Ortho-to-para ratio of interstellar heavy water


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ABSTRACT

Context. Despite the low elemental deuterium abundance in the Galaxy, enhanced molecular D/H ratios have been found in the environments of low-mass star-forming regions, and in particular the Class 0 protostar IRAS 16293-2422.

Aims. The CHESS (Chemical Herschel Surveys of Star forming regions) key program aims to study the molecular complexity of the interstellar medium. The high sensitivity and spectral resolution of the Herschel/HIFI instrument provide a unique opportunity to observe the fundamental para-D2O molecule at 316.8 GHz towards IRAS 16293-2422. The line is seen in absorption with a line opacity of 0.62 ± 0.11 (4σ). From the previous ground-based observations of the fundamental para-D2O transition at 316.8 GHz, we estimate a line opacity of 0.26 ± 0.05 (1σ).

Methods. We detected the fundamental transition of the para-D2O molecule at 607.35 GHz towards IRAS 16293-2422. The line is seen in absorption at 316.8 GHz with a line opacity of 0.26 ± 0.05 (1σ).

Results. We show that the observed absorption is caused by the cold gas in the envelope of the protostar. Using these new observations, we estimate for the first time the ortho-to-para D2O ratio to be lower than 2.6 at a 3σ level of uncertainty, which should be compared with the thermal equilibrium value of 2:1.


1. Introduction

Among all molecules in interstellar space, water is special because of its dominant role in the cooling of warm gas and in the oxygen chemistry as well as for its role in the chemistry of the atmospheres of exoplanets and its potential connection with life. Water abundance in cold molecular gas is very low because it is frozen onto the interstellar grains and forms icy mantles around them. Although water can form theoretically by means of gaseous reactions that first form H2O+ and H2O (e.g. Rodgers & Charnley 2002), no observational evidence of this has yet been found. It is assumed that the major mechanism of water formation occurs on grain surfaces. One observable that helps us to discriminate between the various formation mechanisms is the abundance of single and double deuterated water relative to the normal isotopologue. Another potential discriminant is the ortho-to-para ratio (OPR), namely the ratio of water molecules with different nuclear spins. Since radiative and inelastic collisional transitions between the two ortho and para states are strongly forbidden, the OPR is set at the moment of the water formation and is changed by nuclear spin reactions exchange later on. This can occur in either the gas phase by reactions with ions in which actual nuclei change places, or on the grain surfaces by interaction with electron spins or, perhaps, even other nuclear spins (e.g. Le Bourlot 2000; Limbach et al. 2006). Although little is known about the spin exchange in the gas phase, it is usually assumed that this is a slow process and that the OPR is likely to retain information about the moment of its formation. Empeighenger et al. (2010); Lis et al. (2010) report determinations of the water OPR in several environments based on new Herschel observations. The doubly deuterated isotopologue of water, D2O, consists of two species, ortho and para with a nuclear spin statistic weight 2:1. To date, D2O has only been detected towards the solar-type protostar IRAS 16293-2422 (hereafter IRAS 16293), by observing the fundamental transition of the para-D2O transition at 316.8 GHz (see our Fig. 1 and Butner et al. 2007). The observed line profile (see Fig. 2) shows a component in emission with a deep absorption later on. This can occur in either the gas phase by reactions with ions in which actual nuclei change places, or on the grain surfaces by interaction with electron spins or, perhaps, even other nuclear spins (e.g. Le Bourlot 2000; Limbach et al. 2006). Although little is known about the spin exchange in the gas phase, it is usually assumed that this is a slow process and that the OPR is likely to retain information about the moment of its formation. Empeighenger et al. (2010); Lis et al. (2010) report determinations of the water OPR in several environments based on new Herschel observations. The doubly deuterated isotopologue of water, D2O, consists of two species, ortho and para with a nuclear spin statistic weight 2:1. To date, D2O has only been detected towards the solar-type protostar IRAS 16293-2422 (hereafter IRAS 16293), by observing the fundamental transition of the para-D2O transition at 316.8 GHz (see our Fig. 1 and Butner et al. 2007). The observed line profile (see Fig. 2) shows a component in emission with a deep absorption at the cloud velocity (∼4 km s⁻¹). The emission component has been attributed to heavy water in the hot corino of this source where the grain ices are sublimated and released into the gas phase (Ceccarelli et al. 2000; Bottinelli et al. 2004), based on the detailed analysis of several HDO lines observed in IRAS 16293

* Herschel is an ESA space observatory with science instruments provided by European-led principal Investigator consortia and with important participation from NASA.
Fig. 1. Energy levels for the detected fundamental lines of D$_2$O.

(Parise et al. 2005). The absorption component, whose linewidth is 0.5 km s$^{-1}$, probably originates in the foreground gas (molecular cloud and cold envelope). Therefore, the absorption component provides a straightforward measure of the column density of para-D$_2$O in the cold gas surrounding IRAS 16293.

2. Observations and results

In the framework of the key program CHESS (Ceccarelli et al. 2010), we observed the solar type protostar IRAS 16293 with the HIFI instrument (de Graauw et al. 2010; Roelfsema et al. 2010) onboard the Herschel Space Observatory (Pillbrat et al. 2010). A full spectral coverage of band 1b between 554.5 and 636.5 GHz was performed on 2010 March 2, using the HIFI spectral scan double beam switch (DBS) mode with optimization of the continuum. The fundamental ortho-D$_2$O ($1_1^1$–$0_0^0$) transition lies in this frequency range, at 607.35 GHz (see Fig. 1). The HIFI wide band spectrometer (WBS) was used, providing a spectral resolution of 1.1 MHz ($\sim 0.55$ km s$^{-1}$ at 600 GHz) over an instantaneous bandwidth of 4 $\times$ 1 GHz. We note that the data are acquired at the Nyquist sampling, therefore, with 0.5 MHz steps. The targeted coordinates were $\alpha_{2000} = 16^h32^m22^s.75$, $\delta_{2000} = -24^\circ28'34.2''$. The beam size at 610 GHz is about 35$''$, the theoretical main beam (respectively forward) efficiency is 0.72 (resp. 0.96), and the DBS reference positions were situated approximately 3$'$ east and west of the source. The data were processed using the standard HIFI pipeline up to level 2 with the ESA-supported package HIPE 3.01 (Ott et al. 2010). The 1 GHz chunks are then exported as FITS files into the Generic Spectral Analysis Software (GASS) developed for Molecular Spectroscopy (Müller et al. 2005; Brünken et al. 2007), as reported in Table 1. In the same table, we report also the parameters of the para-D$_2$O ($1_1^1$–$0_0^0$) fundamental line previously observed at the JCMT, published by Butner et al. (2007), at a rest frequency of 316 799.81 MHz. The data were retrieved from the JCMT archive and reduced within CLASS. We performed a 3-component Gaussian fit with CASSIS, which takes into account the ortho and para D$_2$O lines (see Butner et al. 2007)).

In this work, the para-D$_2$O line in emission has an intensity of $0.16 \pm 0.01$ K in main beam temperature, and a linewidth of $4.01 \pm 0.77$ km s$^{-1}$. The bright line at a $V_{LSR}$ of 10.1 km s$^{-1}$ is likely due to CH$_3$OD (see Butner et al. 2007) with an intensity of $(0.16 \pm 0.01)$ K and a linewidth of $(4.6 \pm 0.5)$ km s$^{-1}$. The parameters for the resulting fit of the para-D$_2$O absorption line are quoted in Table 1.

3. Determination of the D$_2$O OPR

Crimier et al. (2010) used the JCMT SCUBA maps of IRAS 16293 at 450 $\mu$m and 850 $\mu$m (and other data)
and a density lower than about $5 \times 10^6$ cm$^{-3}$, the HIFI beam is much larger than the continuum emitting region. Owing to the low temperatures found in the cold envelope surrounding IRAS 16293 as well as the parent cloud, much more extended than the continuum emitting region. We note that, as long as the sizes of the absorbing layer are larger than the sizes of the region emitting the continuum, the line-to-continuum ratio does not depend on the sizes of the telescope beam used for the observations. Therefore, we can compute the D$_2$O OPR directly from the line-to-continuum ratios of the JCMT and Herschel observations, with no further correction. We note also that the para-D$_2$O line has an emission component that Butner et al. (2007) attributed to the hot corino region, whereas here we consider an absorption component only. In contrast, the ortho-D$_2$O line reported here exhibits absorption only because the emission component is probably diluted in the 35'' HIFI beam, which is much larger than the 15'' JCMT beam at 316 GHz.

Adopting the density and temperature profiles of the envelope of IRAS 16293 (Crimier et al., 2010), the gas at a distance larger (in radius) than 900 AU has a temperature lower than 30 K and a density lower than about $5 \times 10^6$ cm$^{-3}$ (see Fig. 3). Thus, given the temperature of the gas absorbing the D$_2$O lines, we consider only the first two levels of each D$_2$O form. We use computed collisional rates for the two fundamental deexcitation transitions of ortho- and para-D$_2$O with para-H$_2$ in the 10–30 K range of 2.3 $\times$ 10$^{-11}$ and 3.8 $\times$ 10$^{-11}$ cm$^{-3}$ s$^{-1}$ respectively (Wiesenfeld et al., in prep.). At the low temperatures found in the cold envelope, it is likely that H$_2$ is mainly in its para form (Pagani et al., 2009; Troscompt et al., 2009). For the collisional rates given above, the critical densities of the ortho- and para-D$_2$O fundamental transitions are $1 \times 10^8$ and $2 \times 10^7$ cm$^{-3}$ respectively, and the upper levels of the two transitions are only moderately subthermally populated for a density of $5 \times 10^6$ cm$^{-3}$. For a two-level system, the species column density can be computed as

$$N_{\text{tot}} = \frac{8\pi n_3^3}{A_{\text{ul}}} \Delta V \sqrt{\pi} \int_0^\infty \frac{Q(T)}{v_{\text{abs}}} \exp\left[\frac{E_u}{k/T_{\text{abs}}} - 1\right] \frac{1}{h\nu} \frac{\exp(\hbar\nu/kT_{\text{mon}})}{\exp(\hbar\nu/kT_{\text{abs}}) - 1},$$

where $A_{\text{ul}}$ is the Einstein coefficient (2.96 $\times$ 10$^{-3}$ s$^{-1}$ for the ortho transition and 6.3 $\times$ 10$^{-4}$ s$^{-1}$ for the para transition), $E_u$ is the upper energy ($E_u/k = 15.2$ K for the para transition and $29.2$ K for the ortho transition), $g_u$ is the upper statistical weight (3 for the ortho transition, 6 for the para transition), $\nu$ is the frequency (316.79981 GHz for the para transition and 607.349449 GHz for the ortho transition), $\Delta V$ is the linewidth (cm s$^{-1}$), and $\tau$ is the opacity at the line center. The parameter $T_{\text{ex}}$ is the excitation temperature and $Q(T)$ its corresponding partition function. In the approximation of the escape probability formalism, $T_{\text{ex}}$ is defined by the equation

$$T_{\text{ex}} = \frac{h\nu/k}{h\nu/kT_{\text{mon}} + \ln(1 + A_{\text{ul}}/C_{\text{ul}})},$$

where $C_{\text{ul}} = \gamma_{\text{ul}} \times n_{\text{collission}}$, $n_{\text{collission}}$ being the density of the collision partner (in this case para-H$_2$) and $\gamma_{\text{ul}}$ being the collisional rate in cm$^3$ s$^{-1}$ (values given above). The $\beta$ parameter represents the probability that a photon at some position in the cloud escapes the system. For a static, spherically symmetric, and homogeneous medium, Osterbrock & Ferland (2006) derive this parameter as a function of the optical depth $\tau$ in the direction of the observer (see their Appendix 2). The opacity at the line center is expressed as a function of the line depth ($T_{\text{abs}} = T_{\text{C}} - T_{\text{L}}$) and the continuum ($T_{\text{C}}$)

$$\tau = -\ln\left(1 - \frac{T_{\text{abs}}}{T_{\text{C}} - J_\nu(T_{\text{ex}}) + J_\nu(T_{\text{C}})}\right),$$

where $J_\nu(T_{\text{ex}}) = (h\nu/k)/(\exp(h\nu/k) - 1)$ and $T_{\text{C}}$ is the cosmic microwave background radiation temperature (2.73 K). In the limit of $\tau \gg 1$, $T_{\text{C}} - T_{\text{abs}} \approx J_\nu(T_{\text{ex}}) - J_\nu(T_{\text{C}})$, and $T_{\text{ex}} \approx 5$ K. Since the D$_2$O transitions are probably optically thin, we can reasonably assume that $T_{\text{ex}}$ is lower than 5 K and $J_\nu(T_{\text{ex}}) - J_\nu(T_{\text{C}})$ is negligible.

As discussed above, we assume that the absorbing layer is much larger than the continuum emitting region. Owing to the uncertainty in the H$_2$ density (lower than $5 \times 10^6$ cm$^{-3}$)
and the kinetic temperature (lower than 30 K), we applied the method described above to determine the column densities with $n_{\text{H}_2} = 10^5 \, \text{cm}^{-3}$ and $T_{\text{kin}} \sim 20 \, \text{K}$. Table 1 lists the computation of the optical depths for both lines and their corresponding uncertainties. Since $\tau = -\ln(T_L/T_C)$, the uncertainty in the line optical depth is given by $\delta\tau = \exp(\tau) \times \delta(T_L/T_C)$. Our computation yields an OPR equal to $1.1 \pm 0.4$ with the corresponding column densities $N_{\text{ortho}} = (8.7 \pm 2.1) \times 10^{11} \, \text{cm}^{-2}$ and $N_{\text{para}} = (7.8 \pm 2.6) \times 10^{11} \, \text{cm}^{-2}$. All errors here are $1\sigma$. Both lines are optically thin and their $T_{\text{ex}}$ are lower than 5 K. We note that decreasing the density and/or the kinetic temperature does not change the OPR by more than 10%. Therefore, the OPR is lower than 2.4 at a $3\sigma$ level of uncertainty (where we added the $3\sigma$ statistical error and the mentioned 10% to the 1.1 value). We assumed (see Sect. 2) that the relative gains to the lower and upper sidebands are equal. Since we do not have any information about the sideband ratio at the frequency of the D$_2$O line, we can only introduce a maximum uncertainty of 16%, corresponding to the overall calibration budget for band 1b. The resulting upper limit to the OPR is therefore increased to about 2.6. Figure 4 shows the measured OPR interval and the thermal equilibrium as a function of the gas temperature.

### 4. Conclusions

As discussed in Sect. 3, the gas absorbing the D$_2$O line is located more than 900 AU from the center and has a temperature that is lower than 30 K. The comparison between the upper value of the measured D$_2$O OPR and the thermal equilibrium value shows that they are consistent with a gas at a temperature of higher than about 15 K (at a $3\sigma$ level of confidence), and, therefore, with the assumed absorbing gas location. On the other hand, the D$_2$O gas may have formed in a previous phase, where the gas was colder, and, in this case, it means that it had the time to thermalise to the Boltzmann value. Unfortunately, given our poor knowledge of the mechanisms that can exchange the D$_2$O spins (see the Introduction), it is difficult here to infer the timescale for this change and, consequently, to provide a lower limit to the object age. On the other hand, the relatively large uncertainty in the OPR derived here does not allow either to exclude a non-thermal equilibrium situation. Higher signal-to-noise ratio observations will be needed to lower the uncertainty in the OPR value and derive a more robust result.

Using the density and temperature profiles of the envelope of IRAS 16293 by Crimier et al. (2010), the column density of the gas colder than 30 K is about $1 \times 10^{22} \, \text{cm}^{-2}$. Therefore, the D$_2$O abundance (with respect to H$_2$) is about $2 \times 10^{-11}$. An estimate of the water abundance profile will soon be available with the HIFI observations with a much higher spatial and spectral resolution than the one provided by the ISO observations (Ceccarelli et al. 2000). The D$_2$O molecules might form with one OPR, but then could freeze out on grain surfaces that could modify the ratio and then become desorbed. Owing to the high uncertainty in the H$_2$O abundance, we cannot at the time being completely exclude or confirm that formation can be described by grain surface chemistry. A modeling of the OPR evolution is beyond the scope of the present letter. With an improved calibration and better understanding of the instrumental effects, a more accurate determination of the D$_2$O OPR in this source and potentially other sources will be possible. ALMA may also hopefully yield an answer in the near future with the observation of cold D$_2$O with a higher spatial resolution.

To summarize, this Letter presents the first tentative estimate of the OPR for the D$_2$O molecule, demonstrating the outstanding capabilities of the HIFI instrument. The poor knowledge of the exchange mechanisms of the nuclear spins and the relatively large error in the derived OPR prevent us from drawing firm conclusions about the formation of heavy water at that time.

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### References
