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IDENTIFICATION OF A LIKELY RADIO COUNTERPART TO THE RAPID BURSTER

CHRISTOPHER B. MOORE,¹ ROBERT E. RUTLEDGE,² DEREK W. FOX,³ ROBERT A. GUERRIERO,⁴
WALTER H. G. LEWIN,³ ROBERT FENDER,⁵ AND JAN VAN PARADIJS^{5,6,7}

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ABSTRACT

We have identified a likely radio counterpart to the low-mass X-ray binary MXB 1730–335 (the Rapid Burster). The counterpart has shown 8.4 GHz radio on/off behavior correlated with the X-ray on/off behavior as observed by the *RXTE*/ASM during six VLA observations. The probability of an unrelated, randomly varying background source duplicating this behavior is 1%–3% depending on the correlation timescale. The location of the radio source is R.A. 17^h33^m24^s.61, decl. –33°23′19″.8 (J2000), $\pm 0'.1$. We do not detect 8.4 GHz radio emission coincident with type II (accretion-driven) X-ray bursts. The ratio of radio to X-ray emission during such bursts is constrained to be below the ratio observed during X-ray-persistent emission at the 2.9 σ level. Synchrotron bubble models of the radio emission can provide a reasonable fit to the full data set, collected over several outbursts, assuming that the radio evolution is the same from outburst to outburst but given the physical constraints the emission is more likely to be due to ~ 1 hr radio flares such as have been observed from the X-ray binary GRS 1915+105.

Subject headings: radiation mechanisms: nonthermal — radio continuum: stars — stars: individual (MXB 1730–335) — X-rays: stars

1. INTRODUCTION

Discovered in 1976 (Lewin et al. 1976), the Rapid Burster (MXB 1730–335, hereafter RB) is located in the highly reddened globular cluster Liller 1 (Liller 1977), which has a distance modulus of 14.68 ± 0.23 , corresponding to 8.6 ± 1.1 kpc, determined by main-sequence fitting (Frogel, Kuchinski, & Tiede 1995). Liller 1 has a small optical core radius (about 6''.5).

Radio observations of transient X-ray binaries have found several to be radio transients as well (see Hjellming & Han 1995 for a review), both black hole candidates (A0620–00, GS 2000+25, GS 2023+33, GS 1124–68) and neutron stars (Aql X-1, Cir X-1, Cen X-4). The outbursts of X-ray transients are due to a sudden turn-on of accretion onto the compact object in a binary lasting from several days to months (Tanaka & Lewin 1995; Chen, Shrader, & Livio 1997). The radio spectral and temporal behavior in some of these objects is described by a synchrotron bubble model (Van der Laan 1966), which indicates that there are plasma outflows associated with the X-ray outburst. In particular, the black hole candidate and superluminal jet source GRS 1915+105 has been seen to exhibit correlated behavior at X-ray, infrared, and radio wavelengths (Fender et al. 1997; Mirabel et al. 1998; Eikenberry et al. 1998a, 1998b; Fender & Pooley 1998) that is readily understood in terms of the synchrotron emission of expanding plasmoids that have been ejected from the inner regions of the system

(Fender et al. 1997; Mirabel et al. 1998; Fender & Pooley 1998).

Despite its frequent X-ray outbursts (average interval, ~ 220 days; Guerriero et al. 1999), several previous studies have not detected radio emission from the RB (Johnson et al. 1978; Lawrence et al. 1983; Grindlay & Seaquist 1986; Johnston, Kulkarni, & Goss 1991; Fruchter & Goss 1995). Improvements in radio sensitivity since some of those studies were performed made the detection of an RB radio counterpart at previously unattainable flux levels practical. Furthermore, the advent of the *RXTE* all-sky monitor made it possible to correlate radio observations with a well-sampled X-ray light curve. We therefore undertook a search for a radio counterpart of the Rapid Burster in X-ray outburst. Since the type II X-ray bursts of the Rapid Burster are thought to be caused by the same phenomenon (spasmodic accretion) as the X-ray outbursts (van Paradijs 1996), one might expect that they are accompanied by simultaneous radio emission. The observations reported here marginally exclude (at the 2.9 σ level) simultaneous radio burst emission from the likely radio counterpart.

1.1. X-Ray Behavior

The RB is a transient X-ray source that has been observed during the past few years to go into outburst approximately every 220 days for a period of ~ 30 days (Guerriero et al. 1999). It is the only low-mass X-ray binary (LMXB) that produces two different types of X-ray bursts (Hoffman, Marshall, & Lewin 1978). Type I bursts, which are observed from ~ 40 other LMXBs, are due to thermonuclear flashes on the surface of an accreting neutron star. Type II bursts, which have been observed from only one other LMXB (GRO J1744–28; Kouveliotou et al. 1996; Lewin et al. 1996), are sudden releases of gravitational potential energy resulting from accretion instabilities. For a detailed review of type I and type II bursts, see Lewin, van Paradijs, & Taam (1993, 1995).

The RB does not fit easily into the “Z/atoll” paradigm of LMXBs (Rutledge et al. 1995), in which the correlated X-ray fast-timing and spectral behavior of LMXBs divides the

¹ Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, MS 78, Cambridge, MA 02138; cmoore@cfa.harvard.edu.

² Space Radiation Laboratory, California Institute of Technology, MC 220-47, Pasadena, CA 91125; rutledge@srl.caltech.edu.

³ Massachusetts Institute of Technology, Department of Physics, 37-627, Cambridge, MA 02139; derekfox@space.mit.edu, lewin@space.mit.edu.

⁴ United States Military Academy, Department of Physics, West Point, NY 10996; hr6958@exmail.usma.army.mil.

⁵ Astronomical Institute Anton Pannekoek, University of Amsterdam, Center for High-Energy Astrophysics, Kruislaan 403, 1098 SJ Amsterdam, Netherlands; rpf@astro.uva.nl.

⁶ University of Alabama, Huntsville, AL; jvp@astro.uva.nl.

⁷ Deceased.

population into two classes (Hasinger & van der Klis 1989; van der Klis 1995). However, the RB is known to exhibit periods of behavior characteristic of an atoll source. During an outburst in 1983, the RB exhibited strong persistent emission (PE) and type I bursts with no type II bursts, which is behavior typical of the atoll sources (Barr et al. 1987). This behavior has also been seen in outbursts observed more recently with *RXTE*. For the first ~ 17 days of these outbursts, strong PE is accompanied only by type I X-ray bursts. Type II bursting behavior begins after this period and continues while the PE decreases (details of these observations are given in Guerriero et al. 1999).

In spite of 20 years of theoretical modeling, no satisfactory model exists for the disk instability that drives the type II bursts (for a review, see Lewin et al. 1995). Models that require weak magnetic fields ($\lesssim 10^9$ G) were effectively excluded with the discovery of GRO J1744–28, which contains a pulsar with a magnetic field of $\approx 10^{11}$ G (Finger et al. 1996; Cui 1997) and also exhibits type II bursts (Kouveliotou et al. 1996; Lewin et al. 1996).

1.2. Previous Radio Observations

There have been several radio studies of fields of view containing Liller 1, some with the goal of identifying a radio counterpart of the Rapid Burster. Some have focused on finding radio variability in hopes of catching bursts in the radio band correlated with type II X-ray bursts, while others have sought persistent radio sources. We summarize the most stringent limits on the flux density of a persistent radio source in Table 1; additional, less stringent limits are summarized by Lawrence et al. (1983).

Fruchter & Goss (1995, hereafter FG) discovered a radio source in three bands (0.33, 1.5, and 4.5 GHz), with a spectral slope of ~ -2 . They observe the flux density of this source at 1.5 GHz above the upper limit derived from an observation at another time (Johnston et al. 1991). FG interpret this discrepancy as the result of beam dilution in the Johnston et al. observations ($0''.5$ beam obtained with the VLA in A-array) over resolving a large population of radio pulsars in Liller 1. Thus, the interpretation of this radio source was as an integration over a population of radio pulsars in Liller 1. To date, no radio pulsations have been detected from the direction of Liller 1 (R. Manchester 1999, private communication).

Prior to the present work, there have been two studies that produced limits on radio emission simultaneous with

type II X-ray bursts of the RB. Johnson et al. (1978) observed simultaneously in the radio and X-ray bands during a period when a total of 64 type II bursts were observed in the X-rays and placed upper limits on the simultaneous flux density of radio bursts (at 2.7 and 8.1 GHz) of ~ 20 mJy. Rao & Venugopal (1980) observed at 0.33 GHz during two X-ray bursts and placed an upper limit of 0.2 Jy on the radio flux density during the bursts.

There are claims of radio burst detections from the RB without simultaneous X-ray observations (Calla et al. 1979, 1980a, 1980b; O. P. N. Calla, private communication cited in Johnson et al. 1978). Calla et al. report observing approximately 14 radio bursts on nine different days with peak flux densities of 400–600 Jy at 4.1 GHz and durations of 10–500 s. This phenomenon has not, however, been confirmed at other observatories, and the reported flux densities are substantially above the limits placed by Johnson et al. (1978) during their simultaneous X-ray/radio observations. If the reported radio bursts are real, they do not appear to be correlated with type II X-ray bursts (Lawrence et al. 1983). Using the 16.7 ks of observations at 8.4 GHz obtained in the course of this work, we find a 3σ upper limit of 250 mJy on the flux density of our likely radio counterpart during any single 3.3 s integration. Thus, we see no evidence for radio flares of the type reported by Calla et al. during either the X-ray active or quiescent periods.

1.3. Objectives

The goals of the present work are first to search for a radio counterpart of the Rapid Burster, detectable either in X-ray outburst or in quiescence, and second to determine whether the radio emission can be tied to the active accretion during an outburst.

In § 2 we describe the radio and X-ray observations and present some general results. In § 3, we present the results of our analyses of the radio observations in the context of the “standard model” of radio emission from X-ray transients: the synchrotron bubble model. We evaluate the likelihood of the counterpart identification in § 4, discuss the results of these observations in § 5, and list our conclusions in § 6.

2. OBSERVATIONS

X-ray observations were made with two instruments on the *Rossi X-Ray Timing Explorer* (*RXTE*; Bradt, Rothschild, & Swank 1993): the all-sky monitor (ASM; Levine et al. 1996) and the proportional counter array (PCA; Zhang

TABLE 1
PREVIOUS RADIO OBSERVATIONS OF PERSISTENT SOURCE

Reference	Frequency (GHz)	Measurement (μ Jy)	Epoch	Rapid Burster State	Note
Fruchter & Goss 1995	0.33	9000 ± 1000	1990 Aug 1–3	Unknown	VLA DnC
	1.5	280 ± 50	1993 25 May	Unknown	VLA CnB
	4.5	95 ± 14	1990 Jul –1991 Jan	Unknown	11 observations; VLA C, B
Johnston et al. 1991	1.5	<180 (4σ)	1990 Apr 9–14	Unknown	VLA A
Grindlay & Seaquist 1986	4.5	<380 (3σ)	1982 Jun 21–22	Unknown	VLA A
Johnson et al. 1978	2.7	$(3 \pm 2) \times 10^3$	1977 Apr 23, 24, 27	Bursting	NRAO Greenbank
	8.1	$(5 \pm 3) \times 10^3$	1977 Apr 23, 24, 27	Bursting	
Lawrence et al. 1983	15.5	$<93 \times 10^3$ (3σ)	1980 Apr 20	Not Bursting	Haystack Observatory
	14.5	$<294 \times 10^3$ (3σ)	1980 Apr 17 –Jul 28	Not Bursting	University of Michigan
	8	$<636 \times 10^3$ (3σ)	1980 Apr 17 –Jul 28	Not Bursting	
	5	$<60 \times 10^3$ (3σ)	1980 Jul 31–Aug 4	Not Bursting	Parkes, CSIRO

NOTE.—Additional observational limits on variable radio emission may be found in Lawrence et al. (1983); the periods when the RB was “not bursting” may have had persistent emission, which was difficult to distinguish from the nearby source MXB 1728–34.

et al. 1993; Jahoda et al. 1996). The ASM consists of three 1.5–12 keV X-ray proportional counters with coded-aperture masks. It obtains observations of approximately 80% of the sky every 90 minutes and is sensitive to persistent X-ray sources down to ~ 5 mcrab (3σ detection) in a typical day's cumulative exposure (Remillard & Levine 1997). The standard data products of the ASM include the daily average count rate of a known catalog of X-ray sources including the RB.⁸ A summary of our radio observations and the contemporaneous ASM X-ray flux measurements is shown in Table 2. The PCA is a collimated array of gas-filled proportional counters, sensitive in the 2–60 keV energy range, with a field of view (FOV) of $\sim 1^\circ$ and a geometric area of ~ 6500 cm². In total, five simultaneous PCA and radio observations were performed; they are discussed in detail below. When the PCA is pointed directly at the RB, there is a nearby ($\sim 0.5^\circ$) persistent X-ray source (4U 1728–34) that contributes to the measured X-ray flux. During part of these observations, the PCA is pointed offset from the RB at a reduced collimator efficiency, excluding 4U 1728–34 entirely from the field of view. The uncertainty in the aspect correction of the PCA dominates our uncertainty in the RB count rates, which throughout this paper are given as the count rate for the RB only (aspect corrected, background subtracted, with the count rate from 4U 1728–34 assumed constant over the 1 hr observation, subtracted).

The radio observations were performed at three different observatories. The Very Large Array (VLA)⁹ observed every outburst discussed in this work and made the first detection of the radio counterpart reported here. VLA observations were obtained in two closely spaced bands, each of 50 MHz nominal bandwidth. Each band was

observed in both right-circular and left-circular polarization, making the total observed bandwidth 100 MHz for each polarization. The complex antenna gains were set using observations of a nearby compact radio source (1744–312), and the flux density scale was determined using observations of 3C 286.

The Rapid Burster was observed on two occasions with the Australia Telescope Compact Array (ATCA).¹⁰ The complex antenna gains were set using observations of the nearby phase calibrator PKS B1657–261. The observations were made simultaneously in two orthogonal linear polarizations at 4800 and 8640 MHz with a bandwidth of 128 MHz at each frequency. For both observations the array was in the extended 6A configuration with baselines in the range 337–5939 m.

The Submillimeter Common User Bolometer Array (SCUBA; Cunningham et al. 1994) on the James Clerk Maxwell Telescope (JCMT)¹¹ observed on 1998 30 January UT using the 850 μ m system in photometry mode.

To date, there have been five outbursts of the RB observed with *RXTE*/ASM. The progression of four of these outbursts, including the evolution of the RB's bursting behavior, is described elsewhere (Guerriero et al. 1999). During three of these outbursts, we performed radio observations with the VLA, ATCA, or SCUBA, the results of which are listed in Table 2. We describe the observations in detail in the following sections.

2.1. 1996 November Outburst

Radio observations with the VLA in A-configuration were made at 8.4 GHz while the RB was in X-ray quiescence on 1996 October 14. These produced low upper limits in

¹⁰ The Australia Telescope is funded by the Commonwealth of Australia for operation as a National Facility managed by the CSIRO.

¹¹ The James Clerk Maxwell Telescope is operated by The Joint Astronomy Centre on behalf of the Particle Physics and Astronomy Research Council of the United Kingdom, the Netherlands Organization for Scientific Research, and the National Research Council of Canada.

TABLE 2
RADIO AND X-RAY OBSERVATIONS RESULTS

Date (UT)	Observatory (configuration)	ν (GHz)	S_ν (μ Jy)	<i>RXTE</i> /ASM 1 day Avg (counts s ⁻¹)	<i>RXTE</i> /PCA (counts s ⁻¹)
1996 Oct 14.....	VLA (D \rightarrow A)	8.4	45 \pm 30	0.15 \pm 0.31	...
1996 Nov 06.8.....	VLA (A)	8.4	370 \pm 45	9.8 \pm 0.8	2830 \pm 150
1996 Nov 11.9.....	VLA (A)	4.9	190 \pm 45	10.2 \pm 1.5	1660 \pm 100
	VLA (A)	8.4	310 \pm 35	10.2 \pm 1.5	1660 \pm 100
1997 Jun 29.3.....	VLA (C)	4.9	210 \pm 70	17.8 \pm 0.6	3680 \pm 200
	VLA (C)	8.4	310 \pm 45	17.8 \pm 0.6	3680 \pm 200
1997 Jul 24.1.....	VLA (CS)	8.4	41 \pm 30	0.80 \pm 0.73	210 \pm 100
1998 Jan 30.8.....	SCUBA	350	<3000	20.4 \pm 0.4	...
1998 Jan 31.1.....	ATCA (6A)	4.8	<390	16.6 \pm 0.9	...
	ATCA (6A)	8.6	<360	16.6 \pm 0.9	...
1998 Feb 08.1.....	ATCA (6A)	4.8	<480	6.88 \pm 0.31	...
	ATCA (6A)	8.6	<930	6.88 \pm 0.31	...
1998 Feb 19.6.....	VLA (D \rightarrow A)	8.4	13 \pm 30	2.72 \pm 0.35	360 \pm 120
Simultaneous Radio/X-ray Observations of Type II Bursts					
B1997 Jul 24.1.....	VLA (CS)	8.4	-300 \pm 230	...	3240 \pm 100
B1998 Feb 19.6 ^a	VLA (D \rightarrow A)	8.4	42 \pm 120	...	3550 \pm 100
B1998 Feb 19.6 ^b	VLA (D \rightarrow A)	8.4	74 \pm 82	...	2350 \pm 100

NOTE.—Upper limits are 3σ .

^a Using 3.33 s bins with greater than 1700 counts s⁻¹ average PCA count rate.

^b Using 3.33 s bins with greater than 300 counts s⁻¹ average PCA count rate.

⁸ See <http://space.mit.edu/derekfox/xte/ASM.html>.

⁹ The VLA is part of the National Radio Astronomy Observatory, which is operated by Associated Universities, Inc., under cooperative agreement with the National Science Foundation.

both the radio ($45 \pm 30 \mu\text{Jy}$) and X-ray (0.15 ± 0.31 ASM counts s^{-1}) bands (errors are 1σ ; 73 ASM counts $\text{s}^{-1} = 1$ crab).

On 1996 October 29, the RB was detected by the ASM to have begun an X-ray outburst, reaching peak intensity at 1996 October 30 14:30 UT (± 1 hr). VLA observations, with simultaneous *RXTE*/PCA observations, took place 7.2 days after the time of X-ray peak flux (1996 November 6.8 UT), at which time a radio source that was not present in the October 14 observation was detected with a flux density of $370 \pm 45 \mu\text{Jy}$ (8.4 GHz). The new source was consistent with a point source, but the low signal-to-noise ratio places upper limits of $0''.5$ on the size in both dimensions. The source is located at R.A. $17^{\text{h}}33^{\text{m}}24^{\text{s}}.61$, decl. $-33^{\circ}23'19''.8$ (J2000), $\pm 0''.1$ (see Fig. 1; the positions of radio source detections described below are consistent with this position). During the 3.6 ks PCA observation, one type I X-ray burst was observed along with strong persistent emission (PE; 220 ± 10 mcrab; throughout this work, 1 crab = $13,000$ PCA counts s^{-1}), but there were no type II X-ray bursts. For the RB, we find an approximate conversion between PCA count rate and X-ray flux to be 3×10^{-12} ergs $\text{cm}^{-2} \text{s}^{-1}$ per PCA count s^{-1} (2–20 keV).

A second simultaneous VLA/PCA observation took place 12.3 days after the X-ray maximum of the outburst (on 1996 November 11.88 UT) at 4.89 and 8.44 GHz; the radio point source had flux densities of 190 ± 45 and $310 \pm 35 \mu\text{Jy}$, respectively. During this observation, one type I X-ray burst was observed with persistent emission of 128 ± 8 mcrab and no type II X-ray bursts.

2.2. 1997 June–July Outburst

The RB began the following X-ray outburst on 1997 June 25, reaching a peak flux at 1997 June 26 10:30 UT (± 2 hr). Simultaneous VLA/PCA observations took place 2.9 days

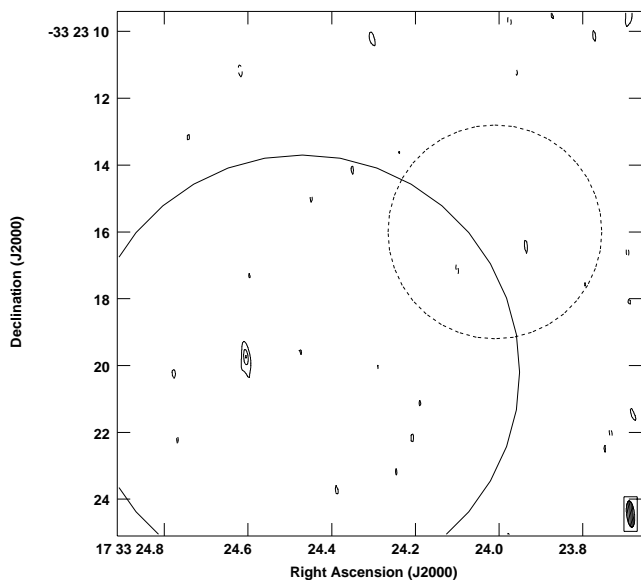


FIG. 1.—Discovery image. VLA map at 8.4 GHz of Liller 1 on 1996 November 6.8. The size of the clean beam is illustrated in the lower right corner. The radio contours are at 3σ , 6σ , and peak flux density. The solid circle is the $6''.5 \pm 0''.5$ core radius at the optical position for Liller 1 (Kleinmann et al. 1976, confirmed by Picard & Johnston 1995; which we adopt over the $3''.5 \pm 0''.5$ found by Grindlay et al. 1984), which has absolute optical astrometric uncertainty of $1''$. The dashed circle is the 2σ error circle of the X-ray localization of the Rapid Burster (Grindlay et al. 1984).

after X-ray maximum while the VLA was in C-array. The radio object was present at both 4.89 and 8.44 GHz, with flux densities of 210 ± 70 and $310 \pm 45 \mu\text{Jy}$, respectively. During this observation, the PE observed by the PCA was 280 ± 15 mcrab, and there were two type I X-ray bursts but no type II X-ray bursts.

On 1997 July 24.08 UT, a second simultaneous VLA/PCA observation took place, 27.6 days after X-ray maximum. The VLA was in CS-array, and the radio point source was not detected at $41 \pm 30 \mu\text{Jy}$ (8.44 GHz). The PE was very weak (< 8 mcrab); over the whole observation, the average intensity (bursts + PE) was 210 ± 100 PCA counts s^{-1} (16 mcrab). A total of seven type II X-ray bursts were observed with the PCA with an average intensity of 2350 ± 100 PCA counts s^{-1} (180 mcrab) during the bursts and a mean duration of 12 s. Of these, six had simultaneous coverage at VLA at 8.44 GHz. The radio data corresponding to these six bursts were extracted to determine whether there is detectable radio emission during the brightest bursts (3240 ± 100 PCA counts s^{-1}). We find an upper limit on the radio emission during these type II X-ray bursts of $690 \mu\text{Jy}$ (3σ).

2.3. 1998 January–February Outburst

The next RB outburst began on 1998 January 27, peaking in X-ray intensity at 1998 January 29 12:30 UT (± 4 hr). During PCA observations 1.2 days after X-ray maximum (1998 January 30 19:19–21:09 UT), strong PE (4000 ± 100 counts s^{-1}) was observed along with two type I X-ray bursts but no type II X-ray bursts. Observations at the JCMT were carried out 1.3 days after the X-ray maximum using the SCUBA system at 350 GHz and failed to detect the proposed radio counterpart with a 3σ upper limit of 3 mJy.

Beginning 1.6 days after the X-ray maximum, the ATCA observed for a 3 hr period (1998 January 31 1:37–4:46 UT) at 4.8 and 8.64 GHz. These observations produced flux density upper limits of 390 and 360 μJy , respectively (3σ). Between 1:33 and 4:28 UT on 1998 February 8, 9.6 days after the X-ray maximum, observations at the ATCA produced 3σ upper limits of 480 μJy (4.80 GHz) and 930 μJy (8.6 GHz).

On 1998 February 19, 20.1 days after the X-ray maximum, we performed a 7.8 ks PCA observation during which a total of 91 type II bursts were observed. Simultaneous VLA observations occurred during the second half of this period, and 39 type II bursts were observed simultaneously by the PCA and the VLA at 8.44 GHz. These bursts had average peak count rates of ~ 6000 counts s^{-1} (due only to type II burst emission, excluding background and PE of 250 PCA counts s^{-1}) and durations of ~ 10 s. Integrated over the full radio observation, the 3σ upper limit on radio emission from a point source at the position of the proposed radio counterpart was $90 \mu\text{Jy}$ (8.4 GHz). This observation is the only radio observation in the present work during which a large number of type II X-ray bursts were observed. Therefore, we can use it to explore the relationship between the type II X-ray bursts and the radio emission.

In order to compare the X-ray flux to the radio flux density, we rebin the PCA data to correspond to the 3.33 s integration periods of the VLA data. To optimize the signal-to-noise ratio for the detection of radio bursts under the assumption that the radio flux density is proportional to

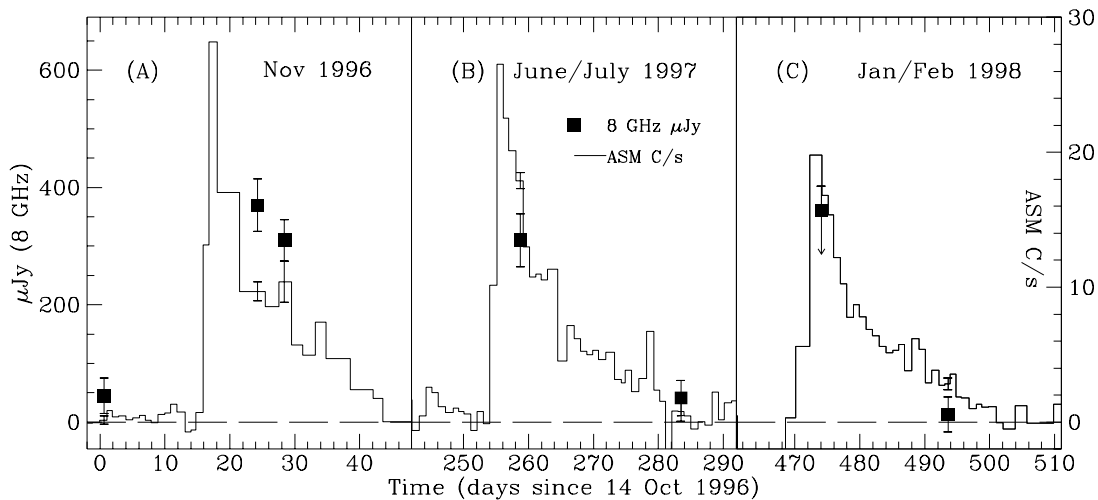


FIG. 2.—8.4 GHz flux density (VLA, and ATCA upper limit below $600 \mu\text{Jy}$) of the proposed radio counterpart vs. the *RXTE*/ASM (2–10 keV) daily average count rate of the Rapid Burster as a function of time, near the (a) 1996 November, (b) 1997 June/July, and (c) 1998 January/February outbursts. The reference date is 1996 October 14, MJD 50370. The scale of the radio flux density (left axis) and the X-ray count rate (right axis) are shown. The radio flux density turns on and off with the 2–10 keV X-ray flux of the Rapid Burster. The probability of a randomly varying radio source duplicating the observed X-ray on/off cycle is 1%–3% (§ 4.2). Radio upper limits are 3σ . The ATCA data point for 1998 February 8 (day 482) is off the figure with a 3σ upper limit of $930 \mu\text{Jy}$.

the X-ray flux, we use only time bins with average count rates greater than $1700 \text{ counts s}^{-1}$. During these time bins, the radio emission is constrained to be less than $360 \mu\text{Jy}$ (3σ) at 8.4 GHz, while the average X-ray count rate was $3550 \pm 100 \text{ counts s}^{-1}$.

One can decrease the noise level in the radio data by simply increasing the integration time. However, if the radio flux density is strictly proportional to the X-ray flux, the signal-to-noise ratio will decrease. Using the radio data observed when the type II burst X-ray flux was greater than $300 \text{ counts s}^{-1}$ yields a 3σ upper limit of $255 \mu\text{Jy}$ on the radio emission during bursts while the average X-ray type II burst intensity was $2340 \pm 120 \text{ counts s}^{-1}$.

2.4. Relationship between X-Ray and Radio Intensity

In Figure 2, we show the 8.4 GHz VLA flux densities measured during three consecutive outbursts of the RB compared with the *RXTE* ASM X-ray intensity measurements. The radio outbursts and quiescent periods correspond well with the X-ray state of the RB, indicating a possible relationship between the radio and X-ray sources.

For the ASM, the best-fit linear relation between the 8.4 GHz flux density measured at the VLA and the ASM X-ray flux (forcing the line through the origin) gives $S_{8 \text{ GHz}} = 27 \pm 1.7 \mu\text{Jy} (1 \text{ ASM counts s}^{-1})^{-1}$. However, this is a very poor fit, with a reduced $\chi^2_{\nu} = 5.6$ (for 6 degrees of freedom

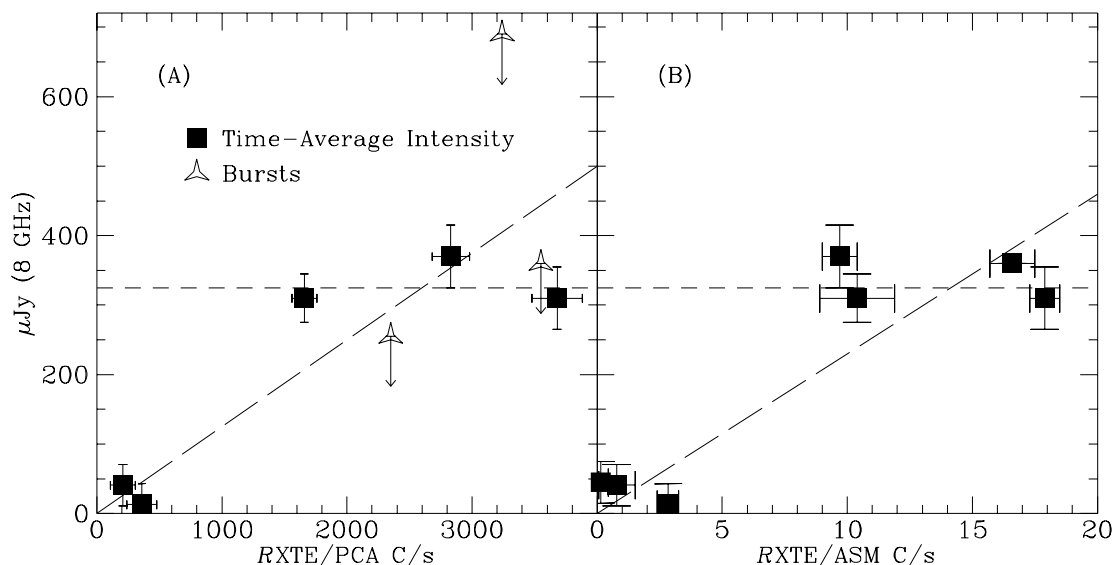


FIG. 3.—8 GHz flux density vs. (a) the simultaneous *RXTE*/PCA count rate and (b) the contemporaneous *RXTE*/ASM count rate. The time-averaged intensity (both X-ray bursts and persistent emission; squares) and the 3σ upper limits derived from data taken only during type II bursts (triangles) are shown. (Note that the two most constraining of these points are derived from alternative analyses of the same data; see § 2.3.) In general, a linear correlation between X-ray and radio intensity does not provide a good fit to the data; a constant radio flux density at high time-averaged X-ray intensity is more consistent. The radio flux densities include VLA (8.4 GHz) and ATCA (8.6 GHz) observations with sensitivity below $600 \mu\text{Jy}$.

[dof]). For the PCA data, the relation is $S_{8\text{ GHz}} = 125 \pm 8 \mu\text{Jy}$ (1000 PCA counts s^{-1}) $^{-1}$, again with a very poor fit ($\chi^2_\nu = 4.0$ for 4 dof). These data and the linear fits are shown in Figure 3.

It seems also plausible (observationally, if not theoretically) that the radio source is “on” at a constant flux density (of $325 \pm 25 \mu\text{Jy}$ at 8.4 GHz) whenever the X-ray source is above a threshold of ~ 3 ASM counts s^{-1} (~ 500 PCA counts s^{-1}). In Figure 3a, we also show the 3σ upper limits for the excess radio emission correlated with the average type II burst intensity. The two most constraining upper limits (which were drawn from the same data; see § 2.3) are still marginally consistent with the radio detections during periods of X-ray persistent emission of comparable X-ray intensity. The most constraining simultaneous radio emission limit during the type II bursts is discrepant at the 2.9σ level.

3. ANALYSIS OF RADIO OBSERVATIONS USING THE SYNCHROTRON BUBBLE MODEL

One possible description of our radio flux density measurements involves a synchrotron bubble model (Van der Laan 1966; Hjellming & Han 1995). In this model, the physical source of the radiation is a dense, expanding bubble of plasma. The relativistic electrons in the plasma (assumed to have a more or less isotropic velocity distribution and a power-law energy distribution) emit synchrotron radiation as their trajectories are deflected by interactions with ambient or entangled magnetic fields. The resulting radio spectrum is strongly peaked with a power-law form at frequencies (much) below or above the peak frequency. This peak frequency decreases as the bubble expands in a well-determined fashion: $\nu_m \sim \rho^{-(4\gamma+6)/(\gamma+4)}$, where ν_m is the peak frequency, ρ is the normalized radius of the bubble, and γ is the power-law index of the electron energy distribution, $N(E)dE \sim E^{-\gamma} dE$ (Van der Laan 1966).

A synchrotron bubble’s time-dependent radio emission is completely determined if we set the time, $t_0(\nu_0)$, and the flux density at spectral peak, $S_0(\nu_0)$, when the spectral peak reaches some nominal frequency ν_0 (which we will take to be the VLA X band, 8.44 GHz). Furthermore, one must define the functional form of the bubble’s radial expansion in time. Here we adopt the convention of previous authors (Van der Laan 1966; Hjellming & Han 1995) and assume a power-law form for this expansion: $\rho \sim t^{1/\alpha}$.

The initial hypothesis that we wish to investigate is whether our radio flux measurements may be explained by the evolution of a single synchrotron bubble initiated at the start of each X-ray outburst. We perform this analysis for two primary reasons: first, models of this nature, with time-

scales of the order of those observed in the radio emission, have proven successful in modeling the radio outbursts of other X-ray binaries (Han 1993; Hjellming & Han 1995); second, even if the model is not successful, we expect it to produce insight into the physical processes responsible for the radio emission we observe.

We determine the X-ray outburst start times by fitting the *RXTE* ASM light curve to a functional form consisting of an onset time, a linear rise in intensity, a peak intensity time, and an exponential decay, similar to the form used by Guerriero et al. (1999), except that we do not include the secondary flares that are found by those authors in two of the outbursts. The results of our fits are consistent with those of Guerriero et al., and parameter uncertainties are probably dominated by the systematics associated with this simplified outburst intensity model. Based on these fits, we determine the time after outburst onset for each of the radio observations.

We have performed fits to our radio observations using the synchrotron bubble model of Van der Laan (1966, eqs. [11] and [12]); note, however, that his equation (11) is incorrectly typeset and should read

$$S(\nu, \rho) = S_{m0} \left(\frac{\nu}{\nu_{m0}} \right)^{5/2} \rho^3 \times \frac{\{1 - \exp[-\tau_m(\nu/\nu_{m0})^{-(\gamma+4)/2} \rho^{-(2\gamma+3)}]\}}{[1 - \exp(-\tau_m)]}, \quad (1)$$

where the variables are as defined in that work. For purposes of this analysis, we assume that the observed RB outbursts generate synchrotron bubbles with identical physical properties and time evolution. The data are not sufficient to adequately constrain all four independent parameters, so we have fixed the expansion index α at three canonical values: $\alpha = 1$, corresponding to free expansion with constant velocity, $\alpha = 2.5$, corresponding to energy-conserving expansion into an ambient medium (as in the Sedov phase of a supernova remnant), and $\alpha = 4$, resulting from a momentum-conserving (but not adiabatic) expansion into an ambient medium (Hjellming & Han 1995). Application of the geometric corrections to the van der Laan model suggested by Hjellming & Johnston (1988) altered the fit parameters slightly but did not significantly improve the fits or change the character of the solutions (the light curves changed by less than 10%). Therefore, we quote results from the simpler model. The resulting synchrotron bubble parameters are given in Table 3, and the models are compared to the observations in Figure 4.

As indicated in Table 3, the models provide reasonable, but statistically unacceptable, fits to the data ($\chi^2_\nu = 1.7\text{--}4.4$,

TABLE 3
SYNCHROTRON BUBBLE MODEL FITS TO THE RADIO OBSERVATIONS

EXPANSION INDEX α (fixed)	χ^2_ν Fit	NORMALIZATION (at 8.44 GHz)		
		$S_{08\text{ GHz}}$ (μJy)	$t_{08\text{ GHz}}$ (days)	ELECTRON ENERGY INDEX (γ)
1	4.4	$37.^{+102}_{-3}$	$27.9^{+1.1}_{-13.7}$	$<1.07^a$
2.5	2.3	$342.^{+35}_{-36}$	$10.0^{+1.5}_{-0.7}$	$2.6^{+4.4}_{-0.4}$
4	1.7	$438.^{+61}_{-63}$	$8.5^{+4.9}_{-3.5}$	$9.0^{+5.0}_{-2.4}$

NOTE.—The models presented here had 9 dof.

^a Fixing $\alpha = 1$ drove the fits to $\gamma = 1$, the minimum physical value. $\gamma < 1.07$ in this model with 90% confidence.

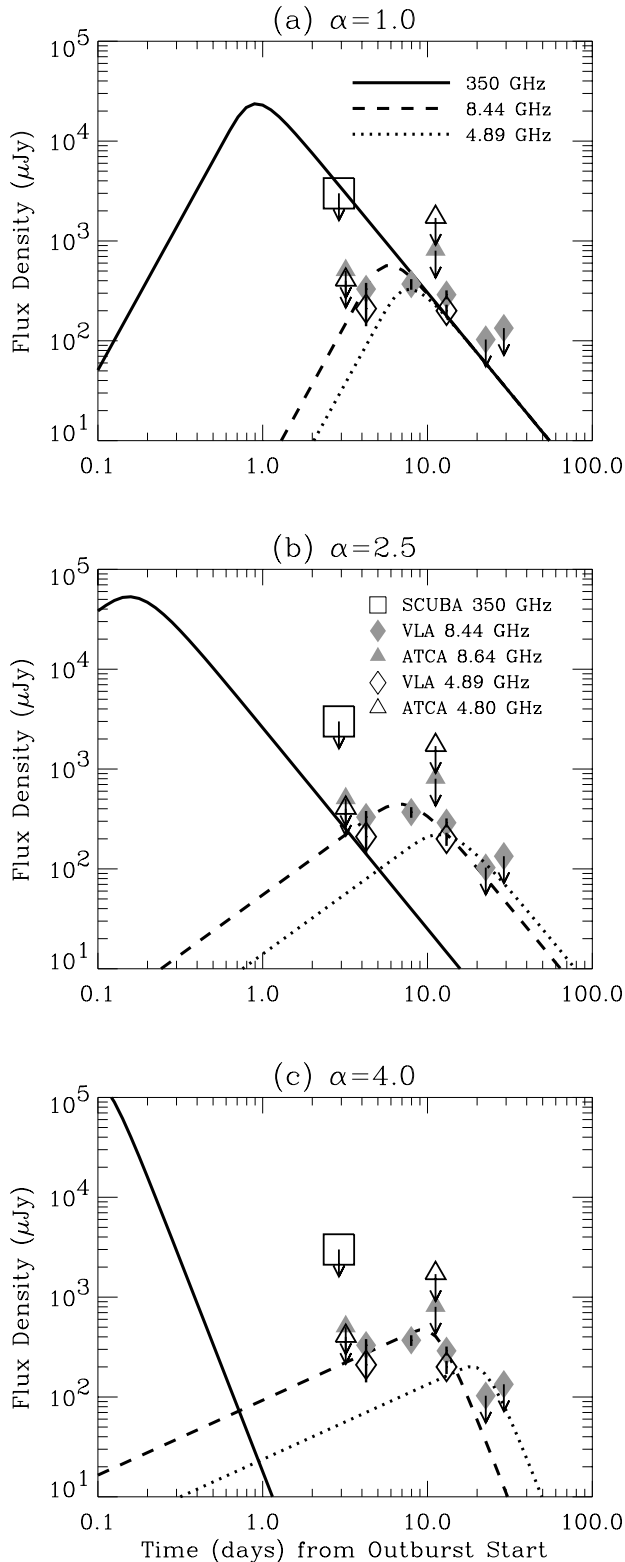


FIG. 4.—Data and upper limits (3σ) from our radio observations, with best-fit synchrotron bubble model superposed for (a) $\alpha = 1$ ($\chi^2_\nu = 4.4$), (b) $\alpha = 2.5$ ($\chi^2_\nu = 2.2$), and (c) $\alpha = 4$ ($\chi^2_\nu = 1.9$). Time is measured from the start of each Rapid Burster outburst, as determined by our model fits to the *RXTE*/ASM data (see § 3). Lines indicate the predicted light curve of the synchrotron bubble event at three frequencies: 350 GHz (solid line), 8.44 GHz (dashed line), and 4.89 GHz (dotted line). The slightly different operational frequencies of the ATCA are effectively indistinguishable on the plot from the VLA frequencies shown. Parameters of the model fits are given in Table 3. Although this one bubble model for the outbursts provides a reasonable fit to the observations, the derived parameters of the bubble's expansion are unphysical (see § 5).

for 9 dof). Given our many simplifying assumptions, particularly the assumption that all outbursts are identical, this is perhaps not surprising. We note that the free expansion ($\alpha = 1$) model fit is significantly worse than the other two and violates the 3σ upper limit imposed by the 1998 January 30 SCUBA observation (Fig. 4a); for these reasons we prefer the models that assume a deceleration of the bubble's expansion by an ambient medium.

Even these models, however, are in the end physically unacceptable, as may be determined by looking at the underlying physical parameters. We can express S_0 in terms of the ambient magnetic field density, H_0 , and the angular extent of the source, θ_0 , at time $t_{0.8\text{ GHz}}$, as follows:

$$S_0 = 0.85h(\gamma)\left(\frac{H_0}{1\text{ mG}}\right)^{-1/2}\left(\frac{\theta_0}{1\text{ mas}}\right)^2\text{ Jy}, \quad (2)$$

where γ is the power-law index of the electron energy distribution ($2.6^{+4.4}_{-0.4}$ for the adiabatic model and $9.09.0^{+5.0}_{-2.4}$ for the momentum-conserving model) and $h(\gamma)$ is a known function that varies from 4.1 to 1.04 as γ increases from 1 to 10:

$$h(\gamma) = \frac{\pi\Gamma[(3\gamma + 19)/12]\Gamma[(3\gamma - 1)/12]\Gamma(3/4)}{\sqrt{6}\Gamma[(3\gamma + 2)/12]\Gamma[(3\gamma + 22)/12]\Gamma(5/4)}, \quad (3)$$

where Γ is the Euler gamma function.

Under the adiabatic assumption, the 8.44 GHz flux density of the source when the spectral peak reaches that frequency is $S_{0.8\text{ GHz}} = 340 \pm 35\ \mu\text{Jy}$, whereas under the momentum-conserving assumption it is $440 \pm 50\ \mu\text{Jy}$. Thus, the indications are that $\theta_0 \sim 10^{-3}$ mas at $t_{0.8\text{ GHz}}$, under both models, for fields of the order of 1 mG (with a weak $H_0^{1/4}$ dependence on the actual field strength). At the distance of the RB (8.6 kpc), that corresponds to a linear size of $d_0 \sim 2\text{ lt-s}$ and expansion velocities at time $t_{0.8\text{ GHz}}$ of the order 1 km s^{-1} , much too slow for the expanding hot plasma that the model requires. We are forced to conclude that more complicated models are required to accurately describe the observations (see § 5).

4. EVALUATION OF COUNTERPART LIKELIHOOD

4.1. Positional Coincidence

First, we consider the probability that an unrelated radio source lies close to the RB on the sky. Microjansky source counts at 8.44 GHz by Windhorst et al. (1993) find a source density of $1.14 \times 10^{-2}\text{ arcmin}^{-2}$ for flux densities above $300\ \mu\text{Jy}$. Therefore, the probability of an unrelated source brighter than $300\ \mu\text{Jy}$ falling within $8''.6$ of the X-ray position of the RB is 7×10^{-4} . It should be noted that the μJy sources are typically associated with faint blue galaxies (Fomalont et al. 1991) and that the survey fields are not representative of the rich environment of a globular cluster near the Galactic center. Because the microjansky source population in an environment comparable to Liller 1 is unknown, we do not use the probability computed above in our evaluation of the counterpart likelihood, which is based exclusively on the correlation between the radio and X-ray light curves.

The best X-ray position for the RB (Grindlay et al. 1984) places the radio source $8''.6$ away from the RB with 1σ X-ray and radio uncertainties of $1''.6$ and $0''.1$, respectively. Taken at face value, this is significant at the 5.4σ level, not including the systematic uncertainty in the radio to optical (i.e., *Einstein* star tracker) reference frame shift. This makes

the radio source identification with the RB seem less likely. However, further investigation reveals that the quoted Gaussian errors are an inadequate description of the true error distribution.

The *Einstein* position for the Rapid Burster was determined from a single pointing (four *Einstein* HRI pointings at Liller 1 were made, but only one found the RB in outburst). The $1''.6$ 1σ uncertainty in this position is estimated from the dispersion of the errors in the single-pointing *Einstein* positions of X-ray sources with known optical counterparts (Grindlay et al. 1984). An important cross-check of the requisite star-tracker calibration (which dealt with a number of complex systematic effects; Grindlay 1981) was provided by the *Einstein* globular cluster X-ray source program, which performed multiple pointings at each of eight globular clusters (GCs) with known bright X-ray sources in order to determine the positions of these sources to better than $1''.6$ accuracy (Grindlay 1981; Grindlay et al. 1984). These eight clusters were NGC 104 (47 Tuc), NGC 1851, Terzan 2, Liller 1, NGC 6441, NGC 6624, NGC 6712, and NGC 7078 (M15).

If the star-tracker calibration was successful in accounting for all significant sources of systematic error, then we would expect the resulting sample standard deviations, σ in the positions of the GC X-ray sources as derived from the multiple pointings at each, to cluster strongly around $\sigma = 1''.6$. Examining the quoted uncertainties of Grindlay et al. (1984), however, we find that the actual σ values deviate from $1''.6$ by up to a factor of 2: $\sigma = 0''.8$ for the X-ray source in NGC 6712 (5 pointings), while $\sigma = 3''.2$ for the X-ray source in Terzan 2 (4 pointings). A χ^2 test shows that in NGC 6712, σ is too small at the 98% confidence level and that for Terzan 2, σ is too large with 99.95% confidence. Together, these deviations are unlikely to be statistical, and they represent two of the seven clusters examined (recall that Liller 1 had only one pointing with a detection, so its σ value is undetermined).

We conclude that there are probably remaining systematic effects in the *Einstein* aspect solutions that have not been accounted for. These effects average into the quoted $1''.6$ error over many pointings but caused substantial non-Gaussian excursions within the context of the GC X-ray source program. Since the multiple pointings at each cluster were typically executed over a time span of weeks to months, the unaccounted-for systematics are not likely to be temporal in nature but rather to relate to position on the sky. They may relate, for example, to the density of suitable stars within the star-tracker fields. In this connection it is worth noting that Terzan 2 and Liller 1 are the two sample GCs nearest the Galactic Center and are only 3° apart on the sky.

4.2. Significance of the X-Ray–Radio Correlation

The probability that the observed radio/X-ray behavior is produced by an unrelated background radio source depends on the model for the variability behavior of that source. We consider first a source model in which the background radio source varies randomly with a duty cycle, p , and has a short autocorrelation timescale allowing us to consider each of our measurements to be statistically independent.

We take the radio source to be “on” or “off” when our 1σ radio sensitivity is greater or less than $70\ \mu\text{Jy}$ and the source is or is not detected, respectively, at the 3σ level. Moreover, we define the Rapid Burster to be on or off (for

these purposes) when its 1 day average *RXTE* ASM count rate is greater or less than $4\ \text{counts s}^{-1}$ (see Fig. 3 and § 2.4). Based on these assumptions, the probability of an unrelated variable background radio source mimicking the X-ray on/off state of the RB is the product of the normalized probability distribution for obtaining the number of radio on and off observations for a given p and the probability distribution of observing the radio and X-ray sources to be on and off simultaneously for a given p , integrated over all values of p :

$$P = \frac{\int_0^1 p^N (1-p)^M p^{N_s} (1-p)^{M_s} dp}{\int_0^1 p^N (1-p)^M dp}, \quad (4)$$

where N is the total number of times the radio source is observed on, M is the total number of times the radio source is observed off, N_s is the number of times the radio source and X-ray source are observed on simultaneously, and M_s is the number of times the radio source and X-ray source are observed off simultaneously. We have assumed a uniform prior distribution for p itself. For our observations (see Table 2), $N = 3$, $M = 3$, $N_s = 3$, and $M_s = 3$, which yields a probability of 1.2% for an unrelated radio source mimicking the observed on-off behavior of the RB. This indicates that it is unlikely that the observed radio/X-ray correlation is produced by a randomly varying background radio source.

More complex models for the radio source might postulate that it occasionally turns on and remains so for T_{on} days (autocorrelation timescale). We can estimate the probability of a chance correlation of such a source with the Rapid Burster, without performing involved Monte Carlo simulations, by making use of the framework developed above. If we speculate that T_{on} is such that the two detections during 1996 November (separated by 5 days) are perfectly correlated and all other observations are uncorrelated, then we have 5 coincident X-ray/radio observations ($N = N_s = 2$, $M = M_s = 3$), and the probability of a chance correlation is 2.6%. Note that if the autocorrelation timescale is shorter than 5 days then all six observations are statistically independent as above. Note also that the autocorrelation timescale cannot be much longer than 5 days because as it approaches ~ 20 days it approaches the observed radio on/off time (that is, the initial nondetection followed by a detection during 1996 November, and the detection followed by a nondetection in 1997 June/July).

One might reasonably be concerned about the possibility of an unrelated radio source that turns on suddenly and remains so for an extended period of time. The radio data do not exclude the possibility that such a source was present and remained on during the unsampled period between the 1996 November and 1997 June/July outbursts of the RB. In this case, we have only three independent measurements of coincidence with the RB. However, the radio source is observed to make the transition from off to on (and the reverse) in $\lesssim 25$ days. The radio data show one turn-on and one turn-off, both of which are well sampled and coincident with the X-ray behavior of the RB. If an unrelated radio source turns on every ~ 300 days, then the probability of the turn-on coincidence observed in 1996 November is $\sim 8\%$. The slow decay of the RB X-ray outburst makes a turn-off coincidence somewhat more probable. If we cannot distinguish turn-offs that are separated by ± 10 days, then the probability of the observed coincidence in the 1997 June/July outburst is $\sim 15\%$. If the two events

are independent, the probability of this type of source mimicking the observed RB X-ray turn-on/off behavior is $\sim 1\%$.

The population of microjansky radio transients and their outburst properties are not well studied. The probability of an unrelated, variable radio source having the time-dependent characteristics above and lying this close to the RB is somewhat less than unity, although we have neglected this in the above calculations. It seems reasonable to assume the most conservative of the above approaches as the upper limit; we therefore find that the upper limit on the probability of a flaring background radio source mimicking the X-ray behavior of the RB is 3%.

4.3. Possible Relationship between the FG Radio Object and the Present Object

The relationship between the variable radio source we observe and the steep spectrum source ($\alpha = -2$) observed by FG is not clear. To within the astrometric uncertainty ($1''.5$, 3σ), the two sources are at the same position on the sky. If FG's interpretation of their source is correct and it represents the integrated emission of a population of radio pulsars, then it seems likely that their radio position is the center of the GC. (The $\sim 2''$ separation between the FG radio source and the optical center of Liller 1 is consistent with the $1''$ uncertainty in the absolute optical astrometry; see Fig. 1). For a $6''.5$ core-radius GC (Kleinmann, Kleinmann, & Wright 1976; Picard & Johnston 1995), the $1''.5$ radio error circle represents 12% of the optical light of the GC. Thus, it is not unlikely that (if the RB is associated with the GC) the radio counterpart to the Rapid Burster would also lie close to the center of the GC.

The three flux-density measurements of the FG source were made over three different (sometimes overlapping) epochs. Combining the three measurements into a single spectrum (which was found to be steep) assumes that the source is not variable. The conflicting 1.5 GHz measurements of FG and Johnston et al. may indicate variability of a factor of 2 or more (2σ) between 1990 April and 1993 May. The X-ray state of the Rapid Burster during these observations is unknown, and it may be that it was X-ray active during FG's 1.5 GHz observation, providing the additional radio flux above that expected for an underlying population of radio pulsars in Liller 1. The relationship between the FG radio source and the RB could be illuminated by radio measurements of the same epoch at 0.33, 1.5, and 4.5 GHz, taken both while the RB is in X-ray outburst and in quiescence.

4.4. Conclusion: a Likely Radio Counterpart

The probability that a serendipitously located variable radio source would mimic the Rapid Burster X-ray state as has been observed is small (1%–3%), but not dismissably so. The number and distribution of faint, variable radio sources toward the Galactic Center is not well known. There has been at least one other recent instance where a radio source, variable on a timescale of several days, was discovered $\sim 3'$ from a bright X-ray source (although the X-ray flux was not correlated; Frail et al. 1996a, 1996b). In addition, globular clusters are known to harbor both millisecond radio pulsars and accreting X-ray sources, so the appearance of an unrelated variable radio source in Liller 1 must be considered more likely than for a random field. Other than the correlated radio and X-ray states and the marginal consistency of

the radio spectrum with synchrotron bubble models, we have not observed any radio behavior that would tie this object uniquely to the Rapid Burster.

The apparent positional discrepancy between the radio source and the RB is likely explained by the non-Gaussian distribution of the X-ray position error as discussed above and acknowledged by Grindlay et al. (1984) and J. E. Grindlay (1998, private communication). Proper motion of the RB will probably contribute only negligibly to the discrepancy, given the association with Liller 1, even though an interval of 20 yr separates the *Einstein* and radio observations (a proper motion of $\sim 0''.5$ yr $^{-1}$ is required). Improvement upon the X-ray localization will be obtained during an approved *AXAF* observation, which should determine the X-ray position to $\lesssim 0''.5$, thus confirming or excluding this radio counterpart.

At present, we identify this radio source as a likely radio counterpart of the Rapid Burster. Apart from the *AXAF* observation, the case could be strengthened by observing bursts from this location (either in radio or IR), by continued observations in X-ray and radio of correlated on/off behavior, by observation of larger swings in the X-ray and radio fluxes, which would permit a definite X-ray/radio correlation to emerge, or by discovering other radio behavior that is correlated with X-ray behavior of the Rapid Burster (such as short timescale variability). For example, a program of 12 short VLA observations reaching a noise level of $45 \mu\text{Jy beam}^{-1}$ taken at 2 week intervals and correlated with contemporaneous *RXTE* ASM observations would reduce the probability of an unrelated source mimicking the RB to 0.03%, even if the RB remains quiescent in X-rays and there are no radio detections.

5. DISCUSSION OF THE RADIO OBSERVATIONS

The radio observations of the proposed radio counterpart on 1996 November 11.88 determine that the radio spectral slope is flat to inverted with $\alpha = 0.9 \pm 0.3$ ($S_\nu \propto \nu^\alpha$).

The JCMT/SCUBA observation at 350 GHz is, to date, the earliest radio observation relative to the beginning of an X-ray outburst. A few hours prior to this observation, a PCA observation measured the RB PE to be 4000 ± 100 counts s $^{-1}$. Using the measured radio spectral slope ($\alpha = 0.9 \pm 0.3$) and the PCA/radio conversion at 8.4 GHz, the extrapolated radio spectrum gives a 350 GHz flux density of 4.5–41 mJy (with uncertainty dominated by the spectral slope). This is well above the 3 mJy 3σ upper limit obtained at the JCMT, indicating that the radio emission is not a simple power-law spectrum, proportional in intensity to the instantaneous X-ray intensity. This nondetection is also the crucial observation in ruling out one class of synchrotron bubble models (see § 3).

Synchrotron bubble behavior, associated with the outburst of an X-ray transient, has been observed on many occasions; Hjellming & Han (1995) show radio data and fits for A0620–00, Cen X-4, GS 2000+25, Aql X-1, GS 2023+338 (V404 Cyg), and GRS 1124–683 (see also Han 1993). In that respect the detection of radio emission from the Rapid Burster during its outbursts is not particularly surprising.

The physical parameters derived from our model fits, however, indicate expansion velocities of ~ 1 km s $^{-1}$ about 10 days after the outburst start, which is well below the physical lower limit set by the sound speed of the hot plasma, $c_s \sim 0.1(T_k)^{1/2}$ km s $^{-1}$ (where T_k is the temperature

of the plasma in kelvins, $\sim 10^7$ in SS 433; Hjellming & Johnston 1988).

Our assumption of a generic bubble event accompanying each outburst is, therefore, likely to be flawed, and we must consider more complex models. For example, radio emission over the course of an outburst may result from the summed emission of a succession of synchrotron bubbles, each of which expands at high speed and therefore brightens and fades (at a given frequency) much more quickly than the emission at large. If the flux levels that we see are produced by $T_k \sim 10^7$ plasma, then the expected rise times for the synchrotron bubbles are ~ 1 hr. Alternatively, each of our radio detections may simply have caught a single fast bubble in the midst of its expansion; in that case, the timing of our detections would indicate how the rate of these bubble ejections changes over the course of an outburst. Relevant to both of these possibilities, it is worth noting that the radio flares from GRS 1915+105 have rise and decay times of ~ 1 hr (Mirabel et al. 1998; Fender & Pooley 1998).

We are therefore motivated to consider whether the type II X-ray bursts of the Rapid Burster might produce individual synchrotron bubbles, with associated infrared and radio emission. The synchrotron bubbles of GRS 1915+105 have been shown to be related to active accretion, as determined by simultaneous X-ray observations (Mirabel et al. 1998; Eikenberry et al. 1998a). Scaling down the brightness of the infrared emission seen in that source by the factor of 10 difference in X-ray flux between it and the RB suggests that there may be microjansky infrared flares ($K \sim 14$) during RB type II X-ray bursts, if synchrotron bubbles are indeed being formed. We are currently pursuing short time resolution IR observations of the RB during its next outburst to test this hypothesis; naturally, observation of any bursting counterpart to the RB will confirm or reject the VLA counterpart proposed here.

The radio detections we report here could not themselves have been produced by the type II X-ray bursts; the radio detections were made prior to day 12 of each outburst, while the type II X-ray bursts were not observed until after day 14 in every case.

The observations of 1997 July 24 were carried out when the VLA was in its new CS (C-short) configuration. The combination of the array configuration with low-elevation observing yielded many short projected baselines with lengths all the way down to the antenna separation (25 m). Heavily tapered maps of these data reveal an extended source with peak flux density of approximately 3 mJy (at 8.44 GHz) that is at least as large as the primary beam. The axis of symmetry of the extended source lies at its closest point approximately $1'$ northeast from the RB, which corresponds to 2.5 pc at the 8.6 kpc distance of the RB. The source is probably unrelated to the RB, but its presence is an important consideration for future observations that include short interferometric baselines.

6. CONCLUSIONS

We have detected a likely radio counterpart for the Rapid Burster, with radio emission correlated with the X-ray out-

bursts. The likelihood of an unrelated variable radio source duplicating the observed correlation between the X-ray flux and radio flux density is low (1%–3%), but not dismissably so. There is an apparent discrepancy between the X-ray position of the RB and the radio counterpart, but it is likely due to the non-Gaussian distribution of the X-ray position errors. Confirmation of the counterpart from additional observations, an already approved *AXAF* observation, further radio observations while the *RXTE* ASM is still operational, or possibly infrared observations while the RB is in outburst, is required. The time and spectral evolution of the radio source, while not physically interpretable as a full-outburst synchrotron bubble (as seen in some other transient X-ray binaries), may be due to approximately hour-long radio flares such as have been seen from the superluminal-jet source GRS 1915+105.

Our lower limit on the time delay between X-ray emission and radio emission from the type II bursts (≥ 1 s) is consistent with the delay expected from a synchrotron bubble. Our observation of a persistent radio source ~ 5 days after the start of active accretion (i.e., an outburst) sets an upper limit for the radio versus X-ray time delay. The correlation of radio flux density with persistent X-ray intensity in this system indicates that the radio flux density is related to active accretion onto the surface of the neutron star, as accretion is responsible for the X-ray outburst. This suggests that excess radio emission may be produced during the type II X-ray bursts, which are themselves accretion driven. However, our observations produce no evidence of simultaneous radio/X-ray bursts, marginally constraining them (2.9σ) to be below the level that we observe during periods of comparable X-ray flux in persistent emission. This may imply that radio bursts at 8.4 GHz do not occur simultaneously with X-ray type II bursts.

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REFERENCES

- Barr, P., White, N. E., Haberl, F., Stella, L., Pollard, G., Gottwald, M., & Parmar, A. N. 1987, *A&A*, 176, 69
- Bradt, H. V., Rothschild, R. E., & Swank, J. H. 1993, *A&AS*, 97, 355
- Calla, O. P. N., Barathy, S., Snagal, A. K., Bhandar, S. M., Deshpande, M. R., & Vyas, H. O. 1980a, *IAU Circ.*, 3458, 1
- Calla, O. P. N., Bhandari, S. M., Deshpande, M. R., & Vats Hari, O. M. 1979, *IAU Circ.*, 3347, 1
- Calla, O. P. N., Sangal, A. K., & Barathy, S. 1980b, *IAU Circ.*, 3467, 1
- Chen, W., Shrader, C. R., & Livio, M. 1997, *ApJ*, 491, 312
- Cui, W. 1997, *ApJ*, 482, L163
- Cunningham, C. R., Gear, W. K., Duncan, W. D., Hastings, P. R., & Holland, W. S. 1994, *Proc. SPIE*, 2198, 638
- Eikenberry, S. S., Matthews, K., Morgan, E. H., Remillard, R. A., & Nelson, R. W. 1998a, *ApJ*, 494, L61
- Eikenberry, S. S., Matthews, K., Murphy, T. W. J., Nelson, R. W., Morgan, E. H., Remillard, R. A., & Muno, M. 1998b, *ApJ*, 506, L31
- Fender, R. P., & Pooley, G. G. 1998, *MNRAS*, 300, 573
- Fender, R. P., Pooley, G. G., Brocksopp, C., & Newell, S. J. 1997, *MNRAS*, 290, L65
- Finger, M. H., Koh, D. T., Nelson, R. W., Prince, T. A., Vaughan, B. A., & Wilson, R. B. 1996, *Nature*, 381, 291
- Fomalont, E. B., Windhorst, R. A., Kristian, J. A., & Kellerman, K. I. 1991, *AJ*, 102, 1258
- Frail, D. A., Kouveliotou, C., van Paradijs, J., & Rutledge, R. 1996a, *IAU Circ.*, 6323, 2
- . 1996b, *IAU Circ.*, 6307, 1
- Frogel, J. A., Kuchinski, L. E., & Tiede, G. P. 1995, *AJ*, 109, 1154
- Fruchter, A. S., & Goss, W. M. 1995, *A&A*, 16, 245 (FG)
- Grindlay, J. 1981, in *X-Ray Astronomy with the Einstein Satellite*, ed. R. Giacconi (Dordrecht: Reidel), 79
- Grindlay, J. E., Hertz, P., Steiner, J. E., Murray, S. S., & Lightman, A. P. 1984, *ApJ*, 282, L13
- Grindlay, J. E., & Seaquist, E. R. 1986, *ApJ*, 310, 172
- Guerriero, R., et al. 1999, *MNRAS*, 307, 179
- Han, X. H. 1993, Ph.D. thesis, New Mexico Inst. Mining and Tech.
- Hasinger, G., & van der Klis, M. 1989, *A&A*, 225, 79
- Hjellming, R. M., & Han, X. 1995, in *X-Ray Binaries*, Vol. 1, ed. W. H. G. Lewin, J. van Paradijs, & E. van den Heuvel (Cambridge: Cambridge Univ. Press), 308
- Hjellming, R. M., & Johnston, K. J. 1988, *ApJ*, 328, 600
- Hoffman, J. A., Marshall, H. L., & Lewin, W. H. G. 1978, *Nature*, 271, 630
- Jahoda, K., Swank, J. H., Giles, A. B., Stark, M. J., Strohmayer, T., Zhang, W., & Morgan, E. H. 1996, *Proc. SPIE*, 2808, 59
- Johnson, H. M., Catura, R. C., Lamb, P. A., White, N. E., Sanford, P. W., Hoffman, J. A., Lewin, W. H. G., & Jernigan, J. G. 1978, *ApJ*, 222, 664
- Johnston, H. M., Kulkarni, S. R., & Goss, W. M. 1991, *ApJ*, 382, L89
- Kleinmann, D. E., Kleinmann, S. G., & Wright, E. L. 1976, *ApJ*, 210, L83
- Kouveliotou, C., van Paradijs, J., Fishman, G. J., Briggs, M. S., Kommers, J., Harmon, B. A., Meegan, C. A., & Lewin, W. H. G. 1996, *Nature*, 379, 799
- Lawrence, A., et al. 1983, *ApJ*, 267, 301
- Levine, A. M., Bradt, H., Cui, W., Jernigan, J. G., Morgan, E. H., Remillard, R., Shirey, R. E., & Smith, D. A. 1996, *ApJ*, 469, L33
- Lewin, W. H. G., et al. 1976, *ApJ*, 207, L95
- Lewin, W. H. G., Rutledge, R. E., & Kommers, J. M., van Paradijs, J., & Kouveliotou, C. 1996, *ApJ*, 462, L39
- Lewin, W. H. G., van Paradijs, J., & Taam, R. 1993, *Space Sci. Rev.*, 62, 223
- . 1995, in *X-Ray Binaries*, Vol. 1, ed. W. H. G. Lewin, J. van Paradijs, & E. van den Heuvel (Cambridge: Cambridge Univ. Press), 175
- Liller, W. 1977, *ApJ*, 213, L21
- Mirabel, I. F., Dhawan, V., Chaty, S., Rodriguez, L. F., Marti, J., Robinson, C. R., Swank, J., & Geballe, T. 1998, *A&A*, 330, L9
- Picard, A., & Johnston, H. 1995, *A&AS*, 112, 89
- Rao, A. P., & Venugopal, V. R. 1980, *Bull. Astron. Soc. India*, 8, 41
- Remillard, R., & Levine, A. 1997, in *All-Sky X-Ray Observations in the Next Decade*, ed. M. Matsuoka & N. Kawai (Waco: Riken), 29
- Rutledge, R., Lubin, L., Lewin, W., Vaughan, B., van Paradijs, J., & van der Klis, M. 1995, *MNRAS*, 276, 753
- Tanaka, Y., & Lewin, W. 1995, in *X-Ray Binaries*, Vol. 1, ed. W. H. G. Lewin, J. van Paradijs, & E. van den Heuvel (Cambridge: Cambridge Univ. Press), 126
- van der Klis, M. 1995, in *X-Ray Binaries*, Vol. 1, ed. W. H. G. Lewin, J. van Paradijs, & E. van den Heuvel (Cambridge: Cambridge Univ. Press), 252
- van der Laan, H. 1966, *Nature*, 211, 1131
- van Paradijs, J. 1996, *ApJ*, 464, L139
- Windhorst, R. A., Fomalont, E. B., Partridge, R. B., & Lowenthal, J. D. 1993, *ApJ*, 405, 498
- Zhang, W., Giles, A. B., Jahoda, K., Soong, Y., Swank, J. H., & Morgan, E. H. 1993, *Proc. SPIE*, 2006, 324