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DISCOVERY OF THE BROAD-LINED TYPE Ic SN 2013cq ASSOCIATED WITH THE VERY ENERGETIC GRB 130427A

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ABSTRACT

Long-duration gamma-ray bursts (GRBs) at $z < 1$ are found in most cases to be accompanied by bright, broad-lined Type Ic supernovae (SNe Ic-BL). The highest-energy GRBs are mostly located at higher redshifts, where the associated SNe are hard to detect observationally. Here, we present early and late observations of the optical counterpart of the very energetic GRB 130427A. Despite its moderate redshift, $z = 0.3399 \pm 0.0002$, GRB 130427A is at the high end of the GRB energy distribution, with an isotropic-equivalent energy release of $E_{\text{iso}} \sim 9.6 \times 10^{53}$ erg, more than an order of magnitude more energetic than other GRBs with spectroscopically confirmed SNe. In our dense photometric monitoring, we detect excess flux in the host-subtracted $r$-band light curve, consistent with that expected from an emerging SN, ~$0.2$ mag fainter than the prototypical SN 1998bw. A spectrum obtained around the time of the SN peak (16.7 days after the GRB) reveals broad undulations typical of SNe Ic-BL, confirming the presence of an SN, designated SN 2013cq. The spectral shape and early peak time are similar to those of the high expansion velocity SN 2010bh associated with GRB 100316D. Our findings demonstrate that high-energy, long-duration GRBs, commonly detected at high redshift, can also be associated with SNe Ic-BL, pointing to a common progenitor mechanism.

Key words: gamma-ray burst: individual (GRB 130427A) -- supernovae: individual (SN 2013cq)

Online-only material: color figures

1. INTRODUCTION

The standard paradigm for long-duration gamma-ray bursts (GRBs) involves a broad-lined Type Ic supernova (SN Ic-BL) with $M_V \sim -19$ mag (Woosley & Bloom 2006; Hjorth & Bloom 2012), such as those predicted by the collapsar model (MacFadyen & Woosley 1999). This fact is based on spectroscopic evidence in SNe from low-luminosity GRBs (Bromberg et al. 2011) such as SN 1998bw accompanying GRB 980425 (Galama et al. 1998), as well as relatively higher-luminosity GRBs, such as SN 2003dh accompanying GRB 030329 (Hjorth et al. 2003; Stanek et al. 2003). Interestingly, for the two low-redshift cases of GRB 060505 and GRB 060614, no associated SN was found to deep limits (Fynbo et al. 2006; Della Valle et al. 2006; Gal-Yam et al. 2006), but since then no similar events have been reported.

GRB 130427A (Maselli et al. 2013; Elenin et al. 2013) is remarkable as it is both extremely energetic and located at a moderately low redshift of $z = 0.3399 \pm 0.0002$ (Levan et al. 2013 and this work). Using the spectral parameters for the prompt emission given by von Kienlin (2013), we derive an isotropic gamma-ray energy $E_{\text{iso}} \sim (9.6 \pm 0.4) \times 10^{53}$ erg in the 1–10000 keV rest-frame energy band. The optical afterglow peaked at $R \approx 7.4$ mag during the prompt emission phase (Wren et al. 2013). Only ~$5\%$ of all GRBs with measured redshifts are at such a small distance (see, e.g., Jakobsson et al. 2012; Figure 1) and those are usually low-luminosity events. In contrast, GRB 130427A was an extremely energetic burst and

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20 We adopt a ΛCDM cosmology with $H_0 = 67.3$ km s$^{-1}$ Mpc$^{-1}$, $\Omega_m = 0.315$, and $\Omega_{\Lambda} = 0.685$ (Planck Collaboration 2013).
hence it has enabled detailed studies of a system analogous to those usually only found at higher redshifts (Fan et al. 2013; Laskar et al. 2013; Tam et al. 2013).

Due to the dearth of such extremely luminous GRBs at low redshift and their faintness at high redshift, we so far have had no spectroscopic evidence for accompanying SNe in very energetic GRBs with $E_{\text{iso}} > 10^{52.7}$ erg (for photometric evidence of an SN associated with the very energetic GRB 080319B, see Tanvir et al. 2010). In Figure 1, we show the gamma-ray energy release of GRBs as a function of redshift. Overplotted are systems with spectroscopic evidence for an SN. GRB 130427A stands out as an exceptional system. It is unclear if progenitor models involving an SN can power such energetic GRBs (see, e.g., Piran 2004). Observationally, it has proven to be difficult to test the SN properties of these events. In this work, we focus on the search for and discovery of an SN accompanying this remarkable burst.

2. OBSERVATIONS AND DATA ANALYSIS

2.1. Photometry

Our first follow-up photometry was carried out at the 1 m optical telescope located in Weihai, Shandong Province, China. The bright optical counterpart of the GRB was well detected in the Cousins $R_C$ filter at $R_C \sim 15.5$ mag 4.178 hr after the burst. Initially, the Sloan Digital Sky Survey (SDSS) filters were not available at the telescope, but two days into our monitoring campaign, SDSS $r$ and $i$ filters were installed and have been available since then.

Our photometric follow-up observations were mainly obtained at the 2.5 m Nordic Optical Telescope (NOT) using the ALFOSC instrument. Photometry was primarily obtained in the $r$ band, complemented by $ugiz$ data useful for monitoring the spectral evolution of the counterpart. As shown in Figure 2, there is indication from the NOT images that the GRB counterpart lies in the northwest part of an extended host galaxy. With our spatial resolution, the afterglow is blended with the host and we thus use a relatively large aperture ($\sim 3.8'$ in diameter) to measure the magnitudes of the counterpart plus host. Using a smaller aperture would provide lower fluxes, but the host contribution would be hard to quantify, especially at late times (i.e., $\sim 10$ days and beyond after the burst). In this way, we consistently include most of the host light.

Additionally, follow-up observations were obtained at the 2.16 m telescope in Xinglong, Hebei Province, China. Photometry was obtained in the Cousins $R_C$ and $I_C$ filters and then transformed to SDSS $r$ and $i$ magnitudes based on our afterglow spectra.

All optical data were reduced in a standard way using IRAF v2.15 in the Scisoft 7.7 package. Magnitudes were calibrated with two nearby bright SDSS field stars, SDSS J113230.55+274420.3 and SDSS J113220.11+274133.5, whose zero-point errors were 0.01 mag and were propagated into the final magnitude measurements. The SDSS $r$-band light curve is presented in Figure 3 and a log of the observations is shown in Table 1.

We fit the NOT multi-color photometry taken in the first night after the trigger ($g = 17.31 \pm 0.01$ mag, $r = 17.06 \pm 0.01$ mag.

extinction law. Within the errors, this value is consistent with the reddening derived assuming an SMC or LMC extinction law because of the small amount of reddening and the wavelength range probed by our observations.

2.2. Spectroscopy

Our first spectrum was obtained using NOT/ALFOSC. The total exposure was 1800 s with a mean time of 0.44 days post-burst. The spectrum covers the range 3200–9100 Å with a resolving power of ~700. We identify prominent absorption lines of Mg ii 2796 & 2803, Mg i 2852, and Ca ii 3934 & 3968, as well as weak emission lines of [O ii] 3727 and Hβ, all at a common redshift of $z = 0.34$.

A second spectrum with intermediate resolution was obtained shortly afterward using the Very Large Telescope (VLT) equipped with the XSHOOTER spectrograph. The continuum was well detected over the full range 3000–24800 Å. A number of absorption features are visible, including Fe ii 2344, Mn ii 2577, Mg ii 2796 & 2803, Mg i 2852, Ti ii 3074, Ca ii 3934 & 3968, Na i 5890 & 5896, and emission lines such as [O ii] 3727, Hβ, [O iii] 5007, and Hα, all at a common redshift of $z = 0.3399 \pm 0.0002$. In the XSHOOTER spectrum, Na i D 5890 & 5896 absorption was detected at the redshift of the host. We measure equivalent widths of 0.18 ± 0.02 and 0.08 ± 0.03 Å for the Na i D1 and D2 components, respectively. Using the relations in Poznanski et al. (2012), we obtain an estimate for $E(B - V)_{\text{host}} = 0.03 \pm 0.01$ mag, but remark that there exists a substantial dispersion of $E(B - V) \sim 0.15$ mag in this relation. Considering different calibrations/systematics involved in the above $E(B - V)_{\text{host}}$ measurements, we adopt $E(B - V)_{\text{host}} = 0.05$ mag for the host extinction.

Given the relatively low redshift, we planned a third spectroscopic observation with the aim of detecting SN signatures. Based on the light curve evolution, we obtained a spectrum of the optical counterpart and host galaxy with the 10.4 m Gran...
Table 1

<table>
<thead>
<tr>
<th>Mid Timea (days)</th>
<th>Magb</th>
<th>Error</th>
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<td>17.06</td>
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<td>NOT/ALFOSC</td>
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Notes.

a The middle time of each epoch, in units of days post-burst, relative to the Fermi trigger time of 07:47:06 UT on 2013 April 27, ~51 s ahead of the Swift trigger time of this burst.

b Magnitudes not corrected for a Galactic reddening of $E(B - V)_{\text{MW}} = 0.02$ mag.


Telescopio Canarias (GTC) 16.7 days after the GRB. This time corresponds to 12.8 days in the GRB rest frame. Observations were 4 × 1200 s, covering the range of 4800–10000 Å with a resolving power of ~600. The slit was oriented to cover both the afterglow position and the host galaxy nucleus.

All spectroscopic data were reduced in a standard way using dedicated IRAF pipelines or European Southern Observatory (ESO) pipelines. The resulting spectra are presented in Figure 4.

2.3. Host Galaxy

We use the cataloged pre-explosion imaging from the SDSS “DRS” (Aihara et al. 2011) to estimate physical properties of the GRB host galaxy and build a physical model of the stellar emission. The SDSS ugriZ photometry was fitted within the LePhare program (Arnouts et al. 1999; Ilbert et al. 2006) using stellar population synthesis models from Bruzual & Charlot (2003), as detailed in Krühler et al. (2011). Based on the model fit to the data (Figure 4), we derive a luminosity $L_B = -19.8 \pm 0.2$ mag, a stellar mass of $M_* = 10^{9.0\pm0.2} M_\odot$, a star-formation rate $SFR_{\text{SED}} = 2 \times 10^{-5} M_\odot$ yr$^{-1}$, and an age of the starburst $\tau = 400^{+300}_{-250}$ Myr for the host of GRB 130427A.

Host galaxy emission lines are detected above the SN continuum, including [N II] 6584 in the GTC spectrum, albeit at low significance. These detections allow us to place constraints on the metallicity of the explosion host environment using the calibrations in Pettini & Pagel (2004). We measure log (O/H) + 12 = 8.43 ± 0.07 and 8.51 ± 0.09 using the O3N2 and the N2 methods, respectively (statistical errors only), which is at the top right of the GRB-SN range in the metallicity–$M_{B,\text{host}}$ plane, similar to the cases of SNe Ic-BL without observed GRBs and SNe Ib+Ib (see Figure 2 in Modjaz et al. 2011). We note that the host galaxy also nicely lies within the 1σ dispersion of the mass–metallicity relation for normal field galaxies (Kewley & Ellison 2008). After including the systematic dispersion of 0.14 and 0.18 dex (Pettini & Pagel 2004) for the two methods, respectively, these results translate to a metallicity of 0.55 ± 0.19 and 0.67 ± 0.25 Z_⊙, respectively (Asplund et al. 2009).

3. SN 2013cq ASSOCIATED WITH GRB 130427A

3.1. Decomposition of the GTC Spectrum

We scaled the GTC spectrum by our simultaneous photometry from the same night, dereddened it by $E(B - V)_{\text{host}} = 0.02$ mag (Schlegel et al. 1998), where the subscript refers to the Milky Way galaxy, brought it to the same resolution, and then subtracted the model host galaxy spectrum. These steps resulted in a “clean” spectrum of the transient. Afterward, the spectrum was dereddened by $E(B - V)_{\text{host}} = 0.05$ mag for the host extinction at $z = 0.34$. Both the original and the final GTC spectra are shown in Figure 4.

Although the resulting spectrum is noisy, it shows clear SN features, with the most prominent being a strong bump peaking at ~6700 Å (observer-frame; ~5000 Å rest-frame). The features are broad (and no H or He can be seen), justifying the classification of SN 2013cq as an SN Ic-BL (de Ugarte Postigo et al. 2013).

SN 1998bw (Patat et al. 2001), associated with GRB 980425, does not provide a good spectral match to SN 2013cq, mainly

![Figure 4](image-url)
because its main peak is located more to the red (rest-frame $\sim 5200\AA$ at similar phases). The same is true for SN 2006aj (Pian et al. 2006; Sollerman et al. 2006), associated with GRB 060218. Instead, we find a better match with SN 2010bh, associated with GRB 100316D, which is known to have high expansion velocities, up to $10,000\text{ km s}^{-1}$ higher than other previous SNe Ic-BL associated with GRBs at all phases (Bufano et al. 2012). In particular, the best match is obtained with the spectrum of SN 2010bh at a rest-frame time of 12.7 days, very close to the rest-frame time of 12.5 days for SN 2013cq here (shown in Figure 4). The similarity with SN 2010bh is striking although there may be small color differences (e.g., the difference redward of $\sim 7500\AA$), which may reflect the diversity in GRB-SN spectra and/or the uncertain extinction correction toward SN 2010bh.

Considering that the P-Cygni feature on the left of the strong spectral features is consistent with SN 2010bh, associated with GRB 060218, $\sim 9.6 \times 10^{53}\text{ erg}$, comparable to that of high-redshift GRBs and much larger than local events. The fact that an SN progenitor model accounts for even very energetic bursts suggests that a common progenitor model, such as the collapsar model, may account for the majority of all long-duration GRBs.

To overcome the challenge of providing enough energy to power the GRB, it is likely that the large observed energy is due to beaming (for the strong beaming of GRB 130427A, see Laskar et al. 2013), making the true energy much lower (typically by two orders of magnitude).

4. DISCUSSION

Our photometric and spectroscopic campaign has led to the unambiguous discovery of an SN Ic-BL, SN 2013cq (de Ugarte Postigo et al. 2013), associated with GRB 130427A at $z = 0.34$. GRB 130427A is one of the most energetic bursts ever detected with $E_{\text{iso}} > 9.6 \times 10^{53}\text{ erg}$, comparable to that of high-redshift GRBs and much larger than local events. The fact that an SN progenitor model accounts for even very energetic bursts suggests that a common progenitor model, such as the collapsar model, may account for the majority of all long-duration GRBs.

Our discovery now suggests that not only core-collapse SNe, but specifically stripped envelope, high-velocity SNe, are almost an inevitable consequence of the deaths of stars that form all GRBs. A common mechanism is therefore at play powering GRBs with very different high-energy properties.

It is worth noting that the comparable peak B-band luminosities of SN 2013cq and SN 1998bw are consistent with the suggestion of Hjorth (2013) that there is an upper envelope to the brightness of GRB-SNe that drops slightly with isotropic luminosity. The origin of such an upper envelope is intriguing, but currently not clear.

As shown in Modjaz et al. (2011), among previous GRB-SN events, the highest oxygen abundance log (O/H) + 12 at the SN position was less than 8.3. These authors furthermore showed that when the oxygen abundance rises toward a seemingly critical value of 8.5, GRB-SN events tend to locate in dwarf galaxies with $M_{B,\text{host}} > -19.0\text{ mag}$. Note that the abundance value of 8.5 is a typical one for SNe Ic-BL, without observed GRBs, which occur in both luminous and dwarf galaxies. This abundance value is also typical for SNe Ib+Ib, which happen in relatively luminous galaxies with $M_{B,\text{host}} < -18.5\text{ mag}$ (Modjaz et al. 2011). With an abundance of 8.43 $\pm 0.07$ for the SN position and $M_{B} = -19.8 \pm 0.2\text{ mag}$ for the host, GRB 130427A/SN 2013cq is consistent with subclasses of core-collapse SNe such as SNe Ic-BL, without observed GRBs and SNe Ib+Ib in the metallicity–$M_{B,\text{host}}$ plane, implying that GRBs do not exclusively explode in low-metallicity dwarf galaxies.

\footnote{Note that we are using and transforming a bolometric light curve of SN 1998bw that has been constructed using observational data in $UBVRIJH$ filters.}
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