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OLD PUZZLE, NEW INSIGHTS: A LITHIUM-RICH GIANT QUIETLY BURNING HELIUM IN ITS CORE

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ABSTRACT

About 1% of giant stars have been shown to have large surface Li abundances, which is unexpected according to standard stellar evolution models. Several scenarios for lithium production have been proposed, but it is still unclear why these Li-rich giants exist. A missing piece in this puzzle is the knowledge of the exact stage of evolution of these stars. Using low- and high-resolution spectroscopic observations, we have undertaken a survey of lithium-rich giants in the Kepler field. In this Letter, we report the finding of the first confirmed Li-rich core-helium-burning giant, as revealed by asteroseismic analysis. The evolutionary timescales constrained by its mass suggest that Li production most likely took place through non-canonical mixing at the RGB tip, possibly during the helium flash.

Key words: asteroseismology – stars: abundances – stars: individual (KIC 5000307) – stars: late-type – stars: oscillations

Online-only material: color figures

1. INTRODUCTION

Lithium nuclei are readily destroyed via proton capture when they are exposed to temperatures exceeding ~2.6 × 10^6 K. As a star leaves the main-sequence phase, its surface Li abundance is expected to decrease due to the inward penetration of the convective envelope called the first dredge-up (FDU; e.g., Salaris et al. 2002). This process carries material from the surface to hotter interior regions where lithium is burned, depleting its amount compared to the initial value (Iben 1967). As a consequence, the expected canonical surface lithium abundance of a 1.5 M⊙ star at the end of the FDU is A(Li) ∼ 1.512 (e.g., Palmerini et al. 2011).

Nevertheless, about 1% of giant stars show an unusual enhancement in their surface Li abundance (cf. Brown et al. 1989), an occurrence that challenges standard stellar evolution models. A variety of scenarios have been proposed to explain this phenomenon (e.g., Sackmann & Boothroyd 1999; Boothroyd & Sackmann 1999; Romano et al. 1999; Charbonnel & Balachandran 2000), commonly featuring Li enrichment during the asymptotic-giant branch (AGB) phase or at the luminosity function bump in the red giant branch (RGB). Evolutionary classifications of many known Li-rich giants, primarily based on their location in the Hertzsprung–Russell diagram (HRD), have been found consistent with this assumption.

However, observations have revealed the existence of Li-rich giants at different luminosities along the RGB (e.g., Monaco et al. 2011), whereas any lithium enhancement taking place at the bump phase is expected to be depleted by the time the star reaches the RGB tip. Moreover, Kumar et al. (2011) found Li-rich giants having atmospheric properties in agreement with them being in the clump phase (see also Martell & Shetrone 2013). Finding Li-rich giants in the core-helium burning phase is evidence of a different enrichment scenario that could be triggered by non-canonical mixing at the RGB tip (e.g., during core helium ignition) or at the clump phase itself.

If we aim at unveiling the processes that enhance lithium in red giants, detailed elemental abundances from high-resolution spectroscopy and a clear classification of the evolutionary phase of these stars are absolutely critical. This is now possible thanks to asteroseismic observations from space-borne missions, which allow discrimination between red giants in the hydrogen-shell or core-helium burning phases (Bedding et al. 2011). A lithium-rich giant below the RGB-bump luminosity has recently been found by Anthony-Twarog et al. (2013) in the Kepler field. In this Letter, we report the discovery of the first confirmed Li-rich clump red giant star, KIC 5000307. We determine the combined spectroscopic and asteroseismic properties of this target, and use these data to investigate the mixing and nucleosynthesis processes that could originate its Li enhancement.

2. CANDIDATE SELECTION AND OBSERVATIONS

As Li-rich giants are extremely rare, large spectroscopic surveys are ideal for their identification. The Large Sky Area Multi-Object Fiber Spectroscopic Telescope (LAMOST; Cui et al. 2012) is well-suited for the discovery of these rare objects, since its low-resolution (R ∼ 2000) spectra has a wavelength coverage containing the 6708 Å Li line. Using synthetic spectra,
spectra were reduced using the FIEStool13 reduction software).

13 http://www.not.iac.es/instruments/fies/fiestool/FIEStool.html

Li 6708 line with an equivalent width Wλ > 200 m Å is detectable. Using this value, we derived a Li abundance of A(Li) = 2.80 in LTE and A(Li) = 2.71 in NLTE, computed following the NLTE corrections of Lind et al. (2009). The [C/Fe] and [N/Fe] ratios, as well as the 12C/13C isotopic ratio, were derived from spectral synthesis of CH and CN lines (B. Plez 2011, private communication) using MOOG under molecular equilibrium. Details of the oscillator strengths and dissociation energies can be found in Ruchti et al. (2011).

The syntheses are shown in Figure 1. The resultant abundances indicate that the star is deficient in carbon, [C/Fe] = −0.8 ± 0.1, while enhanced in nitrogen, [N/Fe] = +0.9 ± 0.15. Features sensitive to the 12C/13C ratio are all very weak. We can thus only place a limit of 12C/13C < 20 using the CN feature at 4208.3 Å. Unfortunately, all oxygen lines are affected by telluric emission which prevented us from determining the oxygen abundance.

3.2. Lithium and CN Abundances

The 7Li abundance was determined following the methodology described in Ruchti et al. (2011). We measured an equivalent width of 250 m Å for the 6708 Å Li line. Using this value, we derived a Li abundance of A(7Li) = 1.4 ± 0.4, consistent with results from the literature. The overall power spectrum shows a Gaussian-shaped envelope (e.g., Chaplin & Miglio 2013), where the frequency of maximum oscillation power is known as νmax, which we derive from the Fe i lines and that from the Fe ii lines. We note that the spectroscopic log g value agrees well with the one derived using asteroseismology (see Section 4 below).

3. SPECTROSCOPIC ANALYSIS

All parameters of KIC 5000307 derived directly from spectroscopic and asteroseismic observations, as well as those determined in combination with stellar models, are listed in Table 1.

### Table 1

<table>
<thead>
<tr>
<th>Property</th>
<th>Value</th>
</tr>
</thead>
<tbody>
<tr>
<td>R.A. (&quot;)d</td>
<td>19:13:39.1</td>
</tr>
<tr>
<td>decl. (&quot;)</td>
<td>+40:11:04.6</td>
</tr>
<tr>
<td>K0</td>
<td>11.23</td>
</tr>
<tr>
<td>FIES Obsdate (yyyymmdd)</td>
<td>20130429</td>
</tr>
<tr>
<td>FIES Spectra S/N2</td>
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<tr>
<td>vmax (μHz)</td>
<td>42.46 ± 0.47</td>
</tr>
<tr>
<td>∆ν (μHz)</td>
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</tr>
<tr>
<td>Period spacing (s)</td>
<td>319.95 ± 0.4</td>
</tr>
<tr>
<td>Rotational splitting δνrot (nHz)</td>
<td>50 ± 10</td>
</tr>
<tr>
<td>Mass (M⊙)</td>
<td>1.536 ± 0.009</td>
</tr>
<tr>
<td>Radius (R⊙)</td>
<td>11.01 ± 0.217</td>
</tr>
<tr>
<td>log gseis</td>
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</tr>
<tr>
<td>L/L⊙</td>
<td>1.819 ± 0.023</td>
</tr>
<tr>
<td>Age (Myr)</td>
<td>2082 ± 417</td>
</tr>
<tr>
<td>[Fe/H]d</td>
<td>−0.29 ± 0.05</td>
</tr>
<tr>
<td>T eff, K</td>
<td>5000 ± 70</td>
</tr>
<tr>
<td>log gpec</td>
<td>2.56 ± 0.1</td>
</tr>
<tr>
<td>vτ (km s⁻¹)</td>
<td>1.4 ± 0.1</td>
</tr>
<tr>
<td>EW Li (6708 Å)</td>
<td>250</td>
</tr>
<tr>
<td>A(Li)</td>
<td>2.80</td>
</tr>
<tr>
<td>A(Li)NLTE</td>
<td>2.71</td>
</tr>
<tr>
<td>[C/Fe]</td>
<td>−0.80 ± 0.10</td>
</tr>
<tr>
<td>[N/Fe]</td>
<td>0.90 ± 0.15</td>
</tr>
<tr>
<td>12C/13C</td>
<td>&lt;20</td>
</tr>
</tbody>
</table>

Notes.

b. Kepler magnitude.
c. Estimated near the Li 6708 line.

degraded to the resolution of LAMOST, we found that a Li 6708 line with an equivalent width Wλ > 200 m Å is detectable. We used the LAMOST observations of the Kepler field to search for potential Li-rich candidates. To confirm the Li abundance and derive more accurate stellar parameters, we obtained high-resolution spectroscopic observations using the Fibre-fed Echelle Spectrograph (FIES) on the Nordic Optical Telescope located at La Palma Observatory in the Canary Islands, Spain. We set FIES to deliver a resolving power of R ~ 46,000 with a spectral coverage of 3700–7300 Å (all spectra were reduced using the FIESStool13 reduction software).

3.3. Lithium and CN Abundances

The spectroscopic stellar parameters (effective temperature, surface gravity, and metallicity) were determined following the iterative methodology described in Ruchti et al. (2013), which uses on-the-fly non-LTE (NLTE) corrections to Fe i lines. Briefly, the spectral analysis was performed with the MOOG program (Sneden 1973), using one-dimensional, plane-parallel Kurucz model atmospheres (Castelli & Kurucz 2004 and references therein), which are computed under the assumption of local thermodynamic and hydrostatic equilibrium. The effective temperature, T eff, was derived from the wings of the Balmer lines through profile fits to Hα and Hβ. The microturbulence was found by minimizing the slope of the relationship between the NLTE-corrected abundance of iron from Fe i lines and the reduced Wλ. The surface gravity, log g, and metallicity, [Fe/H], were then derived by minimizing the difference between the abundance of iron from the NLTE-corrected Fe i lines and that from the Fe ii lines. We note that the spectroscopic log g value agrees well with the one derived using asteroseismology (see Section 4 below).

4. ASTEROSEISMIC ANALYSIS

Stars with outer convective envelopes show stochastically excited oscillations that travel across their interiors, called solar-like oscillations. Observing the frequency of these pulsations yields immediate information about the size of the cavity where these waves propagate, providing stringent constraints on its physical properties. Low- and intermediate-mass red giants exhibit this type of pulsations, clearly visible in the data obtained by the CoRoT and Kepler missions (De Ridder et al. 2009; Huber et al. 2011).

For the asteroseismic analysis we have used nearly three years of Kepler data (Q0–Q12). The Kepler light curve has been extracted using the pixel data following the methods described in S. Mathur et al. (2014, in preparation) and corrected following García et al. (2011). The power spectrum of KIC 5000307 is shown in Figure 2, where the angular degree ℓ describing the geometrical component of each oscillation mode is labeled. Modes of the same degree and consecutive order n are approximately equally spaced in frequency, as can be seen for the ℓ = 0 case in Figure 2. This quantity is called the large frequency separation ∆ν = νℓ,n − νℓ,n−1. The overall power spectrum shows a Gaussian-shaped envelope (e.g., Chaplin & Miglio 2013), where the frequency of maximum oscillation power is known as νmax. These two global oscillation parameters, ∆ν and νmax, have been determined using the method described by Hekker et al. (2010).

Solar-like oscillations originate from two types of standing waves: those whose restoring force is the pressure gradient (called p-modes) and those where buoyancy acts as the restoring
quantity (called g-modes). In main-sequence stars, only pure p-modes reach observable amplitudes at the stellar surface, showing a very regular pattern for every angular degree (e.g., Bedding 2011). When a star leaves the main sequence, hydrogen burning ceases in the center and is restricted to a thin shell outside the inert helium core. The star’s envelope expands and its core contracts, inducing a large increase in the buoyancy frequency in the center. This produces coupling between the cavities where p-modes and g-modes reside, resulting in mixed non-radial (ℓ ≥ 1) modes whose pulsation frequencies behave like p-modes in the outer envelope and g-modes in the deep interior (see, e.g., Chaplin & Miglio 2013, for a review).
Red giants show a very dense spectrum of modes of mixed character, as seen in Figure 2 for the dipole ($\ell = 1$) modes. The underlying $g$-modes are approximately equally spaced in period, and measuring this separation provides strong constraints on the evolutionary stage of the star (see Figure 3 in Bedding et al. 2011). The period spacing of dipole modes ($\Delta \Pi_1$) has been determined by two independently developed methods, those of Mosser et al. (2012a) and A. Datta et al. (in preparation). Briefly, the latter method uses the empirical Lorentzian variation of the observed period spacing around an underlying $p$-mode to determine the $\Delta \Pi_1$, or the vertical stacking in a period-échelle diagram. Mosser et al. (2012a) on the other hand uses an asymptotic expansion of mixed modes. Regardless of the method employed, the resulting period spacing values are consistent within the uncertainties quoted in Table 1.

In Figure 3, we show the $\Delta \nu$ versus $\Delta \Pi_1$ diagram of the Kepler sample analyzed by Mosser et al. (2012a). This figure can be used to effectively discriminate between stars in the ascending red giant branch phase and those burning helium in the core (Bedding et al. 2011; Mosser et al. 2011). Our target sits in the region where clump stars are (period spacing values above $\sim 150$ s), confirming that this Li-rich giant has ignited helium in its core. Combining this with its value of $\Delta \nu$, a proxy for the stellar mean density and thus very sensitive to the stellar radius, ensures that the star has not yet evolved toward the AGB phase (see Stello et al. 2013).

To determine the asteroseismic mass, radius, and age of KIC 5000307, we apply a technique known as the grid-based method (e.g., Stello et al. 2009; Silva Aguirre et al. 2012). Details on the tracks and equations used to determine the theoretical asteroseismic quantities can be found in Section 3 of Silva Aguirre et al. (2013). We use BaSTI isochrones (Pietrinferni et al. 2004) including the effects of core overshooting during the main-sequence and semiconvection during the clump phase, as suggested by observations of dipole modes period spacing in Kepler red giants (Montalbán et al. 2013).

The input parameters fed to the grid-based analysis are the spectroscopic temperature and metallicity, and the asteroseismic global quantities $\Delta \nu$ and $v_{\text{max}}$. The method is applied using the Bayesian scheme described in Serenelli et al. (2013). Knowledge of the evolutionary stage of the target is implemented as a prior on the Bayesian probabilities that allows a precise determination of the stellar age. The asteroseismic mass of $\sim 1.5 M_\odot$ is consistent with a star that violently ignited helium in a flash. Our results are based in a set of isochrones not considering mass-loss in the RGB, and to include systematic effects particularly affecting the age estimate we have added in quadrature to the uncertainties the difference in the central values obtained with a set of isochrones using a Reimers (1975) mass loss rate of $\eta = 0.4$.

Mosser et al. (2012b) have determined mean core rotation periods in red giants observed by the Kepler satellite. Using modes of mixed character, the authors found that stars ascending the red giant branch slightly increase their core rotation periods, while clearly spinning down in the red clump phase. The rotational splitting $\delta \nu_{\text{rot}} = 50 \text{Hz}$ measured in KIC 5000307, corresponding to a period of $\sim 100$ days and common to many other clump stars, is evidence of no particularly fast core rotation (see Figures 6 and 7 in Mosser et al. 2012b). Similarly, the predicted mean envelope rotation from the modulation of the rotational splitting is much slower than the mean core rotation (see Goupil et al. 2013).

5. DISCUSSION

Combining asteroseismic analysis with classical spectroscopic observations, we have confirmed the first Li-rich giant quiescently burning helium in its core. The obtained mass, radius, age, and period spacing give a consistent picture of a star that has gone through the helium flash. In Figure 4, we show its position in the HRD, together with the sample recently identified by Kumar et al. (2011). It can be seen that the lack of asteroseismic information could have resulted in KIC 5000307 being mistakenly identified as a $\sim 2 M_\odot$ star in the RGB phase. Knowledge of the evolutionary stage of the target allows us to test mixing hypotheses for this star.

In order to enhance the surface Li abundance, the $^7\text{Be}$ isotope (produced in the inner H-burning regions) must be quickly transported by deep circulation to the cooler upper stellar layers before decaying into lithium (Cameron & Fowler 1971). Several physical processes have been envisaged as the non-canonical mixing mechanism capable of taking $^7\text{Be}$ to the stellar surface, such as rapid rotation (Drake et al. 2002), the interaction with a companion star (Denissenkov & Herwig 2004), or magneto-thermohaline mixing (Denissenkov et al. 2009).

In RGB stars, it is believed that the presence of a steep mean molecular weight gradient left behind by the bottom of the convective envelope at the end of the FDU prevents any extra mixing between the outer convective envelope and the hot layers where H burning is occurring. However, observations of the carbon isotopic ratio $^{12}\text{C}/^{13}\text{C}$ reveal that this quantity decreases below the canonical FDU value of $\sim 25$ once the star passes the RGB bump and evolves toward the RGB tip (see Lind et al. 2009b; Palmerini et al. 2011). This is a signature of some non-canonical mixing (such as thermohaline mixing, see Charbonnel & Lagarde 2010) associated with the H-burning shell advancing in mass outward and crossing the position of maximum convective penetration.
The low [C/Fe] and enhanced [N/Fe] of our target, combined with a likely non-canonical carbon isotopic ratio, describe a consistent picture of additional mixing after the FDU (possibly at the RGB bump). However, the aforementioned processes also predict a decrease in the Li abundance as the star evolves toward the RGB tip. Thus, even if lithium were produced at the RGB bump by these mechanisms, it should be depleted by the time the star ignites helium in its core.

Denissenkov (2012) proposed a mixing scenario in the RGB phase due to fast internal rotation and enhanced mixing across the radiative zone that could explain the lithium abundances and carbon isotopic ratios of the sample stars presented by Kumar et al. (2011). Under this prescription, after evolving past the RGB bump the star should zigzag in the HRD toward the RGB tip. Thus, even if lithium were produced at the RGB bump by these mechanisms, it should be depleted by the time the star ignites helium in its core.

A hypothesis first proposed by de La Reza et al. (1996) is that Li-rich giants can be identified using their far-infrared color properties. In this scenario, whatever mechanism is responsible for enhancing the surface Li abundance also produces the formation of a circumstellar shell of ejected material, thus affecting the photometric properties of these stars. The only far-IR photometry available for KIC 5000307 comes from the WISE catalog (Cutri et al. 2012), which unfortunately has a large uncertainty in the 22 μm W4 color and is not conclusive about the presence of IR excess. Further observations in this region of the spectrum would be highly valuable to validate this scenario.

Another possibility is that lithium appeared in the stellar surface as a consequence of non-canonical mixing during helium ignition. Previous studies of the He-flash in very metal-poor low-mass stars suggested that the convective zone produced by the huge energy release of He-burning could penetrate the overlying hydrogen-rich layers (see Schlattl et al. 2001 and references therein). The resulting inward migration of protons (H-injection) into high-temperature regions leads to a H-shell flash. As a consequence, when the convective envelope is deepening and merging with this H-flash driven convective zone, the surface is enriched with a large amount of matter that has been processed in hydrogen fusion.

The occurrence of H-injection has been found in hydrodynamic simulations of the core He-flash by Mocák et al. (2011) at solar metallicity. Mimicking the hydrodynamical simulations with a hydrostatic evolutionary code, the authors were able to reproduce H-injection in a 1 M⊙ model with the outcome of a strong pollution of the stellar envelope. Interestingly enough, the Mocák et al. (2011) results show that the surface 12C/13C drops below ∼10 while lithium is enriched to a level of 10 Li ∼ 3.7. However, they predict carbon enrichment in the surface by a factor of two or three, which we do not observe in KIC 5000307.

If non-canonical mixing during the He flash is responsible for the lithium enhancement, Li-rich core-helium burning stars should be mostly concentrated at the early clump stage since any enhancement occurred at the flash is expected to be depleted as the star evolves toward the AGB phase due to the deepening of the outer convection zone. Under that assumption, if KIC 5000307 has only recently gone through the helium flash, remnant processes of that episode could still be adjusting its interior structure and might be visible in the oscillation spectrum. In a future study, we will aim at detailed modeling of this target using individual frequencies of oscillations, as well as extending the sample of Li-rich giants in the Kepler field to get a more detailed view of Li-enrichment from the RGB to the AGB.

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