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# First firm spectral classification of an early-B pre-main-sequence star: B275 in M17

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## Abstract

The optical to near-infrared (300 – 2500 nm) spectrum of the candidate massive young stellar object (YSO) B275, embedded in the star-forming region M17, has been observed with X-shooter on the ESO *Very Large Telescope*. The spectrum includes both photospheric absorption lines and emission features (H I and Ca II triplet emission lines, first and second overtone CO bandhead emission), as well as an infrared excess indicating the presence of a (flaring) circumstellar disk. The strongest emission lines are double-peaked with a peak separation ranging between 70 and 105 km s<sup>-1</sup>, and they provide information on the physical structure of the disk. The underlying photospheric spectrum is classified as B6–B7, which is significantly cooler than a previous estimate based on modeling of the spectral energy distribution. This discrepancy is solved by allowing for a larger stellar radius (i.e. a bloated star) and thus positioning the star above

the main sequence. This constitutes the first firm spectral classification of an early-B pre-main-sequence (PMS) star. We discuss the position of B275 in the Hertzsprung-Russell diagram in terms of PMS evolution. Although the position is consistent with PMS tracks of heavily accreting protostars ( $\dot{M}_{\text{acc}} \gtrsim 10^{-5} M_{\odot} \text{ yr}^{-1}$ ), the fact that the photosphere of the object is detectable suggests that the current mass-accretion rate is not very high.

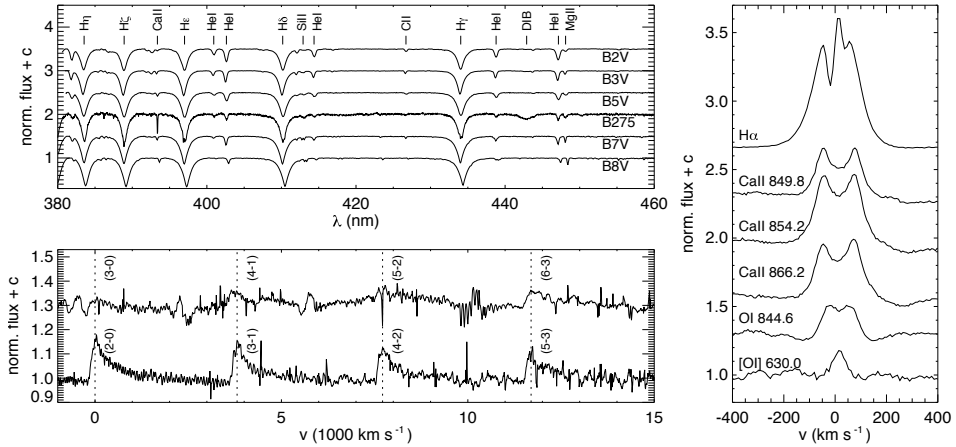
## 6.1 Introduction

Observational and theoretical evidence is accumulating that massive stars form through disk accretion, similar to low-mass stars. Accretion persists despite the strong radiation pressure and ionizing power produced by the massive young stellar object (YSO) that may prevent matter from accreting (e.g., Keto et al. 2006; Krumholz et al. 2009). Given the short main-sequence lifetime of massive stars, the mass accretion rate must be high (up to  $\sim 10^{-3} M_{\odot} \text{ yr}^{-1}$ , Hosokawa et al. 2010) to ensure that the star does not leave the main sequence before the accretion process has finished.

Evidence of accretion must come from the detection of circumstellar disks and bipolar jets, as observed around forming low-mass stars (e.g., Appenzeller & Mundt 1989). Disks and outflows around massive YSO candidates are being reported (Chini et al. 2004; Kraus et al. 2010; Chapter 3) but the physical properties of the forming massive stars remain uncertain. The mass of the central object has to be estimated from the emerging flux, and the direct detection of the photospheric spectrum turns out to be very difficult at this early stage of evolution (e.g., Testi et al. 2010).

Infrared surveys have revealed several hundred candidate massive YSOs, based on luminosity arguments (e.g., Urquhart et al. 2011). A (*K*-band) spectrum has been obtained for only a few of these (Hanson et al. 1997, 2002; Bik et al. 2006), and they show a red continuum, most likely due to hot dust, and an emission line spectrum that includes Br $\gamma$  and, often, CO 2.3  $\mu\text{m}$  bandhead emission. The latter emission can be modeled as being produced by a Keplerian rotating disk surrounding the young, potentially massive star (Bik & Thi 2004; Blum et al. 2004; Wheelwright et al. 2010).

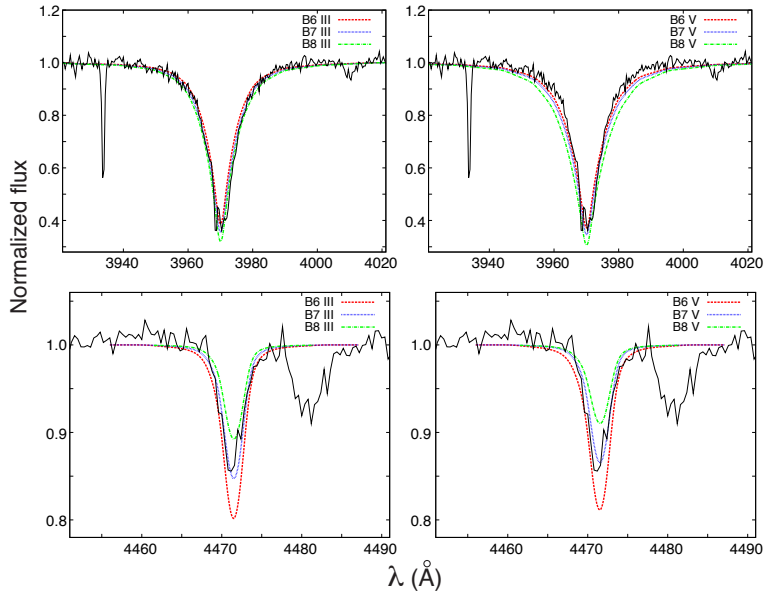
As massive stars show most spectral features in the UV and optical ranges, the study of their photospheric properties would strongly benefit from extending the spectral coverage as far to the blue as possible. Extinction by the surrounding gas and dust makes this an observational challenge. Only in rare cases have spectra of candidate massive YSOs been obtained at optical wavelengths. Hanson et al. (1997) obtained optical and near-infrared spectra of candidate massive YSOs in M17, one of the most massive nearby star-forming regions in the Galaxy (Hoffmeister et al. 2008; Broos et al. 2007; Povich et al. 2009). For the “normal” OB stars Hanson et al. (1997) found a good correspondence between the optical and *K*-band spectra, but the massive YSO optical spectra remained inconclusive. For four massive YSO candidates, they registered the optical spectrum from 400 to 480 nm, indicating a high mass and luminosity. The blue spectrum of the strong CO emission source B275 showed no definite photospheric features other than hydrogen, so that the nature of this source remained uncertain. The spectral energy distribution (SED), though, indicated spectral type late-O or early B, at an adopted distance of 1.3 kpc. We exploit the high efficiency and broad wavelength coverage of the new medium-resolution spectrograph X-shooter on the ESO *Very Large Telescope* (VLT) to (i) detect the photospheric spectrum of B275 in M17, (ii) determine its effective temperature in order to place the candidate massive YSO unambiguously onto recent evolutionary tracks, and (iii) search for ongoing accretion activity and investigate the structure of the disk.



**Figure 6.1:** *Top left:* The blue spectrum of B275 in M17 shown next to B main-sequence-star spectra (Gray & Corbally 2009). *Bottom left:* The 1<sup>st</sup> and 2<sup>nd</sup> overtone CO emission bands. Zero velocity corresponds to the first component in the series (at 2294 and 1558 nm, respectively). *Right:* A sample of the emission line profiles in the spectrum of B275. The Ca II triplet lines and O I 845 nm are superposed on hydrogen Paschen series absorption lines. The flux of the H $\alpha$  line is scaled down by a factor 5; the structure near the peak is a remnant of the nebular-line subtraction.

## 6.2 VLT/X-shooter observations of B275

VLT/X-shooter spectra were obtained of the massive YSO B275 in M17 (CEN 24, RA(2000.0) = 18<sup>h</sup>20<sup>m</sup>25<sup>s</sup>.13, DEC(2000.0) = -16°10'24''56,  $V = 15.55$  mag,  $K = 8.05$  mag, Chini et al. 1980; Skrutskie et al. 2006) on August 11, 2009 at 03h20 UT, during the first science verification run (PI Chini). The observations in the UVB arm (300–600 nm) were binned (2 pixels) in the wavelength direction in order to increase the signal-to-noise ratio of this part of the spectrum, while still oversampling the resolution element. The 1.6" slit was used, inducing a resolving power  $\mathcal{R} = 3300$ . For the VIS (550–1000 nm) and the NIR arm (1000–2500 nm) a 0.9" slit was used ( $\mathcal{R} = 8800$  and 5600, respectively). The total exposure time was 45 minutes, yielding a typical signal-to-noise ratio of 70. For more details on the X-shooter instrument and its performance, see D’Odorico et al. (2006); Vernet et al. (2011). The observing conditions were good (0.6" seeing in  $V$  and 76 % Moon illumination). The spectra were obtained by nodding the star on the slit, allowing for background subtraction. The standard procedures of data reduction were applied using the X-shooter pipeline version 0.9.4 (Goldoni et al. 2006; Modigliani et al. 2010). For flux calibration and telluric absorption correction, the standard stars EG274 and HD180699 were used.



**Figure 6.2:** FASTWIND model profiles of He (*top*) and He I 447.1 nm (*bottom*) lines for B6–B8 giants (*left*) and main-sequence stars (*right*). The B7 III model provides the best fit with the observed profiles.

## 6.3 Results

In the following we present the results for the accurate classification of the photospheric spectrum, analyze the interstellar spectrum to determine the extinction, model the SED using the flux-calibrated X-shooter spectrum, and describe the emission line spectrum produced by the circumstellar disk.

### 6.3.1 Spectral classification

Hydrogen absorption lines were detected by Hanson et al. (1997) in the blue spectrum of B275, but do not allow for an accurate spectral classification. As shown in Fig. 6.1, a number of helium and metal lines can be used to classify the photospheric spectrum. The He I 400.9 nm and C II 426.7 nm lines, prominent down to spectral type B3, are very weak. The He I 447.1 nm / Mg II 448.1 nm ratio is a useful spectral indicator for mid- to late-B stars (Gray & Corbally 2009) as the neutral helium line disappears towards lower temperature (A0) and the magnesium line strengthens. When also considering another line ratio, Si II 412.8 nm / He I 448.1 nm, the spectral type becomes B6 ( $\pm$  one subtype).

The spectral type and luminosity class of B275 are further constrained by comparison of the observed H I and He I line profiles (as well as the shape of the SED, see Sect. 6.3.3), to model profiles produced with FASTWIND (Puls et al. 2005). This code calculates non-LTE line-blanketed stellar atmosphere models and is especially suited

to modeling stars with strong winds, but it can also be used to examine  $T_{\text{eff}}$  and  $\log g$  dependent photospheric lines of H and He. We constructed a grid of models (in varying  $T_{\text{eff}}$  and  $\log g$ ) of B6–B8 dwarf and giant stars. The synthetic H I and He I profiles resulting from the models are convolved with the corresponding instrumental and rotational profiles. We adopt  $v_{\text{r}} \sin i = 100 \text{ km s}^{-1}$ . An acceptable fit is obtained for a B7 V model (Fig. 6.2); however, the best fit is obtained for a B7 III model, with  $T_{\text{eff}} = 13,000 \pm 500 \text{ K}$  and  $\log g = 3.5 \pm 0.3$ . This is the first time that the spectral type of a candidate massive YSO has been accurately determined.

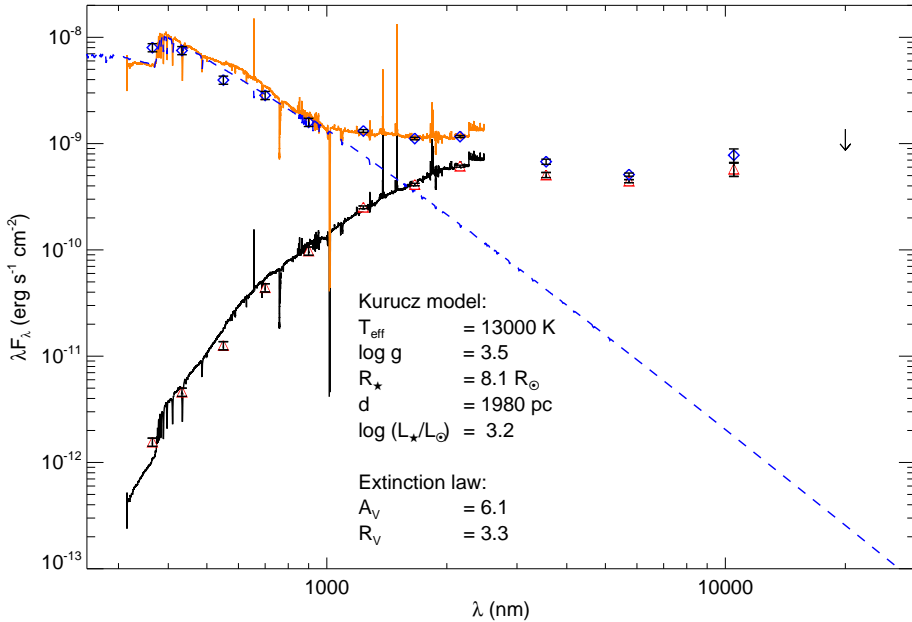
### 6.3.2 Interstellar spectrum

The optical spectrum of B275 includes several interstellar features: atomic resonance transitions (e.g., Ca II H&K, Na I D) and diffuse interstellar bands (DIBs). The DIB strength provides a measure of the interstellar extinction. For the DIBs centered at 578.0, 579.7, and 661.4 nm, we measure an equivalent width of 0.063, 0.014, and 0.021 nm, respectively, with a typical error of 10 %. Using the relations from Cox et al. (2005), we arrive at an E(B-V) of  $1.0 \pm 0.1 \text{ mag}$ . For an average value of  $R_V = 3.1$ , these DIB strengths yield  $A_V \approx 3 \text{ mag}$  of visual extinction. This is less than the determination of  $A_V \approx 6.1 \text{ mag}$  from dereddening the SED (Sect. 6.3.3). Hanson et al. (1997) note that the DIB features in spectra of M17 stars do not show large variations in strength, despite the fairly wide range in total extinction,  $A_V = 3 - 10 \text{ mag}$ . We consider the explanation likely that the DIBs are mostly tracing foreground dust and that the (unidentified) DIB carriers may only exist in the diffuse medium, not in the dark cloud environment of M17.

### 6.3.3 Spectral energy distribution

Figure 6.3 shows the flux-calibrated X-shooter spectrum (300–2500 nm) of B275. The photometric data points demonstrate the accuracy of the spectrophotometric calibration. The long standing debate over the distance to M17 (ranging from 1.3 to 2.1 kpc) has recently been settled by the measurement of the trigonometric parallax of the CH<sub>3</sub>OH maser source G15.03–0.68 (Xu et al. 2011), resulting in a distance of  $1.98_{-0.12}^{+0.14} \text{ kpc}$  so that M17 is most likely located in the Carina-Sagittarius spiral arm.

We deredden the flux-calibrated X-shooter spectrum of B275 (Fig. 6.3) using the parameterization of the extinction law by Cardelli et al. (1989). The dereddened spectrum is a fit to a Kurucz model (Kurucz 1979, 1993) based on an iterative procedure, with fixed parameters  $T_{\text{eff}} = 13,000 \text{ K}$ ,  $\log g = 3.5$ ,  $d = 1.98 \text{ kpc}$  and  $R_V = 3.3$  (an effective value resulting from interstellar and local extinction). This yields independent best-fit values of  $A_V = 6.1 \pm 0.6 \text{ mag}$  and  $R_{\star} = 8.1 \pm 0.8 R_{\odot}$ . Note that this radius is much larger than that of, e.g., a B5 zero-age main sequence (ZAMS) star ( $2.7 R_{\odot}$ , Hanson et al. 1997). An additional constraint is provided by the height of the Balmer jump, which also varies with  $T_{\text{eff}}$  and  $R_{\star}$ : Fig. 6.3 demonstrates that the Balmer jump (as well as the Paschen jump) is nicely fit to the observed spectrum. Thus, with a larger radius the discrepancy between the classification of the photospheric spectrum and the dereddened SED is solved. The consequence is that B275 is not on the main sequence but is a so-called



**Figure 6.3:** The flux-calibrated X-shooter spectrum of B275 from 300–2500 nm (black) along with the photometric data (red triangles, black error bars) from Chini et al. (1980) (UVBRI), 2MASS (Skrutskie et al. 2006, JHK), *Spitzer* GLIMPSE (Benjamin et al. 2003, 3.6, and 5.8  $\mu\text{m}$ ), and Nielbock et al. (2001) (N, Q). When dereddened ( $A_V = 6.1$  mag, orange line, blue diamonds), the SED is described well by a B7 III Kurucz model (blue, dashed line). The excess flux at 500 – 800 nm is an instrumental feature.

bloated star, where the appropriate spectral type would be B7 III. The corresponding luminosity is  $\log L_*/L_\odot = 3.2$ .

### 6.3.4 Accretion signatures

A pronounced, double-peaked emission feature is detected in the strongest H Balmer lines, the Ca II triplet and the O I 844.6 nm line (Fig. 6.1). The measured peak-to-peak separation ranges from  $71 \pm 7 \text{ km s}^{-1}$  (O I 844.6 nm) to  $105 \pm 3 \text{ km s}^{-1}$  (Ca II 849.8 nm), and is centered at the rest-frame velocity of the star. The Ca II triplet lines are probably produced in an optically thick medium, since their strength ratio is not 1:9:5. The strongest lines of the H I Paschen and Brackett series also exhibit a central emission component, though it is single-peaked. The higher series members include a weaker emission component that may be double-peaked. A number of metallic emission lines (e.g., C I and Fe II) are detected throughout the spectrum, albeit very weak.

Prominent CO 1<sup>st</sup>-overtone emission bandheads are detected at 2.3  $\mu\text{m}$ , with clear evidence of a blue shoulder. We also confirm the presence of 2<sup>nd</sup>-overtone CO bandhead emission at 1.5  $\mu\text{m}$  (Hanson et al. 1997). CO is easily dissociated so must be shielded



from the strong UV flux of the young massive star. On the other hand, to produce 1<sup>st</sup> overtone emission, CO must be excited, requiring a temperature in the range between 1500 and 4500 K (Bik & Thi 2004). This temperature might even be higher, considering the unprecedented detection of 2<sup>nd</sup> overtone emission. These conditions can be met in the plane of a dense circumstellar disk where the CO molecules can be formed, excited, and protected from dissociation through “self-shielding”. The relative strength and shape of the CO bandheads can be modeled by an optically thin Keplerian disk (Bik 2004; Blum et al. 2004), where the blue shoulder would imply a relatively high inclination angle of the disk (“edge-on”). Blum et al. (2004) model the CO 2–0 first-overtone rovibrational bandhead at 2294 nm of B275 resulting in  $v \sin i = 109.7 \pm 0.6 \text{ km s}^{-1}$  (at the inner edge of the CO emission zone) and surface density  $N_{\text{CO}} = 3.5 \pm 0.2 \times 10^{21} \text{ cm}^{-2}$ . The double-peaked emission profiles, as shown in Fig. 6.1, are very similar to the emission line profile of a single line obtained by Blum et al. (2004).

We find no evidence for veiling of the optical spectrum or any strong indications of active “heavy” accretion and/or jets, such as those observed in some other systems (Chapter 3). The [O I] 630 nm line very likely has a nebular origin.

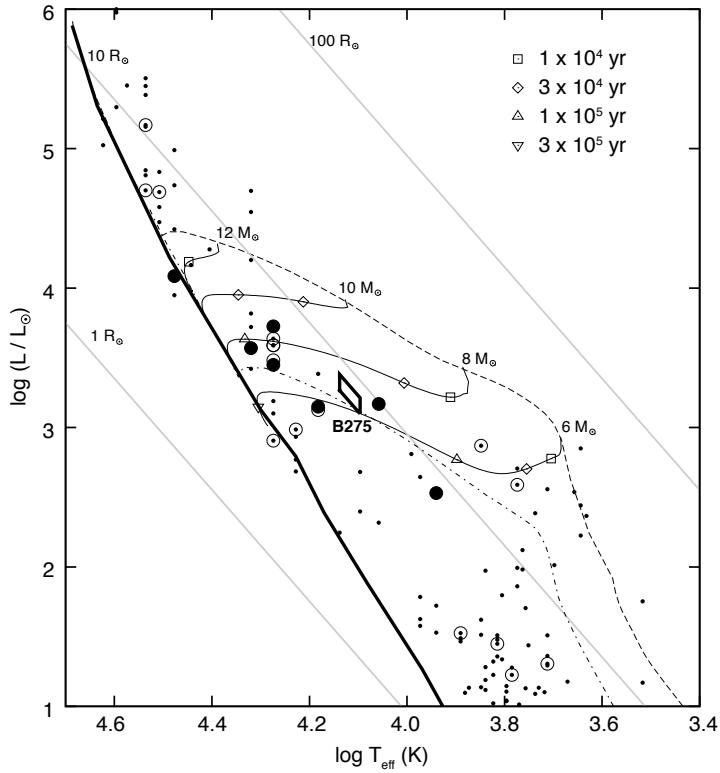
## 6.4 Discussion

The accurate spectral classification and SED fit result in a well-defined position of B275 in the Hertzsprung-Russell diagram (HRD, Fig. 6.4). It is located well above the ZAMS, demonstrating its PMS nature. If B275 is contracting towards the ZAMS, the final ZAMS mass would be 6–8  $M_{\odot}$  (spectral type B1-B2), assuming that no additional mass is accreted.

To be visible at this location in the HRD, the star must have experienced an average accretion rate of at least  $10^{-5} M_{\odot} \text{ yr}^{-1}$  in its recent history. B275 may thus be the long-sought-for example of an early-B PMS star. Figure 6.4 also shows other nearby stars in M17. Hanson et al. (1997) derive an age of  $\sim 1$  Myr for the M17 cluster and Hoffmeister et al. (2008) estimate that the PMS objects are less than  $5 \times 10^5$  yr old. Based on its location on the HRD we estimate the age of B275 at  $10^5$  yr.

B275 bears some resemblance to a classical Be star. However, we note that Be stars do not emit CO 1<sup>st</sup> and 2<sup>nd</sup> overtone emission. In addition, B275 would be classified as a luminous Herbig Be star according to the definition discussed in Carmona et al. (2010), but our analysis of the photospheric spectrum allows for a more quantitative classification.

B275 has a significant amount of infrared excess, starting at  $1 \mu\text{m}$ , and a flat SED between 2 and  $10 \mu\text{m}$  (Nielbock et al. 2001). This indicates the presence of a flaring circumstellar disk in which the dust has not yet settled. However, the visibility of the photosphere, the small number of optical gas emission lines and the absence of a jet lead us to believe that the current mass-accretion rate is not very high. Nevertheless, the CO (2-0) and (3-0) emission originates in a dense and highly excited inner part of the disk. Either the system is in an intermittent phase between accretion episodes, or it is on the verge of photo-evaporating its disk. Either scenario is consistent with its location



**Figure 6.4:** The location of B275 (thick parallellogram) in the HRD next to PMS tracks from Hosokawa & Omukai (2009) with the ZAMS mass labeled and open symbols indicating lifetimes. The thin dashed and thin dot-dashed lines are the birth lines for accretion rates of  $10^{-4} M_{\odot} \text{ yr}^{-1}$  and  $10^{-5} M_{\odot} \text{ yr}^{-1}$ , respectively; the thick solid line is the ZAMS (Schaller et al. 1992). The filled and open circles represent stars in M17 for which a spectral type has been determined (Hoffmeister et al. 2008), within a radius of  $0.5$  and  $1.0$ , respectively; dots are other stars in M17. B275 is on its way to becoming a 6–8  $M_{\odot}$  ZAMS star.

in the HR-diagram: an intermediate-mass, visible star in M17 on its way to becoming an early-B main-sequence star.

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