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Published in: Monthly Notices of the Royal Astronomical Society

DOI: 10.1093/mnras/stw917

Citation for published version (APA):
A non-parametric method for measuring the local dark matter density

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Accepted 2016 April 18. Received 2016 March 28; in original form 2015 August 7

ABSTRACT

We present a new method for determining the local dark matter density using kinematic data for a population of tracer stars. The Jeans equation in the z-direction is integrated to yield an equation that gives the velocity dispersion as a function of the total mass density, tracer density, and the ‘tilt’ term that describes the coupling of vertical and radial motions. We then fit a dark matter mass profile to tracer density and velocity dispersion data to derive credible regions on the vertical dark matter density profile. Our method avoids numerical differentiation, leading to lower numerical noise, and is able to deal with the tilt term while remaining one dimensional. In this study we present the method and perform initial tests on idealized mock data. We also demonstrate the importance of dealing with the tilt term for tracers that sample above the disc plane. If ignored, this results in a systematic underestimation of the dark matter density.

Key words: Galaxy: disc – Galaxy: kinematics and dynamics – dark matter.

1 INTRODUCTION

Dark matter (DM) is an elusive component of the cosmos. Its existence has long been inferred from its gravitational interactions with ordinary matter on scales ranging from dwarf galaxies and galactic rotation curves to lensing galaxies, galaxy clusters, and the Universe as a whole (for reviews see e.g. Rubin 1993; Junghans, Kamionkowski & Friest 1996; Bergström 2000; Bertone, Hooper & Silk 2005; Bertone 2010; Masseny, Kitching & Richard 2010). However its exact nature remains unknown.

Understanding the distribution of DM in cosmological structures, and in particular in the Milky Way, is of great importance to test the standard ΛCDM cosmological model, to make predictions – and allow the interpretation – of experiments that seek to detect DM, and to probe more exotic models for DM (e.g. Read 2014). Here, we will focus our attention on the local DM density – i.e. the average density of DM over a small volume (typically of size ~100–1000 pc) centred around the Sun. The local DM density is a vital input for so-called direct DM experiments, based on the detection of the recoil energy of nuclei struck by DM particles streaming through underground detectors, as the expected event rate is proportional to the local DM density and the DM–nucleon cross-section. In addition, the local DM density is important for indirect searches that look for neutrinos produced by the annihilation of DM particles trapped at the centre of Sun (see e.g. Klasen, Pohl & Sigl 2015, and references therein for a recent update on direct and indirect searches). The results of these searches are also used in explorations of theoretical parameter space, such as those of Supersymmetry (e.g. Allmanach & Lester 2006; Buchmueller et al. 2014; Strege et al. 2014; Bertone et al. 2016; Roszkowski, Sessolo & Williams 2015).

Methods of determining the local DM density can be divided into two categories: those utilizing measurements of stars in a small volume around the Sun (e.g. Kappern 1922; Jeans 1922; Oort 1932; Bahcall 1984; Kuijken & Gilmore 1989b, 1991; Creze et al. 1998; Garbari et al. 2012; Bovy & Tremaine 2012; Smith, Whiteoak & Evans 2012; Zhang et al. 2013; Bienaymé et al. 2014) and those utilizing a host of dynamical tracers – particularly the rotation curve – to constrain more global mass models of the Milky Way (Dehnen & Binney 1998; Catena & Ullio 2010; Salucci et al. 2010; Weber & de Boer 2010; McMillan 2011; Nes et al. 2013; Piffl et al. 2014; Iocco, Pato & Bertone 2015; Pato & Iocco 2015; Pato, Iocco & Bertone 2015). Following the notation of Read (2014) we denote the local DM density derived from local measurements by \( \rho_{\text{DM}} \), and those extrapolated from rotation curves assuming a spherical DM halo as \( \rho_{\text{DM, ext}} \). With the advent of large survey data, local and global methods are beginning to converge (Bovy & Rix 2013; Piffl et al. 2014).

The comparison of \( \rho_{\text{DM}} \) to \( \rho_{\text{DM, ext}} \) can provide insight into the shape of the Milky Way’s DM halo, and thus into the formation of the galaxy (Read 2014). If \( \rho_{\text{DM}} \) from local measurements is smaller than that extrapolated from global measurements, \( \rho_{\text{DM, ext}} \), then this implies a prolate halo, while the opposite would imply a

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squashed, oblate halo. While the former is produced by DM only
N-body simulations, the introduction of baryons into the simulations
produces the latter (Katz & Gunn 1991; Dubinski 1994; Debattista et al. 2008; Read et al. 2009).

If \( \rho_{\text{DM}} \) is derived from tracers distributed vertically in the plane
it is possible to probe not only the local DM density but also its
vertical profile, \( \rho_{\text{DM}}(z) \). This could potentially reveal deviations
from a spherical halo profile, which can form in a number of ways.
First, the DM halo should respond to the formation of the bary-
onic disc by flattening into an oblate halo (Katz & Gunn 1991;
Dubinski 1994; Debattista et al. 2008; Read et al. 2009). Secondly,
the accretion of subhaloes after the formation of the baryonic disc
can also lead to the formation of a ‘dark disc’ (DD) that coro-
tates with the baryonic disc (Lake 1989; Read et al. 2008, 2009;
Purcell, Bullock & Kaplinghat 2009; Ling et al. 2010; Pillepich et al. 2014). Gravitationally these two features, an oblate halo and a
DD, are indistinguishable and degenerate. One method to dis-
tinguish the two is to hunt for the chemically and kinematically
distinct sub-halo stars that would accompany the accreted DD, but
not the contracted DM halo. Recent Gaia-ESO observational data
find no evidence for such stars (Ruchti et al. 2014, 2015), suggest-
ing that the Milky Way has had a rather quiet past since its disc
formed, and therefore does not host a significant accreted DD. If
correct, then any gravitationally detected non-sphericity must imply
a non-rotating locally oblate halo. An enhancement in \( \rho_{\text{DM}} \)
over \( \rho_{\text{DM,loc}} \) is also a feature in some more exotic models, such as
Partially Interacting Dark Matter (PIDM) (Fan et al. 2013),
and some modified theories of gravity (e.g. Read & Moore 2005;
Nipoti et al. 2007).

Measurement of the local DM density dates back almost 100 years
(Kapteyn 1922). The history of this measurement is one of both in-
creasingly precise data and decreasingly strong assumptions. As
such, the error bars on \( \rho_{\text{DM}} \) have not always shrunk with time,
as better data often allows one to discard previous assumptions.

With the advent of Gaia data from 2016 onwards we will have access
to high-precision data on individual stars, and the primary
uncertainty in the determination of \( \rho_{\text{DM}} \) will be systematic model uncertainties. Recent results for the local DM density include
Bienaymé et al. (2014), who used RAVE data to derive a value of
\( \rho_{\text{DM}} = 0.0143 \pm 0.0011 \, \text{M}_\odot \, \text{pc}^{-3} = 0.542 \pm 0.042 \, \text{GeVcm}^{-3} \),
and Zhang et al. (2013) who used Sloan Digital Sky Survey
(SDSS)/SEGUE data to find \( \rho_{\text{DM}} = 0.0065 \pm 0.0023 \, \text{M}_\odot \, \text{pc}^{-3} =
0.25 \pm 0.09 \, \text{GeVcm}^{-3} \). Note that these two results do not overlap
within their stated uncertainties. The significance of this discrep-
ancy is hard to interpret though, as they each use different data
sets and analysis techniques, and both methodologies rely on rather
different assumptions. To make progress we should endeavour
to minimize the number of assumptions made, and apply the same
analysis techniques to both data sets.

In this paper we make progress towards the first goal of reducing
model assumptions. To achieve this we introduce a one-dimen-
sional Jeans analysis method to probe the local DM density using the vertical
motions of tracer stars. We construct a representation of the
tracer density \( v \), and also allow for a DM density profile that is
more complex than simply constant with height as previous local
studies have assumed. Additionally, we deal with the so-called tilt
term, which links radial and vertical motions of the tracer stars. This
term is crucial to understand stellar motions at high-

The broad picture of our problem is this: we have quantities derived
from the motions of the tracer stars, namely \( v \), the tracer density,
\( \sigma_z \), the vertical velocity dispersion, and \( \sigma_R \), the \((R, z)\) velocity
dispersion. Then we have elements of the mass profile, one of
which is unknown – the DM density \( \rho_{\text{DM}} \), and the other which is
known within a band of uncertainty – the baryon density \( \rho_{\text{Baryon}} \).
In this section we first derive the equations to link these quantities,
and then describe how each is modelled.

2.1 Deriving a general 1D jeans method

A population of ‘tracer’ stars moving vertically near the Sun obeys
the collisionless Boltzmann equation:

\[
\frac{df}{dt} = \nabla f \cdot \nabla \Phi + \nabla f \cdot v \nabla \sigma_z + \Phi = 0
\]

where \( f(x, v) \) is the stellar distribution function; \( x \) and \( v \) are 
the positions and velocities, respectively; and \( \Phi \) is the gravitational
potential.

If the system is in dynamic equilibrium (steady state), then we
may neglect the partial time derivative of \( f \). Thus for equilibrium
tracers, if we know their phase space distribution function \( f \), then
we can directly measure the gravitational force from equation (1).
In practice, however, this is hard because \( f \) is six-dimensional (even
a million stars give only 10 sample points per dimension) and we
need to estimate the partial derivatives of \( f \) to solve equation (1).
For this reason, it is common to integrate equation (1) over velocity
to obtain a set of moment equations: the Jeans equations (Binney
& Tremaine 2008). Adopting cylindrical polar coordinates \((R, \phi, z)\)
and focusing on the \(z\)-Jeans equation, we have
\[
\frac{1}{R^2} \frac{\partial}{\partial R} \left( R \frac{\partial V_{\phi}}{\partial R} \right) + \frac{1}{\rho} \frac{\partial}{\partial \phi} \left( \rho \frac{\partial \sigma_z}{\partial \phi} \right) + \frac{v}{\rho} \frac{\partial \sigma_z}{\partial z} = \frac{d \Phi}{dz}
\]
(2)
where \(v\) and \(\sigma_z^2\) are the density and vertical velocity dispersion
profile of a tracer population as a function of height \(z\) above the
disc plane, \(\sigma_{\phi z}\) is the cross term in the vertical dispersion tensor
that couples radial and vertical motions, and \(\sigma_{\phi \phi}\) is the cross-term
coupling vertical and axial motions.

Integrating both sides of equation (2), we derive our key equation
for this paper:
\[
\sigma_z^2(z) = \frac{1}{v(z)} \int_0^z v(z') \left[ K_z(z') - T(z') - A(z') \right] dz' + C
\]
(3)
where \(C\) is a normalization parameter. For \(z = 0\), we have that
\[
\sigma_z^2(0)v(0) = C,
\]
(4)
and so \(C\) simply sets the vertical velocity dispersion at \(z = 0\). [Note
that a similar but less general equation was derived recently in
Smith et al. (2012), equation (10).] We implement the constant \(C\) as
a parameter in our model that we will ultimately marginalize over.
The alternative would be to calculate \(C\) directly from equation (3),
utilizing the fact that as \(z \to \infty, v(z) \to 0\), which gives us
\[
C = - \int_0^\infty v(z') \left[ K_z(z') - T(z') - A(z') \right] dz'.
\]
(5)
However, this would mean that the derived value for \(\sigma_z(z)\) at some
finite \(z\) would depend the properties of our mass model, tilt and
axial terms not just in the range \([0, z]\) but also for the range \([0, \infty]\).

Equation (3) is numerically appealing to solve since, unlike many
previous methods (e.g. Garbanti, Read & Lake 2011; Garbanti et al.
2012), it does not require any numerical differentiation and is therefore
rather robust to noise in the data. Furthermore, equation (3) is valid
for any gravity theory and can therefore be used as a constraint
on alternative gravity models. In this sense, we follow the early pi-
oneering work of Hill (1960) who attempted to also measure \(K_z\)
directly without reference to the Poisson equation.

To connect the vertical acceleration \(K_z\) to the surface mass density
of the disc, however, we must specify a gravitational model. For
standard Newtonian weak field gravity, this is given by the Poisson
equation:
\[
\nabla^2 \Phi = \frac{\partial^2 \Phi}{\partial z^2} + \frac{1}{R} \frac{\partial V_z^2(R)}{\partial R} = 4\pi G \rho
\]
(6)
\text{rotation curve term: } R.

where \(V_z(R)\) is the circular speed (rotation) curve at radius \(R\), and \(\rho\)
is the total matter density. Notice that for a flat rotation curve,
\(V_z = \text{const.}\), and the ‘rotation curve’ term \(R\) vanishes. If \(R\) does
not vanish, then it appears as a shift (potentially as a function of height
\(z\)) to the recovered \(\rho_{\text{DM}}\) that can be corrected for at the end of the
calculation (Garbanti et al. 2012). Equation (6) can be rewritten as
\[
\frac{\partial^2 \Phi}{\partial z^2} = 4\pi G \rho(z)_{\text{eff}}
\]
(7)
where
\[
\rho(z)_{\text{eff}} = \rho(z) - \frac{1}{4\pi G R} \frac{\partial V_z^2(R)}{\partial R}.
\]
(8)

### Table 1. Values of Oort constants

<table>
<thead>
<tr>
<th>Source</th>
<th>Stellar type</th>
<th>(A)</th>
<th>(B)</th>
</tr>
</thead>
<tbody>
<tr>
<td>Mignard (2000)</td>
<td>K-M giants</td>
<td>14.5 ± 1.0</td>
<td>-11.5 ± 1.0</td>
</tr>
<tr>
<td>Branham (2011)</td>
<td>G giants</td>
<td>14.05 ± 3.28</td>
<td>-9.30 ± 2.87</td>
</tr>
</tbody>
</table>

The contribution of this term to the mass density profile can be
quantified via the Oort constants \(A\) and \(B\) (Binney & Tremaine
2008):
\[
\frac{1}{4\pi G R} \frac{\partial V_z^2(R)}{\partial R} = \frac{B^2 - A^2}{2\pi G}.
\]
(9)
Note that \(|B| < |A|\), meaning that the terms in equation (9) are
negative. Thus the effective density \(\rho(z)_{\text{eff}}\) will be higher with the
inclusion of a non-zero rotation curve term. If the rotation curve term
is erroneously neglected it will be absorbed into the density profile,
yielding an over-estimate of the DM density, assuming the baryonic
contribution is well constrained. Taking the most precise values for
\(A\) and \(B\) from Table 1 (Mignard 2000) yields a value for equation (9)
of \(\sim 0.1\, \text{GeV}/\text{c}^2\), which is roughly a third of the expected local DM
density (e.g. Read 2014). For any accurate measure of the local DM density derived from real data we will need to incorporate this
correction, but we leave this for future work.

Neglecting \(R\) and integrating both sides of equation (6), we derive
the familiar result:
\[
\Sigma_z(z) = \frac{\left| K_z \right|}{2\pi G}
\]
(10)
where \(\Sigma_z(z)\) is the surface mass density of the disc.

The overall flow of our method is now apparent. We model the
tracer density \(v(z)\), the mass density distribution \(\rho(z) = \rho_{\text{DM}}(z) + \rho_{\text{baryon}}(z)\) which gives us the \(K_z\) term, the tilt term \(T(z)\), and the
axial term \(A(z)\). These elements each have a number of parameters,
and so in total each model will have an \(N\)-dimensional parameter
space. Specifying values for each of these parameters will then give
us quantitative values for each of these elements, which can then be
inserted into equation (3) to derive \(\sigma_z^2(z)\) for that parameter space
point. We then compare \(v(z)\) and \(\sigma_z^2(z)\) [and \(\sigma_{\phi z}(z)\) as part of
the tilt term model] to data via a \(\chi^2\) test, and then change the values of
our parameters.

Note that the only assumptions that go into equation (3) are that
the motions of stars obey the collisionless Boltzmann equation,
and that the galaxy is in dynamic equilibrium. The assumption
of dynamic equilibrium can be negated somewhat by the presence of
disequilibrium (‘wobbling’) in the disc caused by e.g. the Sagittarius
merger (Purcell et al. 2011; Gómez et al. 2013) or in part by the
presence of spiral arms (Faure, Siebert & Famaey 2014). However,
the impact of these effects on the measurement of the local DM
density is estimated to be approximately 10 per cent (Widrow et al.
2012; Read 2014), less than the corrections arising from the tilt and
rotation curve terms.

In practice, further assumptions are necessary to model the in-
dividual components on the RHS of equation (3). However, this
method can accommodate almost any model for each of these ele-
ments – the only strict requirement is that each element can be
defined at the \(z\)-values corresponding to the centres of the bins used
to calculate the tracer density and velocity dispersion. Hence we
call this method ‘non-parametric’ – we can in principle have many
more parameters for each model than there are data points. In the following subsections we describe how we model each element for this particular study. In Section 2.6 we discuss model selection using the Bayesian evidence, which could potentially allow us to compare alternate models.

As the galaxy is close to being axisymmetric in both the thin disc (Dehnen 1998; Hogg et al. 2005; Aumer & Binney 2009; Bovy, Hogg & Roweis 2009; Pasetto et al. 2012b; Aghajani & Lindegen 2013) and the thick disc (Pasetto et al. 2012a; Sharma et al. 2014), the axial term $\mathcal{A}(z)$ is expected to be small. For this study we assume complete axisymmetry, and thus we take $\mathcal{A}(z) = 0$. If the data show significant non-axisymmetry in the Milky Way, the axial term could be modelled in a similar way to the tilt term as described below.

Note that if the axial term were non-negligible, in the context of the Milky Way, this would imply the presence of significant spiral arms or a bar. Such features are time dependent, implying that if $\mathcal{A}(z) \neq 0$, we should also worry about the time dependence of the gravitational potential $\Phi(t)$ and, therefore, also the partial time derivative of the distribution function $\partial f / \partial t$ (e.g. Monari, Famaey & Siebert 2016). It is not clear if there can exist a regime in which $\mathcal{A}(z) \neq 0$ and such time dependence can be ignored. Here, we simply note that since empirically $\mathcal{A}(z) \ll T(z)\gamma(z)$ (e.g. Pasetto et al. 2012a, b), to leading order we can also safely neglect the time dependence of the distribution function, as we have already done.

2.2 Tracer density model

To apply the Jeans equations we bin data to obtain the tracer density $n$ and the velocity dispersions $\sigma_z$ and $\sigma_R$, and then at a bare minimum we only have to define $n$ and $\sigma_R$, and derive $\sigma_z$ via equation (3), at the bin centres, i.e. at a discrete set of $z$ values.

For this work we model the tracer density as a sum of $N$ exponential discs:

$$n(z) = \sum_j n(0_j) \exp \left( -\frac{z}{h_j} \right),$$

where for the $i$th disc $n(0)$, is the tracer density at $z = 0$ and $h_i$ the scale height. The number of exponential discs can be increased until the Bayesian evidence shows no improvement (see Section 2.6). This method avoids overfitting as each disc is smooth, while still giving freedom to describe more complex data. From this point onwards we consider only $z \geq 0$.

2.3 Baryon parametrization

For this study we use a simple two-parameter model to describe the baryon distribution:

$$\rho_{\text{baryon}}(z) = \frac{1}{4\pi G} \left| \frac{K_{bn}D_{bn}^2}{(D_{bn}^2 + z^2)^{3/2}} \right|,$$

The $K_{bn}$ parameter sets the mass of the disc, and has dimensions of acceleration, while $D_{bn}$ controls the scale height of the disc, and has dimensions of length. Expressed in terms of the $K_z$ parameter of equation (3), the baryon profile becomes

$$K_{z, \text{baryon}} = \frac{K_{bn}z}{\sqrt{z^2 + D_{bn}^2}}.$$

1 In the context of the values of the prior ranges given in Table 2, the gravitational constant equals $G = 4.229 \times 10^{-6}$ km$^2$ kpc$^{-1}$ M$_O^{-1}$ s$^{-2}$.

While this model is not likely realistic for the Milky Way (Flynn et al. 2006; McKee, Parravano & Hollenbach 2015), it has been applied to observational data by Kuijken & Gilmore (1989a,b), and also more recently by Zhang et al. (2013). When applying our method to real data we will consider more sophisticated baryon models, but for this initial study where we are primarily interested in testing our methodology, equation (13) will suffice.

2.4 Dark matter parametrization

The simplest way to parametrize the DM density profile $\rho_{\text{DM}}$ is to assume it is constant with $z$, as done in previous work (e.g. Bahcall 1984; Kuijken & Gilmore 1989a, b; Creze et al. 1998; Garbari et al. 2011, 2012). This assumption works well at low $z$: for a spherical NFW halo with a scale radius of 20 kpc the mid-plane value is correct within 10 per cent up to a height of $z \sim 3$ kpc. For some analyses we also make this assumption, and set $\rho_{\text{DM}}(z) = \rho_{\text{DM, const}}$.

However, this assumption does not allow for the exploration of several interesting effects such as a flattened, oblate halo or a DD. To accommodate such phenomena we add a DD on top of the constant DM. The DD is described using the same disc model as we use for the baryons, yielding a total DM profile:

$$\rho_{\text{DM}}(z) = \rho_{\text{DM, const}} + \frac{1}{4\pi G} \left| \frac{K_{\text{DD}}D_{\text{DD}}^2}{(D_{\text{DD}}^2 + z^2)^{3/2}} \right|.$$

To ensure this disc does not simply become degenerate with the baryonic disc we give the scale height $D_{\text{DD}}$ a higher prior range than that for $D_{bn}$. In terms of the $K_z$ parameter of equation (3), the DM profile is

$$K_{z, \text{DM}} = \frac{2Fz + 2K_{\text{DD}}z}{\sqrt{z^2 + D_{\text{DD}}^2}},$$

where $F = 2\pi G\rho_{\text{DM, const}}$. This disc parametrization of the non-constant DM profile can describe both the accreted DD and a flattened DM halo, and so henceforth we refer only to a ‘dark disc’. Additional DD terms could in principle be added to $\rho_{\text{DM}}(z)$ with parameters $K_{\text{DD}, n}$ and $D_{\text{DD}, n}$, giving a total density profile of

$$\rho_{\text{DM}}(z) = \rho_{\text{DM, const}} + \frac{1}{4\pi G} \sum_n \left| \frac{K_{\text{DD}, n}D_{\text{DD}, n}^2}{(D_{\text{DD}, n}^2 + z^2)^{3/2}} \right|.$$

The prior ranges on the $n$th DD parameters $K_{\text{DD}, n}$ and $D_{\text{DD}, n}$ can be set in relation to those of the $(n-1)$th DD, e.g. requiring $K_{\text{DD}, n} < K_{\text{DD}, n-1}$, meaning each DD is less massive than the previous one. The number of DD to add could be determined from the data via the Bayesian evidence, with DD terms being added up until this degrades significantly. For this study, however, we limit ourselves to one DD term.

2.5 Tilt term

The tilt term from equation (2) links radial and vertical motion. The importance of this term has been noted previously, e.g. Kuijken & Gilmore (1989a), Smith, Wyn Evans & An (2009), Benderbender et al. (2015). Here, we demonstrate that it is possible to deal with this term while remaining within our vertical, one-dimensional framework. Given the data quality currently available for the Milky Way we are required to make a well-motivated assumption about the radial form of $T$, but in the Gaia-era we will be able to directly measure this from the data.

We first assume that the radial profiles of the tracer density and the $(R, z)$-velocity dispersion are exponentials with scale radii of $R_0$
and $R_1$, respectively:

$$v(R, z) = v(z) \exp(-R/R_0),$$

$$\sigma_{R_z}(R, z) = \sigma_{R_z}(z) \exp(-R/R_1).$$

With this assumption the tilt term becomes

$$T(R, z) = \sigma_{R_z}(R, z) \left[ \frac{1}{R} - \frac{1}{R_0} - \frac{1}{R_1} \right].$$

If the disc were observed not to be an exponential then an alternative model could be easily applied at this stage. Indeed, similar but more complex models have featured previously in the literature (see e.g. Binney et al. 2014).

We then model the $(R, z)$-velocity dispersion as a power law at a given galactocentric radius $R$:

$$\sigma_{R_z}(R, z) = A \left( \frac{z}{\text{kpc}} \right)^n R.$$  

We take $R = R_\odot$, and also $R_0 = R_1$, simplifying the tilt term to a model described by the parameters $A$, $n$, and $R_0$:

$$T(R_\odot, z) = A \left( \frac{z}{\text{kpc}} \right)^n \left[ \frac{1}{R_\odot} - \frac{2}{R_0} \right].$$

Note that we are not affected by the assumption $R_0 = R_1$ since these two parameters are trivially degenerate in equation (21). It remains the case that any observational constraints on $R_0$ and/or $R_1$ can be used as priors on equation (21), where these would constrain the term $2/R_0$. The description of $\sigma_{R_z}(z)$ of equation (20) fulfills $\sigma_{R_z}(z = 0) = 0$ by construction and fits remarkably well to different populations, as shown in Fig. 1. This figure presents $(z, \sigma_{R_z})$ data points for high- and low-metallicity populations (blue and red points, respectively) from the SDSS/SEGUE survey as analysed and presented in B"ubenbender et al. (2015), with a sign correction applied.\footnote{Private communication with authors.} The high metallicity sample has [Fe/H] = −0.07 and $\alpha$ = 0.11, while the low-metallicity sample has [Fe/H] = −0.89 and $\alpha$ = 0.34. Metallicity can be used as a proxy for age, with the high-metallicity sample being younger than the older, low-metallicity population (e.g. Loebman et al. 2011).

Taking these data points we fit power-law models as per equation (20), with parameters $A = 123.99$ km s$^{-2}$, $n = 1.16$ for the high-metallicity population (blue) and $A = 180.08$ km s$^{-2}$, $n = 1.44$ for the low-metallicity population (red). The low metallicity population samples further above the disc plane and populates the canonical thick disc of the Milky Way. Populations such as this are more interesting for local DM searches as they allow us to probe higher $z$ regions where the baryon mass contribution begins to drop away leaving behind the DM halo. However as we go higher in the disc the tilt term becomes increasingly important, as illustrated by Fig. 2, which shows the variable $\xi(z)$:

$$\xi(z) = \frac{K_{\nu,\text{DM}}}{K_{\nu,\text{DM}} - T}$$

where $K_{\nu,\text{DM}} = -2F_{\nu,\text{DM}}$, the constant DM density term, with $F = 267.65$ km$^2$ kpc$^{-2}$ s$^{-2}$ and $R_0 = R_1 = 2.5$ kpc, the same values as are used to later generate our mock data in Section 3. Compared to the thin disc population (blue line), the effects of the tilt term become important at much lower $z$ values. In short, to probe local DM we want to use thick disc stars probing higher-$z$ ranges, but the cost is that we must include the tilt term in our analyses.

Ignoring the tilt term will always cause an underestimation of the local DM density, if all other components of the model such

\begin{figure}
\centering
\subfigure[\textit{Figure 1.} $\sigma_{R_z}$ for two populations of G-type dwarf stars in the solar neighbourhood from the SDSS/SEGUE survey (Büdenbender, van de Ven & Watkins 2015). The blue data points are from a younger, high-metallicity population, with $\text{[Fe/H]} = -0.07$ and $\alpha = 0.11$, while the data points in red are from an older, low-metallicity population with $\text{[Fe/H]} = -0.89$ and $\alpha = 0.34$. The blue and red lines are power laws in the form of equation (20) fitted to the blue and red data points, respectively. The older, lower metallicity stars probe further above the disc plane. Dashed lines indicate extrapolation from data.]{
\includegraphics[width=0.5\textwidth]{figure1a.pdf}}
\subfigure[\textit{Figure 2.} The importance of tilt, as quantified by equation (22). The blue and red lines correspond to tilt terms derived from the high- and low-metallicity population fits of Fig. 1. The deviation arising from the tilt term increases more rapidly with height for the low-metallicity population (red line), which probes the high-$z$ region most useful for determining the DM profile. Dashed lines indicate extrapolation from data. For this case $R_0 = R_1 = 2.5$ kpc, and $F = 267.65$ km$^2$ kpc$^{-2}$ s$^{-2}$.]{
\includegraphics[width=0.5\textwidth]{figure2a.pdf}}
\caption{Measuring the local dark matter density}  
\end{figure}
as baryons are held steady. Looking at equation (21), we note that \( R_0 < R_{\odot} \) and \( A > 0 \), meaning the tilt term \( T(R_{\odot}, z) \) is always negative. Then considering equation (3) we see that to fit to \( \sigma_z^2(z) \) in the absence of the tilt term \( T(z^\prime) \), the \( K_\theta(z^\prime) \) term, a negative term, is forced to become less negative in order to compensate. This requires a lower mass density, which, if the baryon density profile is held constant, manifests itself as a decrease in the DM mass density.

### 2.6 Statistical analysis and MultiNest

The model outlined above gives us an \( N \)-dimensional parameter space. To explore this parameter space and derive limits on the various observables we adopt nested sampling (Skilling 2006) as implemented in the publicly available MultiNest code (Feroz & Hobson 2008; Feroz et al. 2009; Feroz et al. 2013). MultiNest is a tool for Bayesian inference and parameter estimation. Bayes theorem is

\[
P(\theta | D, M) = \frac{P(D | \theta, M) P(\theta | M)}{P(D | M)} \tag{23}
\]

where \( M \) is the given model, \( \theta \) is the set of parameters for that model, and \( D \) is the data. The left hand side, \( P(\theta | D, M) \), is known as the posterior, while the three terms on the right are the likelihood \( P(D | \theta, M) = L(\theta) \), the prior \( P(\theta | M) \), and the Bayesian evidence \( P(D | M) \).

The Bayesian evidence, a.k.a. the marginal or model likelihood, is a measure of how well the model performs given the data, and can be expressed as

\[
Z = P(D | M) = \int P(D | \theta, M) P(\theta | M) d\theta. \tag{24}
\]

The performance of two different models given the same data can be compared using the Bayes factor:

\[
B_{01} = \frac{P(D | M_0)}{P(D | M_1)}. \tag{25}
\]

Assuming Gaussian errors it is possible to derive an empirical scale relating the Bayes factor to the strength of evidence for one model over another, as done in Trotta (2008). There, a \([\ln B_{01}] \) value of less than 1 is considered inconclusive, while values of 1.0, 2.5 and 5.0 are considered to give weak, moderate and strong evidence, respectively.

MultiNest takes a given prior probability distribution and likelihood function and calculates the posterior distribution and Bayesian evidence. Our likelihood function is based on the \( \chi^2 \) distribution:

\[
L(\theta) = \exp \left( -\frac{\chi^2}{2} \right). \tag{26}
\]

The value of \( \chi^2 \) is simply

\[
\chi^2 = \chi_{\theta}^2 + \chi_{\sigma_z}^2 + \chi_{K_{\theta}}^2, \tag{27}
\]

where

\[
\chi_{\theta}^2 = \sum_j \left( \frac{(\nu_{\text{data},j} - \nu_{\text{model},j})^2}{\sigma_{\nu,j}^2} \right), \tag{28}
\]

\[
\chi_{\sigma_z}^2 = \sum_j \left( \frac{(\sigma_{\nu,j,\text{data}} - \sigma_{\nu,j,\text{model}})^2}{\sigma_{\nu,j}^2} \right), \tag{29}
\]

\[
\chi_{K_{\theta}}^2 = \sum_j \left( \frac{(K_{\theta,j,\text{data}} - K_{\theta,j,\text{model}})^2}{\sigma_{K_{\theta},j}^2} \right). \tag{30}
\]

Note that if the reconstruction model does not contain a tilt term (e.g. \( T = 0 \)) then \( \chi_{K_{\theta}}^2 \) is set to zero.

### Table 2. Prior ranges for parameters. Gaussian priors are described by a median \( \mu \) and a standard deviation \( \sigma \). Note that \( \nu(0) \) and \( \sigma_z(0) \) are the tracer density and velocity dispersion at \( z = 0 \), while quantities subscripted with \( \mu \) and \( \sigma \) are the values derived from data in the 0th bin, whose bin centre \( z_0 > 0 \). Tracer density \( \nu(0) \) has units of \([\text{stars kpc}^{-3}]\). \( K_{\theta} \) and \( K_{DD} \) terms have units \([\text{km} \cdot \text{pc}^{-1} \cdot \text{s}^{-2}]\).

<table>
<thead>
<tr>
<th>Model parameter</th>
<th>Range or Gaussian ( \mu ) &amp; SD</th>
<th>Type</th>
</tr>
</thead>
<tbody>
<tr>
<td>Baryons</td>
<td>( K_{\theta} )</td>
<td>( \mu = 1500, SD = 150 )</td>
</tr>
<tr>
<td></td>
<td>( D_{\theta} )</td>
<td>( \mu = 0.18 \text{ kpc}, SD = 0.02 \text{ kpc} )</td>
</tr>
<tr>
<td>Constant DM</td>
<td>( \rho_{\text{DM}} )</td>
<td>([2, 20] \times 10^{-3} \text{M}_{\odot} \text{pc}^{-3} )</td>
</tr>
<tr>
<td></td>
<td></td>
<td>([0.076, 0.706] \text{GeV cm}^{-3} )</td>
</tr>
<tr>
<td>DD</td>
<td>( K_{DD} )</td>
<td>([0, 1500] )</td>
</tr>
<tr>
<td></td>
<td>( D_{DD} )</td>
<td>([1.5, 3.5] \text{ kpc} )</td>
</tr>
<tr>
<td>Tilt term</td>
<td>( A )</td>
<td>([0, 200] \text{ km}^2 \text{s}^{-1} )</td>
</tr>
<tr>
<td></td>
<td>( n )</td>
<td>([1.0, 1.9] )</td>
</tr>
<tr>
<td></td>
<td>( R_0 )</td>
<td>( \mu = 2.5 \text{ kpc}, SD = 0.5 \text{ kpc} )</td>
</tr>
<tr>
<td>Tracer density</td>
<td>( \nu(0) )</td>
<td>([0, \nu_0 + 2 \times SD_{\nu,0}] )</td>
</tr>
<tr>
<td></td>
<td>( h )</td>
<td>([0.4, 1.4] \text{ kpc} )</td>
</tr>
<tr>
<td>Velocity disp.</td>
<td>( \sigma_z(0) )</td>
<td>([\sigma_{\nu,0} - 2 \times SD_{\nu,0}, \sigma_{\nu,0} + 2 \times SD_{\nu,0}] )</td>
</tr>
</tbody>
</table>

Table 2 shows the prior ranges used for our analyses. We derive credible regions (CRs) for the DM density parameters by taking its posterior distribution and marginalizing over the remaining parameters. As outlined above our model has several components that can be turned on or off, such as the DD. Using the Bayesian evidence as calculated by MultiNest it is potentially possible to perform model comparison to determine which reconstruction model best fits the data. This idea will be explored in greater depth in subsequent studies.

### 3 MOCK DATA SETS

In this paper we apply our method only to mock data in order to hone and verify it. This mock data is ‘as good as it gets’, in the sense that it has no measurement errors, nor observational biases added to it, and is drawn from relatively simple, known distribution functions. While dynamically realistic ‘N-body’ mocks are preferred, it has already been shown that 1D Jeans analyses are robust to the presence of local non-axisymmetric structure in the disc (Garbari et al. 2011). Furthermore, N-body mocks are expensive; even state-of-the-art simulations do not approach the local sampling expected from Gaia. Finally, the expense of building well-sampled N-body mocks makes it challenging to explore a large parameter space of models, including models with and without tilt, or with and without a DD. For these reasons, we focus here on simpler mock data, similarly to the Read (2014) review. We will consider more dynamically realistic mocks, and mocks that give a faithful representation of the expected Gaia data uncertainties in future work.

We generate here a variety of mock data sets as outlined in Table 3, with different samplings (104, 105, and 106 tracer stars), with and without a tilt term, and also with either no DD, a DD ( núi ), or a more massive ‘big dark disc’ ( h núi ). We assume the axial term \( A \) and the rotation curve term \( \mathcal{R} \) are zero. Our simple baryonic disc model is set up to mimic the real Milky Way, with a scale height parameter of \( D_{\theta} = 0.18 \text{ kpc} \) and a surface density of \( \Sigma_{\theta} = 55.53 \text{ M}_{\odot} \text{pc}^{-2} \), similar to those measured near the solar position (Flynn et al. 2006; Read 2014; McKee et al. 2015). The value of the \( F \) parameter (see equation 15) corresponds to a DM density of \( \rho_{\text{DM, cont}} = 10 \times 10^{-3} \text{M}_{\odot} \text{pc}^{-3} = 0.38 \text{ GeV cm}^{-3} \). For
each scenario we generate 20 mock data sets, allowing us to explore the effects of Poisson noise over a range of realizations.

The mock data consist of a list of stars each with three pieces of data: the position $z$, the vertical velocity $v_z$, and the product of the vertical and radial velocities, $v_R v_z$. The first element, $z$, is generated by drawing randomly from an exponential tracer distribution with scale height $h$:

$$v(z) = \exp \left(-\frac{z}{h}\right).$$

This gives us our list of stellar positions.

To derive the velocities for the mock catalogue we first must define a mass model and tilt. Taking the same parametrizations as described in Sections 2.3, 2.4, and 2.5 for baryons, DM, and tilt, respectively, we set their parameters as per Table 3. This allows us to calculate $K_0$ and $T(z)$. Using equation (5) and its associated assumptions we can calculate $C$. We then use equation (3) to derive $\sigma_R(z)$. For each star we take its position $z'$, find the value of $\sigma_R(z')$, and then draw a velocity from a Gaussian centred on $v_z = 0$ and with variance $\sigma^2_R(z')$.

To generate $v_R v_z$ mock data, when necessary, we take the $A$, $n$, and $R_0$ parameters listed in Table 3 and calculate the $\sigma_R(z)$ profile via equation (20). Taking each star’s position $z'$, we calculate $\sigma_R(z')$, and draw a value of $v_R v_z$ from a Gaussian centred on $v_R v_z = 0$ and with variance of $\sigma^2_R(z')$.

### 4.1 Sampling

Fig. 3 shows reconstructions of the local DM density for using varying numbers of tracer stars. The mock data sets used here are the most simple case, thick_tilt_X_M as described in Table 3, and has no DD or tilt added. As expected the CRs shrink as the number of tracer stars is increased from 10$^3$ stars up to 10$^6$ stars. The SDSS has sampling of the order of 10$^4$ stars (Zhang et al. 2013; Budenbender et al. 2015) while data from Gaia will give upwards of 10$^6$ tracer stars.

### 4.2 Tilt

In Fig. 4 we explore the effects of the tilt term. The left-hand set of CRs in Fig. 4 is the same as the right hand set of CRs from Fig. 3, with no tilt term in the mock data or reconstruction. In the centre set of Fig. 4, however, the mock data thick_tilt_1e6_M has the tilt term turned on, but the analyses are performed with the tilt term set to zero. This illustrates the danger of ignoring tilt as discussed earlier in Section 2.5: the reconstructions return narrow CRs, but as expected they systematically underestimate the value of $\rho_{DM}$, with the true $\rho_{DM}$ lying outside even the 99.7 per cent CRs. This underestimation, however, is remedied by turning on the tilt term in the model, as shown in the right hand set CRs in Fig. 4, where the correct DM density is always at least within the 99.7 per cent CR. The inclusion of extra parameters to describe the tilt term does, however, increase the size of the CRs.

Just as our determination of $\rho_{DM}(z)$ is dependent on the tilt term, the tilt term is in turn dependent on its input parameters, $A$, $n$, and $R_0 = R_1$. While we have been able to fit for $A$ and $n$ using G-type dwarf data (Section 2.5), for $R_0$ we have taken the canonical value of $R_0 = 2.5 \pm 0.5$ kpc from Binney & Tremaine (2008), and further made the assumption that $\sigma_R(R, z)$ has the same radial profile as $v_R(z)$ (i.e. $R_0 = R_1$). When using only a single population, determination of $R_0$ and $R_1$ will be important, as illustrated in Fig. 5. This figure shows the 2D marginalized posterior for $\rho_{DM, const}$ and $R_0$ for a reconstruction of mock data set thick_tilt_1e6_0 with a model including tilt, and demonstrates the degeneracy that
Figure 3. Determination of the DM density profile for varying numbers of tracer stars. These plots show marginalized posteriors for $\rho_{DM}(z) = \rho_{DM,\text{const}}$ for the 20 mock data sets generated for each sampling level. Dark, medium, and light shading show the 68 per cent, 95 per cent, and 99.7 per cent CRs, respectively. Green, purple, and orange colouring indicates that the 68 per cent, 95 per cent, or 99.7 per cent CR, respectively, contains the correct answer. The median value of each posterior is shown by a solid line in green, purple, or orange, while the DM density value used to generate the mock data is shown as a solid black line across the entire plot. The mock data and reconstruction models used contain no tilt term or DD terms. As the number of tracer stars is increased from $10^4$ to $10^6$ the CRs for $\rho_{DM,\text{const}}$ naturally shrink around the mock data value.

Figure 4. Determination of the DM density profile using $10^6$ tracer stars and exploring the effects of including or neglecting the tilt term in the mock data sets and reconstruction models. These plots show marginalized posteriors for $\rho_{DM}(z) = \rho_{DM,\text{const}}$ for the 20 mock data sets generated for each tilt scenario. Dark, medium, and light shading show the 68 per cent, 95 per cent, and 99.7 per cent CRs, respectively. Green, purple, and orange colouring indicates that the 68 per cent, 95 per cent, or 99.7 per cent CR, respectively, contains the correct answer, while pink colouring indicated that the correct answer lies outside even the 99.7 per cent CR. The median value of each posterior is shown by a solid line in green, purple, or orange, while the DM density value used to generate the mock data is shown as a solid black line across the entire plot. The left set of mocks shows the same result as the right set of Fig. 3, with no tilt term in mock data or reconstruction model. The centre set has a tilt term in the mock data but not in the reconstruction model, yielding a systemic underestimation of $\rho_{DM}$. The right set has a tilt term in both mock data and reconstruction model, demonstrating that our method can successfully deal with the tilt term.

exists between $\rho_{DM,\text{const}}$ and $R_0 = R_1$. If using multiple tracer population, each will have different $R_0$ and/or $R_1$, but will all have their motions dictated by the same potential. This will help us break the degeneracy between $R_0$, $R_1$, and $\rho_{DM}$. Further, with the advent of Gaia data we will be able to directly measure the radial profile of $\sigma_R$ and $v(R, z)$ for a given set of tracer stars.

### 4.3 Dark disc

Fig. 6 shows mock data set thick_1e6_0-19, thick_dd_1e6_0, and thick_bdd_1e6_0 reconstructed using models with and without a DD component. The top left panel shows the same CRs as seen in the left hand set of Fig. 4 and the right hand set of Fig. 3.

The left column of Fig. 6 shows the reconstruction of mock data sets with a constant DM density profile: mocks 0–19 in top left and mock 0 in bottom left. The reconstruction in the top panel uses a model with a constant DM density, while the bottom panel uses a model with an additional DD component. The DD reconstruction exhibits a disc structure even though the correct answer is constant $\rho_{DM}$. This is likely due to a bias in the hyper-volume set by the priors – the prior range on the DD parameters goes between no DD ($K_{DD} = 0$) and a maximal DD ($K_{DD} = 1500$), and thus the bulk of the parameter space features at least some DD. There is no ‘negative dark disc’ to counteract this effect and push the mean of the prior range back to no DD.

In the centre and right columns of Fig. 6 we reconstruct mock data sets with a DD of different masses: thick_dd_1e6_0 and thick_bdd_1e6_0 (the ‘big dark disc’). A constant DM density reconstruction (top row) is able to contain the thick_dd_1e6 DD within the 95 per cent CR almost to the last bin, but fails to contain the big dark disc beyond $z = 1.3$ kpc. Adding a DD term allows the reconstruction to fit to the mock data DM profile nicely, as shown in the bottom centre and bottom right panels of Fig. 6.
Measuring the local dark matter density

Figure 5. 2D marginalized posterior of the constant DM density \( \rho_{\text{DM}, \text{const}} \) and the \( R_0 \) parameter of the tilt term (equation 19), generated from mock data set thick_tilt_1e6_0 reconstructed using a model with tilt. Contours show the 68 per cent, 95 per cent, and 99.7 per cent CRs. Marginalization and plotting performed using BRRETT (Sebastian Liem, private communication).

When working with real data, however, we will not have the luxury of knowing the DM profile underlying the data. The black mock data model line will not be there. In subsequent studies we will explore the potential use of the Bayesian evidence, as calculated by MULTINEST, to determine if a DD is justified by observational data.

4.4 Tilt and dark disc

Here, we combine the two elements discussed in previous sections, the tilt term and the DD. Fig. 7 shows reconstructions of the mock data set thick_bdd_tilt_1e6_0. In the top panel the reconstruction model contains neither DD nor tilt term. Again, we see the same effects as we did previously. The missing tilt term yields a consistent underestimation of the DM density, and the constant DM density envelope is too narrow to encompass the density range of the DD. The consistent underestimation can be remedied by adding a tilt term to the reconstruction model, as shown in the middle panel, where the reconstruction model has a tilt term included. Both problems can be resolved in tandem by using a reconstruction model with both tilt and DD terms, as shown in the bottom panel of Fig. 7.

Figure 6. Determination of the DM density profile using \( 10^6 \) tracer stars and exploring the effects of including or neglecting a DD in the mock data sets and reconstruction models. For comparison the top left panel shows the same reconstructions as the right hand set from Fig. 3, i.e. basic mock data sets with no DD, reconstructed with a constant DM density. The remaining panels each show one representative mock and reconstruction, with the full set of mocks and reconstructions given in Appendix B. These panels show marginalized posteriors for \( \rho_{\text{DM}(z)} \), with dark, medium, and light shading indicating the 68 per cent, 95 per cent, and 99.7 per cent CRs, respectively. The median value of the posterior is shown by the solid blue line, while the DM density profile used to generate the mock data is shown as a solid black line. From left to right the columns show reconstructions of mock data sets containing no DD (\( \rho_{\text{DM, const}} \) only), a DD, or a ‘big’ dark disc (BDD). The top centre and top right panels show the determination of a constant DM density from the DD and BDD mock data sets, respectively. The bottom row shows the reconstruction of each of the mock data sets using a model with a constant DM term and a DD term. Light grey vertical lines indicate the bin centres.
5 CONCLUSIONS

In this paper we have presented a new method of determining the vertical DM density profile, and thus the local DM density. The key equation of this method, equation (3), depends only on the assumption of dynamical equilibrium. In practice, further assumptions are made to describe the components going into equation (3); however the scope of possible models is wide, and in principle model selection using the Bayesian evidence can be used to determine the best model. We leave exploration of this last aspect to future work.

The importance of the tilt term has been previously noted (Kuijken & Gilmore 1989a; Smith et al. 2009; Büdenbender et al. 2015), and to derive an accurate value for the local DM density it is vital that our method be able to deal with this term. The baryonic contribution to the galactic density profile drops rapidly as \( z \) increases, leaving DM as the dominant component. Thus to probe the DM density profile sampling thick disc stars that travel higher above the plane are preferable. For these populations the tilt term is more important, as illustrated in Fig. 2. Failure to include the tilt term in the analysis leads to a systematic underestimation of the local DM density, as explained in Section 2.5 and demonstrated in Section 4.2.

One of the novel aspects of our method is that it can deal with the tilt term while remaining within the confines of the one-dimensional \( z \)-direction Jeans equation, which can be seen in Fig. 4. With only the data currently available for the Milky Way, this requires several well-motivated assumptions, as described in Section 2.5. However, with data from \textit{Gaia} we will be able to directly measure the radial profile of the tilt and tracer density, removing the need for such assumptions. While for this paper we have disregarded the rotation curve term (equation 6), we note that an accurate determination of this will be necessary for the implementation of this or other \( z \)-direction methods to real data.

Non-spherical DM density profiles, such as oblate haloes or accreted DDs, can also be fitted using our method by incorporating a DD term into the reconstruction model, which is shown in Fig. 6. Our method can also reconstruct mock data sets containing both a tilt term and a DD, as shown in Fig. 7.

ACKNOWLEDGEMENTS

We thank Michael Feyereisen, Sebastian Liem, Pat Scott, Roberto Trotta, Stephen Feeney, Roberto Ruiz de Austri, Scott Tremaine, Vera Gluscevic, and Juna Kollmeier for very useful discussions. SS acknowledges support from the Swedish Research Council (VR, Contract No. 350-2012-268). JIR would like to acknowledge support from STFC consolidated grant ST/M000990/1 and the MERAC foundation. GB acknowledges support from the European Research Council through the ERC starting grant \textit{WIMPs Kairos}.

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Figure 7. Determination of the DM density profile using \( 10^6 \) tracer stars with a combination of a tilt term and a DD in the mock data (thick\_bdd\_tilt\_1e6) and using a variety of reconstruction models. These plots show marginalized posteriors for \( \rho_{DM}(z) \), with dark, medium, and light shading indicating the 68 per cent, 95 per cent, and 99.7 per cent CRs, respectively. The median value of the posterior is shown as the solid blue line, while the DM density profile used to generate the mock data is shown as a solid black line.
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Klasen M., Pohl M., Sigl G., 2015, Prog. Part. Nucl. Phys., 85, 1
Pasetto S. et al., 2012a, A&A, 547, A70
Pasetto S. et al., 2012b, A&A, 547, A71

**APPENDIX A: VARIATION DUE TO THE NUMBER OF BINS**

Fig. A1 illustrates the effects of changing the number of bins used for this analysis. There we plot the 68 per cent, 95 per cent, or 99.7 per cent CRs for 20 mock data sets (thick_1E6_0-19, no tilt and no DD), and vary the number of bins from five to 30, in increments of five. For only five bins (top left set), the true answer for $\rho_{DM, const}$ is outside the 99.7 per cent CR for all but two of the mocks, for which the true answer is within only the 99.7 per cent CR. The systematic underestimation is due to an overestimation of the baryonic disc. The fifth bin in this scheme covers a range from 1.2 to 2.4 kpc and has its bin centre at $z = 1.6$ kpc, so it is unsurprising that such a low number of bins fail to correctly reconstruct the DM profile, which is a subdominant component, only becomes apparent at higher $z$. Increasing the number of bins to 10, 15, and then to 20 improves the results. The gains from increasing from 20 to 25 and 30 bins are very slight.

**APPENDIX B: DARK DISC RECONSTRUCTIONS WITH MULTIPLE MOCKS**

Here we show the results of reconstructing all mocks data sets generated using a constant DM density only (thick_X_M), constant DM and a regular DD (thick_dd_X_M), or constant DM and a ‘big’ DD (thick_bdd_X_M) (see Table 3 for details). The reconstruction models either have a constant DM density only or a constant DM density plus a DD.
Figure A1. Exploring the effects of binning on the determination of $\rho_{\text{DM}}$. These plots show marginalized posteriors for $\rho_{\text{DM}}(z) = \rho_{\text{DM, const}}$ for the 20 mock data sets thick_1E6_0-19, reconstructed using 5, 10, 15, 20, 25, and 30 bins. Dark, medium, and light shading show the 68 per cent, 95 per cent, and 99.7 per cent CRs, respectively. Green, purple, and orange colouring indicates that the 68 per cent, 95 per cent, or 99.7 per cent CR, respectively, contains the correct answer, while pink colouring indicated that the correct answer lies outside even the 99.7 per cent CR. The median value of each posterior is shown by a solid line in green, purple, or orange, while the DM density value used to generate the mock data is shown as a solid black line across the entire plot.
Figure B1. Marginalized posteriors of $\rho_{DM}(z)$ for mock data sets `thick_1E6_0-19` with $\rho_{DM,\text{const}}$ only (no DD component) reconstructed using a model with $\rho_{DM,\text{const}}$ plus a DD. Dark, medium, and light shading indicate the 68 per cent, 95 per cent, and 99.7 per cent CRs, respectively. The median value of the posterior is shown as the solid blue line, while the DM density profile used to generate the mock data is shown as a solid black line.
Figure B2. Marginalized posteriors of $\rho_{\text{DM}}(z)$ for mock data sets `thick_dd_1E6_0−19` with $\rho_{\text{DM, const}}$ plus a DD component, reconstructed using a model with $\rho_{\text{DM, const}}$ only. Dark, medium, and light shading indicate the 68 per cent, 95 per cent, and 99.7 per cent CRs, respectively. The median value of the posterior is shown as the solid blue line, while the DM density profile used to generate the mock data is shown as a solid black line.
Figure B3. Marginalized posteriors of $\rho_{\text{DM}}(z)$ for mock data sets thick_dd_1E6_0-19 with $\rho_{\text{DM, const}}$ plus a DD component, reconstructed using a model with $\rho_{\text{DM, const}}$ and a DD component. Dark, medium, and light shading indicate the 68 per cent, 95 per cent, and 99.7 per cent CRs, respectively. The median value of the posterior is shown as the solid blue line, while the DM density profile used to generate the mock data is shown as a solid black line.
Figure B4. Marginalized posteriors of $\rho_{DM}(z)$ for mock data sets thick_bdd_1E6_0-19 with $\rho_{DM,\text{const}}$ plus a DD component, reconstructed using a model with $\rho_{DM,\text{const}}$ only. Dark, medium, and light shading indicate the 68 per cent, 95 per cent, and 99.7 per cent CRs, respectively. The median value of the posterior is shown as the solid blue line, while the DM density profile used to generate the mock data is shown as a solid black line.
Figure B5. Marginalized posteriors of $\rho_{\text{DM}}(z)$ for mock data sets thick_bdd_1E6_0−19 with $\rho_{\text{DM, const}}$ plus a ‘big’ DD component, reconstructed using a model with $\rho_{\text{DM, const}}$ and a DD component. Dark, medium, and light shading indicate the 68 per cent, 95 per cent, and 99.7 per cent CRs, respectively. The median value of the posterior is shown as the solid blue line, while the DM density profile used to generate the mock data is shown as a solid black line.
Figure B6. Marginalized posteriors of $\rho_{DM}(z)$ for mock data sets thick_bdd_tilt_X_0–19 with $\rho_{DM, \text{const}}$ plus a ‘big’ DD component and including the effects of tilt, reconstructed using a model with $\rho_{DM, \text{const}}$, a DD component, and incorporating tilt. Dark, medium, and light shading indicate the 68 per cent, 95 per cent, and 99.7 per cent CRs, respectively. The median value of the posterior is shown as the solid blue line, while the DM density profile used to generate the mock data is shown as a solid black line.

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