation of atmospheric neutrinos is performed with the GENHEN package [11] using the Bartol model [12], while the atmospheric muons are produced with MUPAGE [13].

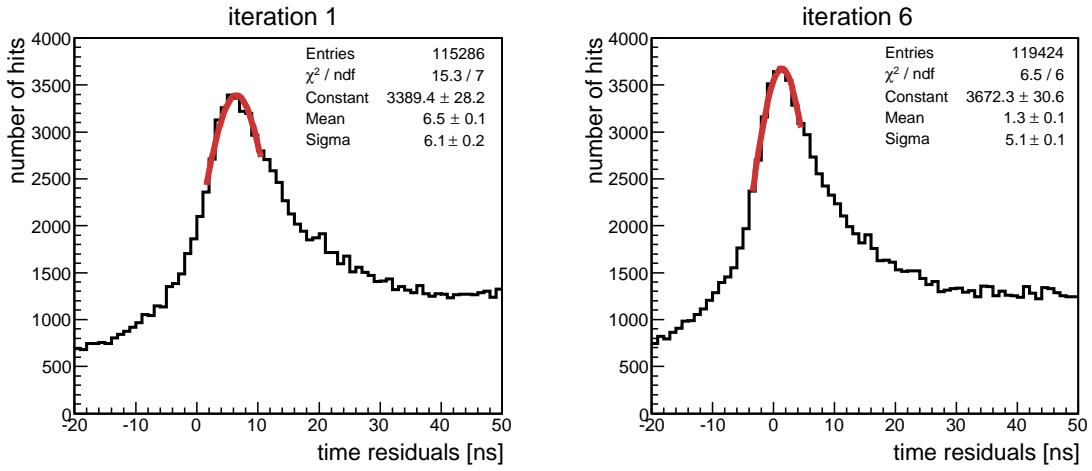


Figure 2: Distributions of the time residuals for hits detected with line 8 after a first iteration of the calibration method (left) and after applying six iterations on the same data sample (right). A Gaussian function is fitted to each distribution around the position of the peak. The RMS of the distribution decreases, and the number of entries increases, with increasing iterations, since the timing corrections applied at each new iteration have a positive effect on the event reconstruction.

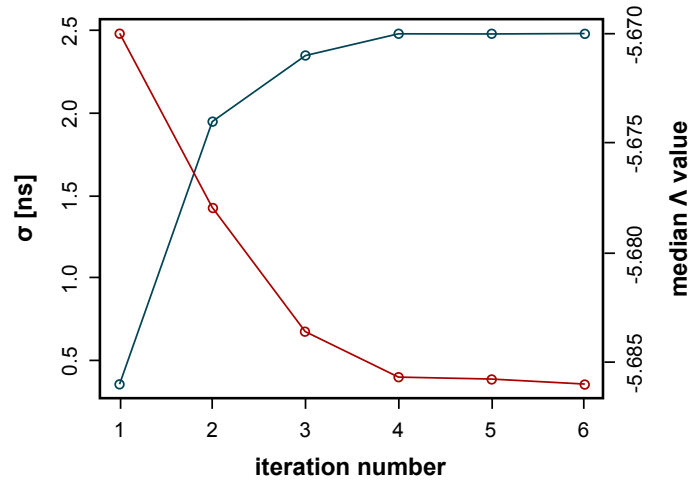


Figure 3: Width of the inter line distribution of time corrections as a function of the number of iterations of the procedure (red line and left-hand side scale). The median value of the fit quality parameter distribution of the events increases with the number of iterations, as expected (blue line and right-side scale).

Line	L1	L2	L3	L4	L5	L6	L7	L8	L9	L10	L11	L12
Offset (ns)	1.3	-3.9	-0.5	-2.1	-3.4	-1.3	0.6	4.9	0.3	-0.5	4.0	2.2

Table 1: Inter line offsets measured with the track residual method. These values are currently used for data processing.

In order to determine the time offsets, the method described in the previous section is here applied to data recorded with 12 detection lines in a few days of data-taking from March 2010. Other periods have been studied (November 2010, April 2011, March 2012) with no significant differences in the time offsets. Good quality events (only tracks fitted with $\Lambda > -6.0$) were used. Fig. 2 shows the time residuals distribution, after one (left) and after six (right) iterations of the method, for hits detected on line 8, which is one of the central lines of the ANTARES detector. The range of the Gaussian fit applied to these histograms was chosen in order to avoid the contribution from scattered photons and to match the most Gaussian-like region of the residuals distribution.

The inter line constants determined for the 12 detection lines are summarised in Table 1. These are the values normally used in ANTARES analyses.

6.1 Impact on the reconstruction and on the angular resolution.

After the application of the corrected inter line timing, an improvement in the quality of the reconstruction was observed with an increase of the number of events with high Λ values. The improvement was particularly important (see Fig. 4) for events reconstructed on several lines, while for vertically down going events, which are mainly single-line reconstructed, the improvement was smaller, as expected. The agreement of the distributions of the Λ parameter for data and simulated events improved after the timing calibration correction by $\sim 15\%$ for well-reconstructed tracks

In order to study the effect of the inter line time shifts on the detector resolution, offsets of the same amounts as the values listed in Table 1 were added to the hit times of simulated events and the reconstructed tracks compared to the results of reconstruction from the usual simulation. Fig. 5 shows the distribution of the angular difference, Ψ , between the reconstructed track direction and the generated one for the default Monte Carlo simulation and for events reconstructed using mis-calibrated lines. Adding the offsets results in a resolution degraded by about 40% degraded resolution for the selected events ($\Lambda > -5.4$, $\cos\theta_{rec} > 0.0$ and $\beta < 1^\circ$). The median value of the distribution worsens from 0.40° to 0.55° for the studied sample.

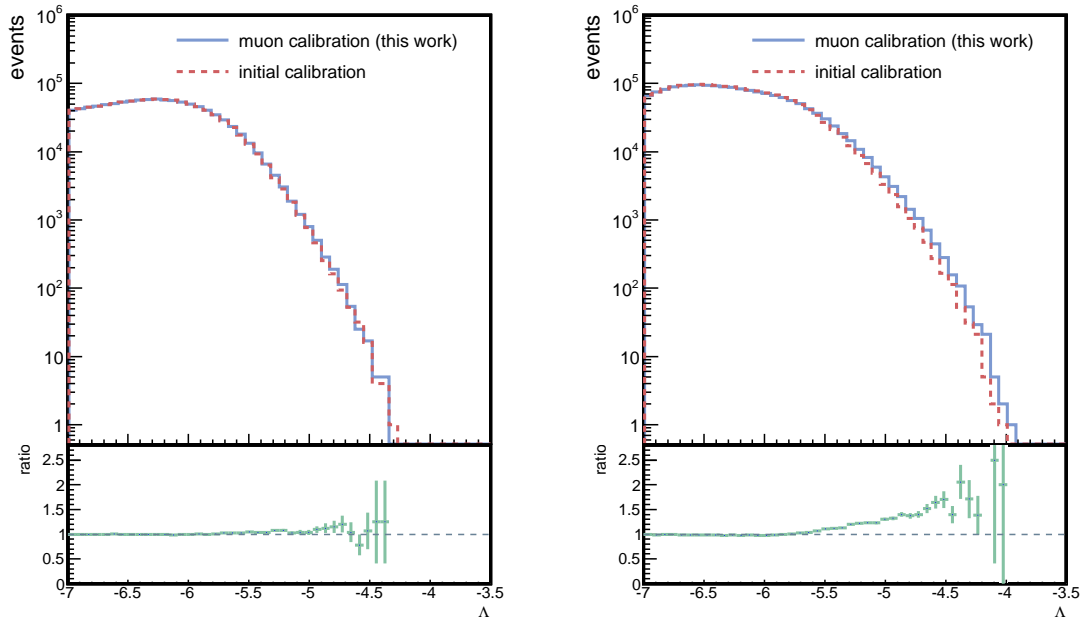


Figure 4: Distributions of the reconstruction quality parameter before and after correcting the inter line offsets for (left) vertically down going tracks, $\cos\theta < -0.9$, and (right) inclined tracks, $\cos\theta > -0.8$. The lower pads show the ratio between the number of events reconstructed after and before applying the inter line timing correction. For inclined tracks, the number of very well reconstructed tracks improves a factor 1.2-2.0 for $\Lambda > -5.3$.

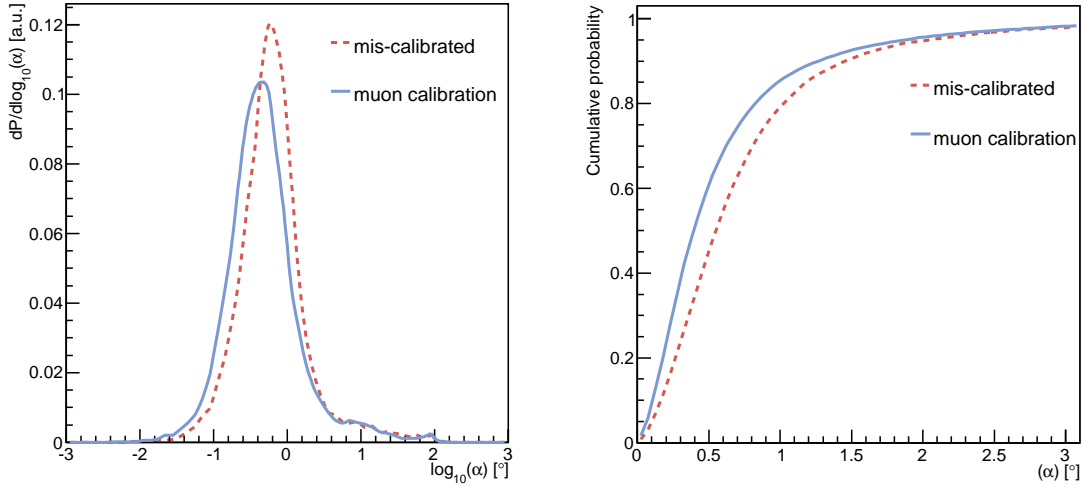


Figure 5: Left: angular error for the simulated tracks with mis-calibrated timing (dashed line) and standard simulation (solid line). Right: the corresponding cumulative distributions.

6.2 Cross-check with simulations

The precision of the muon time residuals method can be tested using simulations by adding inter line time offsets in the event reconstruction chain. The same time offsets as in the values summarised in Table 1 were introduced to distort the timing measured by the lines. The resulting simulated reconstructed tracks had the muon time residuals method applied to determine the inter line timing offsets. Fig. 6 shows the measured time offsets compared with the offsets artificially added to mis-calibrate the lines. The RMS of the distribution is 0.35 ns, which is well within the required precision.

6.3 Cross-check with the optical beacon system

The laser beacon can also be used as well to determine the time offsets between the detector lines. This system provides an independent calibration method to cross-check the results obtained using the information provided by the muon track residuals. For this purpose, a calibration run is taken every month in the ANTARES detector. The laser, at the bottom of line 8, illuminates the detector for about 10 minutes. The data are then analysed following a similar procedure to the one applied for the ARS_TO in-situ calibration [3]. The method is based on the study of the time differences between the emission of the light pulse and the time when light is detected by the OMs. Time residuals are calculated by correcting for the required photon propagation time. Gaussian functions convolved with an exponential are fitted to the time residual distributions. The obtained fit peak values are then plotted as a function of the distance between the OM and the laser beacon position. Only those OMs which are illuminated by the laser at the single photoelectron level are used in the

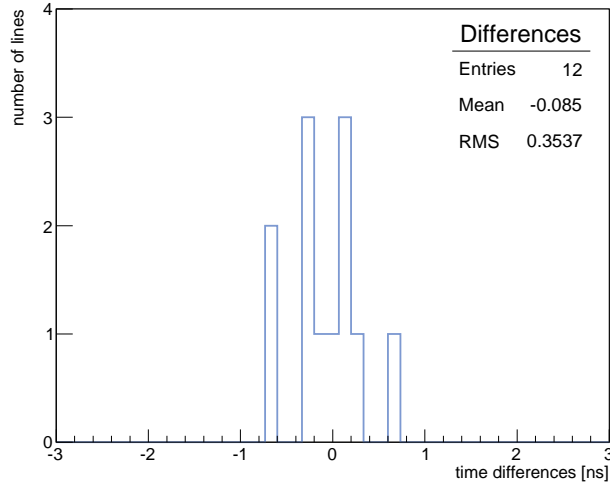


Figure 6: Distribution of the differences between the inter line offsets determined by applying the procedure on Monte Carlo simulations and the values introduced in order to distort the timing of the lines.

calculation, because for them a constant relation between residual peak positions and distance to the light source is expected.

In Figure 7 the inter line offsets measured with the laser optical beacon are compared to the results obtained using the track residuals method. They agree to within 1 ns, except for line 1 and line 8, which both show a discrepancy larger than 2 ns. The discrepancy for line 1 may be due to the large distance from the laser optical beacon (line 1 is the farthest), and for line 8, to a shadowing effect that prevents light from directly reaching the uppermost OMs.

It was found by examining the reconstructed events that the track residuals method is more efficient than the laser optical beacon method, producing a slightly higher number of well reconstructed tracks (see Figure 8).

7 Results on the intra line timing

The intra line calibration with atmospheric muons is performed by producing time residual distributions for every ARS in the detector. In order to reduce the amount of required data, the first step in the scheme explained in section 5 is modified to provide probe-hit time residuals for those ARSs that have registered at least one hit from the current event. Again, the resulting distributions are fit with a Gaussian function, the mean of which is taken as the time offset of the ARS under study. As an example, Figure 9 shows the fitted probe-hit residuals distribution for one particular ARS in line 4 and floor 10.

Since the `ARS_T0` constants provided by the LED beacon calibration system are taken into account in the muon track reconstruction, the procedure discussed in this paper provides corrections to the intra line calibration parameters obtained using the beacons. In Figure 10 the distribution of these corrections is shown including all ARSs with sufficient statistics after one iteration (left) and after

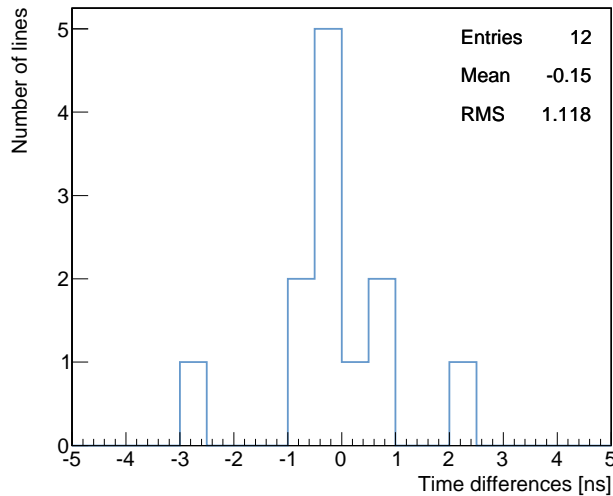


Figure 7: Time difference between the inter line offsets measured using the reconstructed muon tracks and the values obtained using the laser beacon system.

five iterations (right) of the method, when the procedure has converged to a solution.

7.1 Effect on the reconstruction

In order to determine the effect on the reconstruction of the measured ARS time corrections, a small sample of data runs has been analysed. In Fig. 11 a comparison of the reconstruction quality parameter for events reconstructed considering the LED beacon calibration (dashed line) and reconstructed applying the time corrections obtained using the method presented in this paper is shown. A small increase in the number of high quality events is observed, similar to the findings from simulations in Section 6.1, after correcting for artificially introduced time offsets. This validates the intra-line-timing calibration produced with the muon time residuals method. Note also that, while the LED beacons typically calibrate about one thousand ARSs, the muon tracks method allows the timing corrections of almost all of the active channels (nearly 1600 in this study) to be obtained.

8 Conclusions

A method for the time calibration of the ANTARES detector based on the reconstruction of atmospheric muons has been presented. It has been used to measure the relative time offsets between the different detection lines, already measured in the laboratory, with a precision of 1 ns. A calibrated inter line timing results in a factor of 1.2-2.0 improvement in the number of high-quality events reconstructed in the zenith angle range $-0.8 < \cos(\theta) < 0$, i.e. events traversing the detector diagonally. It has also been shown that not accounting for the inter line time corrections would result in a degradation of the expected detector resolution by approximately 40%. A comparison with a timing calibration results obtained with the laser beacon system shows that the two methods

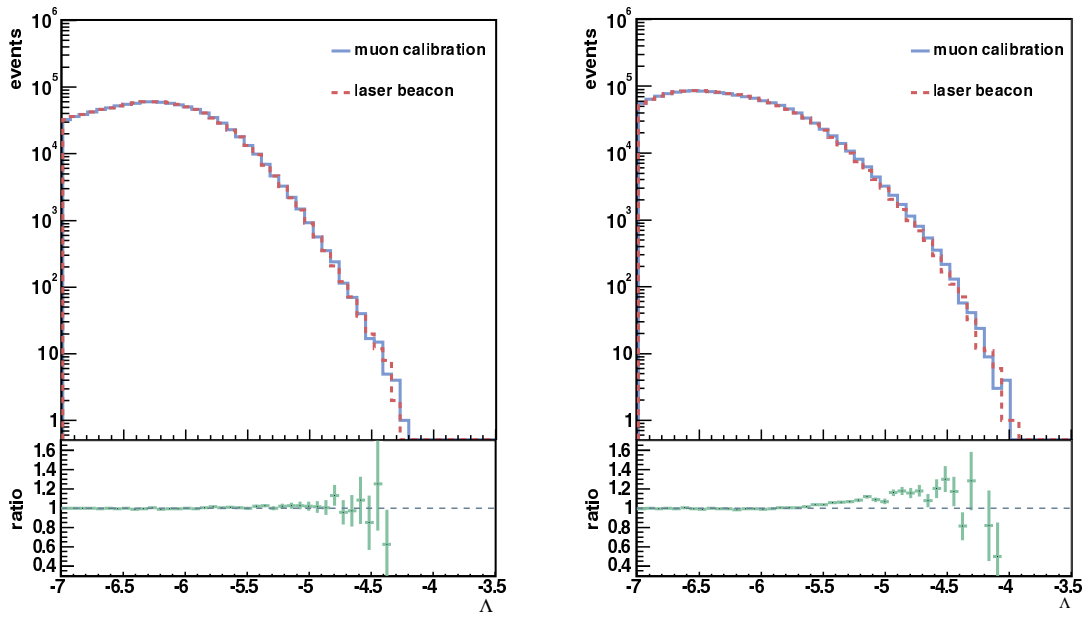


Figure 8: Distribution of the quality of the reconstruction parameter for events that have been reconstructed with left: $\cos(\theta) < -0.9$ and right: $\cos(\theta) > -0.8$, using the inter line offset corrections provided by the track residuals method (solid line) and those obtained using the laser beacon data (dashed line). In each case, the bottom plot shows the ratio between the number of reconstructed events for each case (muons/laser). For $\cos\theta > -0.8$, a $>10\%$ improvement can be seen from $\Lambda > -5.3$.

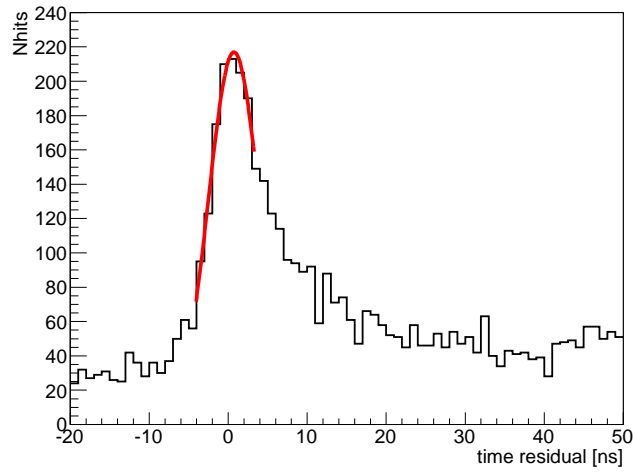


Figure 9: Distribution of the time residuals for ARS number 3 in line 4, floor 10. A gaussian function is fitted around the peak of the distribution, the mean of which is interpreted as the ARS time offset.

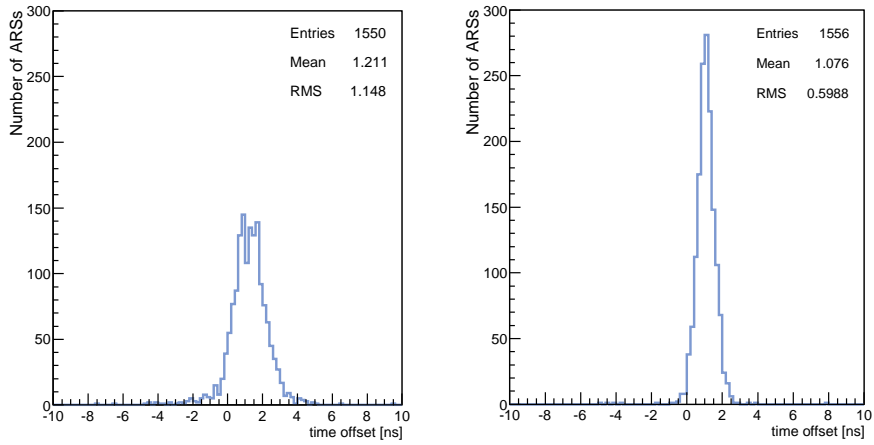


Figure 10: Distribution of the ARS timing corrections after one iteration (left) and after five iterations (right) of the muon time residuals method. The increase in the number of entries is due to a few ARS which are outside the time range in the left figure.

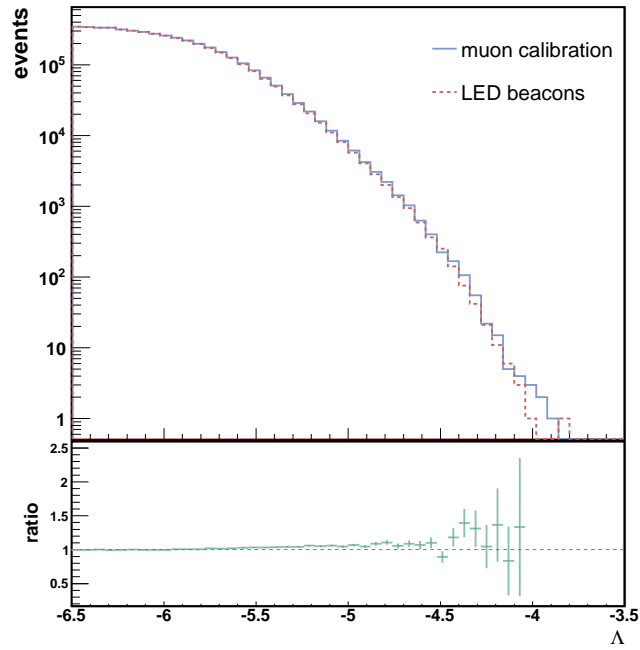


Figure 11: Distributions of the quality of the reconstruction parameter for events reconstructed using the LED beacon intra line calibration constants (dashed line) and reconstructed by taking into account the corrections determined by applying the muon time residuals method (solid line). Shown in the lower pad is the ratio between the events reconstructed by applying the muon time residuals method and the events reconstructed using the LED beacon calibration constants only.

are compatible to within 1 ns. However, the muon time residuals procedure is more efficient for the improvement of the quality of reconstructed events and allows the evaluation of the offsets for two lines that cannot be considered correctly with the laser optical beacon. The muon time residuals method has also been applied to measure corrections for the intra line timing constants first determined using the system of LED optical beacons. The results show a slight improvement in the quality of the reconstruction. The calibration explained here is applied in ANTARES analyses.

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References

- [1] M. Ageron et al., ANTARES Collaboration, “ANTARES: the first undersea neutrino telescope”, *Nucl. Instrum. Methods A* 656 (2011) 11.
- [2] S. Adrián-Martínez et al., ANTARES Collaboration, “Searches for Point-like and extended neutrino sources close to the Galactic Center using the ANTARES neutrino Telescope”, *ApJ*, 786 (2011) L5
- [3] J. A. Aguilar et al, ANTARES Collaboration, “Time calibration of the ANTARES Neutrino Telescope”, *Astrop. Phys.* 34 (2011) 539.
- [4] S. Adrián-Martínez et al, ANTARES Collaboration, “First search for point sources of high-energy cosmic neutrinos with the ANTARES neutrino telescope”, *Astrophys. Phys.* 743 (2011) L14.
- [5] M. Ageron et al, ANTARES Collaboration, “The ANTARES Optical Beacon system”, *Nucl. Instrum. Methods A* 578 (2007) 498.

- [6] J.A. Aguilar et al., ANTARES Collaboration, “Study of large hemispherical photomultiplier tubes for the ANTARES neutrino telescope”, Nucl. Instrum. Methods A 555 (2005) 132.
- [7] P. Amram et al., ANTARES Collaboration, “The ANTARES Optical Module”, Nucl. Instrum. Methods A 484 (2002) 369.
- [8] J. A. Aguilar et al, ANTARES Collaboration, “Performance of the front-end electronics of the ANTARES Neutrino Telescope”, Nucl. Instrum. Methods A 622 (2010) 59.
- [9] J. A. Aguilar et al, ANTARES Collaboration, “The data acquisition system for the ANTARES neutrino telescope”, Nucl. Instrum. Methods A 570 (2007) 107.
- [10] S. Adrián-Martínez et al., ANTARES Collaboration, “Search for cosmic neutrino point sources with four year data of the ANTARES telescope”, ApJ 760 (2012) 53.
- [11] A. Margiotta (for the ANTARES Collaboration), “Common simulation tools for large volume detector”, Nucl. Instrum. Methods A 725 (2013) 98.
- [12] V. Agrawal et al., “Atmospheric neutrino flux above 1 GeV”, Phys. Rev. D 53(1996) 1314.
- [13] G. Carminati et al., “Atmospheric MUons from PArametric formulas: A fast GEnerator for neutrino telescopes.”, Comput. Phys. Commun. 179(2008) 915.