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The CHESS spectral survey of OMC-2 FIR 4*

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ABSTRACT

Context. The freezeout of gas-phase species onto cold dust grains can drastically alter the chemistry and the heating-cooling balance of protostellar material. In contrast to well-known species such as carbon monoxide (CO), the freezeout of various carriers of elements with abundances <10−3 has not yet been well studied.

Aims. Our aim here is to study the depletion of chlorine in the protostellar core, OMC-2 FIR 4.

Methods. We observed transitions of HCl and H2Cl+ towards OMC-2 FIR 4 using the Herschel Space Observatory and Caltech Submillimeter Observatory facilities. Our analysis makes use of state of the art chlorine gas-grain chemical models and newly calculated HCl-H2 hyperfine collisional excitation rate coefficients.

Results. A narrow emission component in the HCl lines traces the extended envelope, and a broad one traces a more compact central region. The gas-phase HCl abundance in FIR 4 is 9×10−11, a factor of only 10−3 that of volatile elemental chlorine. The H2Cl+ lines are detected in absorption and trace a tenuous foreground cloud, where we find no depletion of volatile chlorine.

Conclusions. Gas-phase HCl is the tip of the chlorine iceberg in protostellar cores. Using a gas-grain chemical model, we show that the hydrogenation of atomic chlorine on grain surfaces in the dark cloud stage sequesters at least 90% of the volatile chlorine into HCl ice, where it remains in the protostellar stage. About 10% of chlorine is in gaseous atomic form. Gas-phase HCl is a minor, but diagnostically key reservoir, with an abundance of ≲10−10 in most of the protostellar core. We find the [13Cl]/[17Cl] ratio in OMC-2 FIR 4 to be 3.2±0.1, consistent with the solar system value.

Key words. stars: formation – astrochemistry – ISM: abundances

1. Introduction

The freezeout and desorption of volatiles are amongst the key factors that determine the gas-phase abundances of chemical species in interstellar gas, in particular in protostellar cores (e.g. Bergin & Langer 1997; Caselli & Ceccarelli 2012). The classical example is CO, which is strongly depleted in prestellar cores, but returns to the gas in the warm protostellar stages Class 0 and I (e.g. Caselli et al. 1999; Bacmann et al. 2002; Jørgensen et al. 2005). Similar behaviour, with quantitative differences, is seen or expected for many other species. Here, we investigate the freezeout and main reservoirs of chlorine in a protostellar core.

An ice mantle on a dust grain in a protostellar core mainly consists of a few dominant ice species (e.g. H2O, CO, CO2, CH3OH). However, the ices also contain a large number of minor species, accreted from the gas and formed in the ice matrix. Such minor species very likely include HF and HCl, which are the main molecular carriers of the halogen elements fluorine and chlorine, in well-shielded molecular gas. With increasing temperature, such as seen during infall towards a central protostar, the ice mantle desorbs. For any given ice species, the desorption may have multiple stages, with details influenced by its abundance and interaction with the other species in the ice (Collings et al. 2004; Lattelais et al. 2011). In the hot cores of protostars, where ice mantles largely desorb, most volatile species are expected to be back in the gas phase.

We study the depletion of volatile chlorine towards the OMC-2 FIR 4 protostellar core, using HCl and H2Cl+ as proxies.

* Appendices are available in electronic form at http://www.aanda.org
We employed the Herschel Space Observatory\textsuperscript{1} and the Caltech Submillimeter Observatory (CSO) telescopes and made use of both an updated gas-grain chemical network for chlorine and new calculations of the HCl-H\textsubscript{2} hyperfine collisional excitation rate coefficients. In Sect. 2, we review the interstellar chemistry of chlorine. The source is described in Sect. 3 and the observations in Sect. 4. The analysis separately covers HCl (Sect. 5) and H\textsubscript{2}Cl\textsuperscript{+} (Sect. 6), and the results are discussed in Sect. 7. We present our conclusions in Sect. 8.

2. Interstellar chlorine chemistry

Chlorine has a simple and relatively well characterized interstellar chemistry (e.g. Jura 1974; Dalgarno et al. 1974; Blake et al. 1986; Schilke et al. 1995; Neufeld & Wolfire 2009). In dense molecular gas, the gas-phase formation of HCl begins with the reaction

\begin{equation}
\text{Cl} + \text{H}_2 \rightarrow \text{HCl} + \text{H}.
\end{equation}

Dissociative recombination of H\textsubscript{2}Cl\textsuperscript{+} with e\textsuperscript{−} then produces HCl. A recent study of D\textsubscript{2}Cl\textsuperscript{+} has confirmed the branching ratio of this recombination to be \textasciitilde10\% into HCl, \textasciitilde90\% into Cl (Novotný et al. 2012). Below \textasciitilde100 K, the resulting fraction of chlorine in HCl is \textasciitilde0.1 to 0.3 (Dalgarno et al. 1974; Blake et al. 1986; Schilke et al. 1995; Neufeld & Wolfire 2009). At temperatures of \textasciitilde350 K, all chlorine can be converted into HCl in dense (\textasciitilde10\textsuperscript{6} cm\textsuperscript{-3}) molecular gas by the reaction

\begin{equation}
\text{Cl} + \text{H}_2 \rightarrow \text{HCl} + \text{H},
\end{equation}

on a timescale of \textasciitilde10\textsuperscript{4} yr. At colder temperatures, this reaction is much slower.

In photon-dominated regions (PDRs), a layer can form where the reactions Cl\textsuperscript{+} + H\textsubscript{2} \rightarrow HCl\textsuperscript{+} + H and HCl\textsuperscript{+} + H \rightarrow H\textsubscript{2}Cl\textsuperscript{+} + H lead again via dissociative recombination to HCl. In diffuse PDRs (n\textsubscript{H} \sim 10\textsuperscript{5} cm\textsuperscript{-3}), the HCl and H\textsubscript{2}Cl\textsuperscript{+} column density ratio is \textasciitilde1, while for 10\textsuperscript{7} cm\textsuperscript{-3} it is \textasciitilde10\textsuperscript{2} (Neufeld & Wolfire 2009). At AV \textasciitilde1 however, the ratio can be \textasciitilde0.01, while atomic Cl is the dominant gas-phase reservoir of chlorine. At low densities and AV \textasciitilde1, Cl\textsuperscript{+} dominates.

The abundance of Cl with respect to atomic hydrogen in the solar photosphere is [Cl]/[H] = 3.1\times10\textsuperscript{−7} (Asplund et al. 2009). The meteoritic abundance is 1.8 \times 10\textsuperscript{−7} (Lodders 2003), while in the diffuse interstellar medium (ISM) it is 10\textsuperscript{−7}, indicating a factor of two depletion into refractory grains or volatile ices (Mooney et al. 2012). For our modelling, we refer to the abundance with respect to molecular hydrogen, X(Cl) = [Cl]/[H\textsubscript{2}], and adopt a reference value X(Cl) = 10\textsuperscript{−7}.

The main stable isotopes of chlorine are 33Cl and 35Cl. Their ratio in the solar system is 3.1 (Lodders 2003).

3. The FIR 4 protostellar core

Our target is a nearby intermediate-mass protostellar core, OMC-2 FIR 4\textsuperscript{2} (hereafter FIR 4), located in Orion, at a distance of d \sim 420 pc (Hirota et al. 2007; Menten et al. 2007). Its envelope mass is \textasciitilde30 M\odot and its luminosity within \textasciitilde20″ is \textasciitilde100 L\odot, while the \textasciitilde1″ scale envelope has an estimated luminosity of \textasciitilde400–1000 L\odot (Mezger et al. 1990; Furlan et al. 2014; Crimier et al. 2009, the latter is hereafter referred to as C09). Class 0 suggests an age of \textasciitilde10\textsuperscript{5} yr, and Furlan et al. (2014) propose that FIR 4 is amongst the youngest protostellar cores known. In the far-infrared to millimetre regimes, FIR 4 is undergoing intense study in the Herschel key programmes CHESS (Ceccarelli et al. 2010; Kama et al. 2010, 2013; López-Sepulcre et al. 2013a,b; Ceccarelli et al. 2014) and HOPS (Adams et al. 2012; Manoj et al. 2013; Furlan et al. 2014), as well as in a number of ground-based projects.

The main physical components of FIR 4 are the warm, clumpy inner envelope and the cold, extended outer one: a proposed outflow; and a tenuous, heavily irradiated foreground cloud (Crimier et al. 2009; Kama et al. 2013; López-Sepulcre et al. 2013a,b; Furlan et al. 2014). The inner envelope has been resolved into continuum peaks with different luminosities (Shimajiri et al. 2008; Adams et al. 2012; López-Sepulcre et al. 2013b; Furlan et al. 2014). These sources appear to share a large envelope. The line profiles and excitation of CO and H\textsubscript{2}O show that FIR 4 harbours a compact outflow (Kama et al. 2013; Furlan et al. 2014). There are also two PDRs to consider: the dense outermost envelope (n\textsubscript{H} \sim 6 \times 10\textsuperscript{5} cm\textsuperscript{-3}, Crimier et al. 2009), and the tenuous foreground cloud (n\textsubscript{H} \sim 10\textsuperscript{7} cm\textsuperscript{-3}, López-Sepulcre et al. 2013a). Whether or not they are physically connected is unclear. For simplicity, we treat them as separate entities, even though there may be a smooth transition in physical conditions between the two.

4. Observations

The main neutral and ionized molecular carriers of chlorine in the ISM – such as HCl, HCl\textsuperscript{−} and H\textsubscript{2}Cl\textsuperscript{−} – have strong radiative transitions, although their observations from the ground are hampered by atmospheric water absorption. The observations of HCl and H\textsubscript{2}Cl\textsuperscript{−} towards FIR 4, summarized in Table 1 and shown in Figs. 1 (HCl) and 2 (H\textsubscript{2}Cl\textsuperscript{−}), were carried out with the Herschel Space Observatory and the CSO.

4.1. Herschel/HIFI

As part of the CHESS key programme (Ceccarelli et al. 2010), FIR 4 was observed with the Heterodyne Instrument for the Far-Infrared (HIFI) wide band spectrometer (de Graauw et al. 2010) on the Herschel Space Observatory (Pilbratt et al. 2010), in Dual Beam Switch (DBS) mode at a resolution of dv = 1.1 MHz (R \textasciitilde 10\textsuperscript{6}). Multiple transitions of HCl and H\textsubscript{2}Cl\textsuperscript{−} were covered. The data quality and reduction, carried out with the HIPE 8.0.1 software (Ott 2010), are presented in Kama et al. (2013).

4.2. Caltech submillimeter observatory

Ground-based observations of HCl J = 1–0 were carried out in DBS mode with a chopper throw of 240″, using the 690 GHz facility heterodyne receiver of the CSO, on Mauna Kea, Hawaii, on February 6th and 11th, 2013. The atmosphere was characterized by a 225 GHz zenith opacity of 0.04–0.06 or 1 mm of precipitable water vapour. Typical single sideband system temperatures were 5000–7000 K. The backend was the high-resolution FFT spectrometer, with 4095 channels over 1 GHz of IF bandwidth. The on-source integration time was 51 min, resulting in an RMS noise of \textasciitilde0.15 K at a resolution of 3.7 km s\textsuperscript{−1}. Jupiter was used for pointing and calibration. The beam efficiency was 40%, assuming a 148.5 K brightness temperature for Jupiter.

\textsuperscript{1} Herschel is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.

\textsuperscript{2} Identified on SIMBAD as [MWZ90] OMC-2 FIR 4.
4.3. Overview of the data

Both HCl and H$_2$Cl$^+$, as well as their isotopologs, are detected with Herschel/HIFI. The HCl 1–0 transition is also detected with the CSO. The measured line fluxes, integrated over the hyperfine components, are summarized in Table 1.

The HCl lines show evidence for a broad and a narrow component. Both components are present in the HIFI HCl 1–0 data, which also shows a roughly optically thin hyperfine component ratio (2:3:1, in order of increasing $\nu$) for the narrow component. The signal to noise of the HCl 2–1 and the CSO 1–0 data are insufficient to make firm conclusions about the relative importance of the broad and narrow components, however the broad component seems to contribute substantially to both observations. Unfortunately, the H$^{35}$Cl 2–1 transition is contaminated by a CH$_3$OH transition and is therefore excluded from our analysis.

5. Analysis of HCl

Here, we first disentangle the kinematical components of the hydrogen chloride lines. Then, we use radiative transfer and chemical modelling to determine the gas-phase HCl abundance. The outermost envelope of FIR 4 is strongly externally irradiated, and is considered separately in Sect. 5.4.

5.1. Kinematics

A two-component model of the hyperfine structure in the HCl 1–0 observations strongly constrains the kinematic parameters of the broad and narrow components, given in Table 2.

In Fig. 3, we compare the kinematical components of HCl to those of other species from the FIR 4 HIFI spectrum of Kama et al. (2013). The properties of the narrow component match the large-scale quiescent envelope tracers, such as C$^{18}$O. The broad
Fig. 2. Observed spectra of $\text{H}_2\text{Cl}^+$ from Herschel/HIFI. Only the lines formally detected in absorption are shown, and are velocity-corrected and centred on the strongest hyperfine component. Red dashed lines indicate the hyperfine components of the rotational transitions. The continuum has not been subtracted. See also Fig. B.1.

Fig. 3. Kinematical properties of HCl and $\text{H}_2\text{Cl}^+$ compared to other species in FIR 4, each of which is annotated with the main physical component it traces. The label “PDR” refers here to the tenuous foreground cloud discovered by López-Sepulcre et al. (2013a). For the comparison species, the mean Gaussian fit parameters from Kama et al. (2013) are plotted.

Table 2. Kinematical properties of the observed HCl line profiles.

<table>
<thead>
<tr>
<th>Species</th>
<th>Component</th>
<th>$\nu_{lsr}$ [km s$^{-1}$]</th>
<th>FWHM [km s$^{-1}$]</th>
</tr>
</thead>
<tbody>
<tr>
<td>HCl</td>
<td>Narrow</td>
<td>11.4 ± 0.2</td>
<td>2.0 ± 0.2</td>
</tr>
<tr>
<td>HCl</td>
<td>Broad</td>
<td>12.2 ± 0.2</td>
<td>10.5 ± 0.5</td>
</tr>
</tbody>
</table>

component parameters lie between those of the envelope and the outflow tracers, and match the mean properties of CS.

Contrary to HCl, for which exceptional observing conditions or space observatories are required, isotopologs of CS are readily observed with ground-based instruments. In Fig. 4, we show a map of the C$^{18}$S 3–2 line width, based on Plateau de Bure Interferometer data from López-Sepulcre et al. (2013b). A spatially resolved region $\sim 10''$ across (2100 AU radius) and detected at a 3σ confidence at $F_{C^{18}S} \geq 0.4$ Jy km s$^{-1}$ beam$^{-1}$ has a line width of $\geq 6$ km s$^{-1}$. The spatial resolution of the data is $\sim 5''$, the channel width is 4 km s$^{-1}$, and the flux loss compared to single-dish data is $\sim 60\%$, so we can only conclude that if the C$^{18}$S broad component is equivalent to that of HCl, the latter is centrally peaked and extended on several thousand AU scales.

Fig. 4. Line width map of C$^{18}$S 3–2 (colour map), obtained with the Plateau de Bure Interferometer (López-Sepulcre et al. 2013b). The channel width of the data is 4 km s$^{-1}$. Overplotted are continuum contours (black lines), which start at 3σ with a 3σ step, with 1σ = 0.14 Jy km s$^{-1}$ beam$^{-1}$. North is up and east is left.

Based on the arguments presented above, we attribute the narrow and broad HCl line profile component to the outer envelope and some compact yet dynamic inner region, respectively. Given the large difference in line width between the broad HCl component and the outflow tracers CO and H$_2$O, it is unlikely that they trace the same volumes of gas.

5.2. The gas-phase HCl abundance in FIR 4

To obtain an abundance profile, we modelled the HCl excitation and radiative transfer in FIR 4 using the Monte Carlo code, Ratran$^3$ (Hogerheijde & van der Tak 2000) as described below.

5.2.1. New HCl-H$_2$ collisional excitation rates

We modelled the HCl emission using new hyperfine-resolved collisional excitation rate coefficients, presented in detail in Appendix A. These are based upon the recent potential energy surface and rotational excitation rate coefficients of Lanza et al. (2014b,a), Neufeld & Green (1994) and Lanza & Lique (2012). The latter, scaled to H$_2$ collisions by a mass correction factor of 1.38, differ from the new coefficients by a factor of a few at $T_{\text{kin}} > 50$ K, and by around a factor of ten at $T_{\text{kin}} < 50$ K. A similar difference has been found for HF (Guillon & Stoecklin 2012). We discuss the impact of the new excitation rates on HCl depletion estimates in Sect. 7.5.

5.2.2. The source model

As the basis of our modelling, we adopted the spherically symmetric large-scale source structure with no enhancement of the external irradiation field ($G_0 = 1$ interstellar radiation fields, ISRF) from C09. The density and temperature profiles, as well as the relative populations of the relevant HCl rotational levels, are shown in Fig. 5. Due to the limited spatial resolution of the continuum maps it is based on, the source model is not well

$^3$ http://www.sron.rug.nl/~vdptak/ratran/
Bn

-2

-1

-0

a

b

-c

d

A

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5.3. A full chemical model

We modelled the chlorine chemistry at each radial location in FIR 4 with the Nautilus gas-grain chemical code. Nautilus time-dependently computes the gas and grain chemistry including freezeout, surface chemistry, and desorption due to thermal and indirect processes. The grain surface reactions are described in Her Kant et al. (2009) and Semenov et al. (2010). The binding energy of Cl on H$_2$-O ice is 1100 K, for HCl we adopted 5174 K from Olanrewaju et al. (2011). The gas-phase network is based on kida.uva.2011, from Wakelam et al. (2012), and was updated with data from Neufeld & Wolfire (2009). The full network contains 8335 reactions, 684 species and 13 elements. The H$_2$ and CO self shielding are computed following Lee et al. (1996), as described in Wakelam et al. (2012). We adopted a volatile chlorine abundance of 10$^{-7}$. The chemistry is first evolved to steady state in dark cloud conditions. This yields the initial abundances for the time-dependent chemistry in FIR 4, using the C09 source structure for the density and temperature.

Dark cloud stage. The initial conditions for the FIR 4 calculation are computed for dense, cold cloud conditions: a temperature of 10 K, n$_H$ = 2 x 10$^7$ cm$^{-3}$, A$_K$ = 10 mag, and a cosmic-ray ionization rate of $\zeta$ = 10$^{-17}$ s$^{-1}$. The model is evolved to 10$^6$ yr. The species are initially all atomic, except for H$_2$, with most abundances from Hincelin et al. (2011). The abundance of oxygen is set to 3.3 x 10$^{-4}$, and of chlorine to 10$^{-7}$. At the end of the dark cloud stage, about 93% of elemental chlorine is in HCl ice, the remaining 7% is almost entirely in gas-phase atomic Cl. At the temperature, the density and the age of the cloud influence the fraction of Cl and HCl in the gas versus in the ices, but do not affect the HCl/Cl ratio. This ratio, both in the gas and ices, is mostly influenced by the cosmic ray ionization rate $\zeta$. The resulting abundances are used as the initial conditions for FIR 4.

Protostellar stage. For the protostellar stage, we use the C09 density and temperature structure without enhanced external and the range of models within reduced $\chi^2$ = 3. The log$_{10}$($\chi^2$) surface is shown in Fig. 8.

5.2. Fitting a constant abundance profile

As a first guess, we assume a constant HCl abundance in the source. Assuming a fixed source structure, the model has only two free parameters: $X_c$, the gas-phase HCl abundance; and $R_{\text{in}}$, the radius where the line width switches from broad to narrow. We performed a $\chi^2$ minimization on the HCl 1–0 and 2–1 line profiles from HIFI, and the 1–0 from CSO. The best fit parameters, with a reduced $\chi^2$ = 1.45, are $X_c$ = 9 x 10$^{-11}$ and $R_{\text{in}}$ = 7000 AU. In Fig. 7, we show the data, the best-fit model and the range of models within reduced $\chi^2$ = 3. The log$_{10}$($\chi^2$) surface is shown in Fig. 8.
The gas-phase abundance of HCl from Nautilus is consistent with the observational constraints on \(X(\text{HCl})_{\text{gas}}\), as reported in Sect. 5.2, within a factor of a few. This is with the exception of the high-ionization model, where \(X(\text{HCl})_{\text{gas}}\) exceeds the observed limit by two orders of magnitude. The gas-phase HCl abundance of \(\sim 10^{-10}\) — seen in most of the source in both models — is the tip of the iceberg, as 90% of all chlorine is frozen onto grain surfaces as HCl and likely in a water ice matrix.

### 5.4. The outermost envelope PDR

The dense and heavily irradiated photon dominated region in the outermost envelope requires a specialized treatment. We employ the Meudon\textsuperscript{4} PDR code (Le Petit et al. 2006) for this. The physical structure is a slab with a constant density of \(n_{\text{H}} = 6 \times 10^5\) – the outermost density in the C09 source model — extending to \(A_V = 20\) mag. We used two cosmic ray ionization rates, as before: \(\zeta = 10^{-16} \text{ s}^{-1}\) and \(10^{-14} \text{ s}^{-1}\).

The external irradiation of FIR 4 has previously been inferred to be \(G_0 \sim 500\) interstellar radiation field units, based on [CII] 158 \(\mu\text{m}\) emission (Herrmann et al. 1997). We re-evaluated \(G_0\) using the Herschel/HIFI [CII] line flux from Kama et al. (2013), which was obtained at higher spatial resolution and yields \(G_0 \sim 300\). We further checked \(G_0\) based on an archival Spitzer 8 \(\mu\text{m}\) map, which has a substantial contribution from PAH emission, and this directly counts excitation events by ultraviolet photons. The measured flux density of \(\sim 250\) MJy/sr gives \(G_0 = 415\), using Eq. (1) of Vicente et al. (2013) with a standard [C]/[H] \(\sim 10^{-4}\) and 4% of elemental carbon locked in PAHs (Tielens 2005). All estimates are consistent with each other, and with irradiation by the Trapezium cluster at a projected separation of \(\sim 2\) pc. We adopt the middle ground: \(G_0 = 415\).

The resulting abundance profiles of gas-phase HCl, Cl, Cl\(^+\) and H\(_2\)Cl\(^+\) are shown in the left-hand panel of Fig. 10. For both \(\zeta\) values, HCl peaks at \(\sim 10^{-8}\) at the H/\(\text{H}_2\) transition, then decreases deeper into the PDR. By \(A_V \approx 1\), \(X(\text{HCl})_{\text{gas}}\) drops

---

\(\chi^2\) the reduced \(\chi^2\) surface for the two constant HCl abundance model parameters: \(X_c\) and \(R_{\text{iso}}\).
below $10^{-10}$. The HCl abundance in these models is different from the Nautilus results because of different initial conditions and because the Meudon code solves for steady state, which the time-dependent Nautilus models do not reach by $10^5$ yr. The rapid changes around 0.1 to 0.2 AU are due to numerical issues related to H$_2$ formation and the heating-cooling balance. The effect on HCl is limited to $\leq 40\%$ of the column density, and does not substantially affect our conclusions. The temperature in this region is between $\sim 100$ and $\sim 1000$ K, and the total column density in the outer PDR is $N$(HCl) $\approx 1 \times 10^{12}$ cm$^{-2}$.

Using the $X$(HCl) profile shown in Fig. 10 for the outermost envelope, we re-fitted the HCl abundance model from Sect. 5.2. The two $X$(HCl) profiles were joined at $T_{kin} = 25$ K. The impact of the PDR on the HCl emission is small. The best-fit combined model $- X_c = 7 \times 10^{-11}$, $R_{bin} = 9000$ AU $-$ has a reduced Figure 9. HCl and Cl gas- and ice-phase abundances by $10^4$ (solid) and $10^5$ yr (dashed) after the dark cloud stage, as modelled with Nautilus. The cosmic ray ionization rate is $\zeta = 10^{-16}$ s$^{-1}$ (left panel) and $\zeta = 10^{-14}$ s$^{-1}$ (right panel).

![Graph](image)

Figure 10. Meudon models of the chlorine chemistry in the two PDR layers seen towards FIR 4, showing abundances relative to $n_{H_2}$. Left panel: dense PDR in the outermost envelope of FIR 4, with $n_{H_2} = 6 \times 10^3$ cm$^{-3}$, $G_0 = 415$ ISRF. The cosmic ray ionization rate is $\zeta = 10^{-16}$ s$^{-1}$ (solid lines) and $\zeta = 10^{-14}$ s$^{-1}$ (dotted). The position is given outside in from the right. Right panel: tenuous PDR in the foreground of FIR 4, with $n_{H_2} = 10^3$ cm$^{-3}$, $G_0 = 1500$ ISRF and $\zeta = 3 \times 10^{-16}$ s$^{-1}$. The position is given from the strongly irradiated side on the right.

Table 3. Observational results and model predictions for H$_2$Cl$^+$ in the tenuous foreground PDR.

<table>
<thead>
<tr>
<th></th>
<th>$N$ [10$^{14}$ cm$^{-2}$]</th>
<th>$T_{ex}$ [K]</th>
<th>$\delta v_{abs}$ [km s$^{-1}$]</th>
<th>$\delta v$ [km s$^{-1}$]</th>
</tr>
</thead>
<tbody>
<tr>
<td>Observed</td>
<td>1.3 $\pm$ 0.1</td>
<td>4.3 $\pm$ 0.1</td>
<td>9.3 $\pm$ 0.1</td>
<td>1.8 $\pm$ 0.1</td>
</tr>
<tr>
<td>Modelled</td>
<td>0.11</td>
<td>--</td>
<td>--</td>
<td>--</td>
</tr>
</tbody>
</table>

Notes. Dashes are quantities not obtainable from the chemical model. See Sect. 6 for details.

$\chi^2 = 1.55$. The model matches a tentative narrow peak in the HCl 2--1 line, which the constant abundance model does not, although this does not improve the global $\chi^2$.

In the outer envelope PDR, elemental chlorine is not depleted from the gas in the outermost 300 AU. The PDR model is consistent with the observed HCl emission, and still requires most of the narrow and broad HCl flux to originate deeper in FIR 4. It is also consistent with the upper limit on H$_2$Cl$^+$ absorption at 11.4 km s$^{-1}$. The H$_2$Cl$^+$ data is analysed next, in Sect. 6.

6. Analysis of H$_2$Cl$^+$

The H$_2$Cl$^+$ lines, shown in Fig. B.1, appear in absorption against the weak continuum of the source. As shown in Fig. 3, they are blueshifted by 2 km s$^{-1}$ with respect to the systemic velocity of FIR 4, which is 11.4 km s$^{-1}$. Thus, the H$_2$Cl$^+$ absorption does not originate in the dense outer envelope PDR. Instead, the radial velocity matches that of the tenuous foreground PDR layer recently identified towards FIR 4 by López-Sepulcre et al. (2013a). The distance and relation between the foreground layer and FIR 4 are poorly constrained.

We used the CASSIS$^5$ software to perform Markov Chain Monte Carlo fitting of local thermodynamic equilibrium models to the H$_2$Cl$^+$ and H$_2$Cl$_2$ lines. Not all the lines are detected, but they are still useful as constraints. To minimize baseline issues, we added the continuum fit of Kama et al. (2013) to baseline subtracted spectra. We find a column density of $N$(H$_2$Cl$^+$) = $1.3 \times 10^{14}$ and an isotopic ratio of $[^{35}\text{Cl}] / [^{37}\text{Cl}] = (4.3 \pm 0.8)$. The main fitting results are given in Table 3.

$^5$ http://cassis.irap.omp.eu
To compare the observed \( H_2Cl^+ \) column density with that expected from the chlorine chemical network, we again used the Meudon code (Le Petit et al. 2006). The foreground PDR was determined by López-Sepulcre et al. (2013a) to have a density of \( n = 10^3 \text{ cm}^{-2} \), an FUV irradiation of \( \sim 1500 \text{ G}_0 \) and a cosmic ray ionization rate of \( 3 \times 10^{-16} \text{ s}^{-1} \). The resulting abundance profiles of \( \text{Cl}^+, \text{Cl}^2, \text{HCl} \) and \( \text{H}_2\text{Cl}^+ \) are shown in the right-hand panel of Fig. 10. The PDR chemistry model predicts \( N(\text{H}_2\text{Cl}^+) = 1.1 \times 10^{13} \text{ cm}^{-2} \), an order of magnitude below the observed value.

An order of magnitude excess of the observed \( \text{H}_2\text{Cl}^+ \) column density over the chemical model predictions, such as we find for FIR 4, has been noted since the recent first identification of the species towards NGC 6334I and Sgr B2(S) by Lis et al. (1995; Salez et al. 1996; Neufeld & Green 1994; Peng et al. 2010), who suggested viewing geometry as a possible explanation – PDR model column densities are commonly given for a face-on viewing angle. The discrepancy was discussed in detail for several sources by Neufeld et al. (2012), but no satisfactory explanation surfaced, although suggestions for future work were advocated by F14 and must be reproduced by any model. For a luminosity exceeding the protostellar luminosity of \( 100 \text{ L}_\odot \) found by Furlan et al. (2014, F14). The latter authors, as well as López-Sepulcre et al. (2013b), attributed this to the lower spatial resolution data and larger photometric annuli used by C09. We assess here the potential impact on our results.

While the far-infrared fluxes of C09 were likely contaminated by the nearby source, FIR 3, the (sub-)mm continuum maps were spatially resolved on a scale comparable to that studied by F14 and must be reproduced by any model. For a luminosity and thus cooling rate decrease of 1 dex, the corresponding dust temperature decrease is a factor of 1.5 or \( \sim 30\% \). The fixed mm flux demands a matching increase in the density of the emission region. Recalculating the HCl 1–0 excitation for these conditions implies a change in our derived HCl abundance by at most a factor of two.

The impact of the changed source luminosity on the chemistry is also expected to be small. Lowered temperatures would most of all assist in keeping chlorine locked in ices.

7. Discussion

7.1. Depletion of chlorine

We find a ratio of elemental volatile chlorine to gas-phase HCl, \( X(\text{Cl}\text{vol})/X(\text{HCl})_{\text{gas}} \), of \( \sim 1000 \) in FIR 4. Previous studies of the HCl abundance in molecular gas have found ratios in the range 50 to 640, typically \( \sim 400 \) (e.g. Schilke et al. 1995; Zmuidzinas et al. 1995; Salez et al. 1996; Neufeld & Green 1994; Peng et al. 2010). A depletion factor of \( \sim 100 \) was found for the outflow-shocked region L1157-B1 (Codella et al. 2012).

Our modelling suggests that hydrogen chloride ice is the main chlorine reservoir in protostellar core conditions, containing 90 to 98% of the elemental volatile chlorine. Gas-phase atomic Cl contains most of the remaining 10% of the chlorine.

The above result is independent of our choice of cosmic ray ionization rate, \( \zeta \), although inside of \( \sim 5000 \text{ AU} \), the gas-phase abundance of HCl (a minor, but key part in the chlorine budget) can vary substantially depending on this parameter. For \( \zeta = 10^{-16} \text{ s}^{-1} \), the gaseous HCl abundance stays similar to that of the outer envelope (\( \sim 10^{-10} \)), however for \( \zeta = 10^{-14} \text{ s}^{-1} \) it is \( 10^{-8} \) up to \( 10^{5} \text{ yr} \). This is, at face value, not consistent with our observations, as it would cause extremely strong HCl emission which is not observed. On the other hand, evolving the static \( \zeta = 10^{-14} \text{ s}^{-1} \) model beyond \( 10^{5} \text{ yr} \) leads to a decreasing abundance of HCl in the gas, and it eventually falls below \( 10^{-10} \) in the inner envelope.

Another possibility is that infall influences the abundances in the inner envelope, keeping more chlorine in HCl ice than is seen in our static models. This may require quite rapid infall (within a few thousand years on 5000 AU scales) in order to prevent the buildup of a high gas-phase abundance. The poorly known physical structure on \( \lesssim 2000 \text{ AU} \) scales also impacts our modelling of the chemistry as well as the excitation of HCl, leading to further uncertainty about the HCl abundance on small scales. Another possibility is that HCl is liberated from grain mantles more slowly than the bulk ice. This might also explain the low gas-phase HCl abundance found in the L1157-B1 shock by (Codella et al. 2012), and would point to some – as yet unknown – relatively refractory reservoirs of chlorine on grains.
Table 4. $^{35}\text{Cl}/^{37}\text{Cl}$ determinations towards the Orion star forming region. All uncertainties are 1σ.

<table>
<thead>
<tr>
<th>Source/Region</th>
<th>$^{35}\text{Cl}/^{37}\text{Cl}$</th>
<th>Notes</th>
</tr>
</thead>
<tbody>
<tr>
<td>OMC-2 FIR 4 (HCl)</td>
<td>3.2 ± 0.1</td>
<td>this work</td>
</tr>
<tr>
<td>OMC-2 foreground (H$_2$Cl$^+$)</td>
<td>4.3 ± 0.8</td>
<td>this work</td>
</tr>
<tr>
<td>OMC-1 position 1</td>
<td>2.3$^{+1.5}_{-0.8}$</td>
<td>Peng et al. (2010)</td>
</tr>
<tr>
<td>OMC-1 position 2</td>
<td>2.5$^{+0.9}_{-0.7}$</td>
<td>Peng et al. (2010)</td>
</tr>
<tr>
<td>OMC-1</td>
<td>6.5$^{+2.2}_{-2.0}$</td>
<td>Salez et al. (1996)</td>
</tr>
<tr>
<td>Orion Bar</td>
<td>2.1$^{+0.5}_{-0.3}$</td>
<td>Peng et al. (2010)</td>
</tr>
<tr>
<td>solar system</td>
<td>3.1</td>
<td>Lodders (2003)</td>
</tr>
</tbody>
</table>

7.3. The broad component

The broad component has a combination of relatively large line width (∼10 km s$^{-1}$) and large spatial extent (thousands of AU), while embedded in, or at least projected over, a dense protostellar core. As we discuss below, it is not obvious what the nature of this component is. Kinematical evidence relates the broad HCl component to the CS molecule, mapping of which in turn suggests that this component is extended on a scale of ≥2100 AU. Our constant HCl abundance models suggest a best-fit symmetric radial extent of 7000 AU at first glance, such a large line width and large spatial scale make hypotheses other than an outflow seem unlikely.

There is indeed mounting evidence from kinematical and excitation considerations that FIR 4 indeed hosts a compact outflow (Kama et al. 2013, and prep.; Furlan et al. 2014). Spatially and spectrally resolved data, to be presented in a companion paper, suggest that the outflow axis runs roughly north to south, with lobe sizes of at a few thousand AU. However, the C$^{34}$S map in Fig. 4 shows a significant east-west elongation, which seems difficult to explain with such an outflow, unless outflow-driven gas is spilling over the protostellar core surface where the outflow cone breaks out. The C$^{34}$S velocity map shows a slow rotation around the north-south axis, with a typical velocity an order of magnitude below the linewidth.

It has been proposed that a larger outflow from the nearby Class I source, OMC-2 FIR 3, impacts and shocks the FIR 4 core (Shimajiri et al. 2008). It seems unlikely, however, that the FIR 3 outflow is responsible for the broad line emission in FIR 4, because the broad C$^{34}$S emission peaks on-source, and because the high-velocity wings of the CO and H$_2$O lines in FIR 4 are perfectly symmetric around the local $v_{lsr}$. Spatially resolved studies of the high velocity wings of CO are needed to clarify the issue. Previous interferometric observations have had insufficient sensitivity to probe the outflow gas at several tens of km s$^{-1}$.

7.4. The chlorine isotopic ratio

Studies of this ratio throughout the Galaxy have typically been consistent with the solar system value of 3.1 (Lodders 2003), within large error bars (e.g. Salez et al. 1996; Peng et al. 2010; Cernicharo et al. 2010).

Because of the very small Cl isotope mass difference (6%), minimal chemical fractionation is expected, and we determined the isotope ratio via the HCl and H$_2$Cl$^+$ isotopolog ratios. For HCl, we find a line flux ratio of 3.2 ± 0.1, which is a robust isotopolog ratio indicator, given the low optical depth of the lines suggested by the hyperfine component ratios of the narrow HCl emission (see also Cernicharo et al. 2010). For H$_2$Cl$^+$, we found a ratio of 4.3 ± 0.8. The results, summarized in Table 4, are consistent with the solar system value and with values measured elsewhere in the Orion star forming region.

7.5. Impact of the new HCl-H$_2$ excitation rates

That the difference between the HCl-H$_2$ excitation rate coefficients and the scaled HCl-He ones should impact abundance determinations was noted already by Lanza et al. (2014b). As shown in Fig. 6, the new HCl-H$_2$ hyperfine-resolved collisional excitation rates are roughly a factor of five to ten larger than the previously used, mass-scaled HCl-He rates from Neufeld & Green (1994) and Lanza & Lique (2012). This suggests that previous estimates of the gas-phase HCl abundance in molecular gas must be re-evaluated to be up to an order of magnitude lower, and correspondingly the typical fraction of elemental chlorine in gas-phase HCl must be around a factor of 10$^{-3}$ (a depletion factor of ~1000). Modelling results in Sect. 5.3 show that the strong depletion can be well understood in a framework where elemental chlorine is sequestered into HCl ice, where it remains at least as strongly bound as H$_2$O itself.

Based on the new excitation rates, the critical density of the HCl 1–0 transition is ~10$^7$ cm$^{-3}$. This is accurate within a factor of a few in the temperature range of the new rate coefficients (up to 300 K).

8. Conclusions

We carried out a study of chlorine towards the OMC-2 FIR 4 protostellar core, using Herschel and CSO observations of HCl and H$_2$Cl$^+$. Our main findings are listed below.

1. We detect the HCl 1–0 and 2–1 transitions in emission with Herschel and CSO, and H$_2$Cl$^+$ in absorption with Herschel.
2. The narrow HCl component ($FWHM = 2.0$ km s$^{-1}$) traces the outer envelope, and the broad one ($FWHM = 10.5$ km s$^{-1}$) a compact central region, possibly outflow-driven gas.
3. The HCl data are well modelled with a constant abundance of X(HCl)$_{gas}$ = 9 × 10$^{-11}$ in FIR 4, corresponding to ~10$^{-3}$ of the ISM abundance of elemental chlorine.
4. Chemical models show that HCl ice contains ~90 to 98% of all volatile chlorine in the source. The second largest reservoir is gas-phase atomic Cl, up to 10% of the total. All other species have much lower abundances. In the inner 100 AU, HCl gas may hold up to 100% of volatile chlorine.
5. The external irradiation of the FIR 4 envelope is $G_0 = 415$ ISRF. Elemental chlorine is undepleted in the outermost 300 AU of the resulting dense PDR. Including this PDR in the source model gives a best-fit X(HCl)$_{gas}$ = 7 × 10$^{-11}$ in the rest of the source.
6. H$_2$Cl$^+$ traces a recently discovered diffuse, blueshifted foreground PDR. The observed H$_2$Cl$^+$ column density is 1.3 × 10$^{14}$ cm$^{-2}$, an order of magnitude above the model prediction of 1.1 × 10$^{15}$ cm$^{-2}$.
7. Our best estimate of the $^{35}$Cl/$^{37}$Cl isotopic ratio in OMC-2 FIR 4 is 3.2 ± 0.1 (1σ), consistent with other measurements in the solar system and in the Orion region.
8. Newly calculated HCl-H$_2$ hyperfine-resolved collisional excitation rate coefficients exceed previous HCl-He scaled values by up to an order of magnitude at protostellar core temperatures, suggesting that previous estimates of chlorine depletion from the gas should be revisited.
Appendix A: Hyperfine excitation of HCl by H₂

Rate coefficients for rotational excitation of HCl(1Σ⁺) by collisions with H₂ molecules have been computed by Lanza et al. (2014a) for temperatures ranging from 5 to 300 K. The rate coefficients were derived from extensive quantum calculations using a new accurate potential energy surface obtained from highly correlated ab initio approaches (Lanza et al. 2014b).

However, in these calculations, the hyperfine structure of HCl was neglected. To model the spectrally resolved HCl emission from molecular clouds, hyperfine resolved rate coefficients are needed. In this appendix, we present the calculations of HCl-H₂ hyperfine resolved rate coefficients from the rotational rate coefficients of Lanza et al. (2014a). Note that, for rotational levels, we use here the lowercase j instead of the astronomical J notation used in the main body of the paper.

A.1. Methods

In HCl, the coupling between the nuclear spin (I₁ = 3/2) of the chlorine atom and the molecular rotation results in a weak splitting of each rotational level j₁ into 4 hyperfine levels (except for the j₁ = 0 level which is split into only 1 level and for the j₁ = 1 level which is split into only 3 levels). Each hyperfine level is designated by a quantum number F₁ (F₁ = I₁ + j₁) varying between |I₁ − j₁| and I₁ + j₁. In the following, j₂ designates the rotational momentum of the H₂ molecule.

In order to get HCl–H₂ hyperfine resolved rate coefficients, we extend the Infinite Order Sudden (IOS) approach for diatom-atom collisions (Faure & Lique 2012) to the case of diatom-diatom collisions.

Within the IOS approximation, inelastic rotational rate coefficients kᵢ,j₁→j₂,j₂(T) can be calculated from the “fundamental” rates (those out of the lowest j₁ = 0, j₂ = 0 channel) as follows (e.g. Alexander 1979):

$$k_{i,j_1 \rightarrow j_2,j_2}(T) = \sum_{L_1,L_2} (2F'_1 + 1)(2F''_1 + 1) \left( \begin{array}{ccc} j_1 & L_1 & j_2 \\ 0 & 0 & 0 \end{array} \right)^2 \times \left( \begin{array}{ccc} j_2 & L_2 & j_2 \\ 0 & 0 & 0 \end{array} \right)^2 k_{i,j_1 \rightarrow 0,0,0,L_1,L_2}(T). \tag{A.1}$$

Similarly, IOS rate coefficients amongst hyperfine structure levels can be obtained from the kᵢ,j₁,j₂→p,j₁,j₂,j₂(T) rate coefficients using the following formula:

$$k_{i,j_1,j_2 \rightarrow p,j_1,j_2,j_2}(T) = \sum_{L_1,L_2} (2F'_1 + 1)(2F''_1 + 1) \times (2F'_1 + 1) \left( \begin{array}{ccc} j_2 & L_2 & j_2 \\ 0 & 0 & 0 \end{array} \right)^2 \times \left( \begin{array}{ccc} j'_1 & L_1 & j'_1 \\ 0 & 0 & 0 \end{array} \right)^2 \times \left( \begin{array}{ccc} L_1 & F'_1 & F''_1 \\ I_1 & j'_1 & j_1 \end{array} \right)^2 k_{i,j_1 \rightarrow p,0,0,L_1,L_2}(T). \tag{A.2}$$

where ( ) and [ ] are respectively the “3 − j’” and “6 − j’” Wigner symbols.

The IOS approximation is expected to be moderately accurate at low temperature. As suggested by Neufeld & Green (1994), we could improve the accuracy by computing the hyperfine rate coefficients as:

$$k_{i,j_1,j_2 \rightarrow p,j_1,j_2,j_2}(T) = \frac{k_{i,j_1,j_2 \rightarrow p,j_1,j_2,j_2}(T)}{k_{i,j_1,j_2 \rightarrow 0,0,j_1,j_2}(T)} k_{CC}^{CC}(T) \tag{A.3}$$

using the CC rate coefficients k_{i,j_1 \rightarrow j_2,j_2}(T) of Lanza et al. (2014a) for the IOS “fundamental” rates in Eqs. (A.1)–(A.2).

$$k_{i,j_1 \rightarrow j_2,j_2}(T) = (2L_1 + 1)(2L_2 + 1)k_{i,j_1 \rightarrow 0,0,0,L_1,L_2}(T). \tag{A.4}$$

This procedure is found to significantly improve the results at low temperature due to important threshold effects.

Hence, we have determined hyperfine HCl–H₂ rate coefficients using the computational scheme described above for temperature ranging from 5 to 300 K. We considered transitions between the 28 first hyperfine levels of HCl (j₁, j’₁ ≤ 7) due to collisions with para-H₂ (j₂ = 0) and ortho-H₂ (j₂ = 1). The present approach has been shown to be accurate, even at low temperature, and has also been shown to induce almost no inaccuracies in radiative transfer modeling compared to more exact calculations of the rate coefficients (Faure & Lique 2012).

A.2. Results

The complete set of (de)excitation rate coefficients is available on-line from the LAMDA6 (Schöier et al. 2005) and BASECOL7 (Dubernet et al. 2013) websites. For illustration, Fig. A.1 depicts

![Fig. A.1. Temperature dependence of the hyperfine resolved HCl–para-H₂ (upper panel) and HCl–ortho-H₂ (lower panel) rate coefficients for HCl(j₁ = 2, F → j₁’ = 1, F’₁) transitions.](http://www.strw.leidenuniv.nl/~moldata/)

http://www.strw.leidenuniv.nl/~moldata/  
http://basecol.obspm.fr/
Table A.1. A comparison between present hyperfine rate coefficients for ortho- and para-H$_2$, and those of LL12 for He.

<table>
<thead>
<tr>
<th>$j_1, F_1 \rightarrow j_1', F_1'$</th>
<th>10 K</th>
<th>100 K</th>
<th>300 K</th>
</tr>
</thead>
<tbody>
<tr>
<td></td>
<td>p-H$_2$</td>
<td>o-H$_2$</td>
<td>He $\times$ 1.38</td>
</tr>
<tr>
<td>1, 1.5 $\rightarrow$ 0, 1.5</td>
<td>5.54e-11</td>
<td>1.53e-10</td>
<td>3.97e-12</td>
</tr>
<tr>
<td>2, 2.5 $\rightarrow$ 0, 1.5</td>
<td>2.01e-11</td>
<td>3.35e-11</td>
<td>2.64e-12</td>
</tr>
<tr>
<td>2, 3.5 $\rightarrow$ 1, 1.5</td>
<td>1.41e-12</td>
<td>7.35e-12</td>
<td>3.38e-12</td>
</tr>
<tr>
<td>3, 3.5 $\rightarrow$ 0, 1.5</td>
<td>9.31e-12</td>
<td>1.08e-11</td>
<td>2.28e-12</td>
</tr>
<tr>
<td>3, 3.5 $\rightarrow$ 1, 1.5</td>
<td>6.28e-12</td>
<td>1.58e-11</td>
<td>2.04e-12</td>
</tr>
<tr>
<td>3, 2.5 $\rightarrow$ 1, 1.5</td>
<td>5.90e-12</td>
<td>1.48e-11</td>
<td>1.59e-12</td>
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<tr>
<td>3, 2.5 $\rightarrow$ 2, 3.5</td>
<td>3.85e-13</td>
<td>2.27e-12</td>
<td>1.45e-12</td>
</tr>
<tr>
<td>3, 4.5 $\rightarrow$ 2, 2.5</td>
<td>4.56e-13</td>
<td>2.69e-12</td>
<td>1.84e-12</td>
</tr>
</tbody>
</table>

Notes. The rates are in units of cm$^3$ mol$^{-1}$ s$^{-1}$.

Indeed, collisions with helium are often used to model collisions with para-H$_2$. It is generally assumed that rate coefficients with para-H$_2$($j_2 = 0$) should be larger than He rate coefficients owing to the smaller collisional reduced mass.

As one can see, the scaling factor is clearly different from 1.38. The ratio varies with the transition considered and also with the temperature for a given transition. The ratio may be larger than a factor 10. This comparison indicates that accurate rate coefficients with para-H$_2$($j_2 = 0$) and ortho-H$_2$ ($j_2 = 1$) could not be obtained from He rate coefficients. HCl molecular emission analysis performed with HCl-He rate coefficients result in large inaccuracies in the HCl abundance determination.
Appendix B: CASSIS fitting of H$_2$Cl$^+$

In Fig. B.1, we show the H$_2$Cl$^+$ lines used in the LTE fitting with the CASSIS software and the best-fit model resulting from the $\chi^2$ minimization. The hyperfine components and other lines detected nearby are also shown. The data are all from the HIFI spectral survey of FIR 4 (Kama et al. 2013).

![Graphs showing H$_2$Cl$^+$ transitions](image_url)

**Fig. B.1.** The four H$_2$Cl$^+$ lines (black) used in the LTE model fitting with CASSIS, and the best-fit model (red). All spectra are corrected for the foreground PDR velocity of 9.4 km s$^{-1}$. The dashed red lines show the hyperfine components of H$_2$Cl$^+$ transitions, while dashed blue lines in the top right panel indicate the native frequencies of C$_2$H transitions. The other H$_2$Cl$^+$ transitions do not have any lines nearby that were listed as detections in the spectral survey of Kama et al. (2013).