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Power colours: simple X-ray binary variability comparison

L. M. Heil,1 P. Uttley1 and M. Klein-Wolt2,3

1 Anton Pannekoek Institute, University of Amsterdam, Science Park 904, NL-1098 XH Amsterdam, the Netherlands
2 Department of Astrophysics, Research Institute of Mathematics, Astrophysics and Particle Physics, Radboud University Nijmegen, Heijendaalseweg 135, NL-6525 AJ Nijmegen, the Netherlands
3 Science & Technology, Olof Palmestraat 14, NL-2616 LR Delft, the Netherlands

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ABSTRACT

We demonstrate a new method of variability classification using observations of black hole X-ray binaries. Using ‘power colours’ – ratios of integrated power in different Fourier frequency bands – we can clearly differentiate different canonical black hole states as the objects evolve during outburst. We analyse (~2400) Rossi X-ray Timing Explorer observations of 12 transient low-mass black hole X-ray binaries and find that the path taken around the power colour–colour diagram as the sources evolve is highly consistent from object to object. We discuss how the consistency observed in the power colour–colour diagram between different objects allows for easy state classification based on only a few observations, and show how the power-spectral shapes can be simply classified using a single parameter, the power-spectral ‘hue’. To illustrate the benefits of our simple model-independent approach, we show that the persistent high-mass X-ray binary Cyg X–1 shows very similar power-spectral evolution to the transient black hole sources, with the main difference being caused by a combination of a lack of quasi-periodic oscillations and an excess of low-frequency power-law noise in the Cyg X–1 power spectra during the transitional state. We also compare the transient objects to the neutron star atoll source Aquila X-1, demonstrating that it traces a different path in the power colour–colour plot. Thus, power colours could be an effective method to classify newly discovered X-ray binaries.

Key words: X-rays: binaries – X-rays: general – X-rays: individual: Cygnus X-1 – X-rays: individual: Aquila X-1.

1 INTRODUCTION

The X-ray emission from black hole X-ray binaries shows distinct changes over time, interpreted as evidence that the structure of the accretion flow and the X-ray emitting regions close to the black hole evolve over the course of the outburst. For the transient, black hole low-mass X-ray binaries (LMXBs) this behaviour is classified into a series of common ‘states’, according to both the energy spectral and timing properties averaged over an observation (see reviews by Remillard & McClintock 2006; Belloni 2010b). LMXBs rise in flux out of quiescence in a hard state with the energy spectra dominated by strong power-law emission. This emission is interpreted as originating from Compton scattering in a hot, optically thin component close to the compact object (see e.g. Done, Gierliński & Kubota 2007). In the light curve, variability is observed over a wide range of different time-scales with sharp quasi-periodic features representing strong variations over a narrow range of frequencies and broader, less coherent components (see e.g. Wijnands & van der Klis 1999; Klein-Wolt & van der Klis 2008).

Both the energy spectra and the power spectra show evolution within the hard state as the X-ray flux from the source increases. The energy spectra become softer and emission from a blackbody-like component, associated with an optically thick accretion disc, becomes stronger. In addition, the peak time-scales for the power-spectral features shorten and the variability becomes concentrated in a narrower frequency range (0.1–10 Hz). This is typically close to the peak flux of the outburst and the variability becomes concentrated in a narrower frequency range (0.1–10 Hz). This is typically close to the peak flux of the outburst and is classified as the hard–intermediate state (HIMS). The reduction in time-scales has been interpreted as evidence for the inner edge of the disc moving closer to the black hole. As the disc inner edge moves closer, increased disc photon flux incident on the hot corona causes cooling thus making the power law steeper (see e.g. Done et al. 2007). At some point the broad-band variability disappears leaving dominant quasi-periodic features, often with a fundamental component at ~6 Hz and some weak variability on short time-scales – this marks the transition into a soft–intermediate state (SIMS) (see e.g. Belloni 2010a). Little evidence for distinct differences is found in the energy spectra between the HIMS and the SIMS (Dunn et al. 2010), even though these...
dramatic changes in the broad-band variability are observed. After this point the energy spectra are typically observed to be dominated by the disc blackbody component and the level of variability is very low ($\lesssim 3$ per cent fractional rms). The source flux decays whilst the object is in the soft state and the object eventually transitions back into a hard state before returning to quiescence. The spectral evolution of LMXBs over time is typically illustrated using a hardness–intensity diagram (Homan et al. 2001; Belloni 2004). In this simple plot, the ratio of hard to soft X-ray flux when compared to the total emission follows a hysteresis curve as a LMXB goes through an outburst (see e.g. fig. 1 in Klein-Wolt & van der Klis 2008).

Given the clear evolution of variability properties through the different states, it is interesting to see how timing-based classification can be used, not only to identify states but also to help identify the physical changes linked to the spectral and timing evolution. Recently, Muñoz-Darias, Motta & Belloni (2011) demonstrated that it is possible to describe the evolution of the power spectra using an rms–intensity diagram (Muñoz-Darias et al. 2011; Heil, Vaughan & Uttley 2012), where the rms is calculated as the square root of the integrated area underneath the power spectra. Although the rms–intensity plot clearly demonstrates the evolution of the total power, it does not contain information about the shape of the power spectrum which makes direct comparison between different states in objects difficult, since changes in rms due to changes in power-spectral shape cannot be disentangled from changes due to power-spectral normalization.

An alternative to simple rms variability measurement is to fit the power spectrum with an empirical model, typically with a series of Lorentzians of different widths or bending or broken power laws (see e.g. Nowak 2000). These models do not have a strict physical model underpinning them (although see Ingram, Done & Fragile 2009) but they describe the power-spectral shape fairly well. Common correlations are observed between the peak frequencies of the Lorentzians as the source evolves through different states (Wijnands & van der Klis 1999). In the hard and HIMSs all peak frequencies for the components are observed to increase together as the flux increases (see e.g. Axelsson, Borgonovo & Larsson 2005; Klein-Wolt & van der Klis 2008). However, there are some difficulties in fitting models which are not underpinned by physical descriptors. First, the definition of characteristic frequencies is clearly strongly dependent on the model chosen. Secondly, fitting is complicated in observations when the signal-to-noise ratio is low as it can be difficult to discern which features are present. Finally, whole-scale global changes to the power spectra, such as those commonly observed in transitions around the intermediate states, make identification of individual components difficult to track and identify. All of these features complicate attempts to make direct comparisons between different objects and states.

Therefore, to address the complications with current methods of classifying timing behaviour, and inspired by the use of energy spectral ‘colours’ or hardness ratios we develop in this paper a simple ‘colour-based’ approach to classifying power spectra, by defining ‘power colours’: ratios of integrated power (i.e. variance) measured over different Fourier frequency ranges. We apply this approach to the huge archive of observations of transient black hole LMXBs obtained by the Rossi X-ray Timing Explorer (RXTE). Power colours can be measured as long as the rms-variance can be measured in all set frequency bands, meaning that it is effective even in observations with low signal-to-noise ratio (i.e. even when the power is very low but just above the Poisson noise). We find that the different objects follow a distinct and consistent track in a power ‘colour–colour’ diagram and show that this track lends itself to a single-value parametrization of the power-spectral shape (and thus source state), which we call the ‘hue’ of the power spectrum. Finally, we demonstrate that the method allows for direct comparison of the behaviour of different classes of source, namely the LMXBs and the high-mass black hole X-ray binary Cyg X–1, as well as the accreting neutron star atoll source Aquila X–1, revealing intriguing similarities and differences with the black hole LMXBs.

### 2 DATA ANALYSIS

We use the extensive data base of RXTE observations of 12 transient low-mass black hole X-ray binaries and all observations of Cygnus X-1 and Aql X-1; the selected objects and number of observations for each are shown in Table 1. Power spectra were extracted from RXTE PCA (Proportional Counter Array) light curves binned up to 1/128 s with 512 segment size such that the lowest frequency was 1/512 s. The broad-band energy range was chosen to be that closest to 2–13 keV, dependent on the mode the observation was taken in and correcting the channels chosen for changes in the gain of the instrument over time. The lowest energy bin was excluded due to known calibration issues (Gleissner et al. 2004). Scientific event modes (i.e. Event and Good Xenon modes) were preferentially chosen due to their superior time resolution, but Scientific binned modes were used when these were not available. In general, measurements were taken for single observations, however both in the soft states and the low flux hard states where the variability is close to the Poisson noise level, it was necessary to combine the power spectra for multiple observations in order to reduce the error on the rms-variance measurements. These groupings were chosen carefully primarily on the basis of similarity in position on the HID (each group fell within a 5 per cent range in hardness) and time (within 5 d of each other) but we also avoided grouping observations with obvious differences in power-spectral shape or rms.

<table>
<thead>
<tr>
<th>Source name</th>
<th>Total obs.</th>
<th>Good obs.</th>
</tr>
</thead>
<tbody>
<tr>
<td>GX 339–4</td>
<td>619</td>
<td>457</td>
</tr>
<tr>
<td>XTE J1118+480</td>
<td>65</td>
<td>48</td>
</tr>
<tr>
<td>GS 1354−64</td>
<td>7</td>
<td>7</td>
</tr>
<tr>
<td>4U 1543−475</td>
<td>51</td>
<td>21</td>
</tr>
<tr>
<td>XTE J1550−564</td>
<td>268</td>
<td>201</td>
</tr>
<tr>
<td>XTE J1650−500</td>
<td>97</td>
<td>41</td>
</tr>
<tr>
<td>GRO J1655−40</td>
<td>438</td>
<td>312</td>
</tr>
<tr>
<td>XTE J1720−318</td>
<td>85</td>
<td>10</td>
</tr>
<tr>
<td>H 1743−322</td>
<td>451</td>
<td>179</td>
</tr>
<tr>
<td>XTE J1752−223</td>
<td>143</td>
<td>64</td>
</tr>
<tr>
<td>XTE J1817−330</td>
<td>129</td>
<td>26</td>
</tr>
<tr>
<td>XTE J1859+226</td>
<td>107</td>
<td>44</td>
</tr>
<tr>
<td>Cygnus X–1</td>
<td>1465</td>
<td>1414</td>
</tr>
<tr>
<td>Aquila X–1</td>
<td>529</td>
<td>238</td>
</tr>
<tr>
<td>Total</td>
<td>4454</td>
<td>3062</td>
</tr>
</tbody>
</table>

Table 1. (1) Source name; (2) total number of observations going into the sample; (3) ‘Good’ observations, defined as those where the variance is constrained as non-zero at a 3 $\sigma$ level in all four frequency bands. Combined observations are counted as a single observation.
Figure 1. Power-diagram for the ratios shown for all observations of the transient black holes included within our sample. All objects are individually colour-coded according to the legend. PC1 = variance in 0.25–2.0 Hz/0.0039–0.031 Hz and PC2 = variance in 0.031–0.25 Hz/2.0–16.0 Hz.

Four broad geometrically spaced and contiguous frequency bands were selected, each band spanning a factor of 8 in frequency. The bands cover the range from 0.0039 to 16.0 Hz and are defined as 0.0039–0.031 Hz, 0.031–0.25 Hz, 0.25–2.0 Hz and 2.0–16.0 Hz. The variance was measured in each of these bands by integrating under the power spectra and removing the Poisson noise contribution; errors were calculated in a standard manner based on van der Klis (1989) (see the appendix of Heil et al. 2012, for further details). The Poisson noise level was estimated through fitting the white noise within the power spectra in order to reduce systematic errors caused by deadtime effects in the detector.

Comparison of the rms–intensity diagrams in fig. 4 of Heil et al. (2012) reveals that sources in different states can have similar amounts of fractional rms over a broad frequency band. For this reason we choose to compare ratios of power, therefore evaluating the shape of the power spectra but removing the explicit dependence on a value describing the fractional rms which may vary from object to object. For example, the use of ratios of integrated power ensures that the effect of any constant component within the light curve, which will affect the observed variability amplitudes, is removed.

We then took the ratios between the variance values in different frequency bands effectively measuring power colours. The errors for each point were propagated in quadrature. The frequency bands are chosen to cover equal areas in log-space. Consequently, when the power spectra can be described by a power-law index of 1 across the chosen frequency bands, the power colour ratio will also be 1. The ratios were only calculated for points where the fractional variance was constrained at a 3σ level in all four frequency bands involved (these observations are classified as ‘good’ in Table 1).

3 RESULTS

3.1 Power colour analysis

The power colour–colour plot is a quick, simple way to evaluate the shape of the power spectra without requiring detailed power-spectral fitting. We go on to demonstrate that it has the unique property of showing a well-constrained shape for all black holes (BH) LMXBs allowing easy state classification for new sources with only a limited number of observations required.

Fig. 1 shows the power colour values for two particular frequency ratios, power colour ratio 1 (PC1) is defined as variance in 0.25–2.0 Hz/0.0039–0.031 Hz and ratio 2 (PC2) as variance in 0.031–0.25 Hz/2.0–16.0 Hz. These particular ratios not only compare all four broad frequency bands used in the initial analysis, making the most of the available data, but are also separated in frequency. This plot is colour coded according to object, and the similarity in powerspectral evolution throughout outbursts between different sources can be clearly observed through the similarity in tracks traced. Log axes are used in order to highlight small changes in power-spectral shape during source evolution.

The evolution of objects around the power colour–colour diagram during an outburst takes the form of a loop and can be simply described in terms of a single dimension by measuring an azimuthal coordinate of each data point (plotted in terms of the logarithm of the power colours), relative to some central point. The central point used is the mid-point of the x and y extent of the plot, with position [4.51920, 0.453724] in linear units. The semi-major axis is defined at a 45° angle to the x and y-axes. By analogy with the ‘colour-wheel’ representation of colours, we call this angle around the
power colour–colour plot the ‘hue’. To simplify the classification we group the power-spectra according to their hue, splitting the hue angles into 18 sectors, each corresponding to 20° of arc which are indexed on Fig. 2(b); example power spectra for each segment are given both in Fig. 2(c) and the Appendix. The ranges of hue position for the different states are given in Table 2. For an in-depth discussion of power-spectral evolution for multiple objects over the course of an outburst, see Klein-Wolt & van der Klis (2008).

Defining the radial central point in the power colour diagram in an objective way is difficult due to the radial scatter and relative sampling of different states, with most of the measurements being obtained from the longer lasting ‘pure’ hard and soft states which
This is particularly likely to affect the intermediate states, with strong frequency shifting or particular components appearing/disappearing rapidly within an observation. In this case there may be some intrinsic variation which is not fully captured by the current error bars on points within the power colour–colour plot. However, by testing for intrinsic variation in shape within a large sample of single observations, Heil et al. (2012) demonstrated that this only appears to affect a small sub-set of power spectra from LMXBs. Whilst we could have measured single power colour points per orbit rather than per observation in order to limit this effect, this lowers the signal-to-noise ratio and thus increases the errors on each point. These points are then typically coincident in the diagram anyway, so we choose to continue averaging the power spectra over all orbits within an observation.

We now go on to discuss the position of observations in specific states within the power colour–colour diagram with particular reference to how the gradual evolution of the power spectra causes a loop around the parameter space.

**Hard states:** In this state the variability is strong and visible over a wide range of frequencies, as such the power spectra consist of very broad approximately flat-topped noise, an example of which is shown in Fig. 2(c) plot 1. Due to the broad nature of the noise, the power colour ratios are both approximately 1.0. Consequently, the points appear in the upper left hand corner of the power colour diagram. This is illustrated by the yellow shaded region in Fig. 2(a). As the flux from the source increases as the outburst progresses, the strongest variability becomes concentrated into a smaller frequency range; this behaviour can be seen in power spectra 2–7 in Fig. 2(c). Due to the more peaked nature of the broad-band noise objects transition along the light grey path illustrated in Fig. 2(a) until they reach the HIMs, we note that there is gradual evolution in the power spectra from the hard state into the HIMs with no dramatic differences between the two. The difference is more clearly observed in the energy spectra, which now start to become softer.

**Hard-intermediate states:** In the HIMs (PSDs 8–11 in Fig. 2c) the power spectra reach their most strongly peaked shape, with most power concentrated in the two highest frequency bands. Due to the fact that this is the most strongly concentrated that the broad-band noise becomes, these observations can be found in the bottom right-hand corner of the power colour–colour diagram. The type C quasi-periodic oscillations are particularly strong within this state and, as they are found with a range of different amplitudes in various objects, they can likely explain some of the spread which can be observed in the power colour–colour diagram in this state. We explore this point further in Heil, Uttley & Klein-Wolt (2014).

**Soft-intermediate states:** The transition to the SIMs is characteristically only strongly observed in the variability from the source. The peaked broad-band noise disappears, leaving only type-B QPOs and weak low-frequency power-law-like noise. There is huge variation in variability properties within this state, some of which is illustrated in the power spectra for positions 12–15 in both Fig. 2(c) and the Appendix. Evolution within the SIMs is not straightforward. Sources may pass through the SIMs multiple times, transitioning into a soft state or returning to a HIMs. Consequently objects do not evolve around the SIMs part of the loop marked out in Fig. 2(a) in a linear fashion, but typically pass through this region very rapidly as they transition. The variety in power-spectral shapes within the SIMs explains the wide distribution of points in the power colour–colour diagram within this region. However, we can still use the power colour–colour diagram to interpret what individual observations may look like based on their measured hue. All observations in this state contain a strongly peaked high-frequency component.
with the majority of the power above $\sim 1$ Hz. They also contain a weak power-law-like component at low frequencies. The stronger this low-frequency component the further clockwise observations can be found on the power colour–colour diagram. Fig. 2(c) illustrates this point; moving from position 13 to 15 a large difference in relative normalizations between the low- and high-frequency components can be observed.

**Soft states:** the soft states typically show low-frequency variability with an approximate power law of $-1$ bending to a steeper power law at high frequencies ($\sim 5$ Hz); examples can be found in plots 17 and 18 in Fig. 2(c). Although the strength of the variability in this state is much lower than that of the hard states ($\sim 40$ versus $<3$ per cent) the power colour method removes any dependence on power-spectral normalization, consequently most observations in the soft states appear in an identical area of the power colour–colour diagram to the hardest hard states (Fig. 2c, plot 1). This overlap is in the yellow shaded region in Fig. 2(a). The states of these observations can be easily separated by either total rms or hardness and consequently their true states are easily identifiable.

Once objects begin the return back to the hard state they follow an identical path around the power colour–colour diagram but in an anti-clockwise direction. This is demonstrated in Fig. 4 where all observations follow a very well constrained path both into and out of outburst. Simply put, objects initially transition around the power colour–colour diagram in a clockwise direction from position 1 to 18 in Fig. 2(b), and then return back towards quiescence in an anti-clockwise direction (although the exact path found for each object is clearly dependent upon the range of states sampled by the observations during the outburst).

Other frequency bands or ratios will clearly trace different paths around the power colour–colour plot. Not all paths are necessarily loops, but these particular bands have been selected in order to sample the broad-band changes in the power spectra of black hole X-ray binaries using observations taken with RXTE. The advantage in the ratios used here is that almost all states can be cleanly separated, which is not necessarily the case when other frequency bands are used. In addition we note that the power spectra observed from black hole X-ray binaries are known to show some dependence on energy; this is particularly visible in the intermediate states, with high-frequency variability being stronger at harder energies and some low-frequency noise stronger in the soft bands (see i.e. fig. 6 in Gierliński & Zdziarski 2005; Wilkinson & Uttley 2009). As such the power colour–colour plot will show some dependence on the energy band taken. The current energy band used (2.5–13 keV) is fairly broad, but narrower soft bands may show some shift, for instance in a harder energy band power spectra in the HIMS become even more strongly peaked and consequently the points around $180^\circ$ of hue will shift both farther to the right and down.

Using the assigned 20° hue ranges we can also bin the power colour–colour diagram according to source in order to test how similar the power colour–colour paths are for different objects. It also allows us to identify whether the spread observed in the lower-left-hand corner is purely due to a spread between different objects. Fig. 5 shows the results for all transient black holes studied. Although single objects show significant ‘radial’ scatter in the power colour–colour diagram their mean paths are clearly similar with most points falling within 3 sigma of each other (for the standard error on the mean). There are some notable differences, XTE J1550–564 and H1743–322 appear to follow a different path out of the hard state than other objects. In Heil et al. (2014) we demonstrate that this deviation is caused by the particularly strong QPOs present in the light curves from these sources. Direct comparison of the power spectra at positions 6, 7 and 8 (see both Fig. 2c and the Appendix) shows that these features are clearly much stronger in these objects in this particular state than in GX 339–4. Motta et al. (2014) demonstrate that this difference appears to be due to the inclination angles of the sources: XTE J1550–564 and H1743–322 have higher inclinations than GX 339–4 and stronger QPO amplitudes.

We can also see this deviation, caused by the QPOs, in Fig. 3. In this plot we demonstrate the true power of the power colour–colour diagram in that it allows us to directly compare the parameters of observations with similar power-spectral shape in a model-independent manner. Fig. 3 shows that there is a slight shift in hue for some observations of XTE J1550–564 and H1743–322 between $\sim 150^\circ$ and $200^\circ$, although this shift is not universally observed for all observations. We explore this point in more detail in Heil et al. (2014).

Although outbursts of individual sources follow constrained paths around the power colour–colour diagram, overall the paths are very similar and the power colour–colour diagram is strikingly consistent for a wide range of objects. It is already well known that the
evolution of characteristic frequencies in the power spectrum is highly consistent both between different outbursts of the same source and outbursts of many objects (Klein-Wolt & van der Klis 2008). However, the similarity in power colour evolution further demonstrates in a model-independent way that the relative frequencies and amplitudes of power-spectral components must vary in a consistent way.

3.2 Outlying points

Since the paths followed by all objects as they transition around the power colour–colour diagram are so consistent we should be able to identify unusual observations as they are likely to be ‘out of place’ relative to other points of similar spectral hardness. In order to test this hypothesis, we identify the position of three observations of H1743–322 within its power colour–colour diagram, two of these observations have very unusual low frequency (∼0.01 Hz) QPOs (Altamirano & Strohmayer 2012), which have only been observed from this source, and the third is known to show a hard dip (Homan et al. 2005), which is associated with the high inclination of the system to our line of sight. In all of these cases there is additional power in the lowest frequency band relative to other observations in a similar state. Fig. 6 shows the power colour–colour diagram of H1743–322 with these points identified. As this object is relatively faint there are few observations when the power in well constrained in all bands during the low–hard state, so the full evolution is not visible. However the colour-coding shows that these three observations are clearly separated from their counterparts with similar measures of hardness, the value of the ratio PC1 (containing the lowest frequency band) is smaller. This demonstrates that the power colour–colour diagram can be used to identify power spectra with unusual features without the detailed fitting and investigation often required to find them when they would not be easily identifiable using other methods (see Altamirano & Strohmayer 2012, for a discussion of the points with low-frequency QPOs within the HID).

3.3 Exploring the comparison of other object types to the black hole transients

Power colours can be measured for a range of different accreting objects, from neutron star X-ray binaries and cataclysmic variables to active galactic nuclei. The consistency observed in the power colour–colour diagram for the transient low-mass black holes suggests that it may be a suitable method to classify new objects, particularly given the low levels of signal-to-noise ratio required to measure a power colour ratio. We now go on to make initial explorations into how other X-ray binary types’ power colour–colour diagrams compare to that of the transients. Specifically we make comparisons to a high-mass black hole X-ray binary Cygnus X-1 and an atoll type low-mass neutron star X-ray binary Aquila X-1.

3.3.1 Comparing the transient black holes to Cygnus X-1

Fig. 7 shows the power colour–colour plot for all observations in the LMXB sample (black), with the power colours for the black hole high-mass X-ray binary (HMXB) Cygnus X-1 over-plotted (purple). The example power spectra from observations of Cyg X-1 are shown in the Appendix. The x-axis extent of the power colour–colour path for Cyg X-1 is slightly different to that of the transients (discussed further below); however the overall shape is very similar.

The power colour–colour plot of Cyg X-1 appears to indicate that the power spectra of this source show all states observed in the LMXBs except the most extreme low-flux hard states with very broad noise. As such it is possible to discern between the hardest and softest states at the top of the power colour–colour diagram. Energy spectral analysis suggests that Cyg X-1 does not appear to ever reach a true soft state (Belloni 2010b), where the energy spectrum is dominated by a disc blackbody component and only a weak power-law tail is present. However, the power colour–colour diagram demonstrates that the shapes of the power spectra are similar to those observed from the LMXBs (see power spectra from regions 17 and 18) even though the total variance (i.e. the normalization of the power spectra) is higher for Cyg X-1 in the softest states. It is clear that the PC2 (0.031–0.25 Hz)/(2–16 Hz) power colour of Cyg X-1 rarely reaches values greater than 1.5. This is in contrast to the transient objects, which show observations with ratio values up to 2.3 in the hard state and 3.7 in the soft. Upon further examination we find that the soft state observations from
the LMXBs with y-extent values on the power colour–colour diagram greater than those observed in the soft-state observations of Cyg X–1 are all from an unusual state of GRO J1655–40 which is not observed from Cyg X–1; they are discussed further in Uttley & Klein-Wolt (in preparation). Since Cyg X–1 does not reach the broad low-luminosity hard states observed from the BH transients we neglect these states from the discussion below. This point also illustrates why the power colour–colour loop for Cyg X–1 begins at a larger hue than the transient objects; there are very few in the hard state with the very broad-band noise observed in PSDs from position 1 in Fig. 2(b).

Although the shape of the power colour–colour loop followed by Cyg X–1 is similar to that of the transient black holes we observe that there is some deviation between the parameter space covered by the two loops in the HIMS. The LMXB points are clearly delineated from those of Cyg X–1, moving further down and to the right. In order to create this deviation either the variance at low frequencies must be higher in Cyg X–1 than the transients or there must be higher variance in the LMXBs at high frequencies. As these are the states in which the QPOs are strongest and are in fact the dominant feature within the power spectra, the obvious explanation is that this deviation is primarily caused by the extra variance present in these coherent features, which are not visible in the light curves from Cyg X–1. However inspection of the power spectra reveals that this is not necessarily the complete explanation. Comparing the power spectra of Cyg X–1 in locations 10, 11 and 12 to those of the transients reveals that when the amplitude of the broad-band noise under the QPOs in the highest two frequency bands is similar to that of Cyg X–1, there is consistently more power at low frequencies in Cyg X–1 than in the transients. Axelsson et al. (2005) found that an additional power law was required at low frequencies in order to fit the power spectra of Cyg X–1 in these states, indicating that the deviation between the power colour–colour path of Cyg X–1 and the transients is likely to be due to the stronger low-frequency power-law noise in Cyg X–1 together with the presence of the QPOs in the LMXBs.

### 3.3.2 Comparing the transient black holes to Aql X–1

It has been shown in previous work that the power spectra measured from observations of neutron stars show distinct similarities to those of the black holes (Wijnands & van der Klis 1999; Sunyaev & Revnivtsev 2000; Belloni, Psaltis & van der Klis 2002; van der Klis 2006; Klein-Wolt & van der Klis 2008), with the neutron stars showing an additional high-frequency component extending the broad-band noise above 100 Hz. We therefore take all observations of the atoll source Aql X–1 to make a brief comparison between different objects independent of the total amplitude of the different states. This method also allows for direct comparison between different objects in terms of power colours – ratios of integrated power (rms-variance) in different frequency bands – we can evaluate how the power spectra evolve between different states. This method also allows for direct comparison between different objects independent of the total amplitude of the rms-variability within the light curve. We have found the following.

The differences between the black hole X-ray binaries and the neutron star Aql X–1 in the power colour–colour diagram suggest that this diagram can be used as a diagnostic tool, aiding in identifying whether newly identified objects are neutron stars or black holes in a fast manner.

### 4 CONCLUSIONS

We have demonstrated a simple, model-independent method for evaluating the shape of power spectra. By comparing power colours – ratios of integrated power (rms-variance) in different frequency bands – we can evaluate how the power spectra evolve between different states. This method also allows for direct comparison between different objects independent of the total amplitude of the rms-variability within the light curve. We have found the following.

(i) The power colour–colour diagram shows that, across the LMXBs the power spectra evolve over an outburst in a very similar manner. If the power-spectra are represented in terms of multiple
components (e.g. Klein-Wolt & van der Klis 2008), these must evolve similarly in both frequency and amplitude.

(ii) The general evolution can be simply parametrized in terms of the angle along the power colour–colour track, which we call the ‘hue’. For a given power-spectral hue there remains significant scatter in power colours in the intermediate states for a given object, but the mean power colours are similar for different objects, again indicating a common evolution.

(iii) We can use the power colour–colour diagram to identify unusual or ‘out of place’ observations. These include high-inclination objects with dips in their power spectra and additional low-frequency features.

(iv) The overall shape of the power colour–colour track shown by the persistent black hole HMXB Cyg X−1 is similar to that seen in the transient LMXBs. Cygnus X-1 shows nearly all the broad power-spectral shapes observed in the LMXBs excepting those seen in the low-luminosity hard states which Cyg X−1 never reaches.

(v) The power colours of Cygnus X-1 differ from those observed in the LMXBs when the objects are in HIMSs. The difference is due to a combination of stronger low-frequency power-law noise in the Cyg X−1 power spectra, combined with the presence of strong QPOs in the LMBX power spectra.

(vi) We find that parts of the power colour–colour diagram of the Atoll source Aql X−1 differs from that of the black holes, suggesting that this method could provide a simple new tool to help identify the nature of newly discovered X-ray binaries.

The power colour and power-spectral hue is a purely empirical classification, but due to its simplicity and model-independence, it should prove extremely useful in comparing timing evolution with other observables, which will help to unlock the physical origin of the different states. Furthermore, the broad frequency bands used to determine the power colours can be used for fainter sources where detailed power-spectral measurements may not be possible. Assuming that the Poisson noise power dominates the source power in this range, the minimum rms-squared variance for a 3σ detection of variability in a frequency band of width Δν in an observation of length T_{obs} is:

$$\sigma_{\text{min}}^2 = \frac{6(S + B)}{S^2} \frac{\Delta \nu}{T_{\text{obs}}}$$

where S and B are the source and background photon count rates, respectively. Thus for the 2–16 Hz band, where we expect the lowest signal-to-noise ratio, we expect in 10–100 ks to detect variability from sources through the hard and intermediate states, down to count rates ~10 count s$^{-1}$ (assuming negligible background). This will enable the power colours used here to be measured for bright X-ray binaries in nearby galaxies with current imaging X-ray telescopes. It would also be interesting to extend this technique to study other types of variable object.

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REFERENCES


APPENDIX A

Example power spectra for each of the 18 positions illustrated in Fig. 2(b). Power spectra were chosen randomly from the observations of transient black holes in order to illustrate the range of power-spectral shapes observed from the black hole X-ray binaries. Power spectra of Cyg X−1 are also shown for each position. In the position where hard and soft states overlap (18) hard state spectra are labelled HS and soft state SS.