Hydrides in young stellar objects: Radiation tracers in a protostar-disk-outflow system


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Hydrides in young stellar objects: Radiation tracers in a protostar-disk-outflow system**


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ABSTRACT

Hydrides of the most abundant heavier elements are fundamental molecules in cosmic chemistry. Some of them trace gas irradiated by UV or X-rays.

Methods. W3 IRS5 was observed by HIFI on the Herschel Space Observatory with deep integration (≈2500 s) in 8 spectral regions.

Results. The target lines including CH, NH, H2O+, and the new molecules SH+, H2O+, and OH+ are detected. The H2O+ and OH+ J = 1–0 lines are found mostly in absorption, but also appear to exhibit weak emission (P-Cyg-like). Emission requires high density, thus originates most likely near the protostar. This is corroborated by the absence of shift relative to the young stellar object (YSO). In addition, H2O+ and OH+ also contain strong absorption components at a velocity shifted relative to W3 IRS5, which are attributed to foreground clouds.

Conclusions. The molecular column densities derived from observations correlate well with the predictions of a model that assumes the main emission region is in outflow walls, heated and irradiated by protostellar UV radiation.

Key words. astrochemistry – line: identification – stars: formation – stars: massive – photon-dominated region – submillimeter: ISM

1. Introduction

In interstellar clouds, chemical reactions with hydrogen molecules lead to an elementary class of molecules that represent key species in the chemical evolution to larger molecules. These fundamental molecules, known as hydrides, include OH, CH, NH, SH, H2O, and their ions, OH+, CH+, NH+, SH+, H2O+, and H3O+. The combination of hydrogen atoms with a heavier atom causes large dipole moments and large rotation constants, particularly in diatomic hydrides. This widely separates the excitation levels. Only low-J lines are excited at temperatures relevant to star and planet formation. These lines have now become observable with the Herschel Space Observatory.

Many hydrides have a high activation energy in their formation paths. However, if high-energy photons – far UV (FUV) or X-rays – interact with the molecular gas and heat it, hydrides and particularly their ions are greatly enhanced in abundance.

Herschel is an ESA space observatory with science instruments provided by a European-led Principal Investigator consortia and with important participation from NASA.

** Appendix (page 5) is only available in electronic form at http://www.aanda.org

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Herschel/HIFI: first science highlights

LETTER TO THE EDITOR

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with an abundance four orders of magnitude larger than predicted, in W3 IRS5 and other YSOs with the JCMT (Stäuber et al. 2007) raised the expectation that the effects of irradiation may be more dramatic in asymmetric reality. Bruderer et al. (2009) modeled and interpreted the CO$^+$ anomaly as FUV radiation originating in the YSOs and irradiating the walls of outflow cavities.

Based on these predictions, the “Radiation Diagnostics” observations started with deep integrations of a large number of hydrides that was to be followed by a survey of a few species in many sources of various ages and masses. Here we present exploratory observations towards W3 IRS5, a nearby region (1.83 kpc, Imai et al. 2000) of high-mass star formation, moving at $-38.4$ km s$^{-1}$ in the local standard of rest (LSR). W3 IRS5 resembles the Trapezium cluster in Orion but is considerably younger. At least six radio sources are in the Herschel beam. They represent “hypercompact” H II regions produced by high-mass YSOs, at least two of which are O stars (van der Tak et al. 2005; Rodón et al. 2008). In this paper, we report the observations of major hydrides towards W3 IRS5 apart from H$_2$O, CH$^+$, and OH, which are studied in other projects of WISH (Chavarría et al. 2010; Wampfler et al., in prep.).

2. Observations

The Heterodyne Instrument for the Far Infrared (HIFI, de Graauw et al. 2010) on Herschel observed W3 IRS5 between 1 and 8 March 2010 in the Science Demonstration Phase in eight 4 GHz frequency bands for about 2500 s each. One of them includes [C II] at 1900.5369 GHz. We used the wide band spectrometer which has a spectral resolution of 1.1 MHz, and HIPE 3.0 for pipeline and data analysis. The data were taken by double beam switching (DBS), the high-frequency (HEB) bands in fast DBS mode. The off-source position was at a distance of 3 arcmin in the NE and SW without remarkable IR sources.

The current accuracy of the velocity calibration is estimated to be better than 2 km s$^{-1}$. The antenna temperature was converted to main beam temperature, using pre-flight antenna efficiencies. After visual inspection and defringing, the V polarization was shifted linearly in flux to match the H polarization. The two polarizations were then added. A second observation of equal length was made using a local oscillator frequency shifted by 10 km s$^{-1}$. The two data sets were plotted in both upper and lower sideband presentation. If a line matched velocities in one sideband and was double with 20 km s$^{-1}$ separation in the other, the frequency of the former was assumed. All lines of interest could be attributed to a sideband without ambiguity. The continuum was divided by 2 for double sideband observations, assuming that it is the same in both sidebands, of equal sensitivity.

3. Results

Most lines (except H$_2$O$^+$) are split by fine or hyperfine interaction as indicated in Fig. 1. Table 1 lists the observed lines (strongest only for multiples) and summarizes the quantitative observational results. Molecules here detected for the first time in star-forming regions include H$_2$O$^+$, OH$^+$, and SH$^+$. Having the most prominent line near the H$_2$O para ground-state line at 1113.3 GHz, H$_2$O$^+$ is serendipitously detected in many Herschel observations. H$_2$O$^+$ and OH$^+$ are detected in absorption by the interstellar medium (Bruderer et al. 2010b; Gerin et al. 2010; Falgarone et al. 2010; Neufeld et al. 2010; Ossenkopf et al. 2010; Schilke et al. 2010), but also near the systemic velocity of other high-mass YSOs (Wyrowski et al. 2010b).

Several lines are found in absorption, indicated in Table 1 by negative peak values relative to the continuum. All lines predominantly in absorption (NH, OH$^+$, H$_2$O$^+$) originate in molecules in the ground state with a $J = 1$ level energy exceeding 47 K. Lines of CH and SH$^+$, observed in emission, are transitions from...
Table 1. Frequency, upper level energy, and Einstein coefficient of molecules and lines observed by Herschel/HIFI towards W3 IRS5.

<table>
<thead>
<tr>
<th>Molecule</th>
<th>Transition</th>
<th>Frequency [GHz]</th>
<th>$E_u$ [K]</th>
<th>$A_{ad}$ [s$^{-1}$]</th>
<th>Line peak flux [mK]</th>
<th>Line width [km s$^{-1}$]</th>
<th>Line shift [km s$^{-1}$]</th>
<th>Line flux density [K km s$^{-1}$]</th>
<th>Column density [cm$^{-2}$]</th>
</tr>
</thead>
<tbody>
<tr>
<td>CH</td>
<td>1$_0$–1$_0$</td>
<td>536.7611</td>
<td>25.76</td>
<td>6.4(–4)</td>
<td>740 ± 7</td>
<td>10.6</td>
<td>+1.6</td>
<td>14.8 ± 0.02</td>
<td>3.1(13)</td>
</tr>
<tr>
<td>NH</td>
<td>1$_1$–0$_1$</td>
<td>999.9734$^a$</td>
<td>47.99</td>
<td>5.2(–2)</td>
<td>–46 ± 13</td>
<td>2.7</td>
<td>–0.8</td>
<td>3.20 ± 0.02</td>
<td>–5.0(11)</td>
</tr>
<tr>
<td>OH</td>
<td>1$_1$–0$_0$</td>
<td>1303.1186$^a$</td>
<td>49.58</td>
<td>1.8(–2)</td>
<td>–790 ± 21</td>
<td>5.9</td>
<td>+3.4</td>
<td>14.6 ± 0.03</td>
<td>7.1(13)</td>
</tr>
<tr>
<td>OH$^+$</td>
<td>2$_1$–1$_1$</td>
<td>1892.2271$^a$</td>
<td>140.4</td>
<td>5.9(–2)</td>
<td>&lt;225</td>
<td>&lt;1.13</td>
<td>&lt;0.98</td>
<td>&lt;1.1(11)</td>
<td>&lt;1.1(11)</td>
</tr>
<tr>
<td>NH$^+$</td>
<td>1$_2$–0$_1$</td>
<td>1019.2107$^a$</td>
<td>48.91</td>
<td>5.5(–2)</td>
<td>&lt;0.032</td>
<td>&lt;0.14</td>
<td>&lt;0.2</td>
<td>&lt;5.2(9)</td>
<td></td>
</tr>
<tr>
<td>SH$^+$</td>
<td>2$<em>{12}$–0$</em>{10}$</td>
<td>526.0479$^a$</td>
<td>25.25</td>
<td>9.7(–4)</td>
<td>65 ± 3</td>
<td>4.4</td>
<td>–0.3</td>
<td>0.73 ± 0.1</td>
<td>4.1(11)</td>
</tr>
<tr>
<td>SH$^+$</td>
<td>2$<em>{11}$–1$</em>{10}$</td>
<td>1082.9117$^a$</td>
<td>77.2</td>
<td>9.1(–2)</td>
<td>&lt;37</td>
<td>&lt;0.17</td>
<td>&lt;0.2</td>
<td>&lt;3.9(10)</td>
<td></td>
</tr>
<tr>
<td>SH$^+$</td>
<td>3$<em>{12}$–2$</em>{10}$</td>
<td>1632.5179$^a$</td>
<td>155.6</td>
<td>3.1(–2)</td>
<td>&lt;151</td>
<td>&lt;0.68</td>
<td>&lt;0.2</td>
<td>&lt;9.8(10)</td>
<td></td>
</tr>
<tr>
<td>H$_2$O$^+$</td>
<td>3$<em>{12}$–2$</em>{10}$</td>
<td>999.8213$^a$</td>
<td>223.9</td>
<td>2.3(–2)$^a$</td>
<td>&lt;46</td>
<td>&lt;0.20</td>
<td>&lt;1.7</td>
<td>&lt;1.7(10)</td>
<td></td>
</tr>
<tr>
<td>H$_2$O$^+$</td>
<td>1$<em>{12}$–0$</em>{10}$</td>
<td>1115.2040$^a$</td>
<td>53.52</td>
<td>3.1(–2)$^a$</td>
<td>–285 ± 24</td>
<td>5.1</td>
<td>+39.1</td>
<td>2.21 ± 0.03</td>
<td>4.6(12)</td>
</tr>
<tr>
<td>H$_2$O$^+$</td>
<td>4$<em>{20}$–3$</em>{11}$</td>
<td>1031.2937$^a$</td>
<td>232.2</td>
<td>5.1(–3)</td>
<td>570 ± 10</td>
<td>6.2</td>
<td>+0.5</td>
<td>3.8 ± 0.3</td>
<td>9.7(11)</td>
</tr>
<tr>
<td>H$_2$O$^+$</td>
<td>4$<em>{20}$–3$</em>{11}$</td>
<td>1069.8266$^a$</td>
<td>268.8</td>
<td>9.8(–3)</td>
<td>230 ± 10</td>
<td>4.7</td>
<td>–0.3</td>
<td>1.30 ± 0.03</td>
<td>2.9(11)</td>
</tr>
<tr>
<td>H$_2$O$^+$</td>
<td>6$<em>{20}$–5$</em>{12}$</td>
<td>1454.5625$^a$</td>
<td>692.6</td>
<td>7.1(–3)</td>
<td>&lt;245</td>
<td>&lt;0.66</td>
<td>&lt;3.8</td>
<td>&lt;3.8(11)</td>
<td></td>
</tr>
<tr>
<td>H$_2$O$^+$</td>
<td>2$<em>{12}$–1$</em>{10}$</td>
<td>1632.9010$^a$</td>
<td>143.1</td>
<td>1.7(–2)</td>
<td>145 ± 46</td>
<td>6.7</td>
<td>+0.3</td>
<td>1.20 ± 0.12</td>
<td>3.7(11)</td>
</tr>
</tbody>
</table>

Notes. The numbers in parentheses give the decimal power. Negative peak fluxes signify background minus line temperature of lines in absorption. Line widths refer to the FWHM value of the most intense line peak or absorption and its line shift to the systemic velocity of –38.4 km s$^{-1}$. Non-detected peak fluxes are 5σ upper limits at 1 km s$^{-1}$ resolution, and non-detected line fluxes assume a 5 km s$^{-1}$ line width. The column densities refer to the upper (emission) or lower (absorption) state. Molecular data are taken from: ($^a$) CDMS (Müller et al. 2001), ($^b$) Mürtz et al. (1998), ($^c$) Bruderer (2006), ($^d$) Hübners et al. (2009), ($^e$) JPL catalogue.

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Column densities are given in Table 1 for the upper energy level of transitions in emission, and for the lower energy level of lines in absorption. The flux scales are reduced and shifted for better visibility. The velocity is relative to the systemic velocity of W3 IRS5, shown by a vertical dashed line. The horizontal dashed lines indicate the continuum level.

![Fig. 2. Light hydrides with P Cygni type absorption features compared with [C II]. The flux scales are reduced and shifted for better visibility. The velocity is relative to the systemic velocity of W3 IRS5, shown by a vertical dashed line. The horizontal dashed lines indicate the continuum level.](image-url)
4. Discussion and conclusions

The observed shifts of the line peaks relative to the systemic velocity of the YSO are small and suggest that the lines originate in the star-forming region, not in the foreground interstellar medium. This may not be the case, however, for the components of H2O, H2O−, and OH− shifted by a larger amount to roughly the velocity VLSR = 0.

Additional evidence of the origin comes from the lines in emission, which are produced at critical densities (if known) of order 10^7 cm^−3. Even lines predominantly in absorption, such as OH+ and possibly H2O+, have a small emission feature in the red wing, thus a P-Cyg-like profile, near the systemic velocity. This is the first report of Galactic OH+ and H2O+ in emission. Figure 2 shows the similarity of the profiles of OH+ and H2O+. The absorption in CH may be caused by self-absorption, but is similar to that in [C II], which is found in emission at the off position. Absorption lines in star-forming regions need detailed modeling and radiation transfer calculations beyond the scope of this letter.

The ground-state column densities of the component moving with the YSO are 8.1 × 10^{11} cm^−2 for H2O+, and 9.7 × 10^{12} cm^−2 for OH+. The values for the red-shifted components are 3.8 × 10^{12} cm^−2 and 6.1 × 10^{13} cm^−2, respectively. Both components yield larger OH+/H2O+ ratios than the other observations reported in this special feature.

The measured line widths are generally small (< 7 km s^−1) and show no anisotropies. We thus find no evidence of shocks except possibly in OH+, which needs to be studied in combination with shock tracers.

In Fig. 3, column densities are displayed, derived from integrated line fluxes neglecting re-emission or reabsorption of the final state. The derivation is based on Table 1, except for OH+ and H2O− where only the unshifted component attributed to the YSO is used. For H2O+, the value of the J = 4 level as well as the one from the rotational diagram, integrated over all levels, are shown. The column densities are compared with abundances predicted by Bruderer et al. (2010a) in a two-dimensional “standard” YSO model used here as a template, assuming UV and X-ray irradiation by a central high-mass YSO. It enhances the abundance of diatomic hydrides in the outflow walls by many orders of magnitude, such that the beam-averaged abundance is significantly changed. It is averaged over a radius of 20,000 AU, or 10.9′′ at the distance of W3 IRS5.

The similarity of observations and model abundances in Fig. 3 support the scenario of hydride enrichment in outflow walls heated and irradiated by protostellar far UV.

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Appendix A:

The rotational diagram in Fig. A.1 is complemented with ground-based data for the frequency range 300–400 GHz observed at the CSO with comparable beam size by Phillips et al. (1992). The non-detection at 307 MHz is surprising, but possibly an effect of optical depth. The data in Fig. A.1, except 307 GHz, are well fitted by a single rotational temperature of 239 K, suggesting that the observed levels are populated according to an exponential distribution. The derived temperature and column density infer an optical depth of $\tau < 0.1$ for all lines except at 307 GHz. The fitted line (dashed) corresponds to a column density of $8.5(\pm 2) \times 10^{13}$ cm$^{-2}$, consistent with the value derived by Phillips et al. (1992). This leads to a beam averaged H$_3$O$^+$ abundance of $4.2(\pm 1) \times 10^{-10}$ relative to H, to be compared with the theoretical value of $4 \times 10^{-10}$ reported by Bruderer et al. (2010a).

![Fig. A.1. Rotational diagram of H$_3$O$^+$. Numbers indicate the frequencies in GHz of the observed lines.](image-url)

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