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## A CONNECTION BETWEEN PLASMA CONDITIONS NEAR BLACK HOLE EVENT HORIZONS AND OUTFLOW PROPERTIES

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### ABSTRACT

Accreting black holes are responsible for producing the fastest, most powerful outflows of matter in the universe. The formation process of powerful jets close to black holes is poorly understood, and the conditions leading to jet formation are currently hotly debated. In this paper, we report an unambiguous empirical correlation between the properties of the plasma close to the black hole and the particle acceleration properties within jets launched from the central regions of accreting stellar-mass and supermassive black holes. In these sources the emission of the plasma near the black hole is characterized by a power law at X-ray energies during times when the jets are produced. We find that the photon index of this power law, which gives information on the underlying particle distribution, correlates with the characteristic break frequency in the jet spectrum, which is dependent on magnetohydrodynamical processes in the outflow. The observed range in break frequencies varies by five orders of magnitude in sources that span nine orders of magnitude in black hole mass, revealing a similarity of jet properties over a large range of black hole masses powering these jets. This correlation demonstrates that the internal properties of the jet rely most critically on the conditions of the plasma close to the black hole, rather than other parameters such as the black hole mass or spin, and will provide a benchmark that should be reproduced by the jet formation models.

*Key words:* accretion, accretion disks – black hole physics – galaxies: jets – galaxies: nuclei – ISM: jets and outflows – X-rays: binaries

### 1. INTRODUCTION

Powerful jets of plasma are produced by accreting black holes of all sizes ranging from stellar-mass black holes in Galactic X-ray binaries (XRBs) to supermassive black holes in active galactic nuclei (AGNs). The jets are launched close to the black hole event horizon, but the conditions leading to jet formation are still debated. Several models have been put forward that predict the jet properties are governed by the accretion rate, black hole spin, the magnetic field strength and configuration, and/or the properties and location of the inner accretion flow (Blandford & Znajek 1977; Blandford & Payne 1982; Meier 2001; Tchekhovskoy et al. 2011). Despite the recent advances in our understanding of jets, the total jet power is notoriously hard to measure because rather than being radiated locally, the bulk of the energy is transported to large distances from the black holes as a dark flow. Nevertheless, there are estimates that imply jets can dominate the power output of black holes (Gallo et al. 2005; Ghisellini et al. 2014), and that a large fraction of the mass in the accretion flow can escape via outflows depending on the state of accretion (Neilsen & Lee 2009; Ponti et al. 2012).

The classic signature of a relativistic, compact jet is a flat or slightly inverted ( $\alpha_{\text{thick}} \geq 0$  where  $S_\nu \propto \nu^{\alpha_{\text{thick}}}$ ) radio spectrum composed of overlapping synchrotron spectra from different locations in the jet (Blandford & Königl 1979). The flat/inverted radio spectrum will break at some higher frequency  $\nu_b$  associated with a transition to low optical depth, either at the

base of the jet (e.g., Blandford & Königl 1979; Königl 1981; Ghisellini et al. 1985) or at the location where particles are accelerated due to the presence of a shocked zone (e.g., Markoff et al. 2005; Marscher et al. 2008; Polko et al. 2014). The slope of the optically thin spectrum is usually close to  $\alpha_{\text{thin}} \sim -0.7$  or steeper if the electrons are cooled or have a thermal distribution of energies (Pe’er & Casella 2009). The radio luminosity is often used as a proxy for jet power, but jet powers inferred this way can disagree with estimates derived from jet feedback in the form of large-scale lobes and cavities by orders of magnitude (Körding 2008). While the jet power is often estimated directly from the radio luminosity with either a linear or power-law dependency, the total luminosity of the jet is dominated by the high frequency emission and can only be measured accurately by observing the whole spectrum of the jet.

Up until recently, the reason why only a few jet breaks had been identified from XRBs was due to the companion star or accretion disk dominating the emission around the break frequencies (Gallo et al. 2007; Migliari et al. 2007; Rahoui et al. 2011) and the lack of mid-IR data acquired from these sources. However, the latter has recently been improved with multiwavelength campaigns of black hole outbursts including mid-IR and millimeter data. A recent key discovery shows that the peak flux density of the jet can vary dramatically with the state of accretion of XRBs, while the radio luminosity can remain steady (Russell et al. 2013b, 2014; van der Horst et al. 2013), casting into doubt the reliability of using the radio

luminosity as a proxy for the jet power. Similarly for most AGNs, the break and most of the jet spectrum lie under other non-synchrotron components, e.g., galaxy, accretion disk, and/or torus, and it is only recently, owing to the use of adaptive optics at optical/IR frequencies, that a core jet spectrum has been revealed from a few close-by low-luminosity AGNs (Fernández-Ontiveros et al. 2012).

Continuously launched, flat spectrum jets are commonly observed during hard and intermediate X-ray states of XRBs (Fender & Gallo 2014) when the X-ray spectrum is dominated by a power-law spectral component. This component is generally thought to represent the inverse Comptonization of soft seed photons in a plasma cloud of hot electrons close to the black hole. These hot electrons are thought to be located near the black hole, whether in a thermal Comptonization-dominated accretion flow (Zdziarski et al. 1998), radiatively inefficient accretion flow (ADAF/RIAF; Yuan et al. 2003), or at the base of the jet (Markoff et al. 2005). Thus, they are usually dubbed the corona, a term which embodies all the different possibilities for the origin of the hot electrons. Due to the simple nature of black holes, accretion physics is expected to scale globally with black hole mass. Similar to XRBs, the hard X-ray emission from AGNs is expected to arise from a Comptonizing corona (Haardt & Maraschi 1993) and evidence for its compact size and location close to the central black hole has come from several research avenues including studies of iron line spectra and variability (e.g., Fabian et al. 2009), reverberation (e.g., Uttley et al. 2014), microlensing (e.g., Morgan et al. 2012), and obscuration of the corona by clouds (e.g., Sanfrutos et al. 2013).

As noted above, the existence and power of jets launched via accretion onto stellar-mass black holes in XRBs has been found to be linked with specific X-ray spectral and timing properties, which trace the nature of mass accretion onto the black hole. Thus we can expect a link between accretion and ejection to be present in their constituent components. In this paper we test this link using broadband spectral energy distributions (SEDs) of stellar-mass and supermassive black holes. The structure of the paper is the following. In Section 2, we present in detail the multiwavelength properties of our sample of XRBs and AGNs, and in particular the SEDs from the AGN sample and the reduction and analysis of the X-ray data from XRBs. The results of the analysis, which shows an anti-correlation between the X-ray power-law photon index and the jet break frequency, are presented in Section 3. In Section 4, we discuss the origin of this correlation and whether or not the system parameters have an effect on it. In Section 5, we present our conclusions and discuss the ramifications of our results.

## 2. OBSERVATIONS AND DATA ANALYSIS

We searched for sources that have a clear X-ray view to the central region of the black holes and a well-sampled jet spectrum in systems where the measurements are not likely to be skewed by relativistic beaming. This search resulted in 11 stellar-mass black holes in XRBs and 7 low accretion rate AGNs that all show an unambiguous, isolated jet spectrum in addition to a well defined X-ray spectrum that is quasi-simultaneous for XRBs. All sources in our sample exhibit the classic signature of a relativistic, flat, or slightly inverted compact jet. For each source we fit a broken power law to determine the frequency of the jet break  $\nu_b$  and the flux density

$S_{\nu,b}$  at the break frequency based on a broad range of multiwavelength observations.

### 2.1. XRBs

#### 2.1.1. Radio and Jet Break Data

We searched the literature for jet break frequency values from the SEDs of black hole XRBs that have simultaneous X-ray observations. The search resulted in nine sources in the hard X-ray state from Russell et al. (2013a) and references therein: 4U 1543–47, Cyg X–1, GS 1354–64, GX 339–4, V404 Cyg, V4641 Sgr, XTE J1118+480, XTE J1550–564, and XTE J1752–223, with two additional intermediate X-ray state sources, MAXI J1659–152 (van der Horst et al. 2013) and MAXI J1836–194 (Russell et al. 2014), in which flat spectrum radio jets were still observed but the X-ray spectra were softer (see the referenced papers for the figures of the SEDs). The values for the jet break frequencies and the radio and jet break fluxes from the XRB sample are tabulated in Table 1. In some cases the location of the jet break could not be well constrained. We therefore used multiwavelength data to restrict the jet breaks to lie in a certain frequency range in order to include them in the Monte Carlo bootstrap estimation of the correlation and linear regression (Section 3). For GS 1354–64 and V4641 Sgr the optically thick radio-to-IR spectrum overshoots the X-ray spectrum and thus the jet spectrum has to break before the X-ray regime. For GX 339–4 and MAXI J1659–152 the optically thick radio spectrum overshoots the data points in the IR and thus the jet spectrum has to break before the IR regime. For XTE J1752–223 the optically thin jet spectrum overshoots the upper limit in the radio and thus the jet break has to be between the radio and IR regimes. We use the most conservative ranges possible based on the existing data, i.e., we do not assume any particular slope for the optically thin spectrum. In the case of XTE J1752–233 we assume that the spectral index of an optically thick, self-absorbed spectrum can not be more than 5/2 based on the synchrotron theory.

#### 2.1.2. X-Ray Data

We used the *Rossi X-ray Timing Explorer (RXTE)* and *Swift* X-ray observatories to select pointings from the High Energy Astrophysics Science Archive Research Center (HEASARC) archive which are as contemporaneous with the radio observations as possible (within a day for all sources with the exception of GS 1354–64, where the nearest pointing was found to be within two days). The only exception was V404 Cyg, which has no *RXTE* or *Swift* data available, and thus we collected the relevant spectral modeling values from the literature (Zycki et al. 1999). Each *RXTE* pointing was individually reduced by the standard method as described in the *RXTE* cookbook using HEASOFT 6.16. The Proportional Counter Array (PCA) spectrum was extracted from all available proportional counter units (PCUs) in each pointing to maximize the photon counts in the spectra, excluding PCU–0 and PCU–1 after their propane loss in 2000 May 13 and 2006 December 25, respectively. The High Energy X-ray Timing Experiment (HEXTE) spectrum was extracted from both clusters A and B when available. After 2009 December 14 when cluster B stopped rocking we used cluster A for the source data and estimated the background using cluster B. PCA spectra were then grouped to a minimum of  $5.5\sigma$  significance per bin and bins below 3.5 keV and above 20 keV were ignored. In

**Table 1**  
Literature Values for the Jet Break Frequencies and Radio and Jet Break Fluxes in XRBs

| Source         | Time<br>(MJD) | $\log \nu_b$<br>(Hz)    | $S_{\nu,b}$<br>(mJy) | $S_{\nu,5 \text{ GHz}}$<br>(mJy) | $\log L_b/L_{5 \text{ GHz}}$ | References      |
|----------------|---------------|-------------------------|----------------------|----------------------------------|------------------------------|-----------------|
| 4U 1543–47     | 52490.00      | $13.98 \pm 0.20$        | $9.2 \pm 2.8$        | $4.00 \pm 0.05$                  | $4.64 \pm 0.34$              | R13a            |
| Cyg X–1        | 53513.00      | $13.45 \pm 0.02$        | $16.8 \pm 1.2$       | $15.6 \pm 0.2$                   | $3.78 \pm 0.06$              | R13a,Ra11       |
| GS 1354–64     | 50772.00      | [14.13–18.00]           | $2.3 \pm 0.1$        | $2.8 \pm 0.1$                    | [4.31–8.25]                  | R13a            |
| GX 339–4       | 50648.00      | $14.26 \pm 0.12$        | $10.7 \pm 0.9$       | $14 \pm 3$                       | $4.44 \pm 0.25$              | R13a,G11        |
|                | 55266.00      | $13.65 \pm 0.24$        | $115 \pm 11$         | $9.1 \pm 1.1$                    | $5.05 \pm 0.33$              | R13a,G11        |
|                | 55617.00      | [12.63–14.26]           | $21.0 \pm 1.0$       | $2.54 \pm 0.04$                  | [3.83–5.51]                  | C13             |
| MAXI J1659–152 | 55467.10      | [10.34–14.67]           | $10 \pm 1$           | $10.5 \pm 0.8$                   | [0.54–5.03]                  | vdH13           |
|                | 55467.90      | [10.63–14.26]           | $11.2 \pm 0.6$       | $9.9 \pm 0.3$                    | [0.95–4.65]                  | vdH13           |
|                | 55470.10      | $10.34^{+0.09}_{-0.11}$ | $8.35 \pm 0.45$      | $9.75 \pm 0.30$                  | $0.57 \pm 0.14$              | vdH13           |
|                | 55473.90      | [11.54–14.68]           | $10.5 \pm 3.2$       | $3.65 \pm 0.09$                  | [2.13–5.57]                  | vdH13           |
|                | 55476.80      | $10.34^{+0.09}_{-0.11}$ | $0.41 \pm 0.07$      | $0.63 \pm 0.03$                  | $0.45 \pm 0.20$              | vdH13           |
|                | 55488.80      | $9.69^{+0.08}_{-0.10}$  | $0.23 \pm 0.03$      | $0.23 \pm 0.03$                  | $-0.01 \pm 0.20$             | vdH13           |
| MAXI J1836–194 | 55807.12      | $11.37^{+0.11}_{-0.27}$ | $415^{+700}_{-190}$  | $29 \pm 1$                       | $2.83 \pm 0.55$              | R14             |
|                | 55821.97      | $11.16 \pm 0.55$        | $64 \pm 16$          | $34.5 \pm 0.9$                   | $1.73 \pm 0.67$              | R14             |
|                | 55830.95      | $11.98^{+0.27}_{-0.21}$ | $260^{+140}_{-45}$   | $14.1 \pm 0.4$                   | $3.55 \pm 0.34$              | R14             |
|                | 55846.01      | $12.74^{+0.13}_{-0.02}$ | $185^{+30}_{-15}$    | $5.5 \pm 0.4$                    | $4.57 \pm 0.15$              | R14             |
|                | 55861.00      | $13.71^{+0.37}_{-0.01}$ | $27^{+18}_{-5}$      | $2.5 \pm 0.3$                    | $5.04 \pm 0.40$              | R14             |
|                | V404 Cyg      | 47728–9                 | $14.26 \pm 0.06$     | $178 \pm 16$                     | $17.0 \pm 0.7$               | $5.58 \pm 0.12$ |
| V4641 Sgr      | 52857.00      | [14.67–18.00]           | $93 \pm 46$          | $621 \pm 2$                      | [3.85–7.65]                  | R13a            |
| XTE J1118+480  | 51649.00      | $13.43 \pm 0.09$        | $290 \pm 65$         | $4.7 \pm 0.7$                    | $5.52 \pm 0.25$              | R13a            |
|                | 53386.00      | $12.65 \pm 0.08$        | $170 \pm 19$         | $4.4 \pm 0.2$                    | $4.54 \pm 0.15$              | R13a            |
| XTE J1550–564  | 51697.00      | $13.68 \pm 0.33$        | $38 \pm 27$          | $0.9 \pm 0.1$                    | $5.61 \pm 0.76$              | R13a            |
| XTE J1752–223  | 55378.00      | [10.99–14.26]           | $1.3 \pm 1$          | $<0.3$                           | $<5.16$                      | R13a, R12       |

**Note.** The columns are: (1) source name, (2) starting time of the observation in MJD, (3) logarithm of the jet break frequency, (4) flux density at the jet break, (5) radio flux density at 5 GHz, (6) the excess luminosity over the radio luminosity up to the jet break  $L_b/L_{5 \text{ GHz}} = \nu_b S_{\nu,b}/\nu_{5 \text{ GHz}} S_{\nu,5 \text{ GHz}}$ , and (7) references for the jet break values (R13a: Russell et al. 2013a, Ra11: Rahoui et al. 2011, G11: Gandhi et al. 2011, C13: Corbel et al. 2013, vdH13: van der Horst et al. 2013, R14: Russell et al. 2014, C11: Chaty et al. 2011, R12: Russell et al. 2012).

addition, we added a 0.5% systematic error to all channels. In a similar fashion HEXTE spectra were grouped to a minimum of  $5\sigma$  significance per bin ( $20\sigma$  per bin after 2009 December 14), and bins below 18 keV and above 200 keV were ignored. In addition, we added a 1% systematic error to all channels of spectra taken after 2009 December 14. For MAXI J1659–152 and MAXI J1836–194 we selected additional simultaneous *Swift* spectra (within a day from *RXTE* pointings) to better gauge the effect of the disk component on the X-ray spectra. The X-ray Telescope (XRT) windowed timing mode data were processed using *XRTPIPELINE* in *HEASOFT* 6.16 and subsequently the source and background spectrum and response files were extracted using *XRTPRODUCTS*. Exposure maps were generated for each pointing and the pile-up was taken into account by excluding a circular region at the source position with the region size depending on the count rate (Reynolds & Miller 2013).

## 2.2. Active Galactic Nuclei

For the majority of AGNs, both the break and most of the jet spectrum lie under other non-synchrotron components (e.g., emission from the galaxy and the star-forming regions). To isolate the true core emission we need observations based on high-angular resolution techniques: the *Hubble Space Telescope* (*HST*) in the optical, ground-based adaptive optics observations in the near-IR, and ground-based diffraction-limited observations in the mid-IR. Our sample of seven AGNs consists of four low-luminosity AGNs, one FR-I, one FR-II, and Sgr A\*. The four low-luminosity AGNs are the brightest and nearest low-luminosity AGNs ( $L_{\text{bol}} \lesssim 10^{42} \text{ erg s}^{-1}$ )

accessible from the southern hemisphere and correspond to those targets with successful adaptive optics observations and the best *HST* coverage in the optical and UV range. These targets were extracted from the project “The central parsecs of the nearest galaxies” (Prieto et al. 2010; Reunanen et al. 2010), a high-spatial resolution study of the brightest and nearest AGNs carried out at sub-arcsecond scales with the Very Large Telescope, using the NaCo and VISIR instruments in the near- and mid-IR ranges, respectively. Three of the targets are canonical references for the definition of the low-luminosity AGN class: NGC 1052 (Heckman 1980), NGC 1097 (Keel 1983), and M87 (Fabian & Rees 1995). Together with the Sombrero galaxy (NGC 4594), this sample is the best representation of the low-luminosity class in the nearby universe, also in terms of host galaxies (Sa, SB(s)b, E, and S0 for NGC 1052, NGC 1097, M87, and Sombrero, respectively). Furthermore, the SEDs of these four objects are in agreement with those in previous works (Elvis et al. 1994; Ho 1999; Eracleous et al. 2010) in the common wavelength ranges covered (X-rays, optical/UV, and radio), but the lack of high-angular resolution near- and mid-IR observations in the past prevented the identification of a jet-dominated continuum in these sources. Similarly, 3C 120 and Cygnus A are representatives of the radio galaxy population and were found to have isolated jet spectra in their high resolution data. 3C 120 belongs to the FR-I class (core-dominated radio galaxies) and Cygnus A to the FR-II class (lobe dominated). They are accreting at a relatively low accretion rate and show unambiguous flat/inverted synchrotron radio spectra and a spectral break. In addition to the above-mentioned AGNs, we

add Sgr A\*, which is accreting matter in a very low accretion rate, to the sample. The jet break frequency and radio measurement values are taken from the literature (Becker & Falcke 2002).

The small number of low-luminosity AGNs in the sample is mainly due to the faintness of their nuclei, which hinders the use of adaptive optics in these objects. Moreover, we discarded those AGNs affected by internal (torus) or external obscuration (dust lanes in the host galaxy), since it is not straightforward to identify the presence of a jet-dominated continuum in those cases, e.g., Centaurus A (Meisenheimer et al. 2007).

### 2.2.1. Radio and Jet Break Data

The data set consists of sub-arcsecond measurements ( $<0''.4$  apertures) from radio to ultraviolet (Canalizo et al. 2003; Lee et al. 2008; Fernández-Ontiveros et al. 2012; Doi et al. 2013; Asmus et al. 2014; Lopez-Rodriguez et al. 2014), in addition to low-angular resolution ( $>1''$  apertures) measurements from the NASA/IPAC Extragalactic Database, the *Wide-field Infrared Survey Explorer* at the IPAC archive, the Akari Point Source Catalog, and the 2MASS Point Source Catalog. All data have been corrected for Galactic extinction (Schlafly & Finkbeiner 2011). For 3C 120 we also measured fluxes from images taken from the *HST* Legacy Archive. The sub-arcsecond photometry ensures the extraction of the nuclear continuum, minimizing the possible contribution of extended components, e.g., the underlying galaxy and extended dust emission. The low-energy part of each SED of the sample can be fit with a broken power law representing a self-absorbed synchrotron spectrum from the jet with a flat or inverted spectrum below the break frequency, as is the case for black hole XRBs. For the fits we allowed the optically thick index to also be negative to avoid/amend the influence of optical depth effects in the determination of the break frequency and the associated flux. The spectral slope at higher frequencies than the jet break in some of the AGNs is extremely steep ( $\alpha_{\text{thin}} \sim -4$ ), which could indicate the presence of a thermal particle distribution or fast cooling in the inner region of the jet. In our steepest case, in Cygnus A the mid-infrared flux is strongly polarized (Lopez-Rodriguez et al. 2014). This supports the synchrotron nature of the spectrum, and thus we can assume that the turnover is jet-related even in the sources with the steepest spectra. While the photometric errors in the optical/IR are typically lower than  $\sim 5\%$ , we considered a minimum of 10% error on all the measured fluxes to account for variability (Anderson & Ulvestad 2005; Maoz et al. 2005). To estimate the errors on the fitted parameters we used a bootstrapping method to generate synthetic data sets. The variations on the original data set are based on the size of the flux errors. Each one of the synthetic spectra is fitted and in the end the variance of the large number of fit results is used as the error for the fit parameters. We only take the sub-arcsecond resolution measurements into account when fitting the data, with the exception of NGC 1052. In this source the spectrum is clearly dominated by the active nucleus below  $\sim 3 \times 10^{13}$  Hz, and thus we included the low-spatial resolution data at these frequencies for a better coverage of the spectral break. The SEDs, with their best-fit models, are shown in Figure 1. The jet break frequencies and radio and jet break fluxes from the AGN sample are tabulated in Table 2.

### 2.2.2. X-Ray Data

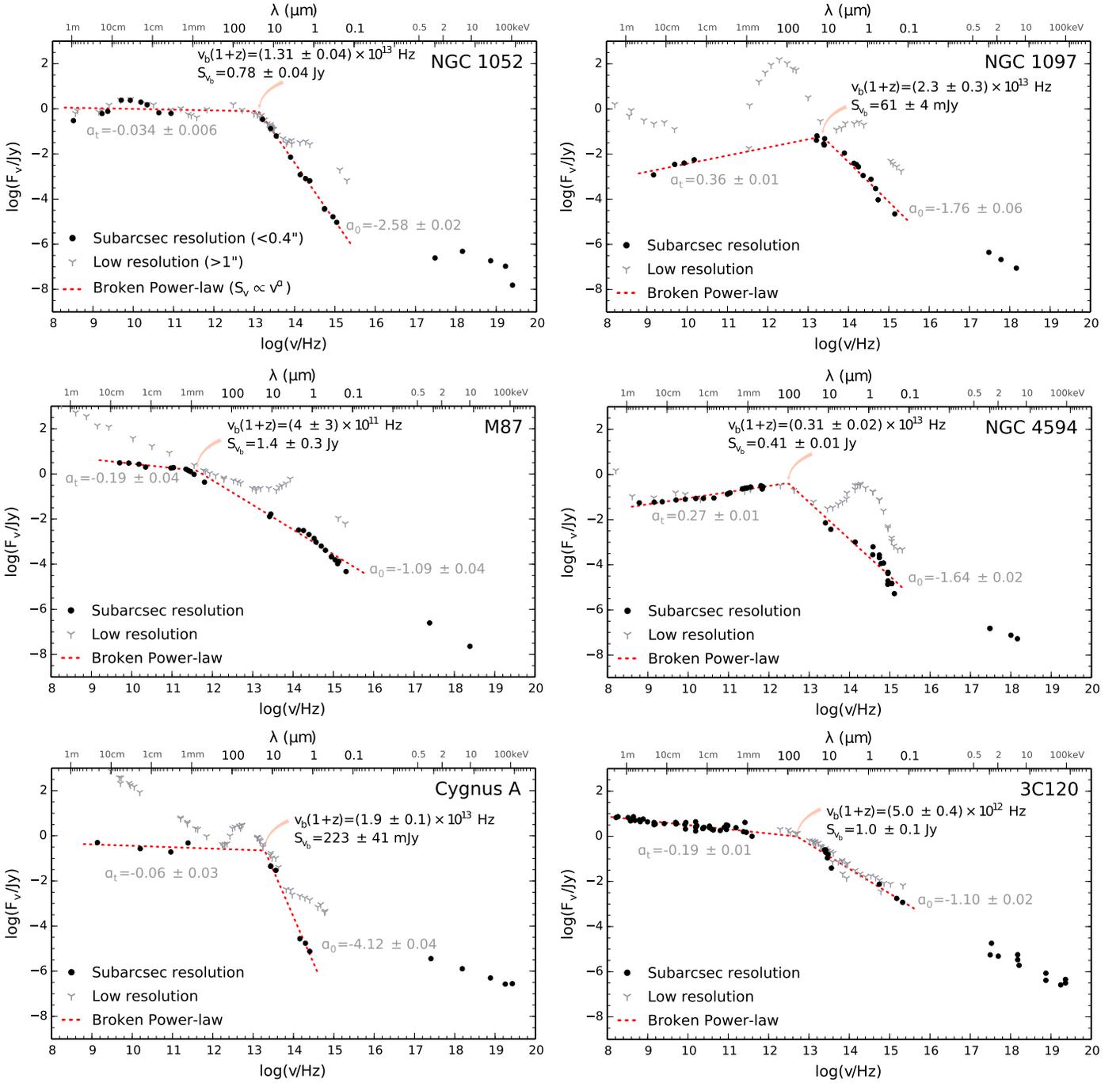
The literature values of the power-law photon indices and luminosities in Eddington units typically measured from the X-ray band 2–10 keV for low-luminosity AGNs (Terashima et al. 2002; González-Martín et al. 2009) and 1–100 keV for Sgr A\*, Cyg A, and 3C 120 (Zdziarski & Grandi 2001; Young et al. 2002; Barrière et al. 2014) of our AGN sample are tabulated in Table 2. The selected X-ray power-law photon indices are less than  $1\sigma$  away from the mean value as calculated from several other values found in the literature. Thus, the non-simultaneity of the X-ray observations with the radio/optical/UV does not present a sizeable correction to the power-law indices. Unlike for XRBs, the bolometric correction factors to turn the X-ray luminosities to accretion rates in Eddington units are larger as most of the accretion luminosity is radiated in lower wavelength regimes instead of X-rays and can range from 10 to 30 (NGC 1097, NGC 4594, M87; Ho 1999) to 1000 (Sgr A\*, Barrière et al. 2014).

### 2.3. X-Ray Spectral Fitting

We analyzed the X-ray data of our black hole XRB sample taken within a few days of each radio spectrum and systematically fit each X-ray spectrum using standard phenomenological models that include a disk blackbody from an accretion disk and a power law with a high energy cut-off. The *RXTE*/*PCA*, *RXTE*/*HEXTE*, and *Swift*/*XRT* (when available) spectra were fitted in *ISIS* (Houck & Denicola 2000) simultaneously with a suitable spectral model (see Table 3) with a constant offset between the spectra from different detectors to correct for calibration differences (if both *RXTE*/*HEXTE* clusters were present, these were added as separate data sets with individual constants). The best fit and its 90% errors on the parameters were determined simultaneously by *ISIS* until the  $\chi^2$  sufficiently converged. We also estimated the X-ray luminosity in Eddington units ( $L_X/L_{\text{Edd}}$ ) by integrating the unabsorbed X-ray flux from 3.5 to 200 keV (normalizing the *RXTE*/*HEXTE* spectra to the level of the *RXTE*/*PCA* spectra) with estimates of the distances and masses of the black holes (Russell et al. 2013a). For both MAXI J1659–152 and MAXI J1836–194 we adopt a distance of 8 kpc and the mass of the black hole as  $10 M_\odot$  (in Russell et al. 2014 they estimated the mass of the black hole as  $4\text{--}15 M_\odot$ , and the distance as  $4\text{--}10$  kpc).  $L_X/L_{\text{Edd}}$  is proportional to  $\dot{m}$  given a bolometric correction factor that relates  $L_X$  to bolometric luminosity of the whole accretion process. The bolometric correction factor is usually unknown, but in the case of XRBs most of the accretion luminosity is radiated in the X-ray regime and thus the value of the bolometric correction factor is likely small (1–5). Obviously, the obtained  $L_X/L_{\text{Edd}}$  values are crude estimates and should be taken as accurate to an order of magnitude. The X-ray model parameters for the AGN sample were taken from the literature (see Table 2).

## 3. RESULTS

Comparing the jet properties with the properties of the plasma close to the black hole, we found a relationship among the X-ray power-law photon index, the jet break frequency, and the “excess luminosity” over the radio luminosity (see below). The photon indices for black hole systems are plotted against the break frequencies in Figure 2 (circles and vertical bars representing XRBs with one



**Figure 1.** SEDs from our sample of AGNs. The low-energy part of each SED is fitted with a broken power-law model (red dashed line) with four free parameters, which are the spectral slope for the optically thin part of the spectrum,  $a_0$ , the spectral slope for the optically thick part of the spectrum,  $a_t$ , the break frequency,  $\nu_b$ , and the flux at the break frequency,  $S_b$ . Low-angular resolution measurements (>1" apertures) are depicted as gray spikes and high-angular resolution measurements (<0".4 apertures) as black dots. Only the high-angular resolution data are used to fit the model, with the exception of NGC 1052, where the low-angular resolution data are clearly dominated by the active nucleus below  $\sim 3 \times 10^{13}$  Hz.

individual XRB, MAXI J1836–194, highlighted in orange and green triangles representing AGNs). There is a clear anti-correlation in the data, encompassing the sample of XRBs, an individual XRB, and the sample of AGNs, spanning five orders of magnitude in the jet break frequency. To better quantify this anti-correlation we calculated the correlation coefficient using Monte Carlo methods (Curran 2014) (Table 4). The Monte Carlo method involves creating  $M$  (here  $M = 10^6$ ) new data sets based on the original data that

are composed of  $N$  data pairs of  $\Gamma$  and  $\nu_b$ . Each new data set consists of randomly chosen pairs from the original data, such that some of the original pairs may appear more than once or not at all. In addition, the pairs in a given new data set are randomly perturbed by random sampling from the normal distribution with means and standard deviations according to the original pairs or from the uniform distribution in the case of a range of frequencies as described above. Taking into account the whole data set, the correlation coefficient

**Table 2**  
The Jet Break Values, Radio and Jet Break Fluxes, and Literature Values for Power-law Photon Indices and X-Ray Luminosities of the AGN Sample

| Source Name | Type      | $z$   | $\log \nu_b(1+z)$<br>(Hz) | $S_{\nu,b}$<br>(mJy) | $S_{\nu,5 \text{ GHz}}$<br>(mJy) | $\log L_b/L_{5 \text{ GHz}}$ | $\Gamma$        | $L_X/L_{\text{Edd}}$ | References |
|-------------|-----------|-------|---------------------------|----------------------|----------------------------------|------------------------------|-----------------|----------------------|------------|
| NGC 1052    | LINER 1.9 | 0.005 | $13.12 \pm 0.01$          | $780 \pm 40$         | $2360 \pm 236$                   | $2.94 \pm 0.08$              | $1.67 \pm 0.40$ | $2 \times 10^{-5}$   | T02, W02   |
| NGC 1097    | LINER 1   | 0.004 | $13.36 \pm 0.06$          | $61 \pm 4$           | $3.3 \pm 0.3$                    | $4.93 \pm 0.12$              | $1.66 \pm 0.12$ | $3 \times 10^{-6}$   | T02, L06   |
| M87         | LINER 1   | 0.004 | $11.60^{+0.25}_{-0.60}$   | $1400 \pm 300$       | $3160 \pm 316$                   | $1.55 \pm 0.56$              | $2.40 \pm 0.11$ | $2 \times 10^{-6}$   | G09        |
| NGC 4594    | LINER 2   | 0.003 | $12.49 \pm 0.03$          | $410 \pm 20$         | $74 \pm 7$                       | $3.53 \pm 0.09$              | $1.89 \pm 0.17$ | $10^{-6}$            | T02        |
| Sgr A*      | ...       | 0     | $11.90 \pm 0.30$          | $4000 \pm 1000$      | $750 \pm 150$                    | $2.93 \pm 0.28$              | $2.35 \pm 0.18$ | $7 \times 10^{-12}$  | B14        |
| Cyg A       | FR-II     | 0.056 | $13.28 \pm 0.02$          | $223 \pm 41$         | $373 \pm 37$                     | $3.36 \pm 0.14$              | $1.52 \pm 0.12$ | 0.001                | Y02, T03   |
| 3C 120      | FSRQ      | 0.03  | $12.71 \pm 0.03$          | $1000 \pm 100$       | $3620 \pm 362$                   | $2.45 \pm 0.12$              | $1.85 \pm 0.05$ | 0.009                | Z01, P14   |

**Note.** The columns are: (1) source name; (2) AGN classification; (3) redshift; (4) logarithm of the redshift-corrected jet break frequency; (5) flux density at the jet break; (6) radio flux density at 5 GHz; (7) the excess luminosity over the radio luminosity up to the jet break  $L_b/L_{5 \text{ GHz}} = \nu_b S_{\nu,b}/\nu_{5 \text{ GHz}} S_{\nu,5 \text{ GHz}}$ ; (8) X-ray power-law photon index and its 90% error; (9) 2–10 keV unabsorbed X-ray luminosity in units of Eddington luminosity; and (10) references for the power-law photon indices, X-ray luminosities, and black hole masses (T02: Terashima et al. 2002, W02: Woo & Urry 2002, L06: Lewis & Eracleous 2006, G09: González-Martín et al. 2009, B14: Barrière et al. 2014, Y02: Young et al. 2002, T03: Tadhunter et al. 2003, Z01: Zdziarski & Grandi 2001, P14: Pozo Nuñez et al. 2014). The jet break values and radio and jet break fluxes of Sgr A\* are taken from Beckert & Falcke (2002).

**Table 3**  
X-Ray Spectral Fits of XRBs

| Source         | <i>RXTE</i> ObsID             | <i>Swift</i> ObsID | Time (MJD)        | Model     | $n_H$           | $\Gamma$        | $\chi^2_{\text{red}}/\text{dof}$ | $L_X/L_{\text{Edd}}$ |
|----------------|-------------------------------|--------------------|-------------------|-----------|-----------------|-----------------|----------------------------------|----------------------|
| 4U 1543–47     | 70124-02-12-00                | ...                | 52490.12          | PL+G      | $0.25^c$        | $1.73 \pm 0.03$ | 0.70/39                          | $2 \times 10^{-3}$   |
| Cyg X–1        | 91096-01-06-00                | ...                | 53513.04          | EE(CPL+G) | $0.6^c$         | $1.70 \pm 0.01$ | 1.34/121                         | $8 \times 10^{-3}$   |
| GS 1354–64     | 20431-01-03-00                | ...                | 50774.39          | E(CPL+G)  | $0.9^c$         | $1.39 \pm 0.01$ | 1.35/111                         | $6 \times 10^{-1}$   |
| GX 339–4       | 20181-01-05-00                | ...                | 50636.34          | E(PL+G)   | $0.6^c$         | $1.55 \pm 0.01$ | 1.08/70                          | $2 \times 10^{-2}$   |
|                | 95409-01-09-03                | ...                | 55266.78          | E(PL+G)   | $0.6^c$         | $1.57 \pm 0.01$ | 1.83/52                          | $2 \times 10^{-1}$   |
|                | 96409-01-09-00                | ...                | 55617.54          | E(PL)     | $0.6^c$         | $1.62 \pm 0.04$ | 0.78/50                          | $7 \times 10^{-3}$   |
| MAXI J1659–152 | 95358-01-02-00 <sup>a</sup>   | 00434928005        | 55467.04/55467.30 | D+PL+3G   | $0.30 \pm 0.01$ | $1.93 \pm 0.01$ | 1.63/403                         | $7 \times 10^{-2}$   |
|                | 95358-01-02-01                | 00434928007        | 55468.08/55468.22 | D+PL+3G   | $0.33 \pm 0.01$ | $2.08 \pm 0.01$ | 1.42/237                         | $5 \times 10^{-2}$   |
|                | 95358-01-03-00                | 00434928009        | 55470.24/55470.24 | D+PL+3G   | $0.32 \pm 0.01$ | $2.17 \pm 0.01$ | 1.49/292                         | $6 \times 10^{-2}$   |
|                | 95108-01-05-00                | 00434928011        | 55472.07/55472.11 | D+PL+3G   | $0.34 \pm 0.01$ | $2.20 \pm 0.03$ | 1.05/203                         | $5 \times 10^{-2}$   |
|                | 95108-01-11-00 <sup>a</sup>   | 00434928013        | 55474.57/55474.12 | D+PL+3G   | $0.32 \pm 0.01$ | $2.24 \pm 0.03$ | 1.32/295                         | $5 \times 10^{-2}$   |
|                | 95108-01-18-01                | 00434928017        | 55477.00/55477.12 | D+PL+G    | $0.39 \pm 0.01$ | $2.15 \pm 0.02$ | 1.59/237                         | $6 \times 10^{-2}$   |
|                | 95118-01-06-00 <sup>a</sup>   | 00031843003        | 55489.26/55489.04 | D+PL      | $0.33 \pm 0.01$ | $2.15 \pm 0.03$ | 1.65/269                         | $2 \times 10^{-2}$   |
| MAXI J1836–194 | 96371-03-03-00 <sup>a</sup>   | 00032087002        | 55806.48/55805.23 | D+PL+G    | $0.32 \pm 0.03$ | $2.13 \pm 0.02$ | 1.15/137                         | $1 \times 10^{-2}$   |
|                | 96438-01-02-00 <sup>a</sup>   | 00032087013        | 55821.84/55821.69 | D+PL+G    | $0.33 \pm 0.01$ | $2.38 \pm 0.06$ | 1.53/189                         | $1 \times 10^{-2}$   |
|                | 96438-01-03-05                | 00032087017        | 55830.88/55830.18 | D+PL+G    | $0.23 \pm 0.03$ | $1.82 \pm 0.02$ | 1.28/125                         | $2 \times 10^{-2}$   |
|                | 96438-01-05-05                | 00032087024        | 55846.55/55846.84 | D+PL+G    | $0.16 \pm 0.06$ | $1.56 \pm 0.02$ | 1.40/70                          | $9 \times 10^{-3}$   |
|                | 96438-01-07-04                | 00032087029        | 55861.51/55861.21 | PL+G      | $0.19 \pm 0.02$ | $1.63 \pm 0.02$ | 0.79/55                          | $9 \times 10^{-3}$   |
| V4641 Sgr      | 80054-08-01-01 <sup>a,b</sup> | ...                | 52857.37          | EE(PL+G)  | $0.4^c$         | $0.93 \pm 0.03$ | 1.10/28                          | $3 \times 10^{-2}$   |
| XTE J1118+480  | 50137-01-06-00                | ...                | 51649.04          | PL+G      | $0.01^c$        | $1.72 \pm 0.01$ | 1.10/74                          | $1 \times 10^{-3}$   |
|                | 90011-01-01-08                | ...                | 53386.34          | PL        | $0.01^c$        | $1.76 \pm 0.01$ | 0.95/56                          | $5 \times 10^{-4}$   |
| XTE J1550–564  | 50135-01-12-00                | ...                | 51696.47          | E(PL+G)   | $0.65^c$        | $1.59 \pm 0.01$ | 0.72/55                          | $4 \times 10^{-3}$   |
| XTE J1752–223  | 95702-01-11-01 <sup>a</sup>   | ...                | 55377.21          | PL        | $0.65^c$        | $1.87 \pm 0.06$ | 1.12/29                          | $6 \times 10^{-3}$   |

**Notes.** The columns are: (1) source name, (2) the ID number of the *RXTE* pointing, (3) the ID number of the *Swift* pointing, (4) starting time of the pointings in MJD (*RXTE/Swift*), (5) the model fitted to the data, (6) the column density  $n_H$  of the absorption component PHABS in units of  $10^{22} \text{ cm}^{-2}$ , (7) the best fit value of the X-ray power-law photon index and its 90% confidence interval, (8) the reduced  $\chi^2$  value and the number of degrees of freedom of the fit, and (9) 3.5–200 keV unabsorbed X-ray luminosity in units of Eddington luminosity. Model key: D = DISKBB, PL = POWERLAW, CPL = CUTOFFPL, G = GAUSSIAN, E = EDGE.

<sup>a</sup> Only the *RXTE*/PCA spectrum is used for the fit.

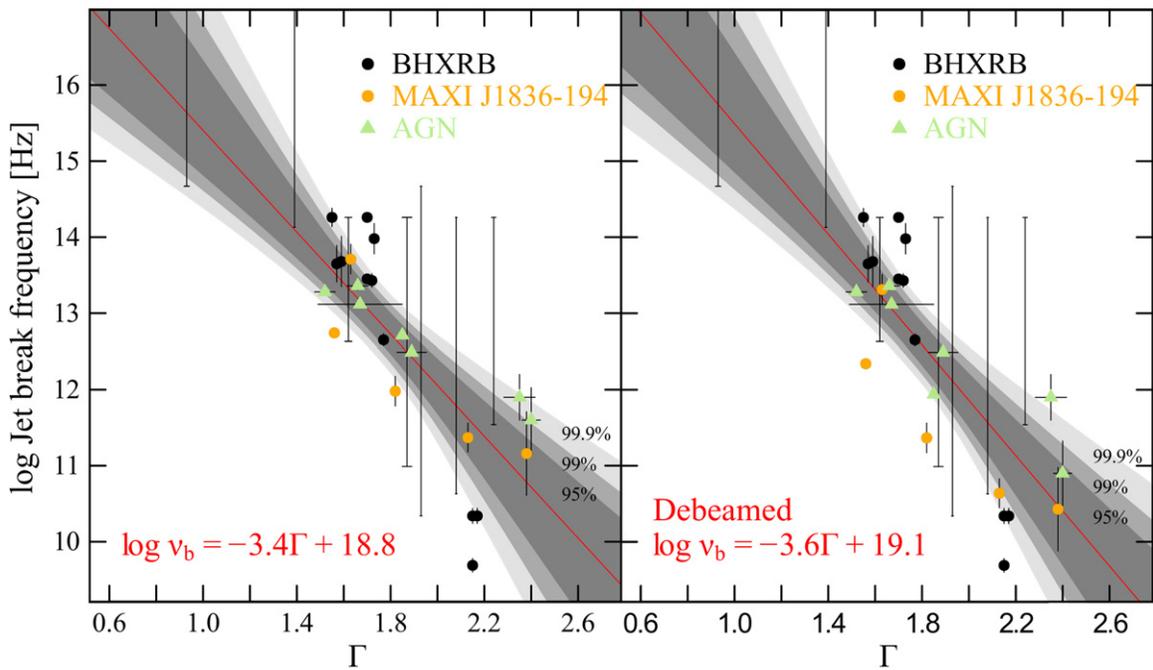
<sup>b</sup> The first part of this pointing exhibits strong flares and rapid spectral variability (Maitra & Bailyn 2006) and thus we select the second part of the pointing, which is more stable, for fitting.

<sup>c</sup> Fixed value in the model fitting.

amounted to  $-0.75$  with a significance of  $4.6\sigma$  (see Table 4 for statistics of subsets). We also performed a linear least-squares regression which is calculated for all randomly perturbed samples and we note the slope and intercept of each fit, and produce 95%, 99%, and 99.9% limits on the possible regressions:  $\log(\nu_b/\text{Hz}) = -3.4^{+0.9}_{-1.4}\Gamma + 18.8^{+2.5}_{-1.6}$  (the errors are the 95% confidence interval; see also Figure 2). This anti-correlation clearly shows an intimate connection between the

region near the black hole (responsible for the X-ray emission) and the jet emission (responsible for the  $\nu_{\text{break}}$ ).

With the knowledge of the frequency and the flux density of the jet break, we can estimate the excess luminosity caused by the variable break frequency over the radio luminosity by measuring  $L_b/L_{5 \text{ GHz}} = \nu_b S_{\nu,b}/\nu_{5 \text{ GHz}} S_{\nu,5 \text{ GHz}} = (\nu_b/\nu_{5 \text{ GHz}})^{1+\alpha_{\text{inick}}}$ , which varies between sources by six orders of magnitude (see Tables 1 and 2). This result clearly



**Figure 2.** Left: the jet break frequency as a function of the X-ray power-law photon index for black hole systems. Right: the same as in the left panel but the jet breaks have been corrected for beaming. In both panels black circles and vertical bars represent the data from XRBs and green triangles are the data from AGNs. One individual XRB, MAXI J1836–194, is highlighted in orange with multiple observations across a state change from hard to intermediate X-ray. The red solid line shows the median of the Monte Carlo bootstrap samples with its formula in the lower left corner of the figure. The shaded gray regions show the 95%, 99%, and 99.9% confidence intervals on the linear regressions. The XRB data with only upper or lower limits on the jet break frequency available are shown as vertical bars depicting the range of the frequency we are considering as a conservative estimate of the break frequency, with the horizontal bars at the ends of the vertical bar depicting the  $1\sigma$  error on the X-ray power-law photon index.

**Table 4**

Spearman Rank Coefficients for the Break Frequency vs. X-Ray Power-law Photon Index Correlation

|          | $R$   | 95% Conf.       | Null (%)           |
|----------|-------|-----------------|--------------------|
| ALL      | -0.76 | (-0.87)–(-0.52) | $2 \times 10^{-4}$ |
| ALL*     | -0.76 | (-0.88)–(-0.48) | $4 \times 10^{-3}$ |
| ALL(db)  | -0.75 | (-0.86)–(-0.50) | $8 \times 10^{-4}$ |
| ALL(db)* | -0.75 | (-0.88)–(-0.47) | $6 \times 10^{-4}$ |
| XRb      | -0.77 | (-0.89)–(-0.46) | $8 \times 10^{-3}$ |
| XRb*     | -0.76 | (-0.91)–(-0.39) | 0.08               |
| XRb(db)  | -0.76 | (-0.89)–(-0.44) | 0.01               |
| XRb(db)* | -0.76 | (-0.91)–(-0.37) | 0.1                |
| AGN      | -0.86 | (-1.00)–(-0.32) | 0.6                |
| AGN(db)  | -0.86 | (-0.96)–(-0.29) | 0.7                |

**Note.** We use 10 different samples to study the correlation: XRbS, AGNs, and both combined (denoted as “ALL”) including the observations with ranges on the break frequency, the samples that do not include the ranges (marked with a star [\*]), and the samples that are debeamed (marked with letters “db”). The correlation coefficient and the best linear fit and their confidence limits are determined by Spearman’s rank Monte Carlo bootstrapping, where  $10^6$  sets of simulated data are created by the bootstrap function and we note the most likely correlation coefficient, 95% confidence interval on the correlation coefficient, and the percentage of cases where the null hypothesis is valid (see the text for a more detailed discussion). The columns are: (1) sample name, (2) the Spearman rank correlation coefficient, (3) 95% confidence limit on the correlation coefficient, and (4) the percentage of cases where the null hypothesis is valid for the Monte Carlo bootstrap method.

demonstrates that the jet luminosities should be recalibrated taking the break frequency into account. Due to the relation between the break frequency and the X-ray power-law photon index, the excess luminosity depends on the value of the

X-ray power-law photon index and can be estimated as  $\log(L_b/\text{erg s}^{-1}) = \log(L_{5\text{GHz}}/\text{erg s}^{-1}) - 3.5^{+0.9}_{-1.0}\Gamma + 9.8^{+2.0}_{-1.6}$  (the errors are the 95% confidence interval) by performing a linear least-squares regression using the above Monte Carlo methods between the excess luminosity and X-ray power-law photon index.

### 3.1. The Effect of Beaming

Generally, it is thought that the bulk Lorentz factors of the jets in XRbS should be around 2 (Fender et al. 2004; Casella et al. 2010). However, direct observational evidence is largely missing in this regard. Together with the inclinations that are generally  $>30^\circ$  the Doppler factors of our XRb sample can be estimated as  $\delta \sim 1$  apart from one source. MAXI 1836–194 is suspected to have a slightly higher bulk Lorentz factor ( $\Gamma_b = 1\text{--}4$ ) and a jet angle oriented close to our line-of-sight ( $4^\circ\text{--}15^\circ$ , Russell et al. 2015). It is likely that the Doppler factor is variable depending on the X-ray flux of the source (see Figure 10 in Russell et al. 2015). We use Russell et al.’s estimation of the bulk Lorentz factors depending on the X-ray flux and the jet angle of  $10^\circ$  and correct for the effect of Doppler boosting ( $\delta = 2\text{--}5$ ) in Figure 2 (right panel) and Table 4.

As in XRbS, we consider the effect of beaming to be small for our sample of AGNs. The only beaming candidates in our sample are M87 and 3C 120, where the Doppler factors have been estimated to be 2–5 (Wang & Zhou 2009) and 5.9 (Hovatta et al. 2009), respectively. Similar to the case of MAXI 1836–194 mentioned above, we correct for the effect of Doppler boosting for M87 and 3C 120 in Figure 2 (right panel)

and Table 4. In general, the debeaming does not have a big impact on the correlation.

#### 4. DISCUSSION

In the above we have shown that an intimate connection exists between the jet break frequency and X-ray power-law photon index. The observed optically thick radio spectra can be produced by many models ranging from non-thermal, hybrid, or thermal distributions of electrons in a single acceleration episode, distributed acceleration along the jet, or an internal shock model (Blandford & Königl 1979; Falcke & Markoff 2000; Stawarz & Ostrowski 2002; Pe’er & Casella 2009; Böttcher & Dermer 2010; Malzac 2013). Thus, more observables are clearly needed to single out a favorable scenario and observations around and above the jet break are crucial in this regard. With the observed correlation we can state that the conditions dictating the jet break and jet spectrum are set by the X-ray emitting region or that they are both driven by an underlying parameter. This idea is further supported by the correlation found earlier between the radio and X-ray luminosities (Hannikainen et al. 1998; Corbel et al. 2003; Gallo et al. 2003; Fender & Gallo 2014). The hard X-ray emission is most likely produced by inverse Compton scattering, although a synchrotron origin has been suggested for some individual XRBs in the hard X-ray state for very low accretion rates ( $\dot{m} \sim 10^{-4}$ ; Russell et al. 2010); even in this case, the X-ray power-law photon index is similar to that produced by inverse Compton scattering. In the future, broadband observations of quiescent, low accretion rate black hole XRBs would be an ideal test of the validity and extent of the correlation. The inverse Compton spectrum depends on the energy distribution of the electrons, optical depth of the medium, and the energy density of the seed photons. Qualitatively, the correlation could be explained by an increase in the amount of seed photons that would produce more Compton scattering. The increased scattering cools down the electron population close to the black hole that is then translated to decreasing values of the jet break in a similar fashion to the blazar sequence (Ghisellini et al. 1998).

Different semi-analytical models and magnetohydrodynamical simulations of jet formation assume different initial conditions, e.g., magnetically arrested disk (Tchekhovskoy et al. 2011) simulations assume that jet launching is magnetically dominated whereas RIAF simulations assume it is not (Yuan et al. 2003). How magnetically dominated the jets are, what they contain, and how particles are accelerated, are all uncertain at this stage. Thus, the initial conditions for jet formation are currently something that all forms of modeling, semi-analytical as well as simulations (e.g., particle-in-cell simulations; Sironi & Spitkovsky 2014), essentially have to insert by hand. Our result provides an observational connection between the characteristics of the particle distribution and properties of the jet and thus may narrow down the initial condition parameter space for jet formation models.

##### 4.1. The Effect of the Black Hole Mass

Due to the mass-scaling properties of black holes, it has been suggested that accretion physics scales globally with black hole mass (Heinz & Sunyaev 2003). Such a scaling is supported by the discovery of the Fundamental Plane of black hole activity, an observed relationship between the X-ray luminosity, radio

luminosity, and black hole mass in the hard X-ray state (i.e., compact jet producing) black hole XRBs and AGNs (Merloni et al. 2003; Falcke et al. 2004; Plotkin et al. 2012). According to theoretical scaling relations (Heinz & Sunyaev 2003), the break frequency between the flat/inverted power law and optically thin power law scales with the dimensionless accretion rate  $\dot{m}$  (defined as the mass accretion rate divided by the Eddington rate) and black hole mass  $M_{\text{BH}}$  as  $\nu_b \propto M_{\text{BH}}^{-(p+2)/(2p+8)} \dot{m}^{(p+6)/(2p+8)}$  for sources with  $\dot{m} \lesssim$  a few percent, which reduces to  $\nu_b \propto M_{\text{BH}}^{-1/3} \dot{m}^{2/3}$  assuming the electron distribution is a power law with index  $p=2$  (the above scaling relation is not very sensitive to the value of  $p$ ). In this case, the difference between AGNs and XRBs in  $\nu_b$  is predicted to be three orders of magnitude in frequency (for nominal values of  $M_{\text{BH,AGN}} \sim 10^9 M_{\odot}$  and  $M_{\text{BH,XRB}} \sim 10 M_{\odot}$ ) for the same  $\dot{m}$ . The difference would be even larger if we take the mass accretion rate into account and assume that the XRBs have systematically larger  $\dot{m}$  than AGNs, which is likely the case as the XRBs in our sample are fairly luminous, with  $\dot{m} > 10^{-3}$ . Such a large difference is not observed for black hole masses in the range from  $10 M_{\odot}$  to  $6 \times 10^9 M_{\odot}$  covered by our sample, so alternative explanations are required.

##### 4.2. The Effect of the Mass Accretion Rate

It is known that the mass accretion rate does not vary substantially over the state transition for XRBs—the X-ray flux, which is a proxy for the mass accretion rate, stays at a similar level when XRBs make a transition from the hard X-ray state to the soft X-ray state and vice versa (Körding et al. 2006). However, according to the correlation between the jet break and X-ray power-law photon index, the largest change in both parameters occurs specifically during the transition with  $\Gamma$  changing from  $\sim 1.6$  to  $\sim 2.4$  and  $\nu_b$  from  $\sim 10^{14}$  to  $\sim 10^{11}$  Hz, whereas during the hard state ( $\Gamma \sim 1.6$ ) the break frequency and X-ray power-law photon index remain approximately constant while the mass accretion rate is changing by orders of magnitude. In addition, one XRB in the sample (MAXI J1836–194) shows a change in the frequency of the jet break by three orders of magnitude over state transitions (orange dots in Figure 2), demonstrating that the effect of the black hole mass and spin on the jet properties including the jet power is negligible compared to the accretion state changes. Therefore the observed correlations could provide evidence for the internal physics and/or accretion mode being the jet power driver for both XRBs and low accretion rate AGNs.

##### 4.3. Estimating the Total Jet Luminosity

Estimating the total jet luminosity (ignoring beaming) from the jet spectral energy distribution also requires the knowledge of the flux above the jet break frequency. Above the break the optically thin spectrum reveals the underlying particle distribution: either a power law ( $\alpha_{\text{thin}} \sim -0.7$  or steeper depending on the cooling mechanisms) or a quasi-thermal distribution ( $\alpha_{\text{thin}} < -1.0$ ). If the optically thin spectrum is shallow enough ( $\alpha_{\text{thin}} > -1.0$ ) the jet luminosity then depends on the emission above the break frequency, up to the electron cooling break frequency. In some black hole XRBs the optically thin spectrum is characterized with  $\alpha_{\text{thin}} > -1.0$ , however, in only one case is there reliable evidence of how far the optically thin spectrum extends (Russell et al. 2014),

corresponding roughly to  $10^3\nu_b$ . As a conservative estimate, the total luminosity normalized to the 5 GHz radio luminosity for XRBs would be less than one order of magnitude greater than  $L_b/L_{5\text{ GHz}}$  when taking the mean optically thin spectral index ( $\alpha_{\text{thin}} \sim -0.85$ ) from our sample of XRBs when constrained, and assuming that the cooling break is located at  $\leq 10^4\nu_b$ . Thus, we can consider that the excess luminosity gives an order of magnitude estimate of the total jet luminosity.

## 5. CONCLUSIONS

We have collected an unprecedented data set of multi-wavelength SEDs from the core of the compact jet in stellar-mass and supermassive black holes in addition to (near-simultaneous in the case of XRBs) X-ray observations. We have discovered a correlation between the X-ray power-law photon index of the corona and the jet break frequency and a resulting correlation between the X-ray power-law photon index and the “excess luminosity” over the radio luminosity, suggesting an intrinsic connection between the plasma close to the black hole and the outflow properties. Further considerations are needed to determine the nature of the jet break and the spectral slope of the optically thin part of the jet spectrum, which can be achieved by detailed modeling of the SEDs of the sources. Our results indicate that the jet production and properties (possibly coupled to the X-ray spectral state changes) are closely related to changes in the corona/hot flow. Therefore this result will serve as a benchmark that should be reproduced by jet formation models and provides observational clues about the connection between particle acceleration and properties of the jet. Our results support the notion that the coronae are black-hole-mass- and spin-independent features of black hole accretion whose presence is essential in producing powerful jets on all scales.

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