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HIGH-CONTRAST NEAR-INFRARED IMAGING POLARIMETRY OF THE PROTOPLANETARY DISK AROUND RY TAU

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ABSTRACT

We present near-infrared coronagraphic imaging polarimetry of RY Tau. The scattered light in the circumstellar environment was imaged at the $H$ band at a high resolution (~0′′05) for the first time, using Subaru/HiCIAO. The observed polarized intensity (PI) distribution shows a butterfly-like distribution of bright emission with an angular scale similar to the disk observed at millimeter wavelengths. This distribution is offset toward the blueshifted jet, indicating the presence of a geometrically thick disk or a remnant envelope, and therefore the earliest stage of the Class II evolutionary phase. We perform comparisons between the observed PI distribution and disk models with (1) full radiative transfer code, using the spectral energy distribution (SED) to constrain the disk parameters; and (2) monochromatic simulations of scattered light which explore a wide range of parameters space to constrain the disk and dust parameters. We show that these models cannot consistently explain the observed PI distribution, SED, and the viewing angle inferred by millimeter interferometry. We suggest that the scattered light in the near-infrared is associated with an optically thin and geometrically thick layer above the disk surface, with the surface responsible for the infrared SED. Half of the scattered light and thermal radiation in this layer illuminates the disk surface, and this process may significantly affect the thermal structure of the disk.

Key words: polarization – protoplanetary disks – stars: individual (RY Tau) – stars: pre-main sequence

Online-only material: color figures

1. INTRODUCTION

Optical and near-infrared observations have revealed structures in protoplanetary disks at the highest angular resolutions currently available. This has provided powerful tools for investigating the possibility of ongoing planet formation and to test the related theories. In particular, the technique of coronagraphic imaging has been extensively used to suppress the stellar flux and detect scattered light from the disk surface with high sensitivities (see Watson et al. 2007 for a review). Without this technique, resolved images of protoplanetary disks can only be
obtained in limited circumstances at optical and near-infrared wavelengths, via an edge-on view or silhouette against bright background nebular emission (e.g., McCaughrean et al. 2000 for a review).

Using coronagraphy, the ongoing survey “Strategic Explorations of Exoplanets and Disks with Subaru” (SEEDS; Tamura 2009) with Subaru/HiCIAO (Tamura et al. 2006) and AO188 (Hayano et al. 2004) has recently discovered structures in a number of protoplanetary disks (Thalmann et al. 2010), in particular in the polarized intensity (hereafter PI) distribution in the near-infrared (Hashimoto et al. 2011, 2012; Muto et al. 2012; Kusakabe et al. 2012; Tani et al. 2012; Mayama et al. 2012; Dong et al. 2012a; Grady et al. 2013; Follette et al. 2013). PI imaging has been used for observations of most of the disks in the SEEDS program, since this suffers significantly less contamination from the stellar flux than the normal intensity $I$. Some disks are associated with spiral structures, disk holes, or azimuthal ring-like flux distributions, which are potential signatures of ongoing planet formation. The scattered light from the disk is also useful for probing grain growth (e.g., Min et al. 2012; Tani et al. 2012) which may be related to formation of rocky cores.

The goal of the SEEDS program for protoplanetary disks is to observe a large number of objects with different stellar masses and ages and understand the evolution of disk structures and grain growth, and therefore the environment of possible ongoing planet formation. In this paper we present near-infrared coronagraphic imaging of RY Tau from the SEEDS program, the first publication of the near-infrared scattered light associated with the disk around this star. RY Tau is an active pre-main-sequence star with a stellar mass of 2 $M_\odot$ (Calvet et al. 2004; Isella et al. 2009). The estimated age ranges from 0.5 (Isella et al. 2009) to 8 Myr (Calvet et al. 2004). The star is associated with a relatively massive disk ($3 \times 10^{-3}$ to $10^{-1} M_\odot$; Isella et al. 2009, 2010) with a large infrared excess at near-to-far infrared wavelengths (e.g., Robitaille et al. 2007), an optical jet (St-Onge & Bastien 2008; Agra-Amboage et al. 2009), a scattering nebulosity due to the remnant of the envelope (e.g., Nakajima & Golimowski 1995; St-Onge & Bastien 2008), and a large time variation in optical photometry and spectroscopy (e.g., Petrov et al. 1999). These indicate a relatively young evolutionary phase. Despite this, the Hz equivalent width observed over the last 30 yr is relatively low (8–20 Å; see Chou et al. 2013 and references therein), similar to more evolved pre-main-sequence stars.

Recent observations with millimeter interferometry seem to show evidence of a hole in the disk seen at 1.3 mm, with a radius of 15 AU (Isella et al. 2010). Such a hole can be made by tidal interaction between the inner disk and protoplanets (e.g., Papaloizou 2007; Zhu et al. 2011). These disks are called “transitional disks” (e.g., Hughes et al. 2009; Brown et al. 2009, and references therein; Mayama et al. 2012). However, RY Tau is different from the other transitional disks because it still drives a jet and has no evidence for a deficit of warm thermal emission near 10 µm (e.g., Robitaille et al. 2007).

The rest of the paper is organized as follows. In Section 2, we summarize our observations and data reduction. In Section 3, we show the observed PI distribution and polarization. We then perform comparisons with simulations of scattered light using conventional disk and dust models with the two approaches described below. In Section 4, we use the full radiative transfer tools developed by Whitney et al. (2003b) and Robitaille et al. (2006, 2007) with a disk geometry obtained by fitting the spectral energy distribution (SED). The model SEDs include the processes of scattering, absorption, and re-radiation on dust grains at all the wavelengths from UV to radio. In Section 5, we use monochromatic simulations of scattered light with our own dedicated code to attempt to better fit the observed PI distribution. Although the thermal structure and re-radiation process are not included, the emission from the disk surface is dominated by scattered light on dust grains, and this simplification allows us to conduct simulations covering a large parameter space in the disk geometry and grain size distribution.

In Section 6, we discuss the implications for the scattering geometry and dust grains, and the possible origins of the non-axisymmetry in the observed PI distribution. Throughout the paper we adopt a distance to the target of 140 pc (Wichmann et al. 1998).

2. OBSERVATIONS AND DATA REDUCTION

Observations were made on 2011 January 27 using HiCIAO and AO 188 at Subaru 8.2 m. As with several other SEEDS observations, the polarization was measured by rotating the half-wave plate to four angular positions (in the order 0°, 45°, 22.5°, and 67.5°) using the PDI (polarization differential imaging) +ADI (angular differential imaging) mode. A single Wollaston prism was used to split incident light into two images, each with a 20′′ × 9′′ field of view and a pixel scale of 9.5 mas pixel$^{-1}$. We obtained 13 full wave plate rotation cycles, taking a 30 s exposure per wave plate position, with a 0′/3 diameter coronagraphic mask. The field rotation was ∼8° during the observations.

The data were reduced using the standard approach for polarimetric differential imaging (Hinkley et al. 2009) as well as the other SEEDS studies (e.g., Hashimoto et al. 2011, 2012; Kusakabe et al. 2012). The reduction was made using the Image Reduction and Analysis Facility (IRAF), pyRAF, and python. Telescope and instrument polarization was corrected following Joo et al. (2008).

We find that the intensity ($I$) distribution observed in the disk region varies between exposures. This implies that the correction of the point-spread function (PSF) with adaptive optics (AO) was not stable during the observations due to the moderate quality observing conditions, resulting in the stellar flux leaking out to a different degree in different images. We therefore use four images (at the wave plate positions of 0°, 45°, 22.5°, and 67.5°) with the minimum intensity distribution (i.e., minimum flux for the halo of the PSF associated with the star) to derive a lower limit for the degree of polarization. Even in these images, the $I$ distribution is centrosymmetric, i.e., very different from that expected for the disk associated with RY Tau (Sections 3–5), suggesting that they are severely contaminated by the stellar flux.

The unstable AO correction also resulted in different PI distributions during the observations. To investigate this effect, we have calculated the co-added PI images in a few different ways: i.e., taking an average or median for all the data sets (i.e., 13 full wave plate rotation cycles), or selecting the best data sets and averaging them. All of the methods provide almost identical results. In the rest of the paper, we use the median images for the 13 full wave plate rotation cycles, with a total integration time of 1560 s.

26 IRAF is distributed by National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.
obscuration by circumstellar dust (Section 6.3). This discrepancy can be attributed to variability due to RY Tau to investigate the physical parameters of the star, disk, and envelope. This fitting is made via comparison between the observed SED and a grid of 200,000 modeled SEDs covering a wide range of parameter space calculated by Robitaille et al. (2006). Once we derive the parameters for the star, disk, and envelope, this region is contaminated by an artifact caused by the variation of the halo in the PSF of the star during the modest observing conditions. From our image we estimate an upper limit for PI/I∗ of 0.4 × 10−7 pixel−1.

The polarization vectors show a centrosymmetric pattern, as observed in several other disks. This is consistent with the scattered flux being dominated by a single scattering, with multiple scatterings being negligible (Section 6.2). As described in Section 2, the degree of polarization shown here is a lower limit due to contamination from the PSF halo of the star in the I flux distribution. We measure a lower limit for the degree of polarization of 9% ± 1% at the peak in the southwest. The actual degree of polarization due to the scattered light would be significantly higher than this limit.

4. CONVENTIONAL FULL RADIATIVE TRANSFER MODELS WITH FITTING SPECTRAL ENERGY DISTRIBUTIONS

Robitaille et al. (2007) developed a fitting tool for the SEDs of young stellar objects, and applied it to sources including RY Tau to investigate the physical parameters of the star, disk, and envelope. This fitting is made via comparison between the observed SED and a grid of 200,000 modeled SEDs covering a wide range of parameter space calculated by Robitaille et al. (2006). Once we derive the parameters for the star, disk, and envelope using the above SED fitter, we can obtain the simulated PI image of the scattered light in the disk and envelope using the Robitaille et al. (2006) code.

In this section, we use this approach to attempt to reproduce the observed PI image. In Section 4.1, we briefly summarize the models based on the description in Robitaille et al. (2006) and Whitney et al. (2003a, 2003b), and our procedure for obtaining the simulated PI images. In Section 4.2, we show the results of the SED fitting and the PI image of the disks based on the fitted parameters.

3. OBSERVED PI DISTRIBUTION

Figure 1 shows the distribution of the PI flux overlaid with the polarization vectors. The PI distribution is elongated along the major axis of the disk observed via millimeter interferometry (P.A. = 24°/204°; Isella et al. 2010). Its angular scale is ∼1′0 (∼140 AU) and ∼0′6 (∼80 AU) along the major and minor axes, respectively. The angular scale for the major axis is similar to that observed by Isella et al. (2010).

The bright part of the PI emission is offset from the star toward the blueshifted jet (P.A. = 294°; St-Onge & Bastien 2008; Agra-Amboage et al. 2009), i.e., the far side of the disk. The PI distribution shows a minimum along the direction of the jet (P.A. = 294°), and increases to maxima at P.A.s of ∼210° and ∼350°. This butterfly-like morphology in the PI distribution is similar to that modeled for some of the flared disks seen at optical wavelengths by Min et al. (2012). While the morphology in the PI distribution is relatively symmetric about the jet (and disk) axis, its brightness is asymmetric. The southwest side is brighter than the northern side by a factor of ∼2 near the peaks (PI/I∗ ∼ 8 × 10−7 and ∼3 × 10−7 per pixel, respectively).

The observed PI distribution, asymmetric with respect to the major axis of the disk, contrasts with the other disks observed in the SEEDS program (Hashimoto et al. 2011, 2012; Muto et al. 2012; Kusakabe et al. 2012; Tanii et al. 2012; Mayama et al. 2012; Grady et al. 2013; Follette et al. 2013), in which the PI distribution is generally symmetric. This indicates that RY Tau is associated with a geometrically thick disk or a remnant envelope, and therefore at a younger evolutionary stage than the others (see, e.g., Fischer et al. 1994; Whitney et al. 2003a). This agrees with the stellar age of ∼0.5 Myr estimated by Isella et al. (2010), i.e., the star is in the earliest stage of the Class II phase, in which the star becomes visible at optical wavelengths (e.g., Stahler & Palla 2005).

Simulations show that such a disk is associated with faint emission at the other side of the disk (Fischer et al. 1994; Whitney et al. 2003a; Section 5). Although Figure 1 shows a similar PI distribution in the northeast to south of the star, this region is contaminated by an artifact caused by the variation of the halo in the PSF of the star during the modest observing conditions. From our image we estimate an upper limit for PI/I∗ of 0.4 × 10−7 pixel−1.

The polarization vectors show a centrosymmetric pattern, as observed in several other disks. This is consistent with the scattered flux being dominated by a single scattering, with multiple scatterings being negligible (Section 6.2). As described in Section 2, the degree of polarization shown here is a lower limit due to contamination from the PSF halo of the star in the I flux distribution. We measure a lower limit for the degree of polarization of 9% ± 1% at the peak in the southwest. The actual degree of polarization due to the scattered light would be significantly higher than this limit.
Note that we use an axisymmetric distribution of the circumstellar material for the simulations in this and following sections. This implies that, in principle, these simulations cannot reproduce the asymmetry at the PI distribution about the rotation axis of the disk (and the jet axis) described in Section 3. Simulations for an asymmetric PI distribution are beyond the scope of this paper. The implications for this asymmetry is discussed in Section 6.3.


The modeled system consists of a pre-main-sequence star with an axisymmetric circumstellar disk, an infalling flattened envelope, and an outflow cavity. The SED is determined using full radiative transfer in the disk and envelope, i.e., including absorption, scattering and re-radiation of light by dust grains in addition to the flux directly observed from the star. The heating sources of the disks are stellar radiation and viscosity in the disk. The stellar spectrum is parameterized by the star’s effective temperature and radius, via comparisons with modeled spectra for stellar photospheres (Kurucz 1993; Brott & Hauschildt 2005).

The density distribution of a standard flared accretion disk (e.g., Shakura & Sunyaev 1973; Lynden-Bell & Pringle 1974) is described in cylindrical coordinate \((r,z)\) by

\[
\rho(r,z) = \rho_0 \left[ 1 - \sqrt{\frac{R_*}{r}} \right]^\alpha \left( \frac{R_*}{r} \right)^{\alpha - 1} \exp \left\{ -\frac{1}{2} \left[ \frac{z}{h} \right]^2 \right\},
\]

where \(\rho_0\) is a constant to scale the density, \(R_*\) is the stellar radius, \(\alpha\) is the radial density exponent, and \(h\) is the disk scale height. The scale height \(h\) increases with radius as \(h = h_0 r^\beta\), where \(\beta\) is the flaring power \((\beta > 0)\). According to more detailed models of disk structures by Chiang & Goldreich (1997) and D’Alessio et al. (1999), the scale height exponent varies with radius in the inner disk but follows a similar power law beyond a few AU (Cotera et al. 2001). In addition to the above parameters the minimum and maximum radii of the disk are included as free parameters. The gas-to-dust mass ratio is assumed to be 100.

Throughout the simulations, Robitaille et al. (2006, 2007) assume \(\alpha = \beta + 1\). This yields the surface density distribution \(\Sigma(r) \propto r^{-1}\), approximately agreeing with that inferred from millimeter interferometry for disks associated with many low-mass pre-main-sequence stars (Andrews et al. 2009, 2010). This surface density power index may not be consistent with the one for the RY Tau disk seen at millimeter wavelengths by Isella et al. (2010). Even so, models with the above assumption fit the observed SEDs well, as shown in Section 4.2.

The density structure for the envelope is given by Ulrich (1976) and Terebey et al. (1984). The envelope is associated with a cavity whose shape varies as \(z \propto r^{1.5}\). Robitaille et al. (2006) contain further details for the envelope and outflow cavity. Although the envelope mass, cavity density, etc., will be calculated using the SED fitter, their contribution to the SEDs of optically visible pre-main-sequence stars such as RY Tau will be negligible (Section 4.2).

The dust composition assumes a mixture of astronomical silicates and graphite in solar abundance without an ice coating. The following two grain size distributions are used based on Wood et al. (2002) and Whitney et al. (2003a, 2003b): (1) the distribution in the denser regions of the disk \((m_{H_2} > 10^{10} \text{ cm}^{-3})\) where one would expect significant grain growth up to 1 mm; and (2) the distribution in the more diffuse regions \((m_{H_2} < 10^{10} \text{ cm}^{-3})\) in which the grain sizes are slightly larger than those in the diffuse interstellar matter. The grain size distribution for the former is described in Wood et al. (2002), and that for the latter is similar to Kim et al. (1994) (hereafter KMH; Whitney et al. 2003a). The former affects the thermal balance in the disk via radiative transfer, and is successful in fitting the SED of the HH 30 disk (Wood et al. 2002; Whitney et al. 2003a, 2003b). The latter grains are responsible for scattered light on the disk surface (and in the envelope) at optical and near-infrared wavelengths. The optical constants for silicate and graphite are taken from Laor & Draine (1993).

The 200,000 SEDs provided by the Robitaille et al. (2006) code include flux and polarization spectra for 250 wavelengths (from 0.01 to 5000 \(\mu\text{m}\)). The authors used 20,000 parameter sets and computed the results for 10 viewing angles from face-on to edge-on at equal intervals in the cosine of the inclination. The SED fitter developed by Robitaille et al. (2007) searches for best-fitting SEDs using the minimum \(\chi^2\) method to fit the observed fluxes at a range of wavelengths. We set a distance to the object of 140 pc, and an acceptable range for the visual extinction \(A_v\) of 2.0–2.4, based on Calvet et al. (2004). See Robitaille et al. (2007) for further details of the fitting process.

To derive the star/disk/envelope parameters for RY Tau we used their SED fitter using the photometric data tabulated in Robitaille et al. (2007) (Table 1). These parameters are used to model the PI images with the Robitaille et al. (2006) radiative transfer code. We used 107 photons for each case. The resultant images for the Stokes parameters \(Q\) and \(U\) were convolved with the PSF of the reference star (Section 2) before obtaining the PI image using \(PI = \sqrt{Q^2 + U^2}\). Then the PI flux is normalized to the stellar \(I\) flux \(I_*\), and scaled to match the pixel size of Subaru/HiCIAO.

4.2. Results

Figure 2 shows the observed and 10 best-fit SEDs. We note that the fluxes at different wavelengths were measured at different epochs (Table 1), and may be highly time variable. In particular, Herbst et al. (1994) reported a variation in the optical fluxes of a factor of \(\sim 8\) from 1961 to 1980 (see also Petrov et al. 1999 for longer monitoring observations). Even so, the fluxes in Figure 2 are well fitted by a single SED except for the 12 \(\mu\text{m}\) flux observed using the Infrared Astronomical
Satellite. A larger excess at this wavelength may be attributed to bright silicate emission (Honda et al. 2006) in the filter coverage (8.5–15 μm).

The parameters for the modeled SEDs are shown in Table 2 along with the χ^2 value and the model ID specified by Robitaille et al. (2006). The constant used to scale the scale height relation (h) is fixed at r = 50 AU from the star. The modeled SEDs are based on four sets of physical parameters for the star/disk/envelope with inclination angles of 57°–76°. The inclination angle for most of the SEDs and the disk mass for all of the models are approximately consistent with those of Isella et al. (2010) obtained using millimeter interferometry (≥30 × 10^3 M☉, respectively).

Despite the similarity in the shape of SEDs, some parameters are significantly different between models. The outer disk radius is ~80 AU for two out of four models, but ~120 and ~400 AU for the remaining. The inner disk radius calculated with regard to the sublimation radius varies by a factor of ~8 between models; the disk mass by a factor of ~4; disk accretion rate by a factor of ~20; and the envelope mass by a factor of ~10^4. The gas+ dust density in the outflow cavity is ~2 × 10^{-21} g cm^{-3} in two out of the four models, but zero for the remaining.

Table 1

<table>
<thead>
<tr>
<th>Wavelength (μm)</th>
<th>Flux (mJy)</th>
<th>Year of Observations</th>
</tr>
</thead>
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<tr>
<td>0.36</td>
<td>43 + 19</td>
<td>1962–1990</td>
</tr>
<tr>
<td>0.44</td>
<td>(1.5 ± 0.6) × 10^2</td>
<td>1962–1990</td>
</tr>
<tr>
<td>0.55</td>
<td>(3.5 ± 1.5) × 10^2</td>
<td>1962–1990</td>
</tr>
<tr>
<td>0.64</td>
<td>(5.1 ± 2.2) × 10^2</td>
<td>1962–1990</td>
</tr>
<tr>
<td>0.71</td>
<td>(8.9 ± 3.9) × 10^2</td>
<td>1962–1990</td>
</tr>
<tr>
<td>1.25</td>
<td>(2.26 ± 0.07) × 10^3</td>
<td>1997–2001</td>
</tr>
<tr>
<td>1.65</td>
<td>(3.8 ± 0.2) × 10^3</td>
<td>1997–2001</td>
</tr>
<tr>
<td>2.2</td>
<td>(4.7 ± 0.1) × 10^3</td>
<td>1997–2001</td>
</tr>
<tr>
<td>3.6</td>
<td>(5.75 ± 0.07) × 10^3</td>
<td>2003–2007</td>
</tr>
<tr>
<td>4.5</td>
<td>(5.30 ± 0.06) × 10^3</td>
<td>2003–2007</td>
</tr>
<tr>
<td>5.8</td>
<td>(4.4 ± 0.2) × 10^3</td>
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</tr>
<tr>
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</tr>
<tr>
<td>12</td>
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<td>1983</td>
</tr>
<tr>
<td>24</td>
<td>(1.8 ± 0.4) × 10^4</td>
<td>2003–2007</td>
</tr>
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<td>25</td>
<td>(2.648 ± 0.005) × 10^4</td>
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<td>60</td>
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<td>1983</td>
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<td>450</td>
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</tr>
<tr>
<td>800</td>
<td>(5.6 ± 0.3) × 10^5</td>
<td>1990</td>
</tr>
</tbody>
</table>

Notes.

a Cited from Robitaille et al. (2007). The L-band (3.5 μm), M-band (4.8 μm), and N-band (10.5 μm) data quoted from Kenyon & Hartmann (1995) are not included, as these have large uncertainties, and are consistent with other observations at similar wavelengths.

b (1) Herbst et al. (1994), and converted to the tabulated wavelengths by Robitaille et al. (2007); (2) 2MASS all-sky survey; (3) Speitler Space Telescope Archive; (4) Weaver & Jones (1992); (5) Andrews & Williams (2005), who quoted the values from Mannings & Emerson (1994).

c Measurements at multiple epochs. The uncertainty is based on the photometric variability.

d The specific year is not clear. We describe the years of operation for the 2MASS all-sky survey.

e The specific year is not clear. We describe the period between the launch of the Spitter Space Telescope and when Robitaille et al. (2007) was published.

Figure 3 shows the simulated PI images for the four physical parameter sets, with different inclination angles. Despite significant differences in the parameters described above, the four parameter sets result in PI distributions that are strikingly similar to each other, except for the fact that model 3007615 produces a slightly fainter distribution than the others, due to the small flaring of the disk (h_{50AU} = 1.6 AU, β = 1.12). This similarity between models, both for the SEDs and PI distribution, can be attributed to the fact that both the infrared radiation and the scattered light at the disk surface are determined by the surface geometry of the disk. The PI distribution shown in Figure 3 is dominated by scattering on the disk surface, and the contribution from the envelope and outflow cavity is significantly smaller. To investigate this, we removed the envelope from the models with a relatively massive envelope and diffuse dust layer in the cavity (models 3000949 and 3012376 according to the SED fitter), and ran the simulations again. We find that the PI distribution is almost identical to the model with the envelope and cavity, in particular within ~100 AU of the star.

The models at inclination angles i = 57° and 76° show two separated bright regions similar to the observations. Furthermore, the modeled PI flux normalized to the stellar I flux (PI/I_*) is of the same order as the HiCIAO observations. However, the offset from the major axis of the disk is significantly smaller for all of the modeled PI images compared to the observations. These offsets for the models and observations are clearly shown in the bottom right plots of Figure 3. While that of the observed image is clearly offset from the major axis, the modeled images are more symmetric about the major axis of the disk. This indicates that the disks inferred from the SEDs are geometrically thinner than that producing the observations, if the shape of the emission is attributed to the disk.

5. MONOCROMATIC SIMULATIONS WITH CONVENTIONAL DISK MODELS

In order to explore the effect of geometry, we relax the density distribution of the disk constrained by the SED fitting and focus on reproducing the PI distribution. To do this we carry out monochromatic radiative transfer calculations using Monte Carlo code developed by us specifically for use with SEEDS observations (the Sprout code). We perform simulations over a large region of parameter space for the disk geometry, and also investigate the effects of different grain size distributions. We describe the details of the simulations in Section 5.1, and the results in Section 5.2.

In this section, we focus on comparisons between the models and observations for the bright part of the PI distribution. Although the modeled PI distribution is associated with faint emission at the other side of the disk, we will not discuss this component. For this component one could easily reproduce a PI flux consistent with observations (i.e., the upper limit), by adding extinction by a flattened envelope not included in the models in this section.

5.1. Models and Simulations

We follow the method described in Fischer et al. (1994). We place a central unpolarized light source equivalent to the star as the starting point for calculating the scattering of photons from the disk. The light path for the next scattering position

27 The “thickness” here does not imply the scale height in Equation (1) but that between the midplane and disk surface determined with the Sprout code developed by us specifically for use with SEEDS observations (the Sprout code). We perform simulations over a large region of parameter space for the disk geometry, and also investigate the effects of different grain size distributions. We describe the details of the simulations in Section 5.1, and the results in Section 5.2.

In this section, we focus on comparisons between the models and observations for the bright part of the PI distribution. Although the modeled PI distribution is associated with faint emission at the other side of the disk, we will not discuss this component. For this component one could easily reproduce a PI flux consistent with observations (i.e., the upper limit), by adding extinction by a flattened envelope not included in the models in this section.

5.1. Models and Simulations

We follow the method described in Fischer et al. (1994). We place a central unpolarized light source equivalent to the star as the starting point for calculating the scattering of photons from the disk. The light path for the next scattering position
Figure 3. Modeled PI distributions for the best-fit SED models using the full radiative transfer code developed by Whitney et al. (2003a, 2003b). The observed PI distribution is also shown at the bottom left and in contour in the modeled images. The color contrast is the same for all the images. The contour levels are 0.75, 1.5, 3, and $6 \times 10^{-7}$ pixel$^{-1}$ scale of HiCIAO relative to the integrated stellar $I$ flux, and they are rotated by 72° from the observations. The dashed lines in the PI images show the positions where we extract the 1D distribution and show in the bottom right. Each profile is made by averaging those at two sides, normalized to the peak PI flux. The observed 1D profile is dashed for the faint side of the disk where it is contaminated by an artifact (Section 3).

(A color version of this figure is available in the online journal.)

Table 2

Best-fit Models Using the SED Fitter

<table>
<thead>
<tr>
<th>Model ID</th>
<th>Inclination</th>
<th>$\chi^2$</th>
<th>Stellar Parameters</th>
<th>Disk Parameters</th>
<th>Others</th>
</tr>
</thead>
<tbody>
<tr>
<td></td>
<td></td>
<td></td>
<td>$R_\star$ ($R_\odot$)</td>
<td>$T_\star$ (K)</td>
<td>$R_{\text{min}}$ ($R_{\text{sub}}$) $R_{\text{max}}$ (AU) $h_{\text{50AU}}$ $\beta$ $M_{\text{disk}}$ ($M_\odot$) $M_{\text{disk}}$ ($M_\odot$ yr$^{-1}$)</td>
</tr>
<tr>
<td>3000949</td>
<td>75.5</td>
<td>1.3 $\times$ 10$^2$</td>
<td>4.9 4.7 $\times$ 10$^3$</td>
<td>1.0 83 2.4</td>
<td>1.13 1.1 $\times$ 10$^{-4}$ 1.2 $\times$ 10$^{-6}$</td>
</tr>
<tr>
<td></td>
<td>69.5</td>
<td>1.6 $\times$ 10$^2$</td>
<td>6.2 2.2 $\times$ 10$^2$</td>
<td>5.3 5.1 $\times$ 10$^3$</td>
<td>7.8 84 2.7 1.16 8.0 $\times$ 10$^{-7}$</td>
</tr>
<tr>
<td></td>
<td>63.2</td>
<td>2.2 $\times$ 10$^2$</td>
<td>63.3 1.8 $\times$ 10$^2$</td>
<td>5.5 5.1 $\times$ 10$^3$</td>
<td>7.8 84 2.7 1.16 8.0 $\times$ 10$^{-2}$</td>
</tr>
<tr>
<td>3012376</td>
<td>69.5</td>
<td>1.6 $\times$ 10$^2$</td>
<td>63.3 1.8 $\times$ 10$^2$</td>
<td>5.5 5.1 $\times$ 10$^3$</td>
<td>7.8 84 2.7 1.16 8.0 $\times$ 10$^{-7}$</td>
</tr>
<tr>
<td></td>
<td>75.5</td>
<td>2.0 $\times$ 10$^2$</td>
<td>56.6 2.3 $\times$ 10$^2$</td>
<td>5.5 5.1 $\times$ 10$^3$</td>
<td>4.6 6.4 $\times$ 10$^3$ 7.0 4.0 $\times$ 10$^2$ 1.6 1.12 4.6 $\times$ 10$^{-2}$ 1.1 $\times$ 10$^{-7}$</td>
</tr>
<tr>
<td>3007615</td>
<td>69.5</td>
<td>2.1 $\times$ 10$^2$</td>
<td>75.5 2.3 $\times$ 10$^2$</td>
<td>5.5 5.1 $\times$ 10$^3$</td>
<td>4.6 6.4 $\times$ 10$^3$ 7.0 4.0 $\times$ 10$^2$ 1.6 1.12 4.6 $\times$ 10$^{-2}$ 1.1 $\times$ 10$^{-7}$</td>
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<tr>
<td>3012808</td>
<td>56.6</td>
<td>2.3 $\times$ 10$^2$</td>
<td>5.0 5.1 $\times$ 10$^3$</td>
<td>4.7 1.2 $\times$ 10$^2$ 2.1 1.17 2.8 $\times$ 10$^{-2}$ 2.4 $\times$ 10$^{-7}$</td>
<td></td>
</tr>
</tbody>
</table>

Notes. $R_\star$, stellar radius; $T_\star$, stellar temperature; $R_{\text{min}}$, minimum disk radius; $R_{\text{max}}$, maximum disk radius; $M_{\text{disk}}$, disk mass; $M_{\text{disk}}$, disk mass accretion rate; $M_{\text{envelope}}$, envelope mass; $\rho_{\text{cavity}}$, density in the outflow cavity; $R_{\text{sub}}$, sublimation radius. The sublimation radius is empirically determined by $R_{\text{sub}} = R_\star(T_{\text{sub}}/T_\star)^{-2.1}$ (Robitaille et al. 2006). The inclination is measured from the polar axis. See the text for definition of $\chi^2$. The masses and mass accretion rates are described for gas+dust assuming the gas-to-dust mass ratio of 100.

a Any temperature listed here does not agree the measured spectral type of RY Tau of F8-G2 (Petrov et al. 1999; Calvet et al. 2004; Mendigutía et al. 2011), corresponding to an effective temperature of 5.8 $\pm$ 0.4 $\times$ 10$^3$ K. The different stellar temperatures listed here do not affect the modeled SEDs and PI image at 1.65 $\mu$m (Section 4).
Wood et al. (2002), and Robitaille et al. (2006). We use the grain in Equation (1) and the dust opacity described below. The is calculated for an opacity distribution based on the disk surface of RY Tau. Honda et al. (2006) conducted mid-infrared SEDs of young stellar objects and evolved stars suggest circumstellar dust (Jäger et al. 1998 and references therein). We still use graphite for the KHM distribution for consistency with the authors as their size distribution is determined assuming graphite for the carbon dust. We use amorphous carbon for C01 and C01 × 15 for the above reason and following Cotera et al. (2001) and Wood et al. (2002). The use of different carbon dusts does not significantly affect the modeled PI distribution and it does not affect the conclusions of the paper (Appendix A).

Calculations for Mie scattering are made using the code developed by Wiscombe (1996). The optical constants for astronomical silicate and graphite are obtained from Draine & Lee (1984): they are identical in the near-infrared with the Laor & Draine (1993) values used for the simulations in Section 4. For amorphous carbon we use the optical constants of Jäger et al. (1998) with a pyrolysis temperature of 600°C. As for the use of different carbon materials (amorphous carbon or graphite), amorphous carbon with different pyrolysis temperatures does not significantly affect the modeled PI distribution and conclusions of the paper (Appendix A).

Table 3 shows the physical properties (opacity, albedo, forward throwing parameter, and the maximum degree of polarization) for the above three dust models. Figure 3 shows the scattering properties of these dust models at 1.65 μm, i.e., I, PI, and the degree of polarization as a function of scattering angle, in the case where the incident light is not polarized. The I and PI fluxes for each scattering angle are derived by normalizing the scattering matrix elements S11 and −S12 by a constant I0 to match the results to the “weighted photon method” described above (i.e., \( \int I/\hat{I}d\Omega = \text{albedo} \)). The forward scattering is more

<table>
<thead>
<tr>
<th>Size Distribution</th>
<th>Mass Fraction of Carbon Dust</th>
<th>( \kappa_{\alpha} ) ( (\text{cm}^2 \text{g}^{-1}) )</th>
<th>( \kappa_{\alpha}/\kappa_{\alpha(0.55 \mu\text{m})} )</th>
<th>Albedo</th>
<th>g</th>
<th>( P_{\text{max}} )</th>
</tr>
</thead>
<tbody>
<tr>
<td>KMH (interstellar medium)</td>
<td>0.38</td>
<td>5.3 \times 10^4</td>
<td>0.16</td>
<td>0.39</td>
<td>0.28</td>
<td>0.59</td>
</tr>
<tr>
<td>C01 (the HH 30 disk)</td>
<td>0.49</td>
<td>1.1 \times 10^4</td>
<td>0.28</td>
<td>0.45</td>
<td>0.54</td>
<td>0.42</td>
</tr>
<tr>
<td>C01 × 15</td>
<td>0.49</td>
<td>5.4 \times 10^3</td>
<td>0.50</td>
<td>0.47</td>
<td>0.66</td>
<td>0.37</td>
</tr>
</tbody>
</table>

Notes.

a See the text for details. The minimum and maximum particle radii are 3 \( \times 10^{-3} \) and 2 \( \mu\text{m} \) for KMH; 1 \( \times 10^{-3} \) and 20 \( \mu\text{m} \) for C01; and 15 times larger for C01 × 15.

b We adopt the mass density 3.3 and 2.26 g cm\(^{-3}\) for silicate and graphite, respectively (Kim et al. 1994), and 1.67 g cm\(^{-3}\) for amorphous carbon (Jäger et al. 1998). This parameter is often defined per gas+dust mass (e.g., Cotera et al. 2001; Wood et al. 2002; Whitney et al. 2003a, 2003b; Dong et al. 2012b), but in this paper we define this per dust mass to discuss the total dust mass in the scattering layer in later sections.
significant for large grain models, but different dust models show a relatively similar distribution of scattered intensity \(I/I_0\) at scattering angles of \(40^\circ-180^\circ\). The polarization shows a similar dependency on scattering angle, peaking at \(\sim 90^\circ\), but the absolute value for the C01 and C01 \(\times 15\) models are \(30\%-40\%\) lower than for the KMH. The PI flux shows a maximum at scattering angles of \(\theta = 75^\circ, 60^\circ,\) and \(28^\circ\) for the KMH, C01, and C01 \(\times 15\) models, respectively, and decreases toward both sides for all these dust models.

We assume an outer radius for the disk of 80 AU, based on our HiCIAO observations and millimeter interferometry by Isella et al. (2010). The inner radius is set to 1 AU. We find that our results are not significantly affected by the choice of inner radius as long as it is within the coronagraphic mask (\(r \sim 30\) AU at a distance to the target of 140 pc) and the disk surface has a flared geometry. Note that the millimeter interferometry of Isella et al. (2010) suggests the presence of a disk hole with a radius of \(\sim 15\) AU, while the presence of near-infrared emission suggests the presence of warm dust within 1 AU (Akeson et al. 2005; Pott et al. 2010).

The remaining free parameters for the density distribution of the disk are \(\rho_0\), \(h_{50\text{AU}}\), \(\beta\), and \(\alpha\). The first three parameters all have a large effect on the offset of the PI distribution from the major axis of the disk. This makes it difficult to search for the parameter sets which best fit the results. To overcome this problem, we set the total optical thickness between the star and the edge of the disk to 1 at a direction \(\theta\) from the midplane of the disk. This angle will be adjusted to fit the degree of the offset of the PI distribution from the major axis of the disk (see Figure 6 for an example). This determines the parameter \(\rho_0\), and as a result, our parameter searches will be made with the remaining two free parameters, i.e., \(h_{50\text{AU}}\) and \(\beta\). For the simulation in Section 5.2 we adopt \(\alpha = \beta + 1\) as we did in Section 4. The use of an independent \(\alpha\) produces almost identical results.

In addition to \(\rho_0\), \(h_{50\text{AU}}\), \(\beta\), and \(\alpha\), Equation (1) also includes the stellar radius \(R_*\), but its uncertainty does not affect the results of simulations. The stellar radius is significantly smaller than the radii for the density distribution we use (i.e., \(>1\) AU), and therefore the term \(\sqrt{R_*/r}\) produces only a very minor contribution to Equation (1). The equation also includes the term \((R_*/r)^\alpha\), but the stellar radius here can be replaced by any radius, depending on the radius at which we define \(\rho_0\). Again, this parameter is not explicitly provided but scaled using the constraint of the optical thickness described above.

Throughout the simulations we assume that re-radiation from the disk is negligible, or that it occurs in very close proximity to the star so that the effect of the extended structure in the emission region far from the star is negligible. To prove the validity of this assumption, we ran this simulation code with the disk parameters in Table 2, and found that the code reproduces the PI distribution in Figure 3 well.

### 5.2. Results

Figure 7 shows examples for the PI distribution using the C01 grain size distribution (i.e., that of the HH 30 disk). The combination of the angle for \(\tau = 1\) (\(\theta = 30^\circ\)) and the viewing angle (\(49^\circ\) from the face-on view) are selected to approximately
fit the observed morphology. As for the observations and models in Sections 3 and 4, the PI flux displayed is normalized to the stellar $I$ flux. The image convolution is made in the same manner as Section 4. The contrast for each figure is adjusted in each image for the best morphological comparison for the bright part.

In Figure 7, all except the case of $h_{50\text{AU}} = 25\text{ AU}$, $\beta = 1.3$ show a butterfly-like morphology similar to the observed PI distribution. The peak PI flux in these models, however, is three to six times larger than the observed value. Increasing $\beta$ produces a brighter PI distribution. In the case of $\beta = 2.0$ and 2.7, the modeled $PI/I_*$ is larger for $h_{50\text{AU}} = 25\text{ AU}$ than $h_{50\text{AU}} = 5\text{ AU}$ by a factor of $\sim 1.5$. This difference is attributed to the fact that, for the latter model, the star suffers extinction from the edge of the disk which decreases $I_*$. In the case of $\beta = 2.0$ and 2.7, increasing $h_{50\text{AU}}$ also produces a larger PI flux at the other side of the disk. The PI distribution is fairly centrosymmetric for $h_{50\text{AU}} = 25\text{ AU}$ and $\beta = 1.3$ for any contrast.

In the same figure we also show the density profiles for the individual disks. The dashed, solid, and dotted curves show the positions at which $\tau = 0.5$, 1, and 2, respectively, measured from the star. These curves indicate that the scattering layer is geometrically thin at the disk surface. Increasing $\beta$ produces a large flaring angle in the outer disk, which explains the brighter PI flux as the surface is more easily illuminated. The curvature of the surface remains similar for the same $\beta$; however, the disk is not flared for $\beta = 1.3$ and $h_{50\text{AU}}$. In this case, the scattering is dominated by the inner part of the disk, producing a centrosymmetric PI distribution as described above.

The viewing angle selected for Figure 7 (49°) is smaller than that determined from millimeter emission by Isella et al. (2010) (>65°). A larger viewing angle does not reproduce the observed PI morphology. Figure 8 shows examples for a viewing angle of 63°. The disk parameters are identical to those for Figure 7, but the optical thickness along the radial direction has been set to $\tau = 1$ at $\theta = 25^\circ$ to fit the offset of the bright PI emission from the major disk axis. The results are similar to those of Figure 7, but the bright part shows a thinner distribution in the vertical direction than the observations. Furthermore, a large $\beta$ results in a larger spatial extension at the bright side. To clearly show the former discrepancy, we extract the one-dimensional (1D) profiles at the positions indicated in Figures 7 and 8 and show them in Figure 9. The peak PI fluxes are comparable to or larger than those in Figure 7, and therefore significantly larger than the observations.

Figure 10 shows the PI images for the same disk parameters as Figures 7 and 8 but with different dust models. Changing the dust model results in little variation in the PI images. This can be explained by the fact that the PI flux at the bright side is dominated by a single scattering with scattering angles $60^\circ$–$135^\circ$, and the PI flux for this range of angles is similar for the three different dust models, as shown in Figure 5. In particular, the PI fluxes shown in Figure 5 for these scattering angles explain the fact that the PI flux for the KMH model is larger than the others in Figure 10 by a factor of $\sim 2$. Table 4 shows the peak $PI/I_*$ for individual models. These are larger than the observations by a factor of 2–12 at the peak.

In summary, the modeled morphology in the PI distribution does not match the observations with the viewing angle inferred from the millimeter observations (>65°). As shown in Figures 7, 8, and 10, the modeled PI distribution from the bright side of the disk does not vary significantly with the value of $h_{50\text{AU}}$, $\beta$, or the choice of dust model once we determine the outer radius of the disk and the optical thickness in a given direction. Some models may also be excluded due to the degree of polarization, the inferred disk mass or extinction toward the disk.
star. However, these do not provide constraints as clearly as those from the PI flux. We discuss these constraints in Appendix B.

6. DISCUSSION

In Section 6.1, we discuss the possible presence of an optically thin scattering layer above the disk which could be responsible for the observed PI distribution. In Section 6.2, we briefly discuss whether the use of more realistic dust models might still fit the observations using a conventional disk model. In Section 6.3, we discuss the implications of the asymmetry of the PI distribution about the jet/disk axis that was shown in Section 3.

6.1. Geometry of the Scattering Layer

In Sections 4 and 5, we found inconsistencies between the models obtained with conventional disk and dust models and the observed PI distribution. The disk geometry obtained from the SED corresponds to a PI distribution with a significantly smaller offset from the major axis of the disk (Section 4). On the other hand, the disk geometries which reproduce this offset require a smaller viewing angle (i.e., close to face-on) than that determined through millimeter observations (Section 5). We have not explored the full parameter space, but once we fix the parameters to reproduce the angular scale along the major axis and the offset of the bright PI emission on the far side of the disk, varying disk and dust parameters has little effect on the resulting PI distribution.

The results of the two sets of simulations (full radiative transfer and monochromatic scattering) may imply that the system consists of (1) a geometrically thin disk which is partially responsible for the infrared SED but does not contribute to the PI flux in the near-infrared and (2) an optically thin and geometrically thick upper layer which is responsible for the
Figure 10. PI distributions with different dust models for $\beta = 2.0$. The other parameters ($\theta$, viewing angle, and $h_{50\text{AU}}$) and contours are the same as Figures 7 and 8. (A color version of this figure is available in the online journal.)

Figure 11. Schematic view of the disk, optically thin layer, stellar radiation, scattered light, and infrared re-radiation observed. See the text for details.

observed PI distribution in the near-infrared and the remaining mid-to-far infrared flux. A similar geometry has been proposed for SEDs observed for some Herbig AeBe stars (Meeus et al. 2001, Group I). Follette et al. (2013) also discuss such a geometry for the SR 21 disk in which the presence of an inner cavity is indicated by submillimeter observations but not seen in the scattered light in near-infrared.

Figure 11 shows a schematic view for the disk, an optically thin layer, stellar radiation, scattered light, and infrared re-radiation. A rough analytic estimate for the optical thickness, density, and mass of such a layer is described in Appendix C. The observed PI image does not show evidence for emission along the major axis without an offset as modeled in Section 4. This can be explained if the thickness of the disk is significantly smaller than those used in Section 4, which provide a PI flux level comparable to the observations.

Such a scattering geometry can easily explain the observed PI flux and offset from the major axis of the disk, by adjusting the column density and vertical distribution of the optically thin layer, respectively. Here we qualitatively demonstrate this using the Sprout code with an optically thin upper layer using Equation (1) and an optically thick disk with a geometrical thickness of zero. To focus on the scattered light outside the coronagraphic mask, we set the density to be zero for disk radii within the mask ($r < 28\text{ AU}$).

Figure 12 shows an example of simulations using the KMH dust model, with a viewing angle of 70°. The butterfly-like bright PI distribution approximately matches the observations. We measure the modeled polarization degree to be 25% at a position corresponding to the PI peak of the observations. This is also consistent with the observations of $\gg 9\%$ measured in Section 3. The faint extended emission at the other side of the disk is brighter than the upper limit of the observations by a factor of $\sim 4$. The modeled flux in this region could be adjusted, e.g., if we included extinction by a remnant envelope, or if we define an exponential cutoff for the outer radii of the disk.
particles in practice appear to be aggregated (e.g., Henning & Stognienko 1996; Dominik 2009), and may also be coated with ice (Malfait et al. 1999; Meeus et al. 2001; Honda et al. 2009). These facts would produce different optical properties. Min et al. (2012) conducted simulations for scattered light from the disk using both spherical and aggregated particles, and show that a lower PI flux is expected for the latter due to the extremely large forward scattering. This would allow us to better match the PI flux we simulated in Section 5 for the disk surface, as our values were significantly larger than the observations for many cases.

However, the disks used in Section 5 would produce significantly different SEDs than those shown in Figure 2. A large flaring disk would produce more stellar photons than the thinner disks we used in Section 4. Furthermore, some disks used in Section 5 have a large extinction toward the star (Appendix B). This would result in a double peaked SED in the optical to near-infrared, and far-infrared, respectively (e.g., D’Alessio et al. 1999; Robitaille et al. 2007). The former could be overcome if the dust particles had a smaller emissivity in mid-to-far infrared, or large albedo at UV-to-IR in order to impede the absorption of the stellar photons. To investigate further, one would need a detailed understanding of the optical properties of the aggregates over a wide range of wavelengths, plus simulations with full radiative transfer.

Alternatively, observations of the PI distribution at different wavelengths would give useful constraints for the nature of the dust grains. We expect that the PI flux is dominated by the photons with a single scattering, even in the case of optically thick disks (such as those used in Sections 4 and 5) in which multiple scattering occurs as well as single scattering. The contribution from multiple scattering would be negligible for the reasons below. In the case of small grains (e.g., KMH), where photons are fairly isotropically scattered, the scattered photons have polarizations with a variety of P.A., canceling each other out. In the case of large grains (e.g., C01 and C01 × 15), in which most of the photons are scattered forward, the polarization of individual photons is significantly reduced after the first scattering.

To demonstrate this, we re-ran the simulations for the disks in Section 5, removing all the photons which experience multiple scatterings. We find that the results are almost identical for the PI flux distribution, with a flux difference within 10% at the peak. This is in contrast with the I flux for which multiple scatterings are responsible for 20%–40% of the flux. Throughout, studies of the PI flux allow a more accurate investigation of the dust properties than studies of the I flux, in particular for optically thick disks.

6.3. Implication of the Asymmetric PI Distribution

As shown in Section 3, the observed PI distribution is asymmetric about the rotation axis of the disk: the southwest side is brighter than the northern side by a factor of ∼2. A similar asymmetry in the distribution of the scattered light has been extensively studied for the HH 30 disk, i.e., a low-mass protostar at a similar evolutionary stage, using the Hubble Space Telescope (e.g., Burrows et al. 1996; Stapelfeldt et al. 1999; Cotera et al. 2001; Watson & Stapelfeldt 2007). These authors attributed such a distribution to uneven illumination by the central object. These observations also show that the distribution of the scattered light is time variable. Watson & Stapelfeldt (2007) summarize possible mechanisms to explain the asymmetric distribution and its time variation. These include

6.2. Dust Grains

While homogeneous spherical particles are widely used for the dust in disks (Whitney & Hartmann 1992; Fischer et al. 1994; Whitney et al. 2003a, 2003b; Robitaille et al. 2006, 2007), dust optically thin layers like that used to reproduce the millimeter flux distribution of the disk (e.g., Andrews et al. 2009; Isella et al. 2010).

The use of different dust models (C01, C01 × 15) provides similar results to Figure 12, but the polarization degree is significantly smaller: 13%–14% at the position corresponding to the PI peak of the observations. Such a polarization may be significantly smaller: 13%–14% at the position corresponding to the PI peak of the observations. Such a polarization may be significantly smaller: 13%–14% at the position corresponding to the PI peak of the observations. Such a polarization may be significantly smaller: 13%–14% at the position corresponding to the PI peak of the observations. Such a polarization may be significantly smaller: 13%–14% at the position corresponding to the PI peak of the observations. Such a polarization may be significantly smaller: 13%–14% at the position corresponding to the PI peak of the observations. Such a polarization may be significantly smaller: 13%–14% at the position corresponding to the PI peak of the observations. Such a polarization may be significantly smaller: 13%–14% at the position corresponding to the PI peak of the observations.
(1) hot (or cool) spots on the star, (2) shadowing by a non-axisymmetric inner disk, (3) obscuration by a companion star, and (4) obscuration by a disk associated with a companion star. The time variability in the scattered light is also observed in the more evolved disk system HD 163296 (Wisniewski et al. 2008). The authors attributed this to the asymmetric shadowing of the inner disk.

Of the explanations described above, shadowing by the inner disk may be the most likely for the asymmetric PI distribution in RY Tau. While uneven illumination by hot (or cool) spots is one of the favored explanations for the HH 30 disk (Wood & Whitney 1998; Stapelfeldt et al. 1999; Cotera et al. 2001; Watson & Stapelfeldt 2007), it is not likely for RY Tau: the optical continuum associated with this star does not show clear evidence for time variability due to such spots (e.g., Petrov et al. 1999; Chou et al. 2013). Obscuration by a secondary star/disk may also be excluded, since there is no clear evidence for the presence of a close companion associated with RY Tau. Near-infrared interferometric observations have ruled out the presence of a companion at 0.35–4 AU from the star and down to a stellar flux ratio of 0.05 (Pott et al. 2010).

The shadowed-disk scenario is corroborated by the UXOR-type variability observed toward this star, i.e., time-variable obscuration by circumstellar dust (e.g., Herbst et al. 1994; Petrov et al. 1999). While many authors attribute this obscuration to the disk, its nature is not clear. It could be clumpy accretion of dust+gas onto the star (see Herbst et al. 1994; Grady et al. 2000, and references therein), the puffed-up inner rim of the disk (Dullemond et al. 2003), or the outer edge of the disk (Thé 1994). Herbst et al. (1994) also state the possibility that such a dust screen is associated with a wind, not a disk or disk accretion. Of the above explanations, the outer edge of the disk may be excluded if the disk associated with RY Tau is geometrically thin, and we are not observing this target close to an edge-on view (Sections 4 and 6.1).

RY Tau showed two abrupt brightening events at optical wavelengths in 1983/1984 and 1996/1997 (e.g., Herbst et al. 1994; Petrov et al. 1999). If we attribute this to a single orbital period around a 2 $M_\odot$ star, the corresponding radius of the disk is 7 AU. This radius is comparable to the radius where Isella et al. (2010) identified emission peaks at two sides of the disk using millimeter interferometry. If obscuration occurs in the inner disk, it would require time variation of the disk structures including warping and precession. This may be possible via tidal interaction with a very low mass companion such as a protoplanet (Hughes et al. 2009).

Synoptic observations of the scattered light of RY Tau would allow us to identify the origin of the asymmetric PI distribution, determine a typical radius where the obscuration occurs, and constrain the disk geometry (and its time variation) within the coronagraphic mask. Bastien (1982) made monitoring observations of optical polarization integrated over the object, revealing variation of the position angle of polarization between $-25^\circ$ and $45^\circ$ over a few years. Sparingly resolved observations like ours would have significant advantages for searching for a periodic variability toward this active and complicated object. We also note that the coronographic observations for the optical $I$ image of RY Tau were made in 2007–2009, and these show a brighter lobe in the north than the southwest (J. McCleary et al. 2007, private communication) in the outer envelope, i.e., the opposite side as we observed in the PI distribution. Although these observations show scattered light in significantly more distant outer regions, this is consistent with the idea of time variation of scattering light associated with this object.

7. CONCLUSIONS

We present near-infrared coronagraphic imaging polarimetry of RY Tau. The scattered light in the circumstellar environment was imaged at $H$ band with a high resolution (≈0″05) using Subaru/HiCIAO. The observed PI distribution shows an angular scale of ≈1″0 (~140 AU) and ≈0″6 (~80 AU) along the major and minor axes of the disk, respectively, exhibiting a butterfly-like shape. The angular scale along the major axis is similar to that of the disk measured using millimeter interferometry. The bright part of the emission is offset from the star toward the direction of the blueshifted jet. Such a distribution can be explained if the object is associated with a geometrically thick disk or a remnant envelope. This agrees with the premise that the system is at the earliest stage of the Class II evolutionary phase.

We perform comparisons between the observed PI distribution and simulations of scattered light with conventional disk and dust models. The simulated images are made using two different approaches. The first approach is to conduct full radiative transfer simulations including absorption, scattering, and re-radiation by dust grains, with disk parameters based on SED fitting (Robitaille et al. 2006, 2007). The second is monochromatic simulations with absorption and scattering only for a significantly larger set of disk parameters to attempt to better fit the observed PI distribution.

The first approach reproduces the PI flux level normalized to the integrated stellar $I$ flux well, but fails to reproduce the offset of the bright PI distribution from the star along the disk axis. The second approach reproduces the butterfly-like morphology in PI distribution well, with a total optical thickness of 1 at 30′ from the midplane of the disk and the viewing angle of ≈50°. However, the model PI distribution does not match the observations with a large viewing angle inferred by millimeter interferometry (>65°). These results for the second approach are relatively independent of the free parameters for the disks and the size distribution of dust grains.

Throughout, we find disagreements between the observed PI distribution and models using conventional disk and dust models. This may imply that the system consists of the following: (1) a geometrically thin disk which is partially responsible for the infrared SED but not the PI flux in the near-infrared and (2) an optically thin and geometrically thick upper layer which is responsible for the observed PI distribution in near-infrared and the remaining infrared flux. Simulations show that this idea approximately explains the observed PI emission with a viewing angle consistent with the observations. Half of the scattered light and thermal radiation in this layer illuminates the disk surface, and this process may significantly affect the thermal structure of the disk.

The PI brightness has an asymmetry about the jet axis by a factor of ≈2 in flux presumably due to uneven illumination caused by obscuration by the dusty environment. Such obscuration could be either due to the accretion of dust+gas onto the star, the puffed-up inner rim of the disk, or the outer edge of the disk. Synoptic studies of the PI distribution would give useful constraints for the geometry of the disk and perhaps its time variation within 5 AU of the star, and a better understanding...
Table 5 shows the dust properties of the C01 and C01 × 15 size distributions (Section 5.1) with different carbon dusts: i.e., amorphous carbon with pyrolysis temperature 400°C–1000°C and graphite. The optical constants for the amorphous carbon and graphite are taken from Jäger et al. (1998) and Draine & Lee (1984), respectively. Figure 13 shows their scattering properties as in Figure 5, but using the above carbon dusts. The simulated PI images using these dust grains show nearly identical morphologies, with differences in PI flux of −20% to +30% compared with the amorphous carbon used in the main text (pyrolysis temperature = 600°C). As discussed in Section 5.2, these similarities and

of the nature of objects with a similar variability at optical wavelengths (UXORs).

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Facility: Subaru (HiCIAO)

APPENDIX A

SIMULATED RESULTS WITH DIFFERENT CARBON DUST MODELS

Table 5 shows the dust properties of the C01 and C01 × 15 size distributions (Section 5.1) with different carbon dusts: i.e., amorphous carbon with pyrolysis temperature 400°C–1000°C and graphite. The optical constants for the amorphous carbon and graphite are taken from Jäger et al. (1998) and Draine & Lee (1984), respectively. Figure 13 shows their scattering properties as in Figure 5, but using the above carbon dusts. The simulated PI images using these dust grains show nearly identical morphologies, with differences in PI flux of −20% to +30% compared with the amorphous carbon used in the main text (pyrolysis temperature = 600°C). As discussed in Section 5.2, these similarities and
differences are attributed to \((PI/I_0)\) at scattering angles 60°–135° in Figure 13.

APPENDIX B
POLARIZATION, DISK MASSES, AND EXTINCTION TOWARD THE STAR FOR GEOMETRICALLY THICK DISK MODELS

Figure 14 shows polarization vectors for the same modeled parameters as Figure 10. For all the models, the vector pattern is centrosymmetric about the star as it is in the observations (Section 3). We measure a degree of the polarization of 10%–28% at the position corresponding to the PI peak of the observations (see Figure 10 for the position). Some seem inconsistent with the observations (> ∼ 9%, Section 3).

Tables 6 and 7 show the disk masses and extinction \((A_V)\) toward the star, respectively, for the models used in Section 5.2. The disk masses are derived assuming a gas-to-dust mass ratio of 100. The extinctions are calculated based on the optical thickness at 1.65 μm and the \(\kappa_{\text{ext},1.65 \mu m}/\kappa_{\text{ext},0.55 \mu m}\) tabulated in Table 3. Table 6 show that the disk mass is significantly smaller for \(h_{50 \text{AU}} = 25 \text{ AU}\) than \(h_{50 \text{AU}} = 5 \text{ AU}\) with a given combination of the dust model, \(\beta\), and \(\theta_\tau = 1\). As a result, the former disks produce brighter PI flux from the other side of the disk (Figures 7, 8, and 10). Disks with a large scale height also provide a larger extinction toward the star.

All the disk masses derived for \(h_{50 \text{AU}} = 25 \text{ AU}\) are significantly smaller than that estimated by Isella et al. (2010) based on interferometry (> ∼ 3 × 10^{-3} \(M_\odot\)). In contrast, a majority of the disk masses derived for \(h_{50 \text{AU}} = 5 \text{ AU}\) are significantly larger than those of pre-main-sequence stars suggested by millimeter observations (< 0.1 \(M_\odot\); e.g., Robitaille et al. 2007; Isella et al. 2009; Williams & Cieza 2011; see also Table 1). However, we do not constrain the models described in Section 5.2 for the following reasons: (1) the disk mass inferred by millimeter
interferometry highly depends on the dust model used to convert the flux to a dust mass (Isella et al. 2010); (2) the modeled disk mass highly depends on the exponent of the radial density distribution $\alpha$, which is assumed to be $\beta + 1$ for the simulations in Section 5.2; and (3) the scattered light results from a small fraction to the total dust mass (Appendix C); therefore, our observations do not directly probe the total disk mass.

Regarding the extinction to the star, some models provide a larger value than that measured by Calvet et al. (2004) ($A_V = 2.2$). However, we note that the measurement of extinction by Calvet et al. (2004) is based on observations at UV–optical wavelengths, adopting an extinction law for cold molecular clouds. Therefore, it may not be directly comparable with the extinctions of Table 7, in particular for the C01 and C01 $\times 15$ size distributions.

**APPENDIX C**

**OPTICAL THICKNESS, DENSITY, AND MASS OF THE OPTICALLY THIN AND GEOMETRICALLY THICK SCATTERING LAYER**

Suppose the dust corresponding to the mass $m$ is located at distance $r$ from the star. The number of PI photons observed at the telescope per second is described as follows:

$$n_{PI} = \kappa_{ext} m \frac{N_*}{4\pi r^2} \frac{(PI)}{I_0} \Omega, \quad (C1)$$

where $\kappa_{ext}$ is the opacity, $N_*$ is the number of the photons at 1.65 $\mu$m ejected from the star in all directions, and $(PI/I_0)$ is the fraction of the PI flux normalized to the incident flux on dust grains (Figure 4). $\Omega$ is the solid angle corresponding to the telescope area, thus $\Omega = A_{tel}/d^2$, where $A_{tel}$ and $d$ are the area of the telescope and the distance to the target, respectively. If we normalize Equation (C1) by the number of stellar photons observed per second, i.e., $n_* = N_*/(4\pi d^2)$, the equation is as follows:

$$\frac{n_{PI}}{n_*} = \kappa_{ext} m \frac{r^2}{A_{tel}} \frac{(PI)}{I_0}. \quad (C2)$$

Here we substitute $n_{PI}/n_* = 8 \times 10^{-7}$, measured at the peak of the PI flux as described in Section 3: $\kappa_{ext} = 1.1 \times 10^4$ cm$^2$ g$^{-1}$ (C01 dust, Table 3), $r = 40$ AU, and $(PI/I_0) = 5 \times 10^{-3}$ str$^{-1}$ (Figure 5). We derive a dust mass of $5 \times 10^{21}$ g. The PI flux measured at each position is based on the HiCIAO pixel scale ($9.48 \times 10^{-3}$ arcsec), and thus this dust mass corresponds to an optical thickness along the line of sight of $\tau = \kappa_{ext} m/A_{pix} \approx 0.15$, where $A_{pix}$ is the area corresponding to each pixel, assuming a distance $d = 140$ pc. Therefore, it is likely that the scattering layer is optically thin along the line of sight.

Such a layer could also be optically thin in the radial direction. The optical thickness along this direction highly depends on the assumed geometrical thickness of the layer. Let us assume a thickness of 30 AU, comparable to the thickness of the disks we used in Section 5 to explain the offset of the PI distribution from the star. One would expect an optical thickness of $\sim 6 \times 10^{-3}$ toward a HiCIAO pixel (1.3 AU) assuming all of the parameters described above. Integrating this over the disk radius of 80 AU, one would expect a total optical thickness of $\sim 0.4$. This is a simple estimate. In reality, the observed PI flux is lower than the maximum at the other positions, and the dependence of $n_{PI}/n_*$ on the distance from the star should be included. Furthermore, the opacity at individual positions depends on the distance from the star (Equation (C2)), and this fact is not included here.

If we assume the same layer thickness, and use the dust mass contained in the pixel with the maximum PI value, we would derive a dust mass density of $2.9 \times 10^{-20}$ cm$^{-3}$. This corresponds to a hydrogen number density of the order of $\sim 10^6$ cm$^{-3}$ assuming a gas-to-dust mass ratio of 100. This is larger than the envelope density at similar radii (50 AU) inferred by the SED fitting in Section 4 by a factor of $\sim 10$ or more.

We estimate the dust mass of the scattering layer as follows, using Equation (C2):

$$M_{dust} = \int m(r)rdr = \int \frac{nP_{PI}(r)}{n_*} \frac{(PI)}{I_0} \frac{1}{\kappa_{ext}} dr. \quad (C3)$$

where $r$ and $d$ are the position and the distance to the star, respectively. Substituting the projected radius for $r$, we derive a dust mass $M_{dust}$ of $3 \times 7 \times 10^{-3} M_{\odot}$. A larger value would be expected if we include the inclination effect for $r$ and the smaller $PI/I_0$, depending on the scattering angle at individual positions. Even so, this estimate would be sufficient to conclude that the optically thin layer above the disk surface has a mass significantly smaller than the total dust mass in the disk inferred by radio observations (10–50 $M_{\odot}$; Isella et al. 2010).

**REFERENCES**


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