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The CHESS spectral survey of star forming regions: Peering into the protostellar shock L1157-B1*

II. Shock dynamics


(Affiliations are available in the online edition)

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ABSTRACT

The outflow driven by the low-mass class 0 protostar L1157 is the prototype of the so-called chemically active outflows. The bright bowshock B1 in the southern outflow lobe is a privileged testbed of magneto-hydrodynamical (MHD) shock models, for which dynamical and chemical processes are strongly interdependent.

Aims. We present the first results of the unbiased spectral survey of the L1157-B1 bowshock, obtained in the framework of the key program “Chemical HERSchel Surveys of star forming regions” (CHESS). The main aim is to trace the warm and chemically enriched gas and to infer the excitation conditions in the shock region.

Methods. The CO 5-4 and o-H2O 1−0 lines have been detected at high-spectral resolution in the unbiased spectral survey of the HIFI-band 1b spectral window (555–636 GHz), presented by Codella et al. in this volume. Complementary ground-based observations in the submm window help establish the origin of the emission detected in the main-beam of HIFI and the physical conditions in the shock.

Results. Both lines exhibit broad wings, which extend to velocities much higher than reported up to now. We find that the molecular emission arises from two regions with distinct physical conditions: an extended, warm (100 K), dense (3×105 cm−3) component at low-velocity, which appears to dominate the flux of the water line at 179−3, which is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.

Key words. stars: formation – ISM: jets and outflows – ISM: individual objects: L1157

1. Introduction

Shocks in protostellar outflows play a crucial role in the molecular cloud evolution and star formation by transferring momentum and energy back to the ambient medium. There is mounting evidence that these shocks often involve a magnetic field, which can be greatly enhanced in shocks, even of moderate velocity. This occurs both from the sputtering of frozen water from grain mantles and through high-temperature sensitive reactions in the gas phase (Elitzur & de Jong 1978; Elitzur & Watson 1978; Bergin et al. 1998). Multi-transition observations of these two enhancements up to several orders of magnitude, as reported for various molecular species in “chemically active” outflows (Bachiller et al. 2001). Conversely, the magnetic field and the ionization fraction play an important role in controlling the size and the temperature of the ion-neutral drift zone. Because of the interplay between the dynamics and chemistry, the physics of MHD shocks requires a comprehensive picture of both the gas and dust physical conditions in the compressed region itself.

Along with H2, H2O and CO are two key-molecules predicted to dominate the cooling of MHD shocks (Kaufman & Neufeld 1996). The abundance of H2O in protostellar regions can be greatly enhanced in shocks, even of moderate velocity. This occurs both from the sputtering of frozen water from grain mantles and through high-temperature sensitive reactions in the gas phase (Elitzur & de Jong 1978; Elitzur & Watson 1978; Bergin et al. 1998). Multi-transition observations of these two

* Herschel is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.
molecules therefore serve as good probes of shock regions with various excitation conditions, and can be used to set stringent constraints on MHD shock models (Flower & Pineau des Forêts, 2001).

The heterodyne instrument HIFI onboard Herschel allows us to study with unprecedented sensitivity the chemical and dynamical evolution of protostellar shocks, at spectral and angular resolutions comparable to the best ground-based single-dish telescopes. This is the main goal of the spectral survey of L1157-B1, carried out in the guaranteed time key-project CHESS.

The source L1157-mm is a low-mass Class 0 protostar located at a distance estimated between 250 pc (Looney et al. 2007) and 440 pc (Vioiti 1969). It drives a spectacular bipolar outflow, which has been studied in detail at millimeter and far-infrared wavelengths. Mapping of the southern lobe of L1157 with the Plateau de Bure Interferometer (PdBI) reveals two limb-brightened cavities (Gueth et al. 1996), each one terminated by a strong bow shock, dubbed “B1” and “B2” respectively (Fig. 1), which are likely the result of episodic ejection events in a precessing, highly collimated jet. The spatial and kinematical structure of B1 has been modelled in great detail by various authors, making it the archetype of protostellar bowshocks in low-mass star-forming regions and the testbed of MHD shock models (Gusdorf et al. 2008a, b).

Here, we report on the emission lines of CO and H$_2$O detected in the low-frequency band of HIFI in the course of the CHESS spectral survey. From comparison with complementary observations, we discuss the origin of the emission, and based on a simple modelling of the source, we derive the water abundance in the shock region.

2. Observations and results

A full coverage of the band 1b at 20$^h$39$^m$10.2$^s$ +68$^\circ$01′10.5″ (J2000) in the bowshock B1 was carried out with the HIFI heterodyne instrument (de Graauw et al. 2010) on board of the Herschel Space Observatory (Pilbratt et al. 2010) during the performance verification phase on 2009 August 1. The corresponding dataset is OBS_1342181160. The HIFI band 1b (from 555.4 to 636.2 GHz) was covered in double beam switching. Both polarizations (H and V) were observed simultaneously. The receiver was tuned in double sideband (DSB) with a total integration time of 140 min. In order to obtain the best possible data reconstruction, the survey was acquired with a degree of redundancy time of 140 min. In order to obtain the best possible data reconstruction, the survey was acquired with a degree of redundancy of 4. The wide band spectrometer (WBS) was used as spectrometer, providing a frequency resolution of 0.5 MHz.

The data were processed with the ESA-supported package HIPE (Ott et al. 2009). Fits file from level 2 were then created and transformed into GILDAS format$^1$ for baseline subtraction and subsequent sideband deconvolution. The spectral resolution was degraded to 1 MHz in the final single sideband (SSB) dataset. The calibration for each receiver (H and V) is better than 2–3%. The relative calibration between both receivers is also rather good, with a difference in intensity of about 4%. The overall calibration uncertainty is about 7%, except for the strong CO line present in the band (see below).

Two strong lines dominate over the molecular transitions detected in the spectral band: the fundamental line of water in its ortho state o-H$_2$O 1$^1$ – 1$^0$ at 556.936069 GHz and the CO 5-4 line at 576.276905 GHz (Fig. 1). The final rms noise is 13 mK. We adopted the theoretical telescope main-beam efficiency $\eta_{mb}$ = 0.723, and a main-beam size of 39$''$ (HPFW) in the whole band. Unless indicated, intensities are expressed in units of antenna temperature $T_A$.

The CO 5-4 transition is detected with an intensity of 9 K ($T_A$) and a linewidth of 5 km s$^{-1}$. We notice a weak absorption feature in the line profile in the redshifted gas over a wide velocity range, which may partly arise from cloud contamination in the reference position. The intensity in the blue wing of the CO line differs by as much as 20% between both polarizations. This effect is not observed towards the H$_2$O line. Its origin is not understood at the moment.

The fundamental o-H$_2$O 1$^1$ – 1$^0$ line is detected with an intensity of 0.9 K ($T_A$) at the peak. It is characterized by a broad linewidth $\approx$15 km s$^{-1}$. The line displays an absorption dip at $v_{lsr} = +2.9$ km s$^{-1}$ and a broad redshifted wing extending up to $+8$ km s$^{-1}$. The broad linewidth of the H$_2$O spectrum could be fit with three Gaussian velocity components, a low-velocity, a high-velocity, and an absorption component. The low (high) velocity component peaks at $v_{lsr} = -0.58$ km s$^{-1}$ ($-7.86$ km s$^{-1}$); the linewidth and peak intensity derived from the fit are 9.56 km s$^{-1}$ and 0.77 K (13.72 km s$^{-1}$ and 0.29 K), respectively. The absorption component was fit by a narrow line Gaussian ($\Delta V = 1.38$ km s$^{-1}$) of amplitude $-0.48$ K centered at $v_{lsr} = +2.9$ km s$^{-1}$. The fit of the individual components and the resulting fit to the water spectrum is displayed in Fig. 1.

Overall, the H$_2$O and CO emission are detected in the same velocity range. The high sensitivity of the HIFI observations permits the detection of emission from the entrained gas up to $v_{lsr} = -30$ km s$^{-1}$, i.e. about 10 km s$^{-1}$ higher than was previously known from ground based observations. However, line profiles differ noticeably and the ratio of the H$_2$O/CO5-4 line intensities increases with increasing velocities from about 0.2 in the ambient gas up to 0.9 at $v_{lsr} = -25$ km s$^{-1}$ (Fig. 3).

\footnote{1 http://www.iram.fr/IRAMFR/GILDAS}
The HIFI main-beam is represented with a black circle. Contours at half-power are drawn in thick.

All the other molecular tracers detected in the HIFI band show a pronounced break in the line profile at $v_{lsr} \approx -7.2$ km s$^{-1}$ (Codella et al. this volume). This is also observed in the CO 6-5 spectrum of B1 obtained by us at the CSO, as part of complementary observations to help analyse the CHESS data. The maps of the whole southern lobe of L1157 in CO 3-2 and 6-5 obtained at the CSO in June 2009, with 24″ and 14.5″ respectively, will be discussed in detail in a forthcoming paper (Lefloch et al. 2010, in prep.).

Below, we define the region with $v_{lsr} < -7.25$ km s$^{-1}$ as the high-velocity component, hereafter HVC (see also Codella et al.), and the region with $v_{lsr} > -7.25$ km s$^{-1}$ as the low-velocity component (LVC). As we discuss below, these two velocity components are characterized by different spatial extents and excitation conditions.

3. Discussion

3.1. Origin of the emission

Due to its relatively high energy above the ground state ($E_{up} = 116$ K) the CO 6-5 transition is a good probe of the warm regions where H$_2$O can evaporate from grain mantles and be released in the gas phase. SiO 2-1, observed at the PdBI at 2.5″ resolution is a particularly good tracer of shocks strong enough to release refractory elements in the gas phase, because it is usually undetected in the cold, quiescent molecular gas.

The overall SiO emission is strongly peaked at the position of B1, which appears as a region of ~15″ size located at the apex of the cavity. Interferometric maps of the southern lobe (Figs. 1–2) reveal extended emission along the eastern wall of the cavity (the low-velocity wing of the bow) and downstream of B1, at velocities close to systemic, both blue and red. By comparing these data with IRAM 30m observations (Bachiller et al. 2001), we checked that unlike the HVC, a fraction of the flux emitted in the LVC is actually missed in the PdBI data, which is direct evidence for extended emission. This is consistent with the CO 6-5 data (Fig. 2). The low-velocity gas emission is located in the wake of B1, reaching ~40″ North from the apex. The area of the LVC amounts to $\approx 1/3$ of the HIFI beam (see Fig. 2). Interestingly, the PACS map of the H$_2$O 179 μm line reveals large-scale emission, spatially coinciding with SiO 2-1 emission in the outflow (Nisini et al. 2010).
component drives the neutral-neutral reactions that efficiently form H$_2$O, and the higher shock velocity can more efficiently remove water from grain mantles (see Melnick et al. 2008), resulting in the much greater ortho-H$_2$O column density than in the LVC. Comparison with NH$_3$ also suggests that the water production in the HVC is strongly dominated by high-temperature reactions (see Codella et al.).

The higher ortho-H$_2$O column density is what produces the higher 179 $\mu$m line flux from this component. Consistent results are obtained by Nisini et al. (2010) based on a 179 $\mu$m PACS map and previous ODIN and SWAS observations of the o-H$_2$O 1$_{0}^{10}$ − 1$_{0}$ line, assuming one single physical component dominates the water line emission in the HIFI beam. Follow-up observations of the higher-excitation lines of CO and H$_2$O with HIFI will help us constrain more accurately the physical conditions of each velocity component (density, temperature) and more generally in the shock.

Both components appear to have the same gas column density N(CO) $\approx$ 10$^{17}$ cm$^{-2}$ (Fig. 4). We found the high-velocity gas component to be unambiguously associated with hot gas ($T > 350$ K) of moderate density $\approx 3.0 \times 10^3$ cm$^{-3}$, whereas the low-velocity component arises from gas at a lower temperature ($\approx 3.0 \times 10^2$ cm$^{-3}$). The temperature estimated for the extended component agree reasonably well with other determinations from NH$_3$ and CH$_3$CN (Tafalla & Bachiller 1995; Codella et al. 2009).

With the physical conditions derived from the CO analysis, we modelled the integrated intensity and the line profile of the o-H$_2$O 1$_{0}^{10}$ − 1$_{0}$ transition as well as the reported PACS measured 179 $\mu$m H$_2$O line intensity ($\approx 10^{-19}$Wcm$^{-2}$, Nisini et al. 2010) to compute the total ortho water abundance in each velocity component. We used a radiative transfer code in the large-velocity gradient approach (and slab geometry) detailed in Melnick et al. (2008), taking into account an ortho to para ratio (OPR) of 1.2 for H$_2$ as derived from Spitzer (Neufeld et al. 2009). Here, we assume the absorption component at +2.9 km s$^{-1}$ is due to foreground gas unrelated to L1157-B1. Together, the two components of the o-H$_2$O 1$_{0}^{10}$ − 1$_{0}$ line produce a total H$_2$O 2$_{12}$ − 1$_{11}$ 179 $\mu$m line flux of 1.1 $\times 10^{-19}$ W cm$^{-2}$. For the temperature range derived from our CO analysis, we estimated ortho-H$_2$O column densities of (4.0 − 5.0) $\times 10^{14}$ cm$^{-2}$ and (2.5 − 3.0) $\times 10^{16}$ cm$^{-2}$ for the LVC and the HVC, respectively. Assuming an OPR of 3, we derived the H$_2$O abundance from comparison with the gas column densities estimated from CO (see Table 1). We obtained an abundance ratio [H$_2$O]/[CO] $\approx$ 0.8 in the high-velocity gas, which is consistent with previous results from ODIN (Benedettini et al. 2002) and agrees reasonably well with the predictions of steady-state C-shock models for this set of physical parameters (shock velocity $V_s$ $\approx$ 20 km s$^{-1}$, pre-shock density $n(H_2)$ $\approx$ 5 $\times 10^3$ cm$^{-3}$; Gusdorf et al. 2008a,b).

An interesting prediction of our model is that the HVC contribution to the 179 $\mu$m flux dominates over the LVC contribution (see last column in Table 1). The higher temperature of this

![Fig. 4. Best-fit solution to the LVG modelling of CO line temperatures in the 3−2, 5−4 and 6−5 transitions for both HVC (left) and LVC (right) components, respectively. The contour of the observed CO 5−4 integrated line area (T$_A$), corrected for main-beam dilution, is drawn in a dashed line, as well as the 5−4/3−2 and 6−5/3−2 line ratios (solid and dotted lines, respectively). Thin lines delineate the uncertainties in the observed values.](http://www.aanda.org)