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Measurement of the $^8$B Solar Neutrino Flux with the KamLAND Liquid Scintillator Detector

S. Abe,1 K. Furuno,1 A. Gando,1 Y. Gando,1 K. Ichimura,1 H. Ikeda,1 K. Inoue,1,2,2 Y. Kibe,1,2 W. Kimura,1 Y. Kishimoto,1 M. Koga,1,2 Y. Minekawa,1 T. Mitsui,1 T. Morikawa,1 N. Nagai,1 K. Nakajima,1 K. Nakamura,1,2 M. Nakamura,1 K. Narita,1 I. Shimizu,1 Y. Shimizu,1 J. Shirai,1 F. Suekane,1 A. Suzuki,1 H. Takahashi,1 N. Takahashi,1 Y. Takemoto,1 K. Tamae,1 H. Watanabe,1 B.D. Xu,1 H. Yabumoto,1 E. Yonezawa,1 H. Yoshida,1 S. Yoshida,1 S. Enomoto,2 A. Kozlov,2 H. Murayama,2,3 C. Grant,4 G. Keefe,4 D. Mckee,4 A. Piepke,2,4 T.I. Banks,3 T. Bloxham,3 J.A. Detwiler,3 S.J. Freedman,2,3 B.K. Fujikawa,2,3 K. Han,3 R. Kadel,3 T. O’Donnell,3 H.M. Steiner,3 L.A. Winslow,3 D.A. Dwyer,5 C. Mauger,5 R.D. McKeown,5 C. Zhang,5 B.E. Berger,6 C.E. Lane,7 J. Maricic,7 T. Miletic,7 M. Batygov,6 J.G. Learned,8 S. Matsuno,8 S. Pakvasa,9 M. Sakai,8 G.A. Horton-Smith,2,9 A. Tang,9 K.E. Downum,10 G. Gratta,10 K. Tolich,10 Y. Efremenko,2,11 Y. Kamyshekov,11 O. Pervozvichkov,11,13 H.J. Karwowski,12 D.M. Markoff,12 W. Tornow,12 K.M. Heeger,2,13 F. Piquemal,14 J.-S. Ricot,14 and M.P. Decowski2,3,15

(The KamLAND Collaboration)

1 Research Center for Neutrino Science, Tohoku University, Sendai 980-8578, Japan
2 Institute for the Physics and Mathematics of the Universe, Tokyo University, Kashiwa 277-8568, Japan
3 Physics Department, University of California, Berkeley and Lawrence Berkeley National Laboratory, Berkeley, California 94720, USA
4 Department of Physics and Astronomy, University of Alabama, Tuscaloosa, Alabama 35487, USA
5 W.K. Kellogg Radiation Laboratory, California Institute of Technology, Pasadena, California 91125, USA
6 Department of Physics, Colorado State University, Fort Collins, Colorado 80523, USA
7 Physics Department, Drexel University, Philadelphia, Pennsylvania 19104, USA
8 Department of Physics, University of Hawaii at Manoa, Honolulu, Hawaii 96822, USA
9 Department of Physics, Kansas State University, Manhattan, Kansas 66506, USA
10 Physics Department, Stanford University, Stanford, California 94305, USA
11 Department of Physics and Astronomy, University of Tennessee, Knoxville, Tennessee 37996, USA
12 Triangle Universities Nuclear Laboratory, Durham, North Carolina 27708, USA and Physics Departments at Duke University, North Carolina Central University, and the University of North Carolina at Chapel Hill
13 Department of Physics, University of Wisconsin, Madison, Wisconsin 53706, USA
14 CEN Bordeaux-Gradignan, IN2P3-CNRS and University Bordeaux I, F-33175 Gradignan Cedex, France
15 Nikhef, Science Park, Amsterdam, Netherlands

We report a measurement of the neutrino-electron elastic scattering rate from $^8$B solar neutrinos based on a 123 kton-day exposure of KamLAND. The background-subtracted electron recoil rate, above a 5.5 MeV analysis threshold is 1.49±0.14(stat)±0.17(syst) events per kton-day. Interpreted as due to a pure electron flavor flux with a $^8$B neutrino spectrum, this corresponds to a spectrum integrated flux of 2.77±0.26(stat)±0.32(syst)×10$^{-1}$ cm$^{-2}$ s$^{-1}$. The analysis threshold is driven by 208Tl present in the liquid scintillator, and the main source of systematic uncertainty is due to background from cosmogenic $^{18}$O. The measured rate is consistent with existing measurements and with Standard Solar Model predictions which include matter enhanced neutrino oscillation.

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Neutrinos from $^8$B beta-decay ($Q\sim 18$ MeV) dominate the high-energy portion of the solar neutrino spectrum and have been a crucial source for solar neutrino experiments. Such experiments, together with reactor anti-neutrino experiments, have established flavor change through Large Mixing Angle (LMA) neutrino-oscillation with matter effects, introduced by Mikheev, Smirnov and Wolfenstein (MSW) [1, 2], as a consistent resolution of the solar neutrino problem [2, 3]. Measurements of neutrino-electron scattering using water-Cherenkov techniques have been achieved in a number of different detectors: Kamiokande [11,12], Super-Kamiokande [5, 6, 13] and SNO [7–10]. An advantage of this detection technique is sensitivity to the neutrino direction of incidence, although analysis thresholds of ~5 MeV are typically imposed. Recently SNO has achieved a threshold of ~3.5 MeV [10]. Liquid scintillator detectors may also be used to measure neutrino-electron elastic scattering; while insensitive to the neutrino direction, these detectors have better energy resolution and have the possibility of lower energy thresholds. A measurement with a threshold of ~ 3.0 MeV, the lowest to date, has been reported by Borexino [14]. In this paper we report an independent, liquid scintillator based measurement of the $^8$B flux with KamLAND.

The KamLAND detector consists of 1 kton of liquid scintillator (LS) confined in a 6.5-m-radius, 135μm-thick, spherical balloon. The LS is 80% dodecane, 20% pseudocumene, and 1.36±0.03 g/l of PPO; the density is 0.78 g/cm$^3$. The balloon is suspended in purified mineral oil within a 9-m-radius stainless-steel sphere (SSS). The scintillation light is recorded by an array of 1879 photomultiplier tubes (PMTs) mounted on the inner surface of the SSS; 554 are reused 20-inch PMTs from Kamiokande, the remainder are new 20-inch PMTs with…
the photo-cathode masked down to 17-inch. The 17-inch tubes have better timing and single-photo-electron (spe) resolution. A 3.2 kton, water-Cherenkov detector surrounds the SSS. It is used as a muon anti-coincidence counter and as shielding against external γ-rays and neutrons.

Event position (vertex) and energy are reconstructed based on the photon hit-time and charge distributions and a detector response model which is calibrated by periodically deploying radioactive sources: $^{203}$Hg, $^{137}$Cs, $^{68}$Ge, $^{65}$Zn, $^{60}$Co, $^{241}$Am$^{9}$Be, $^{210}$Po-$^{13}$C. For the current analysis, the reconstruction omits the 20-inch PMTs due to their poorer timing and spe resolution; in this mode the energy resolution is $\sqrt{7\%E(MeV)}$.

The signal of interest, neutrino-electron elastic scattering from $^{8}$B solar neutrinos, produces recoil electrons with reconstructed energy up to \(\sim 20 \text{MeV}\). We construct a vertex-\(\chi^{2}\) test to select electron-recoil-like events. The test compares the observed PMT charge and hit time distributions as a function of the distance to the event to the expected distributions. The expected distributions and selection criteria were developed and calibrated based on source calibration data and muon-tagged $\beta^{-}$ events from cosmogenic $^{12}$B/$^{12}$N. The efficiency of the selection is $0.99 \pm 0.01$.

The analysis is carried out in reconstructed energy with a threshold of 5.5 MeV, which corresponds to a 5 MeV threshold in physical electron recoil energy. The threshold is chosen to reject background from $^{208}$Tl(Q=5.0 MeV, $\tau_{1/2}$=3.05 min) produced by residual $^{232}$Th contamination present in the LS at a level of $(7.90 \pm 0.25) \times 10^{-17}$ g/g.

An energy-scale model was developed to convert physical energy to reconstructed energy. The model includes the linear response of the LS and particle-dependent non-linearities such as scintillator quenching and Cherenkov light production. The energy scale is constrained by calibration data, tagged $\alpha$-particles from $^{214}$Bi-$^{214}$Po/$^{212}$Bi-$^{212}$Po decays from residual $^{238}$U/$^{232}$Th in the LS and high energy $\beta^{-}$ decays (Q \(\sim 14 \text{MeV}\)) from cosmogenic $^{12}$B/$^{12}$N.

The measurement reported here is restricted to data acquired before the KamLAND LS purification. The data was taken between April 2002 through April 2007, and corresponds to 1432.1 days of run-time. To reduce the external γ-ray background the fiducial volume (FV) is reduced to a 6-m-high, 3-m-radius cylinder centered at the detector center. This shape defines an optimal analysis volume that takes advantage of further shielding by the ultra-pure LS from backgrounds coming from the cavern’s rock walls and a large steel deck which caps the experiment. The FV fraction and systematic uncertainty is determined from short-lived muon spallation products, mainly $^{12}$B(Q=13.4 MeV, $\tau_{1/2}$=20.2 ms), which are assumed to be produced uniformly in the LS. The FV fraction is taken as the ratio of the number of spallation products that reconstruct inside the analysis cylinder to the number that reconstruct in the full LS volume. This fraction combined with the total LS volume in the balloon, which was measured directly during construction to be 1171±25 m$^{3}$, gives a fiducial volume of 176.4 m$^{3}$ with an uncertainty of 3.1%. When all selection cuts, to be described, are applied the remaining exposure is 123 kton-days.

This paper focuses on neutrino-electron elastic scattering, $\nu + e^{-} \rightarrow \nu + e^{-}$ (ES) due to $^{8}$B solar neutrinos. With $3.423 \times 10^{22} e^{-}$ targets per kton of LS, we expect, including oscillation, 1.19 recoil events per kton-day from $^{8}$B in the energy-region of interest (ROI) 5.5–20 MeV. In our calculation we adopt the standard solar model (SSM) of Serenelli et al. [16] which uses the solar abundances of Asplund et al. [17] (AGSS09). This predicts a total $^{8}$B neutrino flux, independent of neutrino flavor, of $\Phi_{B} = 4.85 \pm 0.58 \times 10^{6} \text{cm}^{-2}\text{s}^{-1}$. We use the ES cross section from [18]; the $^{8}$B-\(\nu_{e}\) spectrum of Winter et al. [19]; the oscillation parameters from the global analysis in [4]; and fold in the detector response. ES from hep neutrinos and neutrino interactions on carbon in the LS are treated as a background and are discussed later in the paper. We note that the best choice of solar abundances to use in the SSM is an unresolved question. In AGSS09 the helioseismic discrepancy which emerged when the abundances of Asplund et al. [20](AGS05) were adopted still persists, although it is not as severe. If the older, 1-dimensional, solar abundance evaluation of Grevesse and Sauval [21](GS98) is used, the SSM has excellent consistency with helioseismology and the predicted $^{8}$B flux is $\Phi_{B} = 5.88 \pm 0.65 \times 10^{6} \text{cm}^{-2}\text{s}^{-1}$. Assuming no sterile neutrinos, the total $^{8}$B flux is directly observed in the SNO neutral current measurement to be

\begin{table}[h]
\centering
\caption{Production rate of muon-spallation isotopes that dominate the background (in events per kton-day) [15]. Non-bright muons produce less than \(7 \times 10^{5}\) p.e. in the PMT array.}
\begin{tabular}{|c|c|c|c|c|}
\hline
Isotope & Q\(\text{[MeV]}\) & $\tau_{1/2}$\(\text{[s]}\) & All Muons & Non-Bright Muons \\
\hline
$^{8}$Li & 16 & 0.84 & 15.6±3.2 & 0.7±0.4 \\
$^{8}$B & 18 & 0.77 & 10.7±2.9 & <0.006 \\
$^{11}$Be & 11.5 & 13.8 & 1.4±0.3 & 0.27±0.30 \\
\hline
\end{tabular}
\end{table}

\begin{table}[h]
\centering
\caption{Estimated background contributions for the full exposure and after all cuts. The shorter lived spallation products $^{12}$B, $^{12}$N, $^{9}$C, $^{7}$Li, and $^{4}$He are also considered and their contributions are summarized by “Spallation Other”.

\begin{tabular}{|l|l|}
\hline
Background & Counts \\
\hline
Spallation $^{31}$Be & 89.1 ± 19.1 \\
Spallation $^{8}$Li & 20.5 ± 4.0 \\
Spallation $^{8}$B & 11.0 ± 3.0 \\
Spallation Other & 0.4 ± 0.6 \\
External gamma rays & 25.2 ± 12.6 \\
$^{8}$B CC on $^{13}$C GND & 5.8 ± 1.4 \\
Reactor $\bar{\nu}_{e}$ & 1.6 ± 0.1 \\
$^{8}$B CC on $^{13}$C 3.51 MeV & 1.1 ± 0.4 \\
hep ES & 0.6 ± 0.1 \\
Atmospheric $\nu$ & 2.0 ± 2.0 \\
\hline
Total & 157.3 ± 23.6 \\
\hline
\end{tabular}
\end{table}
$\Phi_{sB} = 5.140^{+0.197}_{-0.207} \times 10^6 \text{ cm}^{-2}\text{s}^{-1} \text{MeV}^{-1}$. 

With the 5.5 MeV analysis threshold the background is dominated by decays of light isotopes produced by muon spallation. An in-depth study of muon activation at KamLAND can be found in $[15]$. A study of light isotope production shows that most (> 80%) light isotope backgrounds are correlated with muons which produce more than 700,000 photoelectrons (p.e.) in the PMT array. We denote these as bright muons. The rate of bright muons is 0.037 Hz, while the total rate of muons passing through the LS is 0.198 ± 0.014 Hz. We apply a number of muon-related cuts to reduce these spallation backgrounds. All events within 200 ms of a preceding muon are rejected. This time veto of the full detector significantly reduces the background from external gamma rays in the ROI within the cylindrical fiducial volume. The uncertainty in the estimate comes from the difference in the observed and simulated attenuation lengths.

As remarked earlier, events from hep solar neutrinos and solar neutrino interactions on carbon are treated as a background. Using the SSM with AGSS09 and the hep spectrum from $[24]$, we estimate, including oscillation, 0.6 ± 0.1 electron recoil events from hep neutrinos in the ROI. In our calculation we use the ES cross section, neutrino oscillation parameters and detector response as was done for the $^8$B ES calculation. The uncertainty is dominated by the difference in the flux prediction of the SSM with the AGSS09 versus with the GS98 solar abundances.

There are $4.30 \times 10^{31}$ carbon nuclei per kton of LS if we assume a natural $^{13}$C of 1.10%. Using the cross sections calculated in $[23]$, we find the largest $\nu_e$-$C$ scattering background contribution to be from charged current (CC) scattering, $^{13}$C+\nu_e \rightarrow ^{13}$N+e$^{-}$ from $^8$B. We estimate, including oscillation, that scattering to the ground state of $^{13}$N produces 5.8 ± 1.4 events in the ROI; and scattering to the 3.51 MeV first excited, which decays by proton emission, contributes 1.1 ± 0.4 electron-recoil-like events in the ROI. In this estimate an additional uncertainty of 30% on the cross section is included. The contribution from higher states of $^{13}$N, neutral current (NC) scattering by $^8$B-$\nu$ and hep-$\nu$ NC and CC interactions on carbon is estimated to be less than 0.13 events in the ROI, and is considered negligible given the other backgrounds. These calculations use the same solar-$\nu$ fluxes, spectra, detector response and oscillation parameters as be-

![Fig. 1: Energy spectrum of $^8$B candidates with the best-fit spectrum and background components from the unbinned energy and rate analysis. The histograms display the results in bins of 0.5 MeV except for the last bin which due to limited statistics extends from 13.5 MeV to 20 MeV.]

### Table III: The systematic uncertainties associated with the unbinned fit to the energy spectrum of the $^8$B candidates. The detection efficiency is dominated by our fiducial volume uncertainty.

<table>
<thead>
<tr>
<th>Source</th>
<th>Uncertainty (%)</th>
</tr>
</thead>
<tbody>
<tr>
<td>$^{11}$Be</td>
<td>10.8</td>
</tr>
<tr>
<td>$^8$Li and $^8$B</td>
<td>3.3</td>
</tr>
<tr>
<td>External gamma rays</td>
<td>6.8</td>
</tr>
<tr>
<td>Other Backgrounds</td>
<td>1.1</td>
</tr>
<tr>
<td>Detection Efficiency</td>
<td>6.3</td>
</tr>
<tr>
<td>Energy Scale</td>
<td>0.8</td>
</tr>
</tbody>
</table>
fore. The effect of using the oscillation parameters from the
global analyses in [13, 26] were investigated but do not change
the result due to the large flux and cross section uncertainties.

In addition to these neutrino interactions, we expect a small
background from atmospheric neutrino interactions. Assuming
a flat spectrum, we use the candidate events in the range of
20 MeV to 35 MeV to extrapolate the atmospheric neutrino
contribution in the signal region, and find 2.0±2.0 events.

A background from positrons from inverse beta decay in-
duced by reactor antineutrinos is tagged by coincidence with
delayed neutron capture γ-rays. The mean capture time in
KamLAND is ~207 µs. The tagging cuts are: a delayed en-
ergy cut, 2.04 MeV < E < 2.82 MeV, to select the 2.2 MeV n-p
capture γ; a timing cut, 0.5µs < ΔT < 660µs; and a position
cut, ΔR < 1.6m. ΔT and ΔR refer respectively to the time and
distance between the prompt positron and delayed capture γ.
The tagging efficiency is 0.940±0.006. The reactor antineu-
trino flux at KamLAND is dominated by Japanese power re-
actors, we calculate the flux and spectrum using data provided
by the Japanese power companies and a simplified reactor model [27]. The residual background is 1.6±0.1, where the
uncertainty in the neutrino oscillation parameters is included.

The total background, broken down in Table III is estimated
to be 157.3±23.6 events. While 6Li has the largest produc-
tion rate, 11Be, due to its long half-life (τ1/2 = 13.8 s), is the
largest contribution to the total background.

After applying all the cuts, 299 electron recoil candidates
are found in the ROI in the 123 kton-day exposure. Subtracting the 157.3±23.6 background events gives a 3B
rate of 1.16±0.14(stat)±0.21(syst) events per kton-day which
is consistent with our prediction including neutrino oscil-
lation. If we neglect neutrino oscillation, assume a pure νe
flux, and correct for the 5.5 MeV threshold, the mea-
ured rate corresponds to a spectrum integrated flux of
ΦES = 2.17±0.26(stat)±0.39(syst)×10^6 cm^{-2} s^{-1}.

A second analysis of this data uses an unbinned maximum
likelihood fit to the 3B candidate energy spectrum. The nor-
malization of the individual backgrounds are allowed to vary
but are constrained by a penalty to the expected value. The
best-fit rate is 1.49±0.14(stat)±0.17(syst) events per kton-
day, with a goodness-of-fit of 49%. Once again assuming a
pure νe flux, this corresponds to a spectrum integrated flux of
ΦES = 2.77±0.26(stat)±0.32(syst)×10^6 cm^{-2} s^{-1}. Fig. 1
shows the energy spectrum of the candidate events with the
best-fit solar neutrino and background spectra. This anal-
ysis does not include the small change in the shape of
the neutrino spectra due to oscillation. Including this ef-
fect does not significantly change the best-fit rate, ΦES =
2.74±0.26(stat)±0.32(syst)×10^6 cm^{-2} s^{-1}, with a goodness
of fit of 57%.

The KamLAND result is compared to other measurements
of the total 8B flux through neutrino-electron elastic scat-
tering in Fig. 2. For this comparison the reported statistical
and systematic errors are added in quadrature. The mean of the experiments, weighted by their uncertainties, is
ΦES = 2.33±0.05×10^6 cm^{-2} s^{-1}. This is dominated by the
very precise Super-Kamiokande water-Cherenkov measure-
ment. The reduced χ^2/NDF is 6.6/8 where the KamLAND
Rate + Energy result and the Borexino E>3 MeV result are
used in the calculation.

A key prediction of the LMA-MSW solution to the sol-
lar neutrino problem is the expected transition from matter-
dominated to vacuum-dominated oscillations, depending on
the neutrino energy and the region of the solar interior probed
by the neutrinos as they journey out of the sun. A general fea-
ture is that the survival probability, defined as the ratio of the
oscillated νe- flux to the unoscillated flux, tends to increase
with decreasing solar neutrino energy. This effect is the re-
sult of neutrino species (8B, 7Be, pp, etc.) either moving off,
or being produced in a region not satisfying the MSW reso-
nance condition. These neutrinos are then described by va-
uum oscillation. The curve in Fig. 3 shows the ratio of the
expected neutrino-electron elastic scattering rates with and
without oscillation calculated using the AGSS09-SSM. The
FIG. 3: Ratio of the measured (threshold corrected) solar neutrino-electron elastic scattering rate and that predicted by the AGSS09-SSM as a function of energy (data points). The average of the water-Cherenkov detector results above 5 MeV is shown. The horizontal error bars indicate the range of energy over which the flux was measured. The central value indicates the flux averaged energy for that range. The solid curve (and ±1σ band) shows the ratio of AGSS09 with neutrino oscillation included to the prediction without oscillation. The discontinuities in the prediction arise as different neutrino sources become significant, the large uncertainty in the CNO neutrinos is evident at ∼1.2 MeV.

data points are the survival probabilities deduced, relative to the SSM prediction, from neutrino-electron elastic scattering measurements at Borexino, KamLAND and the combined water-Cherenkov experiments. The Borexino 8B measurement with a 3.0 MeV threshold [14] and the 7Be flux measurement [28, 29] are consistent with LMA-MSW prediction. The measurements are not yet precise enough to resolve the issue of solar abundances.

This letter reports a measurement of the 8B solar neutrino flux measured through neutrino-electron elastic scattering in the KamLAND LS. Our result is in good agreement with existing measurements from both water-Cherenkov and liquid-scintillator-based detectors. In the data-set presented, 208Tl background limits the KamLAND analysis threshold to 5.5 MeV. We expect to achieve a lower threshold in a future analysis of data taken after purification of the KamLAND LS.

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* Present address: Department of Physics, Tokyo Institute of Technology, Tokyo 152-8551, Japan
† Jointly at: Center for Experimental Nuclear Physics and Astrophysics, University of Washington, Seattle, Washington 98195, USA
‡ Present address: Lawrence Livermore National Laboratory, Livermore, California 94550, USA
§ Present address: Department of Physics, Kansas State University, Manhattan, Kansas 66506, USA
¶ Present address: Department of Physics, Massachusetts Institute of Technology, Cambridge, MA 02139, USA
** Present address: Subatomic Physics Group, Los Alamos National Laboratory, Los Alamos, NM 87545, USA
†† Present address: Department of Physics and Astronomy, Rowan University, Glassboro, New Jersey 08028, USA
‡‡ Present address: SNOLAB, Lively, ON P3Y 1N2, Canada
§§ Present address: Department of Physics, Louisiana State University, Baton Rouge, LA 70803 USA