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Water vapor toward starless cores: The Herschel view*


(Affiliations are available on page 5 of the online edition)

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ABSTRACT

Aims. Previous studies by the satellites SWAS and Odin provided stringent upper limits on the gas phase water abundance of dark clouds ($x$(H$_2$O) < 7 × 10$^{-7}$). We investigate the chemistry of water vapor in starless cores beyond the previous upper limits using the highly improved angular resolution and sensitivity of Herschel and measure the abundance of water vapor during evolutionary stages just preceding star formation. Methods. High spectral resolution observations of the fundamental ortho water (o-H$_2$O) transition (557 GHz) were carried out with the Heterodyne Instrument for the Far Infrared onboard Herschel toward two starless cores: Barnard 68 (hereafter B68), a Bok globule, and LDN 1544 (L1544), a prestellar core embedded in the Taurus molecular cloud complex. Detailed radiative transfer and chemical codes were used to analyze the data. Results. The RMS in the brightness temperature measured for the B68 and L1544 spectra is 2.0 and 2.2 mK, respectively, in a velocity bin of 0.59 km s$^{-1}$. The continuum level is 3.5 ± 0.2 mK in B68 and 11.4 ± 0.4 mK in L1544. No significant feature is detected in B68 and the 3σ upper limit is consistent with a column density of o-H$_2$O $N$(o-H$_2$O) < 2.5 × 10$^{11}$ cm$^{-2}$, or a fractional abundance $x$(o-H$_2$O) < 1.3 × 10$^{-7}$, more than an order of magnitude lower than the SWAS upper limit on this source. The L1544 spectrum shows an absorption feature at a 5σ level from which we obtain the first value of the o-H$_2$O column density ever measured in dark clouds: $N$(o-H$_2$O) = (8 ± 4) × 10$^{12}$ cm$^{-2}$. The corresponding fractional abundance is $x$(o-H$_2$O) = 5 × 10$^{-9}$ at radii > 7000 AU and ≃ 2 × 10$^{-10}$ toward the center. The radiative transfer analysis shows that this is consistent with a $x$(o-H$_2$O) profile peaking at ≈ 10$^{-9}$, 0.1 pc away from the core center, where both freeze-out and photodissociation are negligible. Conclusions. Herschel has provided the first measurement of water vapor in dark regions. Column densities of o-H$_2$O are low, but prestellar cores such as L1544 (with their high central densities, strong continuum, and large envelopes) appear to be very promising tools to finally shed light on the solid/vapor balance of water in molecular clouds and oxygen chemistry in the earliest stages of star formation.

Key words. astrochemistry – line: formation – molecular processes – radiative transfer – stars: formation – ISM: clouds

1. Introduction

The abundance of water vapor in dark clouds is a crucial missing quantity in our understanding of astrochemistry and gas-grain interactions. Before the launch of the Submillimeter Wave Astronomy Satellite (SWAS), pseudo-time-dependent gas phase chemical models predicted water fractional abundances ($x$(H$_2$O) ≈ n(H$_2$O)/n(H$_2$), where n(i) is the number density of species i) around $10^{-7}$ (e.g. Lee et al. 1996). SWAS observations of the o-H$_2$O ($J_{K_a}$=1:110−1:01) line toward the low-mass starless cores B68 and Oph D place stringent upper limits on the abundance of water vapor in cold gas of $x$(H$_2$O) < 3 × 10$^{-8}$ and ≈ 8 × 10$^{-9}$ in B68 and Oph D, respectively (Bergin & Snell 2002).

The Odin satellite provided an even lower upper limit toward the protostellar core Cha-MMS1 (<7 × 10$^{-9}$; Klotz et al. 2008). Water in dark clouds is mostly in the form of ice mantles covering dust grains ($x$(H$_2$O)$_{ice}$ ≈ 10$^{-6}$; e.g., Whittet & Duley 1991). The gas-phase chemical models needed revision. The inclusion of dust grains in the chemistry and the freeze-out of oxygen onto grain surfaces significantly improved the agreement with observations (e.g. Bergin et al. 2000; Viti et al. 2001; Roberts & Herbst 2002; Melnick & Bergin 2005; Aikawa et al. 2005). Hollenbach et al. (2009) also pointed out the importance of illumination by the external far-ultraviolet (FUV) radiation field. They showed that in clouds of constant density illuminated by the local interstellar field (G$_0$ = 1), the water abundance can reach a peak of about 10$^{-7}$ only in a range of local visual extinctions A$_V$ between 1 and 4 mag. At A$_V$ < 1, H$_2$O molecules are photodissociated, whereas at A$_V$ > 4 mag the water is mostly

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files have been measured and successfully modeled (e.g. Tafalla et al. 1998; Alves et al. 2001; Evans et al. 2001; Caselli et al. 2002a; Hotzel et al. 2002; Bergin et al. 2006; Broderick et al. 2007; Crapsi et al. 2007; Kirk et al. 2007; Keto & Caselli 2010). They are representative of the two different classes of starless cores: thermally subcritical (B68) and supercritical (L1544; Keto & Caselli 2008) and they are both far from active sites of star formation. Therefore, they are the ideal sites for studying the water content along the lines-of-sight with dense molecular material not yet affected by protostellar feedback. Observations of the centrally concentrated and contracting L1544 provide insight into the water content just before a protostar is born, thus helping to set the initial chemical conditions in the process of star formation.

2. Observations

As part of the WISH (Water In Star-forming regions with Herschel) key project (van Dishoeck et al., in prep.), we observed the $^{1}O_{H2O}$ ($^{1}O_{H2O}$) line at 556.936.0020 MHz toward B68 (RA(J2000) = $^{17^h22^m38^s20}$, Dec(J2000) = $^{−23^°49′54″0}$, $V_{LSR}$ = 3.4 km s$^{-1}$) and L1544 (RA(J2000) = $^{0^h5^m4^s17^m21}$, Dec(J2000) = $^{25^°10′42″8}$, $V_{LSR}$ = 7.2 km s$^{-1}$) with Herschel/HIFI (de Graauw et al., 2010), on 2010 March 20. The dual beam switch (DBS) mode, with the continuum optimization and fast chop options was used. The wide-band spectrometer (WBS) and the high resolution spectrometer (HRS) were used simultaneously in both H and V polarizations. A beam efficiency of 0.75 was used to convert the antenna temperature scale into main beam brightness temperature ($T_{mb}$).

The individual spectra were reduced using the Herschel interactive processing environment (HIPE, version 3.0.1; Ott 2010). A final analysis was performed in CLASS\(^1\). The H and V spectra were analyzed individually to check for any discrepancy in the calibration. No significant difference was noticed, so the two spectra were averaged. The WBS spectra were smoothed at the nominal 1.1 MHz resolution (0.59 km s$^{-1}$) and the RMS noise level was 2.0 mK for B68 and 2.2 mK for L1544 (1.4 times lower than the RMS reached in the HRS spectra after Gaussian smoothing to the same WBS resolution, the expected difference between the two backends). Only WBS spectra will be considered here.

3. Results

Dust continuum emission in the B68 and L1544 spectra is observed at a level of $3.5 \pm 0.2$ and $11.4 \pm 0.4$ mK ($1.6 \pm 0.1$ and $5.1 \pm 0.2$ Jy), respectively. Figure 1 shows the two spectra after baseline subtraction. No significant feature is seen in the direction of B68, whereas a faint absorption line is present towards L1544, also visible in the HRS spectrum (grey histogram). Figure 2 is a zoom-in of the WBS L1544 spectrum.

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\(^1\) CLASS is the continuum and line-analysis single-dish software within the GILDAS package available at http://www.iram.fr/IRAMFR/GILDAS
in Fig. 1 (empty histogram), plotted with the CO (J=1−0) line observed with the Five College Radio Astronomy Observatory (FCRAO) (shaded histogram; Tafalla et al., in prep.). The O-H$_2$O feature is broad and covers the two velocity components of CO (1−0); the first is centered on the dense core velocity (7.2 km s$^{-1}$; Caselli et al. 2002a), whereas the other (also observed in CO(2−1), but not in high density tracers) peaks at 9.29 ± 0.03 km s$^{-1}$. This is the first evidence of water vapor in the direction of a cold core. Interestingly, the main absorption feature is coincident with the weaker component of the CO(1−0) spectrum, suggesting that most of the absorption is not associated with the dense core material. In the following, we first derive the water vapor abundance in the two starless cores assuming homogeneous conditions. We then present our detailed radiative transfer analysis.

### 3.1. Basic estimates

#### 3.1.1. B68

An upper limit to the fractional abundance of O-H$_2$O can be obtained from the B68 spectrum (Fig. 1). We used the RADEX code (van der Tak et al. 2007), with input parameters appropriate for B68 of kinetic temperature $T_{k} = 10$ K, line width $\Delta v = 0.5$ km s$^{-1}$ (based on CO observations and accounting for the different molecular masses), and H$_2$ volume densities $n$(H$_2$) = 10$^6$ cm$^{-3}$. To reproduce the observed 3$\sigma$ upper limits of 6 mK, with the WBS resolution, the required column density of O-H$_2$O is $N$(O-H$_2$O) = 4.2 × 10$^{19}$ cm$^{-2}$. Within the Herschel beam ($FWHM = 39''$), $N$(H$_2$) = 2 × 10$^{27}$ cm$^{-2}$ (e.g. Alves et al. 2001), thus, for B68, $x$(O-H$_2$O) < 2.1 × 10$^{-6}$, consistent with the SWAS upper limit (Bergin & Snell 2002). However, this analysis implies that: (i) the dust continuum emission at 557 GHz (539 $\mu$m) is negligible; and (ii) there is a uniform physical (excitation) structure. This is not correct (see Sect. 3.2). If the line is in absorption, it cannot be stronger than 1.3$\sigma$ deep because the continuum is 1.3$\sigma$ above the 2.73 K background, and the line is therefore undetectable. Thus, no upper limit can be given.

#### 3.1.2. L1544

The absorption feature toward L1544 allows us to measure, for the first time, the column of absorbing water in front of the far-infrared (FIR)-emitting core material. A Gaussian fit to the feature (dashed curve in Fig. 2) gives an integrated intensity $I = -18.5 ± 4.0$ mK km s$^{-1}$, line width $\Delta v = 2.7 ± 0.6$ km s$^{-1}$, $T_{mb} = -6 ± 2$ mK, and $V_{lsr} = 8.6 ± 0.3$ km s$^{-1}$ (identical, within the errors, to the HRS spectrum fit results). From the continuum level and $T_{mb}$, we obtain an optical depth $\tau = 0.6 ± 0.4$. With $\tau$ and $\Delta v$, it is straightforward to obtain $N$(O-H$_2$O) = $(8 ± 6) \times 10^{12}$ cm$^{-2}$. To arrive at $x$(O-H$_2$O), $N$(H$_2$) is needed. Assuming that CO (1−0) and H$_2$O (1$\_0$−1$\_0$) trace the same material (as suggested by the similar velocities), we estimate $N$(H$_2$) from the CO (1−0) line integrated intensity adopting an X-factor of 2 × 10$^{29}$ cm$^{-2}$ K$^{-1}$ km$^{-1}$ (e.g. Pineda et al. 2008), which gives $N$(H$_2$) = 3 × 10$^{21}$ cm$^{-2}$. Thus, $x$(O-H$_2$O) = $(5 ± 4) \times 10^{-9}$. We note that our $N$(H$_2$) estimate is about 20 times lower than the value found using dust continuum emission at 1.2 mm within the Herschel beam (Ward-Thompson et al. 1999), and the 539 $\mu$m continuum observed by Herschel, assuming a dust opacity at 230 GHz $\kappa_{230} = 0.009$ cm$^{2}$ g$^{-1}$, a dust emissivity index $\beta = 2$ (e.g. Schnee et al. 2010), and a dust temperature of 8 K within the Herschel beam (Crapsi et al. 2007): $N$(H$_2$) = 7 × 10$^{22}$ cm$^{-2}$. This discrepancy is expected given that CO is mostly frozen-out within the central $\approx$7000 AU (e.g. Caselli et al. 1999), where the bulk of the mm and FIR continuum is emitted. Using the $N$(H$_2$) value derived from the continuum, $x$(O-H$_2$O) = $2 \times 10^{-10}$, suggesting that the water vapor abundance is very low toward the core center, like CO. The $x$(O-H$_2$O) values derived here are averages along the line of sight. Detailed radiative transfer modeling and more sensitive observations are needed to place constraints on the abundance profile (see Sect. 3.2).

#### 3.2. Radiative transfer including dust emission

Keto & Caselli (2008, 2010, hereafter KC08 and KC10) used the MOLLIE radiative transfer code and successfully reproduced observed line intensities and profiles in B68 and L1544. From the comparison between observed and simulated spectra, KC08 derived temperature and density profiles of the two starless cores, as well as their velocity structure. The analysis of the present Herschel/HIFI data is based on the KC10 models, which were modified to include water vapor in the code. Two different sets of collisional coefficients are used in the modeling: those from Dick et al. (2010, hereafter DDP) and those from Phillips et al. (1996, hereafter PMG). At temperatures below about 40 K, DDP provide the lowest rates, whereas PMG give the highest rates available, so together they represent the two extreme cases.
values present in the literature. The rates of Dubernet et al. (2009) will be considered in an upcoming paper.

We considered abundance models by Aikawa et al. (2005, hereafter AHRC), Hollenbach et al. (2009), and KC10 (freeze-out, a generalized desorption from Roberts et al. 2007, and photodissociation, but no detailed gas phase and surface chemistry) and we plot those of KC10. The more sophisticated models of AHRC and Hollenbach et al. (2009) predict different abundances that also differ from those of KC10. We will investigate these differences when more sensitive data will be available. Example H2O abundance profiles for the two sources are shown in Figs. 3 and 4, left panels, where the density profiles can also be found. Curves labeled a–c refer to different peak abundances which are compared with the present data.

The dust continuum emission is taken into account by MOLLIE and the observed fluxes at 539 μm are reproduced with the dust properties found by KC08 and KC10. MOLLIE results are consistent with the observed profiles for the three water abundance profiles are shown in Figs. 3, 4, and 9, right panels. For B68, the upper limit is consistent with all three curves, the only exception being curve c, if DDP rates are correct. The o-H2O column density found by integrating curve b along the line of sight, and convolving with the Herschel beam, is N(o-H2O) = 2.5 × 1013 cm−2, about 20 times lower than the upper limit found in Sect. 3.1.1. For L1544, curves b and c in Fig. 4 can explain the absorption features at the dense core velocity (the higher velocity component will be studied with the new Herschel data for this source). With the current sensitivity and spectral resolution, it is impossible to distinguish between the two sets of curves and collisional coefficients. For model c, the corresponding o-H2O column density within the Herschel beam is N(o-H2O) = 3.3 × 1013 cm−2, about 4 times higher than that measured from the absorption feature (Sect. 3.1.2), so it can probably be ruled out. Using model b, we find that N(o-H2O) = 3.3 × 1012 cm−2, which is consistent with the measured value.

4. Conclusions

We have presented Herschel/HIFI observations of o-H2O (1-0→1-0) toward two starless cores embedded in quiescent environments (B68 and L1544). We have interpreted and discussed these data with the help of detailed radiative transfer modeling. In B68, the 3σ upper limit is consistent with a column density of ortho water N(o-H2O) < 2.5 × 1013 cm−2, or a fractional abundance χ(o-H2O) < 1.3 × 10−9, about twenty times lower than the upper limit of L1544 which is consistent with the measured value.

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References
