



UvA-DARE (Digital Academic Repository)

Water in low-mass star-forming regions with Herschel: HIFI spectroscopy of NGC 1333

Kristensen, L.E.; Visser, R.; van Dishoeck, E.F.; Yldz, U.A.; Doty, S.D.; Herczeg, G.J.; Liu, F.-C.; Parise, B.; Jørgensen, J.K.; van Kempen, T.A.; Brinch, C.; Wampfler, S.F.; Bruderer, S.; Benz, A.O.; Hogerheijde, M.R.; Deul, E.; Bachiller, R.; Baudry, A.; Benedettini, M.; Bergin, E.A.; Bjerkeli, P.; Blake, G.A.; Bontemps, S.; Braine, J.; Caselli, P.; Cernicharo, J.; Codella, C.; Daniel, F.; de Graauw, T.; Di Giorgio, A.M.; Dominik, C.; Encrenaz, P.; Fich, M.; Fuente, A.; Giannini, T.; Goicoechea, J.R.; Helmich, F.; Herpin, F.; Jacq, T.; Johnstone, D.; Kaufman, M.J.; Larsson, B.; Lis, D.; Liseau, R.; Marseille, M.; McCoey, C.; Melnick, G.; Neufeld, D.; Nisini, B.; Olberg, M.; Pearson, J.C.; Plume, R.; Risacher, C.; Santiago-García, J.; Saraceno, P.; Shipman, R.; Tafalla, M.; Tielens, A.G.G.M.; van der Tak, F.; Wyrowski, F.; Beintema, D.; de Jonge, A.; Dieleman, P.; Ossenkopf, V.; Roelfsema, P.; Stutzki, J.; Whyborn, N.

Published in:
Astronomy & Astrophysics

DOI:
[10.1051/0004-6361/201015100](https://doi.org/10.1051/0004-6361/201015100)

[Link to publication](#)

Citation for published version (APA):

Kristensen, L. E., Visser, R., van Dishoeck, E. F., Yldz, U. A., Doty, S. D., Herczeg, G. J., ... Whyborn, N. (2010). Water in low-mass star-forming regions with Herschel: HIFI spectroscopy of NGC 1333. *Astronomy & Astrophysics*, 521, L30. DOI: 10.1051/0004-6361/201015100

General rights

It is not permitted to download or to forward/distribute the text or part of it without the consent of the author(s) and/or copyright holder(s), other than for strictly personal, individual use, unless the work is under an open content license (like Creative Commons).

Disclaimer/Complaints regulations

If you believe that digital publication of certain material infringes any of your rights or (privacy) interests, please let the Library know, stating your reasons. In case of a legitimate complaint, the Library will make the material inaccessible and/or remove it from the website. Please Ask the Library: <http://uba.uva.nl/en/contact>, or a letter to: Library of the University of Amsterdam, Secretariat, Singel 425, 1012 WP Amsterdam, The Netherlands. You will be contacted as soon as possible.

UvA-DARE is a service provided by the library of the University of Amsterdam (<http://dare.uva.nl>)

LETTER TO THE EDITOR

Water in low-mass star-forming regions with *Herschel**

HIFI spectroscopy of NGC 1333**

L. E. Kristensen¹, R. Visser¹, E. F. van Dishoeck^{1,2}, U. A. Yıldız¹, S. D. Doty³, G. J. Herczeg², F.-C. Liu⁴, B. Parise⁴, J. K. Jørgensen⁵, T. A. van Kempen⁶, C. Brinch¹, S. F. Wampfler⁷, S. Bruderer⁷, A. O. Benz⁷, M. R. Hogerheijde¹, E. Deul¹, R. Bachiller⁸, A. Baudry⁹, M. Benedettini¹⁰, E. A. Bergin¹¹, P. Bjerkeli¹², G. A. Blake¹³, S. Bontemps⁹, J. Braine⁹, P. Caselli^{14,15}, J. Cernicharo¹⁶, C. Codella¹⁵, F. Daniel¹⁶, Th. de Graauw¹⁷, A. M. di Giorgio¹⁰, C. Dominik^{18,19}, P. Encrenaz²⁰, M. Fich²¹, A. Fuente²², T. Giannini²³, J. R. Goicoechea¹⁶, F. Helmich¹⁷, F. Herpin⁹, T. Jacq⁹, D. Johnstone^{24,25}, M. J. Kaufman²⁶, B. Larsson²⁷, D. Lis²⁸, R. Liseau¹², M. Marseille¹⁷, C. McCoey^{21,29}, G. Melnick⁶, D. Neufeld³⁰, B. Nisini²³, M. Olberg¹², J. C. Pearson³¹, R. Plume³², C. Risacher¹⁷, J. Santiago-García³³, P. Saraceno¹⁰, R. Shipman¹⁷, M. Tafalla⁸, A. G. G. M. Tielens¹, F. van der Tak^{17,34}, F. Wyrowski⁴, D. Beintema¹⁷, A. de Jonge¹⁷, P. Dieleman¹⁷, V. Ossenkopf³⁵, P. Roelfsema¹⁷, J. Stutzki³⁵, and N. Whyborn³⁶

(Affiliations are available on page 5 of the online edition)

Received 31 May 2010 / Accepted 13 July 2010

ABSTRACT

“Water In Star-forming regions with *Herschel*” (WISH) is a key programme dedicated to studying the role of water and related species during the star-formation process and constraining the physical and chemical properties of young stellar objects. The Heterodyne Instrument for the Far-Infrared (HIFI) on the *Herschel* Space Observatory observed three deeply embedded protostars in the low-mass star-forming region NGC 1333 in several H₂¹⁶O, H₂¹⁸O, and CO transitions. Line profiles are resolved for five H₂¹⁶O transitions in each source, revealing them to be surprisingly complex. The line profiles are decomposed into broad (>20 km s⁻¹), medium-broad (~5–10 km s⁻¹), and narrow (<5 km s⁻¹) components. The H₂¹⁸O emission is only detected in broad 1₁₀–1₀₁ lines (>20 km s⁻¹), indicating that its physical origin is the same as for the broad H₂¹⁶O component. In one of the sources, IRAS4A, an inverse P Cygni profile is observed, a clear sign of infall in the envelope. From the line profiles alone, it is clear that the bulk of emission arises from shocks, both on small (≲1000 AU) and large scales along the outflow cavity walls (~10 000 AU). The H₂O line profiles are compared to CO line profiles to constrain the H₂O abundance as a function of velocity within these shocked regions. The H₂O/CO abundance ratios are measured to be in the range of ~0.1–1, corresponding to H₂O abundances of ~10⁻⁵–10⁻⁴ with respect to H₂. Approximately 5–10% of the gas is hot enough for all oxygen to be driven into water in warm post-shock gas, mostly at high velocities.

Key words. astrochemistry – stars: formation – ISM: molecules – ISM: jets and outflows – ISM: individual objects: NGC 1333

1. Introduction

In the deeply embedded phase of low-mass star formation, it is often only possible to trace the dynamics of gas in a young stellar object (YSO) by analysing resolved emission-line profiles. The various dynamical processes include infall from the surrounding envelope towards the central protostar, molecular outflows caused by jets ejected from the central object, and strong turbulence induced within the inner parts of the envelope by small-scale shocks (Arce et al. 2007; Jørgensen et al. 2007). One of the goals of the Water In Star-forming regions with *Herschel* (WISH) key programme is to use water as a probe of these processes and determine its abundance in the various components as a function of evolution (van Dishoeck et al., in prep.).

Spectrally resolved observations of the H₂O 1₁₀–1₀₁ line at 557 GHz with ODIN and SWAS towards low-mass star-forming

regions have revealed it to be broad, ~20 km s⁻¹, indicative of an origin in shocks (e.g., Bergin et al. 2003). Within the large beams (2' and 4'), where both the envelope and the entire outflow are present, outflow emission most likely dominates. Observations and subsequent modelling of the more highly excited H₂O lines with ISO-LWS were unable to distinguish between an origin in shocks or an infalling envelope (e.g., Ceccarelli et al. 1996; Nisini et al. 2002; Maret et al. 2002). *Herschel*/HIFI has a much higher sensitivity, higher spectral resolution, and smaller beam than previous space-based missions, thus is perfectly suited to addressing this question. Complementary CO data presented by Yıldız et al. (2010) are used to constrain the role of the envelope and determine outflow temperatures and densities.

NGC 1333 is a well-studied region of clustered, low-mass star formation at a distance of 235 pc (Hirota et al. 2008). In particular, the three deeply embedded, low-mass class 0 objects IRAS2A, IRAS4A, and IRAS4B have been observed extensively with ground-based submillimetre telescopes (e.g., Jørgensen et al. 2005; Maret et al. 2005) and interferometers (e.g., Di Francesco et al. 2001; Jørgensen et al. 2007). All sources have strong outflows extending over arcmin scales

* *Herschel* is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.

** Tables 2 and 3 (page 6) are only available in electronic form at <http://www.aanda.org>

Table 1. H₂O and H₂¹⁸O emission^a in the NGC 1333 sources^b.

Source	Transition	rms ^c (mK)	Medium		Broad	
			$T_{\text{MB}}^{\text{peak}}$ (K)	$\int T_{\text{MB}} dv$ (K km s ⁻¹)	$T_{\text{MB}}^{\text{peak}}$ (K)	$\int T_{\text{MB}} dv$ (K km s ⁻¹)
IRAS 2A	H ₂ O 1 ₁₁ -0 ₀₀	22	0.22	2.7	0.10	4.1
	2 ₀₂ -1 ₁₁	25	0.25	2.5	0.12	5.6
	2 ₁₁ -2 ₀₂	23	0.16	1.9	0.06	2.7
	3 ₁₂ -3 ₀₃	79	0.14	1.5	0.06	2.7
	3 ₁₂ -2 ₂₁	52	0.13	1.5	0.08	3.8
	H ₂ ¹⁸ O 1 ₁₀ -1 ₀₁	2	...	<0.01	0.01	0.14
IRAS 4A	H ₂ O 1 ₁₁ -0 ₀₀	23	0.36	4.9	0.35	15.3
	2 ₀₂ -1 ₁₁	24	0.34	3.8	0.45	18.0
	2 ₁₁ -2 ₀₂	23	0.14	0.8	0.43	14.0
IRAS 4B	H ₂ O 1 ₁₁ -0 ₀₀	23	0.07	0.4	0.32	13.6
	H ₂ ¹⁸ O 1 ₁₀ -1 ₀₁	3	...	<0.01	0.02	0.43
	H ₂ ¹⁸ O 1 ₁₁ -0 ₀₀	23	...	<0.06	...	<0.09
$FWHM^d$ (km s ⁻¹)			10.7 ± 0.8		42 ± 3	
v_{LSR}^d (km s ⁻¹)			+10.7 ± 0.8		-2.3 ± 3.4	
IRAS 2A	H ₂ O 1 ₁₁ -0 ₀₀	29	1.1	5.1	0.54	14.0
	2 ₀₂ -1 ₁₁	23	0.63	3.5	0.65	17.6
	2 ₁₁ -2 ₀₂	17	0.37	1.8	0.40	10.2
	3 ₁₂ -2 ₂₁	150	0.32	1.3	0.76	17.5
	H ₂ ¹⁸ O 1 ₁₀ -1 ₀₁	3	...	<0.01	0.02	0.43
	H ₂ ¹⁸ O 1 ₁₁ -0 ₀₀	16	...	<0.03	...	<0.04
$FWHM^d$ (km s ⁻¹)			4.6 ± 0.5		24 ± 2	
v_{LSR}^d (km s ⁻¹)			+8.1 ± 0.3		+8.0 ± 0.5	

Notes. ^(a) Obtained from Gaussian fits to each component. In the case of H₂O 1₁₁-0₀₀, this includes extrapolation over the absorption feature. ^(b) The coordinates used are for IRAS2A: 03:28:55.6; +31:14:37.1, IRAS4A: 03:29:10.5; +31:13:30.9, IRAS4B: +03:29:12.0; +31:13:08.1 (J2000). ^(d) Measured in 0.5 km s⁻¹ bins. ^(e) Intensity-weighted average of values determined from Gaussian fits of H₂¹⁶O emission lines. Uncertainties include statistical errors only.

(>15 000 AU). Both IRAS4A and 4B consist of multiple protostars (e.g., Choi 2005). Because of the similarities between the three sources in terms of luminosity (20, 5.8, and 3.8 L_{\odot}), envelope mass (1.0, 4.5, and 2.9 M_{\odot} ; Jørgensen et al. 2009) and presumably also age, they provide ideal grounds for comparing YSOs in the same region.

2. Observations and results

Three sources in NGC 1333, IRAS2A, IRAS4A, and IRAS4B, were observed with HIFI (de Graauw et al. 2010) on *Herschel* (Pilbratt et al. 2010) on March 3–15, 2010 in dual beam switch mode in bands 1, 3, 4, and 5 with a nod of 3'. Observations detected several transitions of H₂O and H₂¹⁸O in the range $E_u/k_B \approx 50$ –250 K (Table 2). Diffraction-limited beam sizes were in the range 19–40'' (4500–9500 AU). In general, the calibration is expected to be accurate to ~20% and the pointing to ~2''. Data were reduced with HIPE 3.0. A main-beam efficiency of 0.74 was used throughout. Subsequent analysis was performed in CLASS. The rms was in the range 3–150 mK in 0.5 km s⁻¹ bins. Linear baselines were subtracted from all spectra, except around 750 GHz (corresponding to the H₂O 2₁₁-2₀₂ transition) where higher-order polynomials are required. A difference in rms was always seen between the H- and V-polarizations, with the rms in the H-polarization being lower. In cases where the difference exceeded 30% and qualitative differences appear in the line profile, the V-polarization was discarded, otherwise the spectra were averaged.

All targeted lines of H₂¹⁶O were detected and are listed in Table 1 and Fig. 1. The 1₁₀-1₀₁ transition at 557 GHz was not observed before the sources moved out of visibility. The H₂¹⁸O 1₁₀-1₀₁ line was detected in all sources (Fig. 2), although

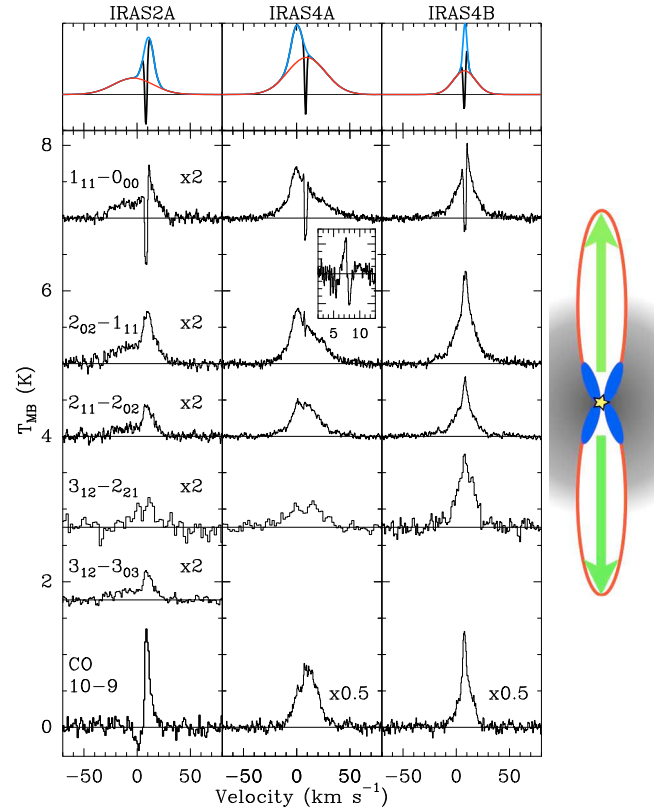


Fig. 1. H₂O spectra of the three NGC 1333 sources. CO 10–9 is shown for comparison (Yıldız et al. 2010); the CO 10–9 emission in IRAS2A is affected by chopping into outflow material. The top panel shows the decomposition into broad (red), medium (blue), and narrow (black) components. The cartoon illustrates the physical origin of each component. The inset shows a zoom on the inverse P Cygni profile in the H₂O 2₀₂-1₁₁ line of IRAS4A, where the other components have been subtracted; the vertical scale ranges from –0.3 to 0.3 K.

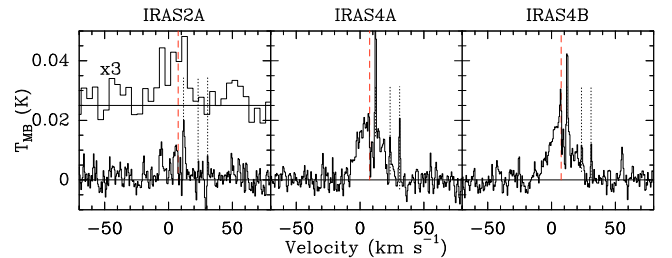


Fig. 2. H₂¹⁸O 1₁₀-1₀₁ spectra of the three NGC 1333 sources along with the CH 536 GHz triplet from the lower sideband (dotted lines). Spectra are shown for a channel size of 0.25 km s⁻¹. The spectrum of IRAS2A has been rebinned to 4 km s⁻¹ to illustrate the detection of a broad feature. The red line shows the source velocity at $v_{\text{LSR}} \sim +7.5$ km s⁻¹.

the detection in IRAS2A was weak ($\sim 5\sigma = 0.13$ K km s⁻¹). This line is superposed on the ground-state CH triplet at 536 GHz, observed in the lower sideband (Fig. 2). Neither the H₂¹⁸O 1₁₁-0₀₀ nor the 2₀₂-1₁₁ line in IRAS2A is detected down to $\sigma < 0.06$ K km s⁻¹.

The H₂O lines exhibit multiple components: a broad emission component ($FWHM > 20$ km s⁻¹) sometimes offset from the source velocity ($v_{\text{LSR}} = +7.2$ – 7.7 km s⁻¹); a medium-broad emission component ($FWHM \sim 5$ – 10 km s⁻¹); and a deep, narrow absorption component ($FWHM \sim 2$ km s⁻¹) seen at the source velocity. The individual components are all reproduced well by Gaussian functions. The absorption is only seen in the H₂O 1₁₁-0₀₀ line and is saturated in IRAS2A and IRAS4A. In IRAS4B, the absorption extends below the continuum level,

but is not saturated. Furthermore, the IRAS4A spectrum of the $2_{02}-1_{11}$ line exhibits an inverse P Cygni profile. The shape of the lines is the same within a source; only the relative contribution between the broad and medium components changes. For example, in IRAS2A the ratio of the peak intensities is ~ 2 , independent of the line, whereas in IRAS4A it ranges from 1 to 2. The H_2^{18}O line profiles compare well to the broad component seen in H_2O , i.e., similar $FWHM > 20 \text{ km s}^{-1}$ and velocity offset. The width is much larger than isotopologue emission of, e.g., C^{18}O ($\sim 1-2 \text{ km s}^{-1}$) and is centred on the source velocity (Yildiz et al. 2010). The medium and narrow components are not seen in the H_2^{18}O $1_{10}-1_{01}$ spectra down to an rms of 2–3 mK in 0.5 km s^{-1} bins.

The upper limits to the H_2^{18}O $1_{11}-0_{00}$ line are invaluable for estimating upper limits to the optical depth, τ . In the following, the limit on τ is derived for the integrated intensity; in the line wings, τ is most likely lower (Yildiz et al. 2010). In the broad component, the limit ranges from 0.4 (IRAS4B) to 2 (IRAS2A), whereas it ranges from 1.1 (IRAS4B) to 2.7 (IRAS2A) for the medium component of the H_2^{16}O $1_{11}-0_{00}$ line. Performing the same analysis to the upper limit on the H_2^{18}O $2_{02}-1_{11}$ line observed in IRAS2A, infers an upper limit to the optical depth of H_2^{16}O $2_{02}-1_{11}$ of 1.5 for the medium component and 1.9 for the broad. Thus it is likely that neither the broad nor the medium components are very optically thick.

3. Discussion

Many physical components in a YSO are directly traced by the line profiles presented here, including the infalling envelope and shocks along the cavity walls. In the following, each component is discussed in detail, and the H_2O abundance is estimated in the various physical components.

3.1. Line profiles

The most prominent feature of all the observed line profiles is their width. All line wings span a range of velocities of $\sim 40-70 \text{ km s}^{-1}$ at their base. The width alone indicates that the bulk of the H_2O emission originates in shocks along the cavity walls, also called shells, seen traditionally as the standard high-velocity component in CO outflow data, but with broader line-widths due to water enhancement at higher velocities (Sect. 3.2 Bachiller et al. 1990; Santiago-García et al. 2009). The shocks release water from the grains by means of sputtering and in high-temperature regions all free oxygen is driven into water. The shocked regions may be illuminated by FUV radiation originating in the star-disk boundary layer, thus further enhancing the water abundance by means of photodesorption. The broad emission seen in the H_2^{18}O $1_{10}-1_{01}$ line arises in the same shocks (see cartoon in Fig. 1).

The medium components ($FWHM \sim 5-10 \text{ km s}^{-1}$) are most likely also caused by shocks, although presumably on a smaller spatial scale and in denser material than the shocks discussed above. For example, the medium component in IRAS2A is seen in other grain-product species such as CH_3OH (Jørgensen et al. 2005; Maret et al. 2005, Fig. 3), where emission arises from a compact region ($<1''$, i.e., $<250 \text{ AU}$) centred on the source (Jørgensen et al. 2007), and the same is likely true for the medium H_2O component in that source. In interferometric observations of IRAS4A, a small (\sim few arcsec) blue-shifted outflow knot of similar width has been identified in, e.g., SiO and SO (Choi 2005; Jørgensen et al. 2007). Small-scale structures exist in the other sources as well, which may produce the medium components.

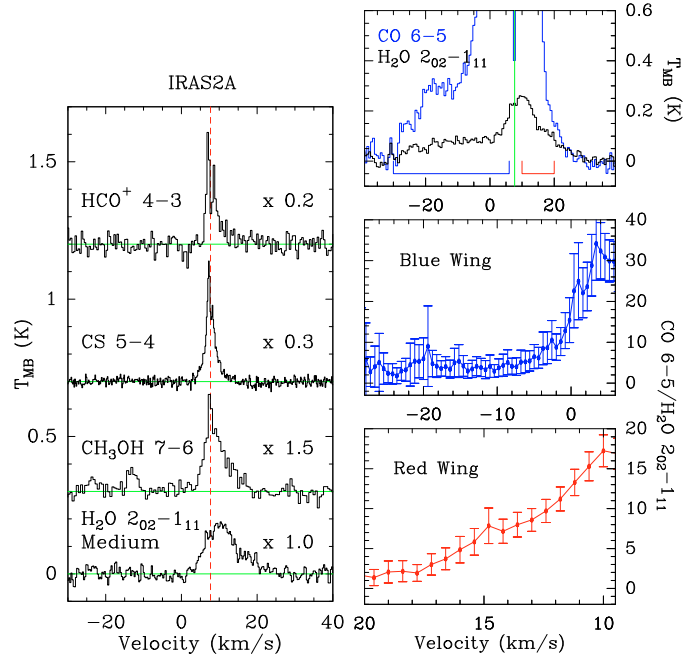


Fig. 3. *Left.* Comparison between the medium component in IRAS2A and other species observed with ground-based telescopes. The broad component has been subtracted for easy comparison. The vertical red line indicates the source velocity at $+7.7 \text{ km s}^{-1}$. *Right.* Comparison between H_2O $2_{02}-1_{11}$ and CO $6-5$ obtained with APEX-CHAMP⁺, and emission ratios for the blue- and red-shifted outflow lobes.

The H_2O $2_{02}-1_{11}$ spectrum of IRAS4A shows an inverse P Cygni profile, a clear sign of infall also detected in other molecular tracers using interferometer observations (Di Francesco et al. 2001; Jørgensen et al. 2007). This infall signature is also tentatively seen in the $1_{11}-0_{00}$ line, but here the absorption from the outer envelope dominates and little is left of the blue emission peak. The signature is not seen in higher-excitation lines. The separation of the emission and absorption peaks is $\sim 0.8 \text{ km s}^{-1}$, whereas it is $\sim 1.5 \text{ km s}^{-1}$ in the observations of Di Francesco et al. (2001) and larger in the observations by Jørgensen et al. (2007), indicating that the infall observed in H_2O $2_{02}-1_{11}$ takes place over larger spatial scales.

The passively heated envelope is seen in ground-based observations of high-density tracers to produce narrow emission, $<3 \text{ km s}^{-1}$, which may be self-absorbed (Fig. 3). For water, this type of emission is not seen in any of the sources; the medium component is broader by a factor of 2–3 with respect to what is expected from the envelope. The absorption seen in all three sources is attributed to cold gas in the outer parts of the envelope. Using interferometric observations, Jørgensen & van Dishoeck (2010) detected compact, narrow ($\sim 1 \text{ km s}^{-1}$) emission in the H_2^{18}O $3_{13}-2_{20}$ line in IRAS4B possibly originating in the circumstellar disk. Scaling the observed emission to the transitions observed here by assuming $T_{\text{rot}} = 170 \text{ K}$ (Watson et al. 2007), the expected emission is typically less than 10% of the rms for any given transition. Hence, when extrapolated to the disks surrounding IRAS2A and 4A, the disk contribution to the H_2O emission probed by HIFI is negligible. The H_2O excitation temperature of the broad component is $220 \pm 30 \text{ K}$, comparable to that found by Watson et al. (2007), but the inferred column density is a factor of 100 higher. Thus, the mid-infrared lines seen by Watson et al. may come from the same broad outflowing gas found by HIFI, provided the mid-infrared lines experience a factor of 100 more extinction.

3.2. Abundances

3.2.1. Shocks: H₂O/CO

The observed broad components are compared directly with HIFI observations of CO 10–9 (Yıldız et al. 2010), because the width and position of the lines are similar and they were obtained using approximately the same beamsizes (22'' versus 19''). The exception is for IRAS2A, where the blue line wing is not observed. The advantage is that no detailed models are required to account for the H₂O/CO abundance, as long as the lines are optically thin, in particular the emission from the wings. The abundance ratio is estimated for various temperatures by using the RADEX escape probability code (van der Tak et al. 2007). The density is assumed to be 10⁵ cm⁻³, appropriate for the large-scale core. If the emission is optically thin, the abundance ratio scales linearly with density resulting in the same line ratio corresponding to a higher abundance ratio. There is little variation in the predicted ratio for $T > 150$ K, the typical temperature inferred by Yıldız et al. (2010). The line ratios and abundance ratios are listed as a function of velocity in Table 3.

The abundance ratio increases with increasing velocity from H₂O/CO of ~0.2 near the line centre to H₂O/CO ≥ 1 in the line wings of all sources for velocity offsets larger than 15 km s⁻¹ with respect to that of the source (Fig. 3). Assuming that the CO abundance is 10⁻⁴, the H₂O abundances are in the range of ~10⁻⁵–10⁻⁴. Only at high velocities is the temperature high enough for oxygen to be driven into water by means of the neutral-neutral reaction $O + H_2 \rightarrow OH + H$; $OH + H_2 \rightarrow H_2O$. The same result was found in the massive outflow in Orion-KL (Franklin et al. 2008), where less than 1% of the gas in the outflow experiences this high-temperature phase. The fraction of gas for which the H₂O/CO abundance is > 1 is ~5–10% for the sources observed here.

For IRAS2A, a deep spectrum of CO 6–5 obtained with CHAMP⁺ on APEX simultaneously with observations of HDO 1₁₁–0₀₀ (Liu et al., in prep.) shows the same morphology in terms of a broad and medium component (Fig. 3). Furthermore, the velocity offset and *FWHM* are the same as for H₂O suggesting that the line profiles are not unique to H₂O, although the broad component is far more prominent in H₂O. The ratio of peak intensities for the two components is ~2–3 in H₂O versus 10 in CO 6–5. Analysing the abundance ratio as a function of temperature shows that H₂O/CO ~ 0.1–1 for $T > 150$ K (Table 3), consistent with what is found for CO 10–9.

3.2.2. Envelope

The simplest way to constrain the H₂O abundance in the outer envelope is with calculations using RADEX on the narrow absorption in the 1₁₁–0₀₀ line. The absorption is optically thick – which requires a para-H₂O column density of $> 10^{13}$ cm⁻² if one assumes typical values for T and $n(H_2)$ of 15 K and 10⁵ cm⁻³. With a pencil-beam H₂ column of ~10²³ cm⁻² (Jørgensen et al. 2002), the total H₂O abundance in the outer envelope is $\geq 10^{-10}$.

For IRAS2A and 4A, the H₂O abundance was further constrained using radiative transfer models. The setup is a spherical envelope with density and temperature profiles constrained from continuum data (Jørgensen et al. 2009), an infall velocity profile $v = (2 \text{ km s}^{-1})(r/r_{in})^{-1/2}$, and a Doppler parameter $b = 0.8 \text{ km s}^{-1}$. Line fluxes were computed with the new radiative transfer code LIME (Brinch & Hogerheijde, submitted). The models constrained the abundance of water in the outer envelope to be ~10⁻⁸. Lower values are insufficient to obtain saturated absorption in the 1₁₁–0₀₀ line, and ~10⁻⁸ is the highest abundance where the resulting narrow emission can be hidden in the

observed higher-excitation H₂O lines. The models predict that the H₂O emission from the warm inner envelope ($r \lesssim 100$ AU) is optically thick, hence no constraints can be obtained from the H₂O spectra on the inner abundance. However, the lack of narrow H₂¹⁸O emission infers an upper limit on the H₂O abundance of ~10⁻⁵ (Visser et al., in prep.).

4. Conclusions

These observations represent one of the first steps towards understanding the formation and excitation of water in low-mass star-forming regions by means of resolved line profiles. The three sources have remarkably similar line profiles. Both the H₂¹⁶O and H₂¹⁸O lines are very broad, indicating that the bulk of the emission originates in shocked gas. The broad emission also highlights that water is a far more reliable dynamical tracer than, e.g., CO. Comparing C¹⁸O to H₂¹⁸O emission and line profiles indicates that the H₂O/CO abundance is high in outflows and low in the envelope. Additional modelling of the emission, should be able to constrain the total amount of water in the envelope and outflowing gas, thus test the high-temperature gas-phase chemistry models for the origin of water. This will be performed for a total sample of the 29 low-mass YSOs to be observed within the WISH key programme.

Acknowledgements. This work is made possible thanks to the HIFI guaranteed time programme. HIFI has been designed and built by a consortium of institutes and university departments from across Europe, Canada and the US under the leadership of SRON Netherlands Institute for Space Research, Groningen, The Netherlands with major contributions from Germany, France and the US. Consortium members are: Canada: CSA, UWaterloo; France: CESR, LAB, LERMA, IRAM; Germany: KOSMA, MPIfR, MPS; Ireland: NUI Maynooth; Italy: ASI, IFSI-INAF, Arcetri-INAF; Netherlands: SRON, TUD; Poland: CAMK, CBK; Spain: Observatorio Astronómico Nacional (IGN), Centro de Astrobiología (CSIC-INTA); Sweden: Chalmers University of Technology – MC2, RSS & GARD, Onsala Space Observatory, Swedish National Space Board, Stockholm University – Stockholm Observatory; Switzerland: ETH Zürich, FHNW; USA: Caltech, JPL, NHSC. HIPE is a joint development by the *Herschel* Science Ground Segment Consortium, consisting of ESA, the NASA *Herschel* Science Center, and the HIFI, PACS and SPIRE consortia. We thank many funding agencies for financial support.

References

- Arce, H. G., Shepherd, D., Gueth, F., et al. 2007, in *Protostars and Planets V*, ed. B. Reipurth, D. Jewitt, & K. Keil, 245
- Bachiller, R., Martín-Pintado, J., Tafalla, M., Cernicharo, J., & Lazareff, B. 1990, *A&A*, 231, 174
- Bergin, E. A., Kaufman, M. J., Melnick, G. J., Snell, R. L., & Howe, J. E. 2003, *ApJ*, 582, 830
- Ceccarelli, C., Hollenbach, D. J., & Tielens, A. G. G. M. 1996, *ApJ*, 471, 400
- Choi, M. 2005, *ApJ*, 630, 976
- de Graauw, Th., Helmich, F. P., Phillips, T. G., et al. 2010, *A&A*, 518, L6
- Di Francesco, J., Myers, P. C., Wilner, D. J., Ohashi, N., & Mardones, D. 2001, *ApJ*, 562, 770
- Franklin, J., Snell, R. L., Kaufman, M. J., et al. 2008, *ApJ*, 674, 1015
- Hirota, T., Bushimata, T., Choi, Y. K., et al. 2008, *PASJ*, 60, 37
- Jørgensen, J. K., & van Dishoeck, E. F. 2010, *ApJ*, 710, L72
- Jørgensen, J. K., Schöier, F. L., & van Dishoeck, E. F. 2002, *A&A*, 389, 908
- Jørgensen, J. K., Schöier, F. L., & van Dishoeck, E. F. 2005, *A&A*, 437, 501
- Jørgensen, J. K., Bourke, T. L., Myers, P. C., et al. 2007, *ApJ*, 659, 479
- Jørgensen, J. K., van Dishoeck, E. F., Visser, R., et al. 2009, *A&A*, 507, 861
- Maret, S., Ceccarelli, C., Caux, E., Tielens, A. G. G. M., & Castets, A. 2002, *A&A*, 395, 573
- Maret, S., Ceccarelli, C., Tielens, A. G. G. M., et al. 2005, *A&A*, 442, 527
- Nisini, B., Giannini, T., & Lorenzetti, D. 2002, *ApJ*, 574, 246
- Pickett, H. M., Poynter, I. R. L., Cohen, E. A., et al. 1998, *J. Quant. Spec. Radiat. Transf.*, 60, 883
- Pilbratt, G. L., Riedinger, J. R., Passvogel, T., et al. 2010, *A&A*, 518, L1
- Santiago-García, J., Tafalla, M., Johnstone, D., & Bachiller, R. 2009, *A&A*, 495, 169
- van der Tak, F. F. S., Black, J. H., Schöier, F. L., Jansen, D. J., & van Dishoeck, E. F. 2007, *A&A*, 468, 627
- Watson, D. M., Bohac, C. J., Hull, C., et al. 2007, *Nature*, 448, 1026
- Yıldız, U. A., van Dishoeck, E. F., Kristensen, L. E., et al. 2010, *A&A*, 521, L40

-
- ¹ Leiden Observatory, Leiden University, PO Box 9513, 2300 RA Leiden, The Netherlands
- ² Max Planck Institut für Extraterrestrische Physik, Giessenbachstrasse 1, 85748 Garching, Germany
- ³ Department of Physics and Astronomy, Denison University, Granville, OH, 43023, USA
- ⁴ Max-Planck-Institut für Radioastronomie, Auf dem Hügel 69, 53121 Bonn, Germany
- ⁵ Centre for Star and Planet Formation, Natural History Museum of Denmark, University of Copenhagen, Øster Voldgade 5–7, 1350 Copenhagen K., Denmark
- ⁶ Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, MS 42, Cambridge, MA 02138, USA
- ⁷ Institute of Astronomy, ETH Zurich, 8093 Zurich, Switzerland
- ⁸ Observatorio Astronómico Nacional (IGN), Calle Alfonso XII 3, 28014 Madrid, Spain
- ⁹ Université de Bordeaux, Laboratoire d’Astrophysique de Bordeaux, France; CNRS/INSU, UMR 5804, Floirac, France
- ¹⁰ INAF – Istituto di Fisica dello Spazio Interplanetario, Area di Ricerca di Tor Vergata, via Fosso del Cavaliere 100, 00133 Roma, Italy
- ¹¹ Department of Astronomy, University of Michigan, 500 Church Street, Ann Arbor, MI 48109-1042, USA
- ¹² Department of Radio and Space Science, Chalmers University of Technology, Onsala Space Observatory, 439 92 Onsala, Sweden
- ¹³ California Institute of Technology, Division of Geological and Planetary Sciences, MS 150-21, Pasadena, CA 91125, USA
- ¹⁴ School of Physics and Astronomy, University of Leeds, Leeds LS2 9JT, UK
- ¹⁵ INAF – Osservatorio Astrofisico di Arcetri, Largo E. Fermi 5, 50125 Firenze, Italy
- ¹⁶ Centro de Astrobiología, Departamento de Astrofísica, CSIC-INTA, Carretera de Ajalvir, Km 4, Torrejón de Ardoz, 28850 Madrid, Spain
- ¹⁷ SRON Netherlands Institute for Space Research, PO Box 800, 9700 AV Groningen, The Netherlands
- ¹⁸ Astronomical Institute Anton Pannekoek, University of Amsterdam, Kruislaan 403, 1098 SJ Amsterdam, The Netherlands
- ¹⁹ Department of Astrophysics/IMAPP, Radboud University Nijmegen, PO Box 9010, 6500 GL Nijmegen, The Netherlands
- ²⁰ LERMA and UMR 8112 du CNRS, Observatoire de Paris, 61 Av. de l’Observatoire, 75014 Paris, France
- ²¹ University of Waterloo, Department of Physics and Astronomy, Waterloo, Ontario, Canada
- ²² Observatorio Astronómico Nacional, Apartado 112, 28803 Alcalá de Henares, Spain
- ²³ INAF – Osservatorio Astronomico di Roma, 00040 Monte Porzio catone, Italy
- ²⁴ National Research Council Canada, Herzberg Institute of Astrophysics, 5071 West Saanich Road, Victoria, BC V9E 2E7, Canada
- ²⁵ Department of Physics and Astronomy, University of Victoria, Victoria, BC V8P 1A1, Canada
- ²⁶ Department of Physics and Astronomy, San Jose State University, One Washington Square, San Jose, CA 95192, USA
- ²⁷ Department of Astronomy, Stockholm University, AlbaNova, 106 91 Stockholm, Sweden
- ²⁸ California Institute of Technology, Cahill Center for Astronomy and Astrophysics, MS 301-17, Pasadena, CA 91125, USA
- ²⁹ University of Western Ontario, Department of Physics & Astronomy, London, Ontario, N6A 3K7, Canada
- ³⁰ Department of Physics and Astronomy, Johns Hopkins University, 3400 North Charles Street, Baltimore, MD 21218, USA
- ³¹ Jet Propulsion Laboratory, California Institute of Technology, Pasadena, CA 91109, USA
- ³² Department of Physics and Astronomy, University of Calgary, Calgary, T2N 1N4, AB, Canada
- ³³ Instituto de RadioAstronomía Milimétrica, Avenida Divina Pastora, 7, Núcleo Central E 18012 Granada, Spain
- ³⁴ Kapteyn Astronomical Institute, University of Groningen, PO Box 800, 9700 AV, Groningen, The Netherlands
- ³⁵ KOSMA, I. Physik. Institut, Universität zu Köln, Zùlpicher Str. 77, 50937 Köln, Germany
- ³⁶ Atacama Large Millimeter/Submillimeter Array, Joint ALMA Office, Santiago, Chile

Table 2. Observed H₂O, H₂¹⁸O and CH transitions^a.

Transition		ν (GHz)	λ (μm)	E_u/k_B (K)	A (10^{-3} s^{-1})	Beam ($''$)	t_{int}^b (min.)
H ₂ O	1 ₁₁ -0 ₀₀	1113.34	269.27	53.4	18.42	19	43.5
	2 ₀₂ -1 ₁₁	987.93	303.46	100.8	5.84	22	23.3
	2 ₁₁ -2 ₀₂	752.03	398.64	136.9	7.06	29	18.4
	3 ₁₂ -3 ₀₃	1097.37	273.19	249.4	16.48	20	32.4
	3 ₁₂ -2 ₂₁	1153.13	259.98	249.4	2.63	19	13.0
H ₂ ¹⁸ O	1 ₁₀ -1 ₀₁	547.68	547.39	60.5	3.59	39	64.3
	1 ₁₁ -0 ₀₀ ^c	1101.70	272.12	52.9	21.27	20	43.5
	2 ₀₂ -1 ₁₁	994.68	301.40	100.7	7.05	22	46.7
	3 ₁₂ -3 ₀₃ ^c	1095.16	273.74	289.7	22.12	20	32.4
CH ^d	3/2, 2 ⁻ -1/2, 1 ⁺	536.76	558.52	25.8	0.66	39	
	3/2, 1 ⁻ -1/2, 1 ⁺	536.78	558.50	25.8	0.23	39	
	3/2, 1 ⁻ -1/2, 0 ⁺	536.80	558.48	25.8	0.46	39	

Notes. ^(a) From the JPL database of molecular spectroscopy (Pickett et al. 1998); ^(b) total on + off integration time; ^(c) observed in the same setting as the main isotopologue; ^(d) observed with H₂¹⁸O 1₁₀-1₀₁.

Table 3. CO 6-5 and CO 10-9/H₂O 2₀₂-1₁₁ line ratios in 5 km s⁻¹ intervals and corresponding abundance ratio for $T > 150$ K and $n = 10^5$ cm⁻³.

dv_{LSR} (km s ⁻¹)	IRAS2A			IRAS4A		IRAS4B		
	CO 6-5/ H ₂ O 2 ₀₂ -1 ₁₁	$x(\text{H}_2\text{O})/$ $x(\text{CO})$	CO 10-9/ H ₂ O 2 ₀₂ -1 ₁₁	$x(\text{H}_2\text{O})/$ $x(\text{CO})$	CO 10-9/ H ₂ O 2 ₀₂ -1 ₁₁	$x(\text{H}_2\text{O})/$ $x(\text{CO})$	CO 10-9/ H ₂ O 2 ₀₂ -1 ₁₁	$x(\text{H}_2\text{O})/$ $x(\text{CO})$
-20--15	5.0	0.34	0.8	1.11
-15--10	3.8	0.45	2.0	0.43	0.9	1.00
-10--5	4.6	0.37	2.8	0.31	1.4	0.64
-5-0	9.3	0.18	2.4	0.36	1.7	0.50
0-5	26.6	0.06	2.9	0.29	2.3	0.37
5-10	17.0	0.10	3.4	0.26	3.9	0.22	2.9	0.29
10-15	11.1	0.15	3.2	0.27	3.3	0.26	2.1	0.42
15-20	3.5	0.48	0.9	1.00	2.4	0.36	1.6	0.53
20-25	1.4	1.25	0.4	2.00	1.0	0.83	1.1	0.77
25-30	0.8	1.11	0.5	1.67
30-35	0.9	2.00	0.9	1.00