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Discovery of twin kHz quasi-periodic oscillations in the low-mass X-ray binary XTE J1701–407

Devraj D. Pawar,¹* Maithili Kalamkar,² Diego Altamirano,² Manuel Linares,³ K. Shanthi,⁴ Tod Strohmayer,⁵ Dipankar Bhattacharya⁶ and Michiel van der Klis²

¹Ramniranjan Jhunjhunwala College, Ghatkopar, Mumbai 400086, India
²Astronomical Institute ‘Anton Pannekoek’, University of Amsterdam, Science Park 904, NL-1098XH Amsterdam, the Netherlands
³Instituto de Astrofísica de Canarias, c/Vía Láctea s/n, E-38205 La Laguna, Tenerife, Spain
⁴UGC Academic Staff College, University of Mumbai, Mumbai 400098, India
⁵Astronomical Institute ‘Anton Pannekoek’, University of Amsterdam, Science Park 904, NL-1098XH Amsterdam, the Netherlands
⁶Inter University Centre for Astronomy and Astrophysics, Pune 411007, India

ABSTRACT

We report the discovery of kHz quasi-periodic oscillations (QPOs) in three Rossi X-ray Timing Explorer observations of the low-mass X-ray binary XTE J1701–407. In one of the observations we detect a kHz QPO with a characteristic frequency of 1153 ± 5 Hz, while in the other two observations we detect twin QPOs at characteristic frequencies of 740 ± 5, 1112 ± 17 Hz and 740 ± 11, 1098 ± 5 Hz. All detections happen when XTE J1701–407 was in its high-intensity soft state, and their single-trial significance is in the 3.1–7.5σ range. The frequency difference in the centroid frequencies of the twin kHz QPOs (385 ± 13 Hz) is one of the largest seen till date. The 3–30 keV fractional rms amplitude of the upper kHz QPO varies between ~18 and ~30 per cent. XTE J1701–407, with a persistent luminosity close to 1 per cent of the Eddington limit, is among the small group of low-luminosity kHz QPO sources and has the highest rms for the upper kHz QPO detected in any source. The X-ray spectral and variability characteristics of this source indicate its atoll source nature.

Key words: binaries: general – stars: individual: XTE J1701–407 – stars: neutron – X-rays: binaries.

1 INTRODUCTION

Low-mass X-ray binaries (LMXBs) can be divided into systems containing a black hole candidate and those containing a neutron star (NS). The accretion properties of these systems can be studied through the timing and spectral properties of the X-ray emission (e.g. van der Klis 1995a; Ford et al. 2000; Wijnands & Miller 2002). On the basis of correlated variations in the X-ray colour–colour diagram (CD) and power density spectra (PDS), the NS LMXBs are classified as Z sources and atoll sources (Hasinger & van der Klis 1989). The Z sources are generally high-luminosity sources (0.5–1.0 of Eddington luminosity $L_{\text{Edd}}$) while the atoll sources are low-luminosity sources (0.01–0.5$L_{\text{Edd}}$; see, e.g., van der Klis 2006b for a review, and Homan et al. 2010 for recent discoveries).

In the CD the atoll sources show three distinct states: the extreme island state (EIS), the island state (IS) and the banana state, the latter is further subdivided into the lower-left banana (LLB), lower banana (LB) and the upper banana (UB). In Fig. 1, we show the CD of the well-known atoll source 4U 1608–52 (e.g. van Straaten, van der Klis & Méndez 2003), where all the different atoll spectral states are seen. Generally, as the source moves from the EIS to the UB through the IS, LLB and LB, the spectrum softens and the soft X-ray intensity increases (see, e.g., Di Salvo, Méndez & van der Klis 2003; Schnerr et al. 2003).

A number of quasi-periodic oscillations (QPOs) and broad-band variability components are often present simultaneously in the PDS of the X-ray light curves of these systems. The timing characteristics of these variability components are related to the spectral state of the source, i.e. to the position of the source in the CD. Generally, the characteristics of the timing features [i.e. frequency, quality factor and fractional root-mean-square (rms) amplitude] vary monotonically as the source moves along the atoll track. This behaviour has been attributed to changes in the accretion rate ($\dot{M}$), the interaction of the matter in the accretion disc and the radiation emitted from the region close to the NS surface which affects the X-ray variability and spectrum (e.g. van der Klis 2006a for a review). QPOs can be seen with frequency between a few mHz to more than a kHz. Broad-band components are only seen up to ~100–200 Hz. Weak band-limited noise is seen in the power spectra of atoll sources when they are in the banana state. This noise becomes stronger as the source spectrum hardens, i.e. as the source moves to the ISs.

*E-mail: devrajdp@gmail.com
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2 OBSERVATIONS AND DATA ANALYSIS

2.1 Light curves and CD

To study the long-term (>days) X-ray variability, we use RXTE Proportional Counter Array (PCA; see Zhang et al. 1993; Jahoda et al. 2006) Galactic bulge scan monitoring observations (Swank & Markwardt 2001); the observations were taken once every ∼3 d; the intensity (in units of cts/sec/5PCUs) is provided in the 2–10 keV energy range.

To study the X-ray spectral variations and the short-term variability (<seconds), we use the 58 pointed observations obtained with the RXTE PCA between 2008 June 9 and 2011 September 20. Each observation is between 0.5 and 17 ks long, adding up to a total of ∼235.7 ks. For details on the observations, see Table 1. We calculated X-ray colours using 16 s time resolution Standard 2 mode data. Light curves were cleaned for instrumental effects like spikes and dropouts, and corrected for background contribution in each band using the standard faint source background model for the PCA (for details of the model, see PCA Digest at http://heasarc.gsfc.nasa.gov/

We define the soft colour as the ratio of average count rates per observation in the energy ranges 3.6–6.4 and 2.0–3.5 keV; the hard colour as the ratio of count rates in the energy ranges 9.7–16.0 and 6.4–9.7 keV; and the intensity as the count rate in the energy range 2–16 keV. The exact count rates in these energy bands were obtained by linearly interpolating between PCU channels. To correct for the difference in sensitivity between different PCUs, we normalized the count rates in each energy band by those of the Crab nebula in the same energy bands (e.g. Kuulkers et al. 1994; van Straaten et al. 2003). The Crab nebula observations used are those closest in time to each XTE J1701−407 observation. Given that the spectra of the source did not vary significantly within an observation (even in the long 17 ks observation), in this paper we report intensities and colours per observation.

2.2 Fourier timing analysis

We used 122 μs time resolution Event mode data available for 39 observations between 2008 June 9 and September 29. For the remaining 19 observations (2010 to 2011), the data are available in the Good Xenon mode which has a resolution of ∼1 μs. These data were binned to also obtain a resolution of ∼122 μs which results in a Nyquist frequency of 4096 Hz.

From the event list of each observation, we calculate the power spectra using fast Fourier transforms (FFTs) of continuous 16 s data segments (leading to a minimum frequency of 1/16 s = 0.0625 Hz) and using data in the 3–30 keV energy range (so as to optimize the signal-to-noise ratio). The 16 s power spectra are then averaged to get one power density spectrum per observation. In order to search for QPOs at very low frequencies, we also calculate FFTs of 1024 s data segments, which gives a minimum frequency of 1/1024 s = 0.976 mHz. In both cases, no dead time or background corrections are done before calculating the power density spectrum. We estimate the dead time corrected Poisson noise spectrum using the analytic function in Zhang et al. (1995). The estimated Poisson noise spectrum is subtracted from the power density spectrum which is then converted to rms normalization (van der Klis 1995b). In the rms normalization, the square root of the integrated power gives the fractional rms amplitude of the source variability in the frequency range over which we integrate.

Figure 1. CD of the atoll source 4U 1608−52 plotted using observations obtained over 14 years. Each point represents the averaged colour per observation (bands are defined as in Section 2.1). The error bars are of the order of the size of the symbols. In this CD all the spectral states are observed. The kHz QPOs are usually observed in the LB states.

Usually, one or two kHz QPOs are detected in the LLB and the LB. No kHz QPOs are generally detected in the UB nor in the harder EIS states (van der Klis 2000, 2006b; Altamirano et al. 2008b). In the hard ISs, two broad features have been suggested as the equivalent of kHz QPOs at low frequencies (<400 Hz; e.g., Psaltis, Belloni & van der Klis 1999; van Straaten, van der Klis & Wijnands 2004, 2005, and references therein).

The Keplerian velocities in the intense gravitational fields near NSs are very high. Accreting matter therefore has very short (milliseconds) orbital periods which may be the cause of the kHz QPOs. In some models (e.g. Miller, Lamb & Psaltis 1998; van der Klis 2000; Lamb & Miller 2003; Lee, Abramowicz & Kluźniak 2004) and observations (e.g. Jonker, Méndez & van der Klis 2002; Markwardt & Swank 2003; Wijnands et al. 2005), the difference in the frequency of the twin kHz QPOs (∆ν) is thought to be related to the NS spin frequency (νs) as ∆ν ≃ v0 or ∆ν ≃ νs/2. However, current data do not allow a definitive statement about this (see, e.g., Méndez & Belloni 2007; Yin et al. 2007; van der Klis 2008; Altamirano et al. 2010).

The Rossi X-ray Timing Explorer (RXTE) has been one of the most successful X-ray astronomy missions. From 1995 to 2012, RXTE has been used for observing known X-ray sources and also for discovering many new ones. An example of the latter is XTE J1701−52, which has been used for discovering many new ones. An example of the latter is XTE J1701−52.
### 2.3 Contaminated observations

Of the 58 RXTE observations, the first 39 observations, performed in 2008 (proposal P93444), were obtained with an RA–Dec. pointing of 255.35, −40.5 and the later 19 observations, performed in 2010–2011 (proposal P95328), were obtained with an RA–Dec. pointing of 255.43, −40.86. During the 2008 observations, the high-mass X-ray binary (HMXB) pulsar OAO 1657−415 was at $\sim 1.17$ from the PCA pointing coordinates, whereas during the 2010–2011 ones the HMXB was at $\sim 0.83$. Therefore, in the latter case, the HMXB is in the 1° FWHM field of view (FoV) of the RXTE-PCA observations of XTE J1701−407. To explore whether OAO 1657−415 affected our observations, we created 1024 s power spectra of all 2010–2011 observations. In observations 95328-01-01-04, 95328-01-02-00, 95328-01-03-00 and 95328-01-04-00, we find a very narrow peak which has a frequency of 27.2 Hz. A period search using the ftools. efitsearch reveals a strong signal at 37.063 s corresponding to a frequency 26.981 Hz. This is consistent with the spin period of OAO 1657−415 (Barnsted, Denis, Bulik & Markinowski 2010). Given that the 2010–2011 observations are clearly contaminated by the flux from OAO 1657−415, we do not use them in the CD or intensity plots but do use them in our timing analysis, as no high-frequency features ($\geq 1.27$ Hz) have ever been detected from HMXBs (Kaur et al. 2007, and references therein). We also analysed all the 86 RXTE observations of OAO 1657–415 for the presence of kHz QPOs and detected none.

3 RESULTS

3.1 Light curve and CD

In the upper panel of Fig. 2, we show the bulge scan long-term light curve of XTE J1701−407. After the outburst onset (detected on 2008 June 8, MJD 546 25), the light curve shows instances of increased emission on top of an average 2–10 keV intensity of $\sim 100$ cts/s/5PCU, during which the intensity increases to nearly two to four times the average. After a 3 yr long outburst since its discovery, the source went below the RXTE Galactic bulge scan monitor detection limit around 2011 August 27 (Degenaar et al. 2011a). A few days later, it rebrightened to $\sim 100$ cts/s/5PCU from 2011 September 16 to 24 (Degenaar et al. 2011b); after that its intensity decreased to $\sim 10$ cts/s/5PCU and remained so till it was last monitored by RXTE on 2011 October 29 (MJD 558 63.4). In the lower panel of Fig. 2, we show the Crab normalized light curve for 2008 observations. These observations sample the source when its intensity was between $\sim 2$ and $\sim 25$ mCrab. In Fig. 3, we show the CD of the 2008 observations. It is clear that the observations are confined to a few distinct regions in the diagram. We refer to the 27
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Figure 2. Top panel: long-term light curve of XTE J1701−407 obtained from the PCA Galactic bulge scan monitoring observations. The arrows and circles mark the approximate times and intensities of the observations in which kHz QPOs are detected. The circles are offset from the PCA scan points as they represent the average count rate per pointed observation. The grey shadow marks the period in 2008 when 39 RXTE pointed observations were performed. The period when the 19 RXTE pointed observations obtained from 2010 January 5 to 2011 September 20 were performed is indicated by the dashed line in the upper-right part of the panel. Bottom panel: normalized light curve of the 39 pointed observations performed in 2008. The error bars are of the order of the size of the symbols. The green circles and blue triangles mark when the source was in the hard and soft state, respectively (see also Fig. 3). We do not show the remaining 19 observations as their intensity/colours are contaminated by the flux of an HMXB, see Section 2.3.

As can be seen in Fig. 2, the first two pointed observations of XTE J1701−407 occur when it was in the hard state (on MJD 546 26 and 546 27). Then the source was detected in the high-luminosity soft state on MJD 546 33 in one observation. As the intensity decreased, the source went to the hard state again where it was observed for ~5 d. After a gap of ~27 d, XTE J1701−407 was observed in the soft state for ~17 d after which its $L_X$ decreased and underwent a transition to the hard state; our data cover ~14 d before they stop. After an ~25 d gap, the source was in the soft state for 1 d (MJD 547 21); within a day XTE J1701−407 transited to the hard state again. Clearly, the soft-state observations track the high-intensity periods; however, there is no clear intensity cut that

Figure 3. Top panel: CD of XTE J1701−407 using the 39 pointed RXTE observations performed in 2008. The arrows indicate observations in which the kHz QPOs are detected. The green circles and blue triangles mark the hard-state (low $L_X$) and the soft-state (high $L_X$) observations, respectively. Comparing with Fig. 1, the hard-state observations lie in the lower region of the $IS$ while the soft-state observations correspond to the banana state. Middle panel: soft colour versus intensity diagram. Bottom panel: hard colour versus intensity diagram. The spectrum becomes softer as the $L_X$ increases. The kHz QPOs are seen when the spectrum is soft and $L_X$ is high. The observations contaminated by a nearby HMXB are not included in the CD (see Section 2.3).
indicates whether the source is in the soft or hard state (as noted by Linares et al. 2009). This is probably due to hysteresis effects, which are typical for this type of systems (Maccarone & Coppi 2003; Meyer-Hofmeister, Liu & Meyer 2005).

### 3.2 Aperiodic variability

We examined all the power spectra in the energy band 3–30 keV individually for variability. We detect significant (>3σ single-trial) QPOs in three different observations, the times of which are marked with arrows in the upper panel of Fig. 2. In Fig. 3, we indicate only the two uncontaminated observations of 2008. It should be noted that all these three QPOs were detected during the high-intensity state. The source is in the soft state during observations 93444-01-02-00 and 93444-01-04-00 (Fig. 3); the spectral state during the observation 95328-01-06-00 is difficult to constrain as the spectra are contaminated by an HMXB in the FoV (OAO 1657−415; see Section 2.3).

In the observation 93444-01-02-00 (MJD 546 33.2), we detect two QPOs at characteristic frequencies 30 ± 0.3 and 1152 ± 5 Hz. In Fig. 4, we show them separately for clarity. The 30 ± 0.3 Hz QPO is accompanied by a broad noise component fitted with a Lorentzian centred at 17.2 ± 5.9 Hz. In observations 93444-01-04-00 (MJD 546 65.1) and 95328-01-06-00 (MJD 554 25.08), we detect twin kHz QPOs at characteristic frequencies νℓ = 740 ± 2 and 738 ± 9 Hz and νu = 1112 ± 17 and 1098 ± 5 Hz, respectively (see Fig. 5).

In the latter case (95328-01-06-00), we also detect a broad bump at 26.3 ± 5.6 Hz and a fractional rms amplitude of ~11 per cent. In the observation 93444-01-04-00, this broad bump is not detected significantly, with a 3σ upper limit of 19 per cent fractional rms amplitude. The difference in the centroid frequencies of the twin kHz QPOs is ∆ν = 385 ± 13 and 360 ± 10 Hz for the two observations, respectively. In Table 2, we report on the best-fitting parameters for the detected features with single-trial significance >3σ.

To calculate the total number of trials, we use the fact that we needed to fix the quality factor Q in our fits (as the best fit gave negative values for these parameters). We performed a Monte Carlo simulation to estimate the significance of the detected QPOs, taking into account the number of trials and the background noise. The results are shown in Table 2. Figure 4 shows an example of the power spectra with the QPOs detected simultaneously in the observation 93444-01-02-00. For clarity, the left-hand panel shows the low-frequency part of the power spectra (QPO is at 30.4 ± 0.3 Hz) and in the right-hand panel the kHz QPO (ν = 1152.7 ± 5.2 Hz). The average count rate in this observation is 39.6 cts/s/PCU after correcting for the background; two detectors were on during this observation.
Figure 5. Left-hand panel: twin kHz QPOs observed simultaneously in the observation 93444-01-04-00 (at 740.8 ± 2.6 and 1112.6 ± 17.0 Hz). The average count rate per detector in this observation is 41.7 cts/s/PCU after correcting for the background. Two detectors were on during this observation. Right-hand panel: twin kHz QPOs observed simultaneously in the observation 95328-01-06-00 (at 740.5 ± 11.7 and 1097.8 ± 5.8 Hz). The average count rate per detector in this observation is 54.5 cts/s/PCU after correcting for the background and two detectors were on during this observation.

Table 2. QPO parameters in the energy range 3–30 keV.

<table>
<thead>
<tr>
<th>Observation</th>
<th>ν max (Hz)</th>
<th>rms (per cent)</th>
<th>Q</th>
</tr>
</thead>
<tbody>
<tr>
<td>93444-01-02-00</td>
<td>17.2 ± 5.9</td>
<td>20.9 ± 3.3</td>
<td>0.09 ± 0.25</td>
</tr>
<tr>
<td>93444-01-04-00</td>
<td>740.8 ± 2.6</td>
<td>11.2 ± 1.6</td>
<td>15.1 ± 9.2</td>
</tr>
<tr>
<td>93444-02-00</td>
<td>1152.7 ± 5.2</td>
<td>26.8 ± 2.2</td>
<td>31.5 ± 10.6</td>
</tr>
<tr>
<td>93444-01-04-00</td>
<td>740.8 ± 2.6</td>
<td>11.2 ± 1.6</td>
<td>15.1 ± 9.2</td>
</tr>
<tr>
<td>95328-01-06-00</td>
<td>26.3 ± 5.6</td>
<td>11.5 ± 1.7</td>
<td>0.7 ± 0.5</td>
</tr>
<tr>
<td>95328-01-06-00</td>
<td>740.5 ± 11.7</td>
<td>14.2 ± 2.3</td>
<td>13.1 ± 7.4</td>
</tr>
<tr>
<td>95328-01-06-00</td>
<td>1097.8 ± 5.8</td>
<td>18.5 ± 1.9</td>
<td>20.7 ± 5.7</td>
</tr>
</tbody>
</table>

values; note that this is common practice when the component is too broad and is consistent with a Lorentzian of νσ = 0 Hz; see, e.g., Belloni et al. 2002; van Straaten et al. 2005); we note that the frequency ν and fractional rms did not change significantly before and after the value of Q was fixed. The averaged power spectrum of the soft-state observations (excluding the contaminated observations and those in which we detect QPOs) has a 0.1–100 Hz integrated fractional rms amplitude of 7.16 ± 1.61 per cent and shows no significant features; we detected only an ~50 Hz signal with a significance of 2.6σ. For this power spectrum, we estimated 7.01, 13.4 and 12.6 per cent confidence upper limits for ~30, ~740 and ~1110 Hz QPOs, respectively, with parameters similar to those in Table 2.

3.3 Correlation of νu and the other characteristic frequencies

The characteristic frequencies of the various power spectral components (so-called $L_b$, $L_{hHz}$, $L_{\ell}$, $L_m$) are correlated to the characteristic frequency νu of the upper kHz QPO (see e.g. van Straaten et al. 2003, 2005; Altamirano et al. 2008b, and references within).

In Fig. 7, we plot the characteristic frequency of the various components versus νu for XTE J1701−407, and for the atoll sources 4U

Table 3. Broad-band noise parameters measured in the averaged power spectrum of the hard-state observations (circles in the CD, Fig. 3); energy range 3–30 keV.

<table>
<thead>
<tr>
<th>ν max (Hz)</th>
<th>rms (per cent)</th>
<th>Q</th>
</tr>
</thead>
<tbody>
<tr>
<td>0.52 ± 0.07</td>
<td>11.6 ± 0.56</td>
<td>0(fixed)</td>
</tr>
<tr>
<td>20.6 ± 5.16</td>
<td>22.47 ± 1.5</td>
<td>0(fixed)</td>
</tr>
</tbody>
</table>
to match the energy range used by Jonker et al. (2001). The flux upper limit used for XTE J1701—407 in Fig. 8 was obtained from Linares et al. (2009) and data for the other sources from Ford et al. (2000) and Jonker et al. (2001).

### 3.5 Correlation of rms and energy

The fractional rms amplitude of QPOs is a measure of the fraction of observed photons which are modulated at the QPO frequency, and so can give additional information useful to understand the physical process that sets the amplitude and/or the frequency of the oscillation (see Cabanal et al. 2010; Gierliński & Zdziarski 2005; Zdziarski et al. 2005, and references within for extensive discussion). The fractional rms amplitude of QPOs increases with energy (the only clear exceptions are the mHz QPOs in three atoll sources thought to be due to marginally stable burning of hydrogen or helium on the NS surface; see, e.g., Altamirano et al. 2008a; Revnivtsev et al. 2001, and references therein). In Fig. 9, we show the fractional rms amplitude versus energy of the kHz QPO at ~1112 Hz in the observation 93444-01-04-00. We plot two points (energy bands 2.06–5.71 and 6.12–31.7 keV; more points lead to larger errors). Points from the observation 93444-01-02-00 are not included because of very low statistics and points from 95328-01-06-00 are not included as it is one of the contaminated observations. For comparison we also include the data for the atoll source 4U 1608—52 (Berger et al. 1996; Méndez et al. 1998). We see that the rms depends weakly on energy for XTE J1701—407, significantly different from the relatively luminous source 4U 1608—52, where it increases monotonically with energy.

### 4 DISCUSSION

In this paper, we report the discovery of kHz QPOs in the LMXB XTE J1701—407. XTE J1701—407 is one of the least luminous LMXBs that show kHz QPOs at an average flux of ~3 × 10^{-10} erg cm^{-2} s^{-1} in the 2–20 keV band (Falanga et al. 2009); from time to time it shows a sudden excursion to the soft state as the source intensity increases. The increased emission is most probably because of an increase in the mass accretion rate (Ṁ). When Ṁ increases, it leads to high L_2 and the spectrum softens. This is a well-known behaviour in compact accreting objects, (see, e.g., Done & Gierliński 2003; Done, Gierliński & Kubota 2007, and references within). Based on the transition tracks between hard and soft states in the CD, Gladstone, Done & Gierliński (2007) have classified atoll sources as diagonals and verticals. Comparison of the first and last panels of Fig. 3 with the second row of fig. 2 in Gladstone et al. (2007) implies that XTE J1701—407 behaves as a vertical. According to Gladstone et al. (2007), this means Ṁ is low enough to allow the magnetic field to emerge from the surface of the NS and affect the accretion flow. Given that the CD of XTE J1701—407 resembles that seen in other low-luminosity sources, it shows thermonuclear X-ray bursts (Markwardt et al. 2008; Falanga et al. 2009; Linares et al. 2009; Chenevez et al. 2010), strong broadband noise in the hard state and kHz QPOs when the source flux is highest and spectra softest, we conclude that XTE J1701—407 can be classified as an atoll source (see also Linares et al. 2009).

### 4.1 Quasi-periodic oscillations

Twin kHz QPOs were detected on two occasions with centroid frequency differences of Δν = 385 ± 13 and 360 ± 10 Hz, i.e. Δν is the same within errors in both cases. This Δν is among
Figure 8. Luminosity versus fractional rms amplitude (5–60 keV) of the upper kHz QPO in various LMXBs from Jonker et al. (2001). We also include the data for 1A 1246−588 from Jonker et al. (2007) and for XTE J1701−407 from our analysis. The $L_X/L_{\text{Edd}}$ for XTE J1701−407 is calculated using a distance of $5.0 \pm 0.4$ kpc from Chenevez et al. (2010); the error bars for $L_X/L_{\text{Edd}}$ are much smaller than the size of the symbols.

Figure 9. Energy dependence of the fractional rms amplitude of the upper kHz QPO in XTE J1701−407 (filled black diamonds) compared with 4U 1608−52, $L_e$ (filled circles) and $L_u$ (filled squares) (Berger et al. 1996; Méndez et al. 1998). The error bars on rms for the filled circles are smaller than the size of the symbols. We overplot reference lines given by $f(x) = ax$, where $a = 1, 1.8, 3$.

4.2 kHz QPO fractional rms amplitude versus luminosity

Jonker et al. (2001) showed that there was an anticorrelation between the kHz QPO fractional rms amplitude and the X-ray luminosity $L_X$ of the source when the QPOs were detected. This was further supported by the results of Jonker et al. (2007), who found $27 \pm 3$ per cent fractional rms amplitude kHz QPOs in the low-luminosity NS 1A 1246−588. As shown in Fig. 8, our discovery of kHz QPOs on XTE J1701−407 further supports the anticorrelation. As can be seen, the anticorrelation has some dispersion, and although kHz QPOs have been detected in sources at lower $L_X$ than XTE J1701−407 (e.g. 4U 0614+109 and 2S 0918−549), the kHz QPO fractional rms amplitude in XTE J1701−407 is the highest rms reported as yet for an upper kHz QPO. To date there is no clear picture that explains this anticorrelation (e.g. Jonker et al. 2001, 2007) and, as suggested by Jonker et al. (2001), it should be considered when modelling the mechanisms producing the kHz QPOs.

Méndez (2006) studied the relation between the maximum fractional rms amplitude of the kHz QPOs observed in a source and the source luminosity, and found that for the upper kHz QPO, the fractional rms amplitude was approximately constant ($\sim 20$ per cent) at $< 0.1 L_{\text{Edd}}$ (2–60 keV), and decreased at $> 0.1 L_{\text{Edd}}$ (see fig. 3 of Méndez 2006). The results on the kHz QPO in 1A 1246−588 by Jonker et al. (2007) cast doubt on this trend, and led those authors to the conclusion that either the relation between kHz QPO frequency and fractional rms is significantly different in 1A 1246−588 from that which is seen in other NS LMXBs or the increase in amplitude and source luminosity does not level off at $\sim 20$ per cent (Méndez et al. 2007).
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