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XMM–Newton and Swift spectroscopy of the newly discovered very faint X-ray transient IGR J17494–3030

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ABSTRACT
A growing group of low-mass X-ray binaries are found to be accreting at very faint X-ray luminosities of \( <10^{36} \text{ erg s}^{-1} \) (2–10 keV). One such system is the new X-ray transient IGR J17494–3030. We present Swift and XMM–Newton observations obtained during its 2012 discovery outburst. The Swift observations trace the peak of the outburst, which reached a luminosity of \( \sim 7 \times 10^{35} \text{ (D/8 kpc)}^2 \text{ erg s}^{-1} \) (2–10 keV). The XMM–Newton data were obtained when the outburst had decayed to an intensity of \( \sim 8 \times 10^{34} \text{ (D/8 kpc)}^2 \text{ erg s}^{-1} \). The spectrum can be described by a power law with an index of \( \Gamma \sim 1.7 \) and requires an additional soft component with a blackbody temperature of \( \sim 0.37 \text{ keV} \) (contributing \( \sim 20 \) per cent to the total unabsorbed flux in the 0.5–10 keV band). Given the similarities with high-quality spectra of very faint neutron-star low-mass X-ray binaries, we suggest that the compact primary in IGR J17494–3030 is a neutron star. Interestingly, the source intensity decreased rapidly during the \( \sim 12 \text{ h} \) XMM–Newton observation, which was accompanied by a decrease in inferred temperature. We interpret the soft spectral component as arising from the neutron-star surface due to low-level accretion, and propose that the observed decline in intensity was the result of a decrease in the mass-accretion rate on to the neutron star.

Key words: accretion, accretion discs – stars: individual: IGR J17494–3030.

1 INTRODUCTION
Low-mass X-ray binaries (LMXBs) are composed of a black hole or a neutron star that accretes material from a (sub)solar companion star that overflows its Roche lobe. A large fraction of LMXBs remain in a dim quiescence state, during which no or very little accretion occurs and the X-ray luminosity \( L_X \) is \( \sim 10^{30–33} \text{ erg s}^{-1} \). However, once in a while, the sources experience outburst events in which their accretion rate increases drastically and consequently their X-ray brightness as well. LMXBs can be classified depending on the maximum 2–10 keV luminosity \( L_X^{\text{peak}} \) that they reach. Those systems that can obtain \( L_X^{\text{peak}} \) of \( \sim 10^{37–39} \text{ erg s}^{-1} \) are called (very) bright systems. The faint ones can reach \( L_X^{\text{peak}} \sim 10^{36–37} \text{ erg s}^{-1} \), and the sources that display \( L_X^{\text{peak}} \) of only \( \sim 10^{34–36} \text{ erg s}^{-1} \) are called very faint X-ray binaries (VFXBs; Wijnands et al. 2006). Despite significant progress over the last few years in our understanding of the behaviour of those VFXBs (Muno et al. 2005; Campana 2009; Degenaar & Wijnands 2009, 2010; Armas Padilla et al. 2011, 2013a; Armas Padilla, Degenaar & Wijnands 2013b), much remains unclear about them (e.g. the mechanism behind their low luminosities is not understood).

One recently discovered VFXB is IGR J17494–3030. The source was first detected during the Galactic Centre INTEGRAL observation performed on 2012 March 17–19 (Boissay et al. 2012). Multiple Swift X-ray telescope (XRT) observations were obtained (Bozzo et al. 2012) but \( \sim 27 \text{ d} \) after the source was first detected, it could not be detected anymore using Chandra (Chakrabarty, Jonker & Markwardt 2013). Inspecting the near-infrared (NIR) images acquired on 2010 July as part of the NIR VVV survey (Minniti et al. 2010), five sources could be identified within the Swift/XRT error circle. Those sources are potential candidates for the quiescent NIR counterpart of IGR J17494–3030 although it cannot be excluded that none of those sources is associated with the source. In this Letter, we present the analysis of four Swift observations and our XMM–Newton observation of IGR J17494–3030 obtained during the outburst decay phase.

2 OBSERVATIONS AND REDUCTION
2.1 XMM–Newton data
IGR J17494–3030 was observed with XMM–Newton (Jansen et al. 2001) on 2012 March 31 for \( \sim 43 \text{ ks} \) (see Table 1). The European
Photon Imaging Camera (EPIC) on board XMM–Newton consists of two MOS cameras (Turner et al. 2001) and one PN camera (Strüder et al. 2001). They were operated in imaging (full-frame window) and timing mode, respectively. We reduced the data and obtained scientific products using the Science Analysis Software (SAS, v. 13.0).

We filtered out an episode of background flaring by only selecting data for which the high-energy count rate was $<0.25$ counts s$^{-1}$ ($>10$ keV) for the MOS cameras and $<0.2$ counts s$^{-1}$ ($10$–$12$ keV) for the PN. The total resulting live time is $\sim 42$ ks for the MOS cameras and $\sim 39$ ks for the PN detector. For the MOS cameras, we extracted the source event file using a circular region centred on the source position and with a radius of $\sim 47$ arcsec; the background was extracted using a circular region with a radius of $\sim 107$ arcsec centred on a source-free region. The MOS source count rate (Table 1) is below the 0.7 counts s$^{-1}$ pile-up threshold.1 After correcting the PN data for the transfer inefficiency that affects data obtained using the timing modes,2 we extracted the source and background events by selecting the data with RAWX columns [32:42] and [5:12], respectively. The PN source count rate (Table 1) is much lower than the pile-up limit of 800 counts s$^{-1}$ for the PN timing mode.1 We generated the spectra and the light curves as well as the response files applying the standard analysis threads. We grouped the spectra to contain a minimum of 25 photons per bin and rebinned the data to not to oversample the intrinsic energy resolution by a factor larger than 3.

### 2.2 Swift data

A total of four observations were obtained of IGR J17494–3030 with the XRT (Burrows et al. 2005) on board Swift (Gehrels et al. 2004). Three observations were performed in photon counting (PC) mode and one in windowed timing (WT) mode (Table 1). We processed the data making use of the HEASOFT v.6.12 software. The data reduction was carried out running the XRTPipeline task. For every observation, spectra, light curves and images were obtained using XSELECT. For the WT data, we used a circular region radius of $\sim 76$ arcsec centred on the source position to extract the source events, and a similar region far enough from the source to extract the background events. For the PC data, we used a circular region of $\sim 47$ arcsec centred on the source position to extract the source events and three circular regions of similar size for the background data. Observations 00032318001 and 00032318003 are affected by pile-up. To mitigate the pile-up effects, we excluded the inner $\sim 11$ arcsec (observation 00032318001) and $\sim 9$ arcsec (00032318003) from the central part of the source regions. We created exposure maps and ancillary response files following the standard Swift analysis threads, and we acquired the last version of the response matrix files from the HEASARC calibration data base (v.14). We grouped the spectra to have at least 20 counts per bin.

### 3 ANALYSIS AND RESULTS

#### 3.1 Spectral analysis

We used XSPEC (v.12.8; Arnaud 1996) to fit the spectra. We incorporated in our models the photoelectric absorption component (PHABS) to account for the interstellar absorption. We simultaneously fitted the spectra obtained with the different XMM–Newton/EPIC detectors; the two MOS spectra were fitted in the energy range $0.5$–$10$ keV and the PN spectrum in the $0.7$–$10$ keV range. We tied all the parameters between the three detectors and introduced a constant factor (CONSTANT) to account for cross calibration uncertainties between the instruments. It was fixed to one for the PN spectrum and allowed to vary freely for the MOS ones. A single power-law (POWERLAW) model returns a hydrogen column density ($N_{\text{HI}}$) of $1.87 \times 10^{22}$ cm$^{-2}$ and a power-law photon index ($\Gamma$) of 2.14. However, with a reduced $\chi^2 (\chi^2_{\nu})$ of 1.28 for 421 degrees of freedom (dof), the fit is not acceptable. Adding a soft blackbody component (BBODYRAD) to the power-law model improves the fit. The resulting spectra with the best-fitting model are shown in Fig. 1.

<table>
<thead>
<tr>
<th>Satellite/Instrument</th>
<th>Mode</th>
<th>Observation ID</th>
<th>Date (yyyy-mm-dd)</th>
<th>Exposure time (ks)</th>
<th>Count rate (counts s$^{-1}$)</th>
<th>Net count rate$^d$ (counts s$^{-1}$)</th>
</tr>
</thead>
<tbody>
<tr>
<td>Swift/XRT</td>
<td>PC</td>
<td>00032318001</td>
<td>2012-03-20</td>
<td>1.0</td>
<td>0.93</td>
<td>0.35</td>
</tr>
<tr>
<td>Swift/XRT</td>
<td>WT</td>
<td>00032318002</td>
<td>2012-03-23</td>
<td>1.0</td>
<td>1.89</td>
<td>1.63</td>
</tr>
<tr>
<td>Swift/XRT</td>
<td>PC</td>
<td>00032318003</td>
<td>2012-03-26</td>
<td>0.6</td>
<td>0.80</td>
<td>0.39</td>
</tr>
<tr>
<td>Swift/XRT</td>
<td>PC</td>
<td>00032318004</td>
<td>2012-03-30</td>
<td>1.1</td>
<td>0.28</td>
<td>0.28</td>
</tr>
<tr>
<td>XMM–Newton/EPIC</td>
<td>(MOS1) Imaging</td>
<td>0694040201</td>
<td>2012-03-31</td>
<td>43.9</td>
<td>0.44</td>
<td>0.38</td>
</tr>
<tr>
<td>XMM–Newton/EPIC</td>
<td>(MOS2) Imaging</td>
<td></td>
<td></td>
<td></td>
<td>0.44</td>
<td>0.38</td>
</tr>
<tr>
<td></td>
<td>(PN) Timing</td>
<td></td>
<td></td>
<td></td>
<td>1.34</td>
<td>0.98</td>
</tr>
</tbody>
</table>

$^d$Count rate after background correction, using the annulus regions to mitigate the pile-up in Swift obs 00032318001 and 00032318003, and after excluding a short episode of background flaring observed during the XMM–Newton observation.

2 See the calibration technical note XMM-SOC-CALTN-0083.
1.4 M⊙ and the radius to 10 km. We used a distance \(D\) of 8 kpc and we set the normalization to 1 (i.e. the whole surface is assumed to emit radiation), which leaves the neutron-star temperature as the only parameter free for the NSATMOS model. The obtained temperature (for an observer at infinity) is \(kT_\infty = 0.182 \pm 0.004\) keV and the soft component contributes \(\sim 22\) per cent to the total 0.5–10 keV unabsorbed flux. The other fit parameters are, within the errors, consistent with the values obtained when using the BBODYRAD model (see Table 2).

All Swift spectra are well described with a single absorbed power-law model (see Table 2). We fixed \(N_H\) to the value obtained from the XMM–Newton fit (1.87 × 10²² cm⁻²). The photon index value is, within the errors, consistent in all observations. The peak unabsorbed flux is \(\sim 1.7 \times 10^{-10}\) erg cm⁻² s⁻¹ (as observed during the second Swift observation) which corresponds to an \(L_X\) of \(\sim 1.3 \times 10^{36}/(D/8\) kpc)² erg s⁻¹ (0.5–10 keV). The lowest detected flux was observed during the XMM–Newton observation. The unabsorbed flux decreased approximately one order of magnitude in \(\sim 8\) d (see also Fig. 2), which gives a luminosity seen during the XMM–Newton observation of \(\sim 8.6 \times 10^{34}/(D/8\) kpc)² erg s⁻¹. We calculated the 0.5–10 keV unabsorbed flux upper limit from the Chandra count rate upper limit reported in Chakrabarty et al. (2013) using WebPIMMS. We assumed an absorbed power-law model with \(N_H = 1.87 \times 10^{22}\) cm⁻² and \(\Gamma = 2.14\) (the values obtained from the XMM–Newton spectra). The calculated 0.5–10 keV unabsorbed flux is <6.8 × 10⁻¹⁴ erg cm⁻² s⁻¹ which results in a luminosity of <5.2 × 10⁻¹⁴/(D/8 kpc)² erg s⁻¹ (see Fig. 2).

### 3.2 Light curve analysis

The count rate curve obtained during the XMM–Newton shows that the source intensity decreased during the observation, from \(\sim 0.65\) to \(\sim 0.25\) counts s⁻¹ within \(\sim 12\) h (see Fig. 2). In order to study the spectral evolution during the observation, we divided the observation in three segments of \(\sim 14.5\) ks each and individually fitted each spectrum (both the PN and MOS spectra simultaneously). We used the absorbed NSATMOS + POWERLAW model with the \(N_H\) fixed to the value obtained from the fit using the data obtained from the whole observation (see Section 3.1). While the power-law index is constant within the errors (\(\Gamma \sim 1.7\)), the temperature decreases from \(kT^\infty = 0.190 \pm 0.004\) to \(170 \pm 0.005\) keV (Fig. 3; the errors are for 1σ). We can fit this decrease adequately with an exponential decay function with an \(\epsilon\)-folding time of \(3.0^{+1.5}_{-1}\) d or with a power-law decay function with an index of \(\sim 0.06 \pm 0.02\) (see

### Table 2. Results from the spectral fits.

<table>
<thead>
<tr>
<th>Instrument/Obs ID</th>
<th>(N_H) (10²² cm⁻²)</th>
<th>(\Gamma)</th>
<th>(kT) (keV)</th>
<th>(T_{fr}) (per cent)</th>
<th>(F_{X,abs}) (10⁻¹¹ erg cm⁻² s⁻¹)</th>
<th>(F_{X,unabs}) (10⁻¹¹ erg cm⁻² s⁻¹)</th>
<th>(L_X) (10³⁵ erg s⁻¹)</th>
<th>(\chi^2) (dof)</th>
</tr>
</thead>
<tbody>
<tr>
<td><strong>Swift/XRT</strong></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>000323180001</td>
<td>1.8 (fix)</td>
<td>1.8</td>
<td>0.2</td>
<td>–</td>
<td>8.6 ± 0.7</td>
<td>14.0 ± 0.2</td>
<td>10.7 ± 0.8</td>
<td>0.99 (15)</td>
</tr>
<tr>
<td>00032318002</td>
<td>1.8 (fix)</td>
<td>1.9</td>
<td>0.1</td>
<td>–</td>
<td>10.2 ± 0.4</td>
<td>17.4 ± 0.1</td>
<td>13.4 ± 0.8</td>
<td>0.98 (79)</td>
</tr>
<tr>
<td>00032318003</td>
<td>1.8 (fix)</td>
<td>1.8</td>
<td>0.3</td>
<td>–</td>
<td>6.9 ± 0.9</td>
<td>11.3 ± 0.3</td>
<td>8.7 ± 0.8</td>
<td>0.53 (9)</td>
</tr>
<tr>
<td>00032318004</td>
<td>1.8 (fix)</td>
<td>2.0</td>
<td>0.3</td>
<td>–</td>
<td>1.8 ± 0.2</td>
<td>3.4 ± 0.2</td>
<td>2.6 ± 0.8</td>
<td>0.92 (13)</td>
</tr>
<tr>
<td><strong>XMM/EPIC</strong></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>0694040201</td>
<td>1.80 ± 0.07</td>
<td>1.76</td>
<td>0.08</td>
<td>0.37 ± 0.03</td>
<td>17.3</td>
<td>0.649 ± 0.005</td>
<td>1.13 ± 0.05</td>
<td>0.86 ± 0.05</td>
</tr>
<tr>
<td></td>
<td>1.89 ± 0.03</td>
<td>1.74</td>
<td>0.07</td>
<td>0.182 ± 0.004</td>
<td>22.9</td>
<td>0.649 ± 0.003</td>
<td>1.18 ± 0.01</td>
<td>0.90 ± 0.01</td>
</tr>
</tbody>
</table>

*Note.* Quoted errors represent 90 per cent confidence levels. The fifth column reflects the fractional contribution of the thermal component to the total unabsorbed 0.5–10 keV flux. \(F_{X,abs}\) and \(F_{X,unabs}\) represent the absorbed and unabsorbed fluxes (0.5–10 keV), respectively. The luminosity \(L_X\) (0.5–10 keV) was calculated adopting a distance of 8 kpc. The \(kT\) for the NSATMOS model is for an observer in infinity.
the presence of the soft component (and its associated temperature) in the X-ray spectrum of IGR J17494–3030 is reminiscent of what is seen in neutron-star X-ray binaries at similar luminosities (e.g. Fridriksson et al. 2010; Armas Padilla et al. 2013b; Degenaar, Wijnands & Miller 2013). In those systems, the thermal component is thought to originate from the surface of the neutron star. Given the strong similarities, we tentatively identify the accretor in IGR J17494–3030 as a neutron star.

If the soft component is indeed originating from the neutron-star surface, then it is likely due to low-level accretion on to the surface. Such low-level accretion will indeed produce a soft spectrum (Zampieri et al. 1995) and in principle one should use the so-called Zamp model (Soria et al. 2011) to fit the soft component in the X-ray spectra. However, that model is not publicly available so the next best things are the neutron-star atmosphere models. Although those models assume that the emission is from a cooling neutron star and therefore incorporate different microphysics compare to the Zamp model, typically the results are very similar between the models (i.e. the inferred surface temperature; Soria et al. 2011). Therefore, we will discuss the results obtained using neutron-star atmosphere models which also allows for direct comparison with previous work in which only neutron-star atmosphere models were used.

During the XMM–Newton observation we see a steady decrease of the source intensity. Together with the intensity, the temperature of the soft component is decreasing as well. Although such a temperature decrease with decreasing intensity would be expected irrespectively if the soft component is due to a cooling accreting disc or due to decreasing accretion on to the neutron-star surface, it is very similar to what recently has been found for the bright neutron-star X-ray transient XTE J1709–267 by Degenaar et al. (2013). Those authors found that a similar (albeit at slightly lower luminosities and lower inferred surface temperatures) decay during the end stages of the 2012 outburst of XTE J1709–267. They favoured an interpretation that the observed temperature decrease was due to the cooling of the neutron-star crust which was heated by the accretion during the outburst.

However, the outburst of IGR J17494–3030 was much shorter (3–4 weeks versus 10 weeks) and much less luminous (peaking at \(~7 \times 10^{35}\) erg s\(^{-1}\) versus \(~5 \times 10^{37}\) erg s\(^{-1}\); 2–10 keV) than that of XTE J1709–267. Within our understanding of crustal heating and cooling models, IGR J17494–3030 should have been heated to a much lesser extent (if at all). It would not be expected to display similarly strong signs of crustal cooling as XTE J1709–267. Therefore, this likely is not the correct interpretation of the observed rapid intensity decay during the XMM–Newton observation for IGR J17494–3030. Since the behaviour of this source is as similar as what has been observed for XTE J1709–267, likely also in that source we did not observe the crust cooling.

Degenaar et al. (2013) alternatively suggested that the temperature decrease observed for XTE J1709–267 was not due to a cooling crust but maybe the thermal component was due to low-level accretion on the neutron-star surface and the accretion rate decreased causing the surface temperature to decrease as well. This scenario could more naturally account for the similarities in the decay seen for the two sources, despite their different outburst properties. Therefore, we favour a low-level accretion scenario to explain the thermal component in both sources. If this is indeed the correct interpretation for the decay seen in both sources, then there is no need for an additional heat source in the neutron-star crust, as was proposed by Degenaar et al. (2013).

Moreover, investigating the literature for more detections of soft thermal components at such low-accretion luminosities of

\[^{3}\] As reference time we used the start of the XMM–Newton observation, but we found that the e-folding time or index is not very sensitive to the exact value of the reference time.
accreting neutron-star systems (or candidate neutron-star systems), we find several additional sources which are consistent with the hypothesis that the soft thermal component in the luminosity range of $10^{34}–10^{35}$ erg s$^{-1}$ is due to low-level accretion on to the neutron-star surface. For example, for the suspected neutron-star very faint X-ray transient XTE J1719–291 and for two confirmed persistent neutron-star VFXBs (1RXH J173523.7–354013, 1RXS J171824.2–402934), we found in previous work that their spectra also required a thermal component in addition to the power-law component (Armas Padilla et al. 2011; Armas Padilla et al. 2013b). We refitted their spectra with the NSATMOS model (we previously used black-body models) and obtained temperatures of $0.162 \pm 0.004$ keV, $0.184 \pm 0.004$ keV and $0.194 \pm 0.003$ keV, respectively, with associated luminosities of $3.3 \times 10^{34}$, $4.4 \times 10^{34}$ and $6.7 \times 10^{34}$ erg s$^{-1}$.

In addition, Fridriksson et al. (2010) reported on a brief accretion flare during the quiescent state of the neutron-star X-ray transient XTE J1701–462 during which the temperature was $0.159 \pm 0.02$ keV for a luminosity of $2.6 \times 10^{34}$ erg s$^{-1}$.

Uncertainties in the distances affect the inferred temperatures, but there appears to be a tendency that for higher luminosities, the temperature is higher, as one would expect for the low-level accretion scenario. This provides strong evidence that in all these sources the thermal component indeed arises from the neutron-star surface as a result of low-level accretion and that the inferred temperature is determined by the instantaneous mass accretion rate.

The decrease of surface temperature while the luminosity (and thus the inferred accretion rate on to the neutron-star surface) decreases can only continue (in the transient sources) as long as the temperature due to the low-level accretion is higher than the interior (crust) temperature of the neutron star. If the accretion rate drops below a certain value, then the light-curve evolution will not be governed anymore by the decay in accretion rate but instead it will be determined by how fast the crust cools down (even small outbursts will have a slightly heated crust) and eventually by the core cooling rate. Therefore, in the light curve we would expect a break at a certain luminosity from a rapid decay of the luminosity to a much slower decay rate. Such a break has been observed in several systems (e.g. XTE J1701–462 by Fridriksson et al. 2010; MAXI J0556–332 by Homan et al., in preparation; Aql X-1 by Campana et al. 1998; Campana et al., in preparation) and indeed has been interpreted as the onset of the crust cooling (e.g. Fridriksson et al. 2010).