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Broad-band modelling of short gamma-ray bursts with energy injection from magnetar spin-down and its implications for radio detectability

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ABSTRACT
The magnetar model has been proposed to explain the apparent energy injection in the X-ray light curves of short gamma-ray bursts (SGRBs), but its implications across the full broad-band spectrum are not well explored. We investigate the broad-band modelling of four SGRBs with evidence for energy injection in their X-ray light curves, applying a physically motivated model in which a newly formed magnetar injects energy into a forward shock as it loses angular momentum along open field lines. By performing an order of magnitude search for the underlying physical parameters in the blast wave, we constrain the characteristic break frequencies of the synchrotron spectrum against their manifestations in the available multiwavelength observations for each burst. The application of the magnetar energy injection profile restricts the successful matches to a limited family of models that are self-consistent within the magnetic dipole spin-down framework. We produce synthetic light curves that describe how the radio signatures of these SGRBs ought to have looked given the restrictions imposed by the available data, and discuss the detectability of these signatures with present-day and near-future radio telescopes. Our results show that both the Atacama Large Millimeter Array (ALMA) and the upgraded Very Large Array are now sensitive enough to detect the radio signature within two weeks of trigger in most SGRBs, assuming our sample is representative of the population as a whole. We also find that the upcoming Square Kilometre Array will be sensitive to depths greater than those of our lower limit predictions.

Key words: gamma-ray burst: general – stars: magnetars.

1 INTRODUCTION
Gamma-ray bursts (GRBs) are extreme outbursts of electromagnetic radiation, releasing energies of the order of $10^{48}$–$10^{52}$ erg in a relativistic outflow, when collimation of this outflow is accounted for (e.g. Cenko et al. 2011). They are divided into two classes: long and short GRBs (LGRBs and SGRBs, respectively), sitting at either side of a $T_{90} \sim 2$ s divide (Kouveliotou et al. 1993), where $T_{90}$ is the time in which the cumulative counts increase from 5 to 95 per cent of the background level. SGRBs typically emit a short ($T_{90} < 2$ s) prompt spike of gamma radiation which decays away into an X-ray afterglow, but $\sim$50 per cent of cases (e.g. Rowlinson et al. 2013) exhibit a flat X-ray plateau which lasts for around a thousand seconds. In some SGRBs, the initial decay is interrupted by a prolonged, softer, rebrightening in the high-energy light curves. This extended emission (EE) was discovered in $\sim 1/3$ of SGRBs (Norris & Bonnell 2006). It usually begins $\lesssim 10$ s after the trigger, and while it often has a lower luminosity than the prompt emission, it can last for a few hundred seconds, implying that the total energy contained is comparable (Perley et al. 2009). Those bursts that were believed to exhibit EE were catalogued by Norris, Gehrels & Scargle (2010), and that sample was updated and expanded in Gompertz et al. (2013).

One of the leading models for SGRBs is the merger of two compact objects: some combination of black holes, neutron stars (NS), and/or white dwarfs (Paczynski 1986; Rosswog, Ramirez-Ruiz & Davies 2003; Belczynski et al. 2006; Chapman et al. 2007). Two possible post-merger remnants have been proposed as the central engine of SGRBs: a black hole (Woosley 1993; Kumar, Narayan & Johnson 2008), or a rapidly-rotating, highly-magnetized NS known as a magnetar (Gao & Fan 2006; Metzger, Quataert & Thompson 2008; Bucciantini et al. 2012; Cheng & Yu 2014; Lü & Zhang 2014; Metzger & Piro 2014). In this paper, we consider the magnetar case. These magnetars can have dipole fields of the order of $10^{15}$ G at birth (Duncan & Thompson 1992; Thompson &
Duncan 1995; Kouveliotou et al. 1999; Esposito et al. 2010) and spin at periods of around 1 ms (Lattimer & Prakash 2004; Metzger et al. 2011; Giacomazzo & Perna 2013). In this context, the prompt emission is considered to be due to relativistic jets powered by the initial merger, and the X-ray plateau seen in ~50 per cent of bursts (Rowlinson et al. 2013) is the result of energy injection into the radiating shock via magnetic dipole spin-down. The luminosity and duration of the X-ray plateau is then just a function of the magnetar’s dipole field strength and spin period (Zhang & Mészáros 2001), and the observed anticorrelation between them (Dainotti, Cardone & Capozziello 2008) has been suggested as a natural prediction of the model (Rowlinson et al. 2014). The magnetar model has been applied to LGRBs (Zhang & Mészáros 2001; Troja et al. 2007; Lyons et al. 2010; Dall’Osso et al. 2011; Metzger et al. 2011; Bernardini et al. 2012), SGRBs (Fan & Xu 2006; Rowlinson et al. 2010, 2013), and EE GRBs (Metzger et al. 2008; Bucciantini et al. 2012; Gompertz et al. 2013; Gompertz, O’Brien & Wynn 2014).

Most applications of the magnetar model on GRB light curves have focused on the X-ray regime. In these efforts, bolometric light curves are constructed based on the X-ray and soft gamma-ray light curves, and they are modelled with a time-dependent luminosity that reflects the magnetic dipole spin-down of a rapidly rotating magnetar (Zhang & Mészáros 2001). It is typically assumed that the spin-down luminosity evolution is converted to an observed light-curve evolution with a certain efficiency factor. This method does not provide information on light curves at lower frequencies, although extending the spectrum to the optical bands has been attempted (e.g. Rowlinson et al. 2013). In this paper, we aim to perform broad-band modelling of SGRB afterglows, using the available light curves at X-ray, optical, and in some cases radio bands, all within the magnetar model. Instead of assuming a direct conversion of magnetic dipole spin-down into light curves at different frequencies, we treat this spin-down as energy injection into the shocks at the front of the relativistic outflow. These shocks emit broad-band synchrotron radiation, and we calculate the light curves across the spectrum as is typically done for GRB afterglows (e.g. Sari, Piran & Narayan 1998; Wijers & Galama 1999) but with a time-varying energy term governed by magnetic dipole spin-down. We apply this model to four SGRBs to illustrate the method, and we give ranges for the physical parameters of the shocks and the magnetar for these four sources. We also discuss the detectability of the range of light curves our models predict for radio emission in the context of current and next generation radio telescopes.

In Section 2, we describe the GRBs in our sample and the broad-band data we collected for all four of them. The model and methods we used to construct the broad-band light curves are given in Section 3. The results of modelling the light curves are shown in Section 4, followed by a discussion and conclusions in Sections 5–7.

### 2 DATA SAMPLE

We collected a sample of four SGRBs with good X-ray observations exhibiting a clear plateau, for which there were contemporaneous optical observations in at least one filter and an identified redshift. Radio observations were not essential, but were a welcome bonus. Our sample represents those SGRBs with the best data availability to test the analysis and introduce the model, but is not an exhaustive list of all SGRBs that satisfy the selection criteria. The classification of GRB 060614 is uncertain (Gehrels et al. 2006), but we include it as an EE GRB here due to the lack of an associated supernova to deep limits (Della Valle et al. 2006; Gal-Yam et al. 2006) as would be expected for an LGRB. Analysis by Zhang et al. (2007) also suggests that this burst is linked to the short class. The X-ray data used here were taken by the Swift X-Ray Telescope (XRT; Burrows et al. 2005) and were downloaded from the UK Swift Science Data Centre (UKSSDC) archives (Evans et al. 2007, 2009). We used the 0.3–10 keV flux light curves, corrected for absorption using a ratio of (counts-to-flux unabsorbed)/(counts-to-flux observed). Details of the data reduction process can be found in Evans et al. (2007, 2009). The 0.3–10 keV flux light curves were then compressed into flux density light curves at 1.73 keV (the bandpass logarithmic mid-point) using the equations in Appendix A.

Table 1. UV, optical, IR, and radio data used. Photon indices $\Gamma$ are for the X-ray data, and come from the UKSSDC spectrum repository (Evans et al. 2007, 2009) which gives 90 per cent confidence interval errors. $E(B-V)$ values are from Schlafly & Finkbeiner (2011).

<table>
<thead>
<tr>
<th>GRB</th>
<th>$\Gamma$</th>
<th>$z$</th>
<th>$E(B-V)$</th>
<th>Reference</th>
</tr>
</thead>
<tbody>
<tr>
<td>051221A</td>
<td>1.95$^{+0.18}_{-0.17}$</td>
<td>0.5465$^{a}$</td>
<td>0.069</td>
<td>[1]</td>
</tr>
<tr>
<td>060614</td>
<td>1.78$^{+0.08}_{-0.09}$</td>
<td>0.1254$^{b}$</td>
<td>0.019</td>
<td>[2,3,4]</td>
</tr>
<tr>
<td>070714B</td>
<td>1.76$^{+0.28}_{-0.24}$</td>
<td>0.9224$^{c}$</td>
<td>0.141</td>
<td>[5] (A,B,C,D)</td>
</tr>
<tr>
<td>130603B</td>
<td>1.98$^{+0.15}_{-0.14}$</td>
<td>0.356$^{d}$</td>
<td>0.02</td>
<td>[6,7,8] (E)</td>
</tr>
</tbody>
</table>

outflow, which expands with time. Internal processes such as shocks between expanding shells of ejecta (Goodman 1986; Paczynski 1986) or magnetic turbulence (Zhang & Yan 2011) convert some of the kinetic energy of the blast wave into electromagnetic radiation, which is observed as the SGRB prompt emission. The blast wave sweeps up ambient particles as it expands into the interstellar medium (ISM), and eventually starts to slow down once it has accumulated sufficient mass. This deceleration radius, $R_{\text{dec}}$, marks the outer boundary for emission processes to be considered ‘internal’.

The interaction between the blast wave and the ISM forms a strong shock at the head of the ejecta, and a synchrotron emission spectrum is set up by the action of electrons traversing the shock front. This is the emission site of the afterglow. For simplicity, we consider here an adiabatic expansion, and invoke a homogeneous ambient medium, as is expected in the vicinity of an NS binary. However, for LGRBs or binaries with pulsar winds the local density profile can be different (see e.g. Starling et al. 2008; Curran et al. 2009).

3.1 Synchrotron emission

The relativistic blast wave accelerates electrons, which subsequently radiate synchrotron emission in accordance with their respective Lorentz factors, which are assumed to have a power-law distribution. There are two distinct emission regimes, dubbed ‘fast cooling’, in which the cooling time-scale of the electrons is less than the lifetime of the source, and ‘slow cooling’, in which the majority of electrons cool on longer time-scales than the source lifetime. (Sari et al. 1998; Wijers & Galama 1999).

The synchrotron spectrum has three characteristic break frequencies: the peak frequency, $\nu_{\text{peak}}$; the self-absorption frequency, $\nu_{\text{sa}}$; and the absorption frequency, $\nu_{\text{abs}}$, at which the medium changes from being optically thin to being optically thick. These breaks are not static in time, but change and evolve with the hydrodynamical expansion of the blast wave. Their position and evolution determines the phenomenology of the corresponding light curve at a given observation time. For simplicity, we consider here an adiabatic expansion, and invoke a homogeneous ambient medium, as is expected in the vicinity of an NS binary. However, for LGRBs or binaries with pulsar winds the local density profile can be different (see e.g. Starling et al. 2008; Curran et al. 2009).

3.2 Energy injection

The magnetar formed by the merger is initially rapidly spinning, with a spin period of the order of $1$ ms. After birth, it loses angular momentum in the form of magnetic dipole spin-down (Zhang & Mészáros 2001), resulting in energy being injected into the outflow and the forward shock for a sustained period, typically of the order of $1000$ s. This was investigated for LGRBs by Dall’Osso et al. (2011). The total energy injected into the shock at a time $t$ after merger is given by (cf. Zhang & Mészáros 2001)

$$E_{\text{d}}(t) = \frac{\eta L_0 t}{\left(1 + t/T_{\text{em}}\right)^2}. \quad (2)$$

The parameter $\eta$ accounts for our ignorance in the efficiency of the transfer of energy from the dipole to the forward shock, both in terms of radiative losses and beaming fraction. $L_0$ is the luminosity of the dipole plateau in erg s$^{-1}$ and $T_{\text{em}}$ is the point at which the plateau turns over, known as the characteristic spin-down time-scale. $L_0$ and $T_{\text{em}}$ are both derived from the underlying physical parameters of the magnetar:

$$L_{0,49} = B_{15}^2 P_{-5}^{-4} P_{6}^6, \quad (3)$$

$$T_{\text{em},3} = 2.05 I_{45} B_{15}^{-2} P_{-5}^{-3} P_{6}^6, \quad (4)$$

where $L_{0,49}$ is $L_0$ in units of $10^{49}$ erg s$^{-1}$ and $T_{\text{em},3}$ is $T_{\text{em}}$ in units of $10^3$ s. $L_0$ is the moment of inertia in units of $10^{35}$ g cm$^2$, and is $\approx 1$ (2) for a 1.4 (2.1) M$\odot$ NS. $P_6$ is the spin period in ms and $B_{15}$ is the dipole field strength in units of $10^{15}$ G. We set the NS radius $R_6 = 1$, since this is consistent with most equations of state (Lattimer & Prakash 2004). These relations place limits on the values of $L_0$ for a given $T_{\text{em}}$, principally through the break-up spin period for an NS (e.g. $P = 0.66$ ms for a 2.1 M$\odot$ NS; Lattimer & Prakash 2004). The upper limit placed on $L_0$ by $P$ is given by

$$L_{0,49} \leq 2.05 I_{45} T_{\text{em},3}^{-1} P_{-5}^{-2} P_{6}^6 \quad (5)$$

because $T_{\text{em}}$ is a fixed quantity for a given GRB.

The two EE bursts in our sample, GRB 060614 and GRB 070114B, are likely to also inject energy into the shock during the EE phase, although without a clear model for what EE is, it is difficult to say how much. To represent EE, we use the energy profile from Gompertz et al. (2014), who used a magnetic propeller to describe the emission feature. These magnetic propellers accelerate infalling material to super-Keplerian velocities, ejecting it from the system at relativistic speeds, where it subsequently shocks to produce electromagnetic radiation. The exact physics behind these models is largely irrelevant for our needs, but the accurate luminosity profile provides a convenient way to introduce EE injection to the system. The total energy in the forward shock at a time $t$ is then given by

$$E_{\text{EE}}(t) = E_k + E_{\text{EE}}(t) + E_{\text{d}}(t). \quad (6)$$

Here, $E_k$ represents the impulsive energy of the blast wave, and is tied to the prompt emission isotropic equivalent energy $E_{\gamma,\text{iso}}$ through a prefactor accounting for beaming and efficiency. $E_{\text{EE}}$ is the energy injected during EE, representing the luminosity profile from Gompertz et al. (2014) multiplied by another prefactor $\kappa$, again to account for beaming and efficiency. $E_{\text{d}}$ is the energy injected by dipole spin-down, given by equation (2). These energies are varied to obtain fits to the data, and the physical implications that the obtained values have for the central engine are discussed in Section 5.
3.3 X-ray and optical fitting

To perform least-squares fitting for broad-band GRB afterglows, one normally requires well-sampled light curves in the X-ray and optical bands, as well as at least two radio bands. Without radio observations, it is very difficult to locate \( v_s \), since this break is normally found at radio frequencies, and \( v_m \) and \( v_{\text{em},\text{max}} \) can only be constrained as a combination, rather than individually. Additionally, if \( v_c \) lies above the X-ray frequency then it too becomes poorly constrained. Because of this, large degeneracies can occur where the observed X-ray and optical light curves give combinations of \( v_m \) and \( v_{\text{em},\text{max}} \) that can be recreated by many different physical parameter values, each having very different implications for the positions of \( v_c \) and \( v_s \). Thus, any fitting can result in parameter uncertainties spanning several orders of magnitude. For our sample, the available data consist of a well-sampled X-ray light curve, as well as a sparsely sampled optical light curve (sometimes in multiple bands) and just one or two radio observations or limits at best per burst. This is insufficient for fitting in the traditional way, so we conduct an order of magnitude search of the parameter space within reasonable parameter limits.

Synthetic light curves are created through a combination of nine free parameters. Three are well constrained by the data: the characteristic spin-down time-scale \( T_{\text{em}} \), the jet break time \( t_b \), and the power-law index of the electron Lorentz factor distribution \( p \). The most constrained; this parameter sets the spectral slope, so the power-law index of the electron Lorentz factor distribution \( p \) can be used for both parameters; values of \( p \) in the range are unable to provide simultaneous fits to the X-ray and optical bands) and just one or two radio observations or limits at best per burst. This is insufficient for fitting in the traditional way, so we conduct an order of magnitude search of the parameter space within reasonable parameter limits.

The remaining six parameters are less constrained. They are \( \epsilon_e \), \( \epsilon_B \), \( n_0 \), \( L \) (where \( L = \eta L_0 \)), \( \kappa \) and \( E_0 \). We apply constraints to the range of allowed values for these parameters. \( \epsilon_B \) has been found to be as low as \( 10^{-8} \) (Barniol Duran 2014; Santana, Barniol Duran & Kumar 2014) and as a fraction can be as high as 1. In practice, \( \epsilon_e \) tends towards higher values than \( \epsilon_B \). We set an upper limit of 1, noting that \( \epsilon_e \) actually refers to the electron population that is emitting synchrotron radiation, rather than the electron population as a whole, and set a lower limit of \( 10^{-3} \) (Kumar 2000), \( n_0 \) is limited between \( 10^{-3} \) and \( 10^{0.7} \), in line with what has been found in these sources (Cenko et al. 2011). The upper limit of \( L \) is set by the argument in equation (5), and values of this parameter below \( \sim 10^{47} \text{ erg s}^{-1} \) are never energetic enough to match the data, so we set the lower limit as \( 10^{52} \text{ erg s}^{-1} \). Within these limits for \( L \), we find that EE ceases to have any influence on the light curve if \( \kappa \lesssim 10^{-2} \). If EE is isotropic, and the observed luminosity is only 1 per cent of the true energy (i.e. the conversion efficiency of kinetic to potential energy in the internal shocks is 1 per cent), then the energy delivered to the synchrotron shock front could be up to 100 times higher than observed in the light curve. In practice, however, the emission is (a) unlikely to be fully isotropic, (b) likely to shock more efficiently than 1 per cent, and (c) certain to be less than 100 per cent efficient at delivering its energy to the synchrotron shock front. For these reasons, we set the upper limit of \( \kappa \) at a still fairly generous factor of 10. Finally, we limit the energy in the shock from prompt emission to \( 10^{48} \text{ erg} < E_0 < 10^{52} \text{ erg} \). The arguments for these limits are identical to those used for \( \kappa \), except that the prompt emission is known to be beamed (Sari, Piran & Halpern 1999; Frail et al. 2001) so the upper limit is lower, and because the injected energy at early times is negligible, \( E_0 \) dominates the early light curve so the lower limit can be much less energetic before its influence vanishes. These limits are summarized in Table 3.

Each combination of parameters creates a synthetic light curve, and the match to the data is assessed by calculating the \( \chi^2 \) value for the X-ray observations, as well as observations in the R band since this is always the best sampled optical light curve. The \( \chi^2 \) values for the two light curves are assessed separately to avoid a situation where an excellent fit to the X-rays but a poor fit to the optical is indistinguishable from a good fit to both, since the statistics will be dominated by the much better sampled X-ray light curve. Upper limits are not included in the \( \chi^2 \) calculations, but were subsequently inspected for violations (see Section 4). Since there are often fewer R-band data points than free parameters, we are not able to calculate the reduced \( \chi^2 \) for the individual bands, but do calculate the overall reduced \( \chi^2 \) by summing the \( \chi^2 \) contribution and dividing by the combined degrees of freedom. The X-ray band \( \chi^2 \) is obtained for data points in the X-ray plateau and later, excluding the preceding

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Minimum</th>
<th>Maximum</th>
</tr>
</thead>
<tbody>
<tr>
<td>( \epsilon_e )</td>
<td>( 10^{-3} )</td>
<td>1</td>
</tr>
<tr>
<td>( \epsilon_B )</td>
<td>( 10^{-8} )</td>
<td>1</td>
</tr>
<tr>
<td>( n_0 ) (cm(^{-3}))</td>
<td>( 10^{-5} )</td>
<td>100</td>
</tr>
<tr>
<td>( L ) (erg s(^{-1}))</td>
<td>( 10^{47} )</td>
<td>( 10^{49} )</td>
</tr>
<tr>
<td>( \kappa )</td>
<td>( 10^{-2} )</td>
<td>10</td>
</tr>
<tr>
<td>( E_0 ) (erg)</td>
<td>( 10^{48} )</td>
<td>( 10^{52} )</td>
</tr>
</tbody>
</table>

Notes. \( *10^{49} \text{ erg s}^{-1} \) is typical, but the real value depends on equation (5).

Table 3. Limits on parameters used in the order of magnitude parameter space search.

<table>
<thead>
<tr>
<th>GRB</th>
<th>( p )</th>
<th>( T_{\text{em}} ) (s)</th>
<th>( t_b ) (d)</th>
</tr>
</thead>
<tbody>
<tr>
<td>051221A</td>
<td>2.4</td>
<td>( 8.0 \times 10^3 )</td>
<td>( \geq 4.0 )</td>
</tr>
<tr>
<td>060614</td>
<td>2.6</td>
<td>( 2.5 \times 10^4 )</td>
<td>1.10</td>
</tr>
<tr>
<td>070714B</td>
<td>2.9</td>
<td>( 2.0 \times 10^3 )</td>
<td>( \geq 0.7 )</td>
</tr>
<tr>
<td>130603B</td>
<td>2.5</td>
<td>( 8.0 \times 10^2 )</td>
<td>0.35</td>
</tr>
</tbody>
</table>
steep decay. This region is believed to be due the curvature effect (Kumar & Panaitescu 2000).

4 MODELLING RESULTS

The order of magnitude parameter search returned a variety of viable combinations across the four GRBs. Each was inspected by eye to ensure that no upper limits were violated and that the model was consistent with (i.e. fainter than) the early X-ray emission, since neither of these things were factored into the \( \chi^2 \) value. The fit each model gave to other optical and UV observations was also inspected for consistency, and those that violated upper limits or provided a poor match to the data were rejected. 16 models were found for GRB 051221A, 6 models were found for GRB 060614, 21 models were found for GRB 070714B, and 17 models were found for GRB 130603B. Example fits for each GRB are shown in Fig. 1. The X-ray rebrightening at around \( 10^{-2} \) d in GRB 070714B is interpreted as an X-ray flare (e.g. Gompertz et al. 2013) due to its short time-scale and apparent discrepancy with the R-band light curve. This is supported by a spectral hardening shown by the photon index fit on the UKSSDC burst analyser\(^1\) (Evans et al. 2007, 2009).

For each parameter combination, the \( \chi^2 \) values are calculated separately for the X-ray and R-band light curves. These are plotted against each other, and we make \( \chi^2 \) cuts at both frequencies that return a sample of the best fits for each GRB. This method prevents the much better sampled X-ray light curve from dominating the selection threshold, as would be the case for a combined reduced \( \chi^2 \) cutoff. The reduced \( \chi^2 \) limits that result from the combination of \( \chi^2 \) cutoffs for each burst are shown in Table 4. The large variations in these limits are a reflection on how constraining the available X-ray and R-band data are to the models; since the fitting procedure is a simple order of magnitude search rather than a least-squares fit, light curves with larger numbers of data points will be much less forgiving on the models applied. A finer parameter search would reduce \( \chi^2 \). The reduced \( \chi^2 \) limits for all four bursts could also be made more uniform with least-squares fitting; however, this approach leads to very large parameter uncertainties, as previously discussed.

Three of the four GRBs (051221A, 070714B, and 130603B) also feature radio detections and upper limits. The presence of radio observations helps to narrow the parameter space, with varying degrees of severity depending on how constraining the observation is. They are shown in Fig. 2, plotted alongside the region described by the various light curves from the surviving models, shown in grey. GRB 051221A is the most constrained by radio observations;

\(^{1}\) www.swift.ac.uk/burst_analyser
Table 4. The range of physical parameters and energy factors found in the models that successfully matched the data (including radio observations). No value for $\kappa$ is shown for GRB 051221A and GRB 130603B because these bursts do not contain EE. The reduced $\chi^2$ thresholds are also shown.

<table>
<thead>
<tr>
<th>GRB</th>
<th>Reduced $\chi^2$ limit</th>
<th>$\epsilon_e$</th>
<th>$\epsilon_B$</th>
<th>$n_0$ (cm$^{-3}$)</th>
<th>$L$ (erg s$^{-1}$)</th>
<th>$E_{\delta}$ (erg)</th>
<th>$\kappa$</th>
</tr>
</thead>
<tbody>
<tr>
<td>051221A</td>
<td>2.8</td>
<td>0.1–1</td>
<td>$10^{-5}$–$10^{-3}$</td>
<td>$10^{-4}$–$10^1$</td>
<td>$10^{47}$–$10^{48}$</td>
<td>$10^{48}$–$10^{51}$</td>
<td>–</td>
</tr>
<tr>
<td>060614</td>
<td>15</td>
<td>0.1–1</td>
<td>$10^{-5}$–$10^{-3}$</td>
<td>$10^{-4}$–$10^2$</td>
<td>$10^{47}$–$10^{48}$</td>
<td>$10^{48}$–$10^{50}$</td>
<td>$10^{-2}$–$10^{-1}$</td>
</tr>
<tr>
<td>070714B</td>
<td>10</td>
<td>0.1–1</td>
<td>$10^{-6}$–$10^{-2}$</td>
<td>$10^{-4}$–$10^2$</td>
<td>$10^{47}$–$10^{49}$</td>
<td>$10^{48}$–$10^{52}$</td>
<td>$10^{-2}$–$10^0$</td>
</tr>
<tr>
<td>130603B</td>
<td>8</td>
<td>0.1–1</td>
<td>$10^{-5}$–$10^0$</td>
<td>$10^{-4}$–$10^1$</td>
<td>$10^{47}$–$10^{49}$</td>
<td>$10^{48}$–$10^{51}$</td>
<td>–</td>
</tr>
</tbody>
</table>

Figure 2. The radio detections and upper limits available for our sample of four GRBs. The light grey region shows the range of fluxes described by the light curves of the model fits that are consistent with the broad-band data. See Section 5 for a discussion in the apparent upper limit violations in GRB 130603B. The dark grey line in GRB 051221A shows the model that comes closest to matching the observations at 8.46 GHz, and its inconsistency with the upper limits illustrates the probable need for reverse shock emission at early times. The dark grey region in GRB 070714B shows the additional range of predictions resulting from the models with no jet break that are consistent with observations. The vertical black dotted line shows the position of the required jet break for GRB 051221A and GRB 130603B, and the position of the earliest jet break allowed by the data for GRB 070714B. The vertical black dashed line marks 2015 January 1 for reference.
eight more models were ruled out due to the consecutive upper limits at 8.46 GHz, including all models not featuring a jet break. The dark grey line shows the model that best matches the radio detection for this burst, but it is at odds with the upper limits. In GRB 130603B, the most luminous models appear to violate some of the upper limits; however, these can be retained due to the possible influence of radio scintillation (Frail et al. 1997; Goodman 1997) which can explain discrepancies in isolated cases. Forward shock emission appears to have some difficulty in matching the radio detections in this burst and GRB 051221A, and possible reasons for this are discussed in Section 5. In GRB 070714B, the upper limit is not at all constraining to the physical parameter space.

For each GRB, a fairly wide range of parameters was found. It is immediately obvious from Table 4 that a high value of $\epsilon_e$ is required in all cases, otherwise the model emission is too faint to match what is observed in both X-rays and optical bands. It should be noted that while we give broad ranges for the physical parameter values, these values only work to reproduce the data in specific combinations. Two of the four GRBs (060614 and 070714B) have values for $p$ that are consistent within the 90 per cent confidence interval with the late-time photon index $\Gamma$ from the UKSSDC spectrum repository (Table 1). The other two lie between the values gained with the late-time photon index $\Gamma$ from the UKSSDC spectrum repository (Table 1). The other two lie between the values gained with the late-time photon index $\Gamma$ from the UKSSDC spectrum repository (Table 1). The other two lie between the values gained with the late-time photon index $\Gamma$ from the UKSSDC spectrum repository (Table 1). The other two lie between the values gained with the late-time photon index $\Gamma$ from the UKSSDC spectrum repository (Table 1).

### 5 MODELLING DISCUSSION

In some cases, most notably GRB 130603B, we have retained models that appear to violate radio upper limits. This is based on radio scintillation arguments (Frail et al. 1997; Goodman 1997), where isolated detections and upper limits could be located at scintillation maxima/minima, and therefore have larger errors than quoted. For this reason, the handful of models that do not obey the upper limit in GRB 130603B have been left in the predictions. In cases like the 8.46-GHz light curve for GRB 051221A, however, consecutive limits are unlikely to have all occurred at scintillation minima, and so the models that passed above more than one of them (including all the no jet break models) were rejected.

This causes a problem when trying to explain the earlier radio detection with forward shock emission alone; the only model that comes close in GRB 051221A, shown by the dark grey line in Fig. 2, is inconsistent with three of the four radio upper limits. While it could be argued that the single detection in GRB 051221A is itself due to scintillation, the situation is even worse in GRB 130603B, where consecutive detections at 6.7 GHz and a further observation at 4.9 GHz cannot be matched by models without rising above multiple upper limits. The natural explanation for this is the presence of a reverse shock propagating backwards through the ejected material. A reverse shock could produce a radio flare, providing a match to the data while still being masked beneath the forward shock emission at higher frequencies (Kulkarni et al. 1999; Sari & Piran 1999; Nakar & Piran 2005; van Eerten 2014). Reverse shocks are believed to have been observed in both LGRBs (e.g. Akerlof et al. 1999; Chandra et al. 2010; Anderson et al. 2014) and SGRBs (e.g. Soderberg et al. 2006 for GRB 051221A).

Our model includes only the most basic features of the magnetar central engine; we have not incorporated reverse shock emission or other sources of radiation (e.g. kilonova emission; Piran, Nakar & Rosswog 2013; Tanvir et al. 2013). Our aim was to show that a physically motivated, self-consistent central engine, in which a newly formed magnetar injects energy into an expanding forward shock as it loses angular momentum, can be reconciled with the longer wavelength (optical, IR, radio) observations of SGRBs, as well as just the X-ray light curves as is usually done. For this reason, and given the roughness of our fitting routine, the fact that the light curves at all frequencies are well recreated by this bare-bones model is encouraging.

Soderberg et al. (2006) modelled the afterglow of GRB 051221A. Our results are in agreement with theirs, except that we find a much wider range in $F_{\nu}$ ($10^{-4}$–$10^{-1}$ in this work, compared to 0.12–1/3 in Soderberg et al. 2006) and $n_0$ ($10^{-4}$–$10^{-3}$ cm$^{-3}$ in this work, compared to (0.5–2.4) × $10^{-3}$ cm$^{-3}$ in Soderberg et al. 2006). This narrow range is likely due to the inclusion of a reverse shock in their modelling, and indeed their forward shock only parameter ranges are much broader, although still narrower than what we find. GRB 051221A was also modelled by Burrows et al. (2006), who obtain a low- and high-density fit, giving a range of $10^{-4}$–$10^{-3}$ cm$^{-3}$ which is in agreement with our findings, and similar to the forward shock only results of Soderberg et al. (2006). Both studies find narrow jets, consistent with our range (Section 5.1), and jet break times of 4–5 d. Fan & Xu (2006) also fitted the magnetar model to the broad-band observations of GRB 051221A, finding a family of physical parameters within our range.

No broad-band modelling has been done on GRB 070714B, but Xu et al. (2009) fitted a model featuring power-law energy injection to GRB 060614, and found a fit with $\epsilon_e \sim 0.12$, $\epsilon_B \sim 2 \times 10^{-4}$, and $n_0 = 0.04$ cm$^{-3}$, in agreement with our range of parameters. By fitting power-law models to the R-band light curves (Della Valle et al. 2006) and a combination of X-ray and optical bands (Mangano et al. 2007), two previous studies have found a jet break at $\sim 1.3$ d in GRB 060614, consistent with what we find at 1.1 d.

The broad-band afterglow of GRB 130603B was modelled by Fong et al. (2014). As in 051221A, our derived range of density values is wider, extending two orders of magnitude) because $\epsilon_B$ range also extends down an order of magnitude further. These ranges highlight the large degeneracies in the parameters; $\epsilon_e$ is confined to a relatively small range ($\sim$one order of magnitude) because $v_m$ and $F_{\nu, max}$ are well constrained by the data, whereas $\nu_t$ and $\nu_s$ are often unconstrained, leading to a variety of acceptable parameter combinations. Fong et al. (2014) find a jet break at $\approx 0.47$ d, and a jet opening angle in the range $4^\circ$–$14^\circ$, both of which are consistent with our own findings. Finally, the magnetar spin period and dipole field values calculated by Fong et al. (2014) for the dipole spin-down injection case intersect with the line for GRB 130603B shown in Fig. 3. Fan et al. (2013) also showed that the magnetar model was capable of reproducing the broad-band emission observed in GRB 130603B for one combination of physical parameters that lies within our range.
5.1 Energetics

The radiative efficiency of a GRB is defined as (cf. Zhang 2007)

$$\zeta = \frac{\epsilon_{\gamma,\text{iso}}}{\epsilon_{\gamma,\text{iso}} + \epsilon_{k,\text{iso}}}$$

(7)

and gives a direct measure of how efficiently the total energy is converted into EM radiation. In our work, we calculate $E_k$ for each burst, which is the energy delivered to the afterglow emission site by the prompt impulse, and makes no assumption on geometry. The lower limit of $E_k$ is not at all constraining; the fit to the plateau emission depends much more on the luminosity of the dipole spin-down injection, $L$. Values for $E_k$ of $10^{58}$ erg and below are indistinguishable from one another, and for a given $E_{\gamma,\text{iso}}$, will just represent an asymptotic approach to a radiative efficiency of 1, which is unphysical. The upper limits of $E_k$ are far more important, since they are constrained by observation in that too much energy contribution will drive the model fluxes up above what is observed, and will mask the plateau feature in cases where $E_{\gamma}$ is negligible in comparison to $E_k$. The approximate maximum value of $E_{\gamma,\text{iso}}$ is given by assuming that the upper limit of $E_k$ came from a strongly beamed geometry with a beaming factor of $\sim 10^{10}$, i.e. the upper limit of $E_{\gamma,\text{iso}}$ is as much as a thousand times greater than the upper limit of $E_k$. The radiative efficiency can then be used to calculate the implied opening angle (cf. Racusin et al. 2009):

$$\theta_j = 0.057 \left[ \frac{3.5}{1 + z} \right]^{3/8} \left[ \frac{\zeta}{0.1} \right]^{1/8} \left[ \frac{n_0}{E_{\gamma,\text{iso},53}} \right]^{1/8}$$

(8)

The range of calculated efficiencies and opening angles are shown in Table 5. The derived efficiencies are consistent with Zhang (2007), who found typical values of $<10$ percent in their sample. Note that for the two EE GRBs, these calculations may be affected by the energy contribution of EE. We find that GRB 060614 tends to demand higher values of $\epsilon_{\gamma}$ and $\zeta$ than the other bursts, which is symptomatic of its more luminous and longer lasting afterglow plateau putting extra demands on the available energy. The derived opening angles are consistent with the results of Ryan et al. (2015). Their results (in degrees) are 26.0$^{+1.40}_{-2.20}$ for GRB 051221A, 17.0$^{+7.08}_{-4.93}$ for GRB 060614, and 19.1$^{+6.38}_{-6.38}$ for GRB 070714B, where the errors are 1σ. These values were obtained by fitting to the X-ray light curves only.

5.2 Magnetar properties

The well constrained value of $T_{\text{em}}$ and the results for $L$ mean that the magnetar properties can be approximated. These approximations assume 100 per cent efficiency and isotropic emission, which is unlikely to be the case; however, in energetic terms a reduction in one compensates for a reduction in the other, and the large starting uncertainty associated with a simple order of magnitude search for $L$ means that this assumption is sufficiently accurate in the context of other sources of error. The range of magnetar spin periods and field strengths is illustrated in Fig. 3. These properties are not well constrained in three out of four bursts due to the degeneracy created by the wide range of physical parameters, as well as uncertainties in measuring the dipole plateau due to contamination from the prompt and EE components. The normalization of the $B$–$P$ relation is set by the value of $T_{\text{em}}$ for each burst, and the suitable combinations run from the minimum spin break-up period up to the point at which the plateau becomes too faint for a good fit, at around $L = 10^{57}$ erg s$^{-1}$.

The implications for EE in GRB 060614 and GRB 070714B are not well defined. While a range of energies that work in the context of the light curves can be found, the physical interpretation is not constrained in terms of beaming or efficiency, save that the results lie in the region found here. One central engine capable of providing such a result is a magnetic propeller (Gompertz et al. 2014). The EE profile used here borrowed the luminosity curve for a 40 per cent efficient isotropic propeller (without the prefactor $k$), so the EE contribution is energetically consistent with the propeller model; however, in this context it was used as a simple indicator of luminosity, and the restrictions it imposes on $P$ and $B$ of the underlying magnetar were not applied.

6 IMPLICATIONS FOR RADIO EMISSION

The models that successfully match the available broadband observations in Section 4 are used to create synthetic light curves in a variety of radio frequencies: 60 and 150 MHz, then 1.4, 15 and 100 GHz. The light curves combine to give a region of predicted flux densities, showing the bounds of what the radio afterglow should have looked like for each GRB at each frequency, given the imposed restrictions of our specific physical model. This is plotted in Fig. 4.

---

**Table 5.** Calculated minimum radiative efficiencies and ranges of opening angles.

<table>
<thead>
<tr>
<th>GRB</th>
<th>$E_{\gamma,\text{iso}}$ (erg)</th>
<th>$\zeta$</th>
<th>$\theta_j$ (deg)</th>
</tr>
</thead>
<tbody>
<tr>
<td>051221A</td>
<td>$1.5 \times 10^{51,4}$</td>
<td>$\geq 1.5 \times 10^{-3}$</td>
<td>2.37–22.7$^*$</td>
</tr>
<tr>
<td>060614</td>
<td>$2.5 \times 10^{51,9}$</td>
<td>$\geq 2.0 \times 10^{-2}$</td>
<td>1.62–19.7</td>
</tr>
<tr>
<td>070714B</td>
<td>$1.6 \times 10^{51,5}$</td>
<td>$\geq 1.6 \times 10^{-4}$</td>
<td>0.87–14.4$^*$</td>
</tr>
<tr>
<td>130603B</td>
<td>$1.0 \times 10^{51,4}$</td>
<td>$\geq 1.0 \times 10^{-3}$</td>
<td>1.01–10.0</td>
</tr>
</tbody>
</table>

Notes: $^a$Soderberg et al. (2006); $^b$Mangano et al. (2007); $^c$Graham et al. (2009); $^d$Fong et al. (2014). $^e$Calculated from a minimum $f_\text{gb}$; value rises with increasing $f_\text{gb}$.
Figure 4. (a) Predicted flux density light curves at 60 and 150 MHz for the four GRBs in our sample. The dark grey line in GRB 051221A shows the model that comes closest to matching the observations at 8.46 GHz in Fig. 2. The dark grey region in GRB 070714B shows the additional range of predictions resulting from the models with no jet break that are consistent with observations. The vertical black dotted line shows the position of the required jet break for GRB 051221A, GRB 060614, and GRB 130603B, and the position of the earliest jet break allowed by the data for GRB 070714B. The vertical black dashed line marks 2015 January 1 for reference. Selected limits from Table 6 are overplotted. (b) Predicted flux density light curves at 1.4 and 15 GHz. (c) Predicted flux density light curves at 100 GHz.
Table 6 shows the sensitivity thresholds for modern-day and future radio telescopes that observe at the frequencies plotted, and a selection of these are superimposed on the light curves. We assess the detectability of each GRB radio afterglow. The flux densities are in general modest, typically peaking in the µJy range; however, the results for the anomalously bright GRB 060614 do extend up to mJy. The signal from each GRB is suppressed by the jet break, which curtails the initial brightening of the emission early on in the
Table 6. Detection sensitivities for different instruments at the frequencies for which we calculate flux density prediction light curves. Limits are 5σ and assume a 12-h integration time. [A] – almascience.eso.org/proposing/sensitivity-calculator.

<table>
<thead>
<tr>
<th>Telescope</th>
<th>Sensitivity (µJy)</th>
<th>Reference</th>
</tr>
</thead>
<tbody>
<tr>
<td><strong>60 MHz:</strong></td>
<td></td>
<td></td>
</tr>
<tr>
<td>LWA1</td>
<td>38000</td>
<td>Ellingson et al. (2013)</td>
</tr>
<tr>
<td>LOFAR</td>
<td>5000</td>
<td>van Haarlem et al. (2013)</td>
</tr>
<tr>
<td><strong>150 MHz:</strong></td>
<td></td>
<td></td>
</tr>
<tr>
<td>MWA</td>
<td>1200</td>
<td>Tingay et al. (2013)</td>
</tr>
<tr>
<td>LOFAR</td>
<td>300</td>
<td>van Haarlem et al. (2013)</td>
</tr>
<tr>
<td><strong>1.4 GHz:</strong></td>
<td></td>
<td></td>
</tr>
<tr>
<td>GMRT</td>
<td>150</td>
<td>Ghirlanda et al. (2014)</td>
</tr>
<tr>
<td>WSRT/Aperif</td>
<td>50</td>
<td>Ghirlanda et al. (2014)</td>
</tr>
<tr>
<td>ASKAP</td>
<td>50</td>
<td>Ghirlanda et al. (2014)</td>
</tr>
<tr>
<td>MeerKAT phase 1</td>
<td>9</td>
<td>Ghirlanda et al. (2014)</td>
</tr>
<tr>
<td>MeerKAT phase 2</td>
<td>6</td>
<td>Ghirlanda et al. (2014)</td>
</tr>
<tr>
<td>SKA phase 1</td>
<td>1</td>
<td>Ghirlanda et al. (2014)</td>
</tr>
<tr>
<td>SKA phase 2</td>
<td>0.15</td>
<td>Ghirlanda et al. (2014)</td>
</tr>
<tr>
<td><strong>15 GHz:</strong></td>
<td></td>
<td></td>
</tr>
<tr>
<td>AMI</td>
<td>70</td>
<td>Zwart et al. (2008)</td>
</tr>
<tr>
<td>VLA</td>
<td>5</td>
<td>Ghirlanda et al. (2014)</td>
</tr>
<tr>
<td><strong>100 GHz:</strong></td>
<td></td>
<td></td>
</tr>
<tr>
<td>CARMA</td>
<td>900</td>
<td>Bock (2006)</td>
</tr>
<tr>
<td>ALMA</td>
<td>6</td>
<td>[A]</td>
</tr>
</tbody>
</table>

light curve in most cases. The region either side of this break usually represents the best opportunity to observe the radio afterglow.

At the lower frequencies (60 and 150 MHz), only the Low Frequency Array (LOFAR) at 150 MHz gets close to being within an order of magnitude of our predictions. The picture is slightly better moving to higher frequencies; in the near future at 1.4 GHz, the Westerbork Synthesis Radio Telescope (WSRT)/Aperif and the Australian Square Kilometre Array Pathfinder (ASKAP) will be sensitive enough to be capable of observing the brighter models in GRB 060614, and graze the upper limits of the GRB 070714B predictions. MeerKAT would have been capable of detecting at least the upper portion of all four bursts, and could have resolved the entire predicted region of GRB 060614 if observations had been made around the time of the jet break.

At 15 GHz, the Arcminute Microkelvin Imager (AMI) is capable of observing the upper reaches of the predictions for all but the highest z burst (GRB 070714B) for around a week, possibly even a month for the brighter portion of GRB 060614. The Very Large Array (VLA), in its expanded capacity (Perley et al. 2011), would have been able to go deeper than our lower limits in each burst except GRB 130603B in the first week, and provide meaningful limits on the evolution of the radio afterglow for up to a year after trigger. Finally, at 100 GHz the Combined Array for Research in Millimetre-wave Astronomy (CARMA) may have been able to detect the brightest models in GRB 060614 and GRB 130603B, and the Atacama Large Millimetre Array (ALMA) would have been able to provide limits similar to those mentioned for the VLA, with a window of weeks in GRB 051221A and GRB 070714B, and months in GRB 060614, where the entire predicted region lay above its sensitivity threshold.
Our model fluxes show that previous radio observations, while able to limit some of the physical parameter space, were not deep enough to place serious constraints on the magnetar model. However, the recently upgraded VLA (Perley et al. 2011) and ALMA are now at $\mu$Jy sensitivity, deep enough to probe even the faintest predicted models. Either telescope can now provide meaningful and highly constraining restrictions on a central engine invoking dipole spin-down injection into a forward shock by making observations within the first week or two after trigger, assuming the four GRBs discussed here are representative of the sample as a whole. Since our sample contains the highest recorded spectroscopic SGRB redshift ($z = 0.9224$; GRB 070714B) and the results in Table 4 show ISM densities at or near the observed lower limit $n_0 \sim 10^{-5}$ cm$^{-3}$, we suggest that our sample does represent SGRB and EE GRB radio fluxes as a whole, rather than the most luminous cases.

The Square Kilometre Array (SKA) paints a rather brighter picture for the future; our results suggest that even at phase 1, we should expect to see magnetar-injection driven 1.4 GHz afterglows for months after trigger if the model is to be believed. All four GRBs shown here would be observable for months, in some cases up to a year after trigger, with only the very faintest models in GRB 070714B and GRB 130603B lying below the sensitivity threshold. By phase 2, all four of the radio afterglows in our sample would have been visible for a year or more, and the entire predicted flux density region could be explored for each with the correct observing strategy. Our findings are in agreement with Feng, Vaulin & Hewitt (2014), who simulated radio afterglow light curves for compact object mergers at the advanced Laser Interferometer Gravitational-Wave Observatory (aLIGO) horizon. We consider here the simplest case of merger followed by injection; however, the radio signal from these mergers may be further enhanced by other processes such as macronovae (Piran et al. 2013).

7 CONCLUSIONS

We have performed order of magnitude fitting to the broad-band afterglows of a sample of four GRBs. We use a physically motivated central engine, invoking energy injection into a forward shock from a magnetar as it rapidly loses angular momentum along open magnetic field lines. By imposing the limitations of a self-consistent central engine for the energy profile of each GRB, we are able to narrow the available parameter space for the physics underlying the evolution of the blast wave as it expands into the ambient medium. Combinations of these parameters are tested against the data, resulting in a family of models that accurately recreate observations. These models are then used to predict the radio signature from the central engine, and are assessed for detectability.

Our results show that current broad-band observations are consistent with the magnetar injection model, as we find physical parameters that lie within the allowed ranges for all bursts. Some discrepancies exist at radio frequencies, suggesting that previous early detections captured emission from a reverse shock propagating backwards through the ejecta, rather than a forward shock moving outwards into the ISM. We find that while recent observational detection thresholds are not constraining to the magnetar model, state-of-the-art facilities such as the upgraded VLA and ALMA are now capable of observing to depths greater than our predicted flux density range if observations are made in the first few weeks, and to maximum sensitivity. We also show that SKA will be capable of observing to depths in excess of our model predictions, and hence is expected to observe these signatures, or impose strict limits on the physical parameters.

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