



**UvA-DARE (Digital Academic Repository)**

**The ISO/SWS spectrum of IRC+10216: the vibrational bands of C<sub>2</sub>H<sub>2</sub> and HCN**

Cernicharo, J.; Yamamura, I.; Gonzalez-Alfonso, E.; de Jong, T.; Heras, A.M.; Escribano, R.; Ortigoso, J.

*Published in:*  
Astrophysical Journal

*DOI:*  
[10.1086/312360](https://doi.org/10.1086/312360)

[Link to publication](#)

*Citation for published version (APA):*

Cernicharo, J., Yamamura, I., Gonzalez-Alfonso, E., de Jong, T., Heras, A. M., Escribano, R., & Ortigoso, J. (1999). The ISO/SWS spectrum of IRC+10216: the vibrational bands of C<sub>2</sub>H<sub>2</sub> and HCN. *Astrophysical Journal*, 526, L41-L44. DOI: 10.1086/312360

**General rights**

It is not permitted to download or to forward/distribute the text or part of it without the consent of the author(s) and/or copyright holder(s), other than for strictly personal, individual use, unless the work is under an open content license (like Creative Commons).

**Disclaimer/Complaints regulations**

If you believe that digital publication of certain material infringes any of your rights or (privacy) interests, please let the Library know, stating your reasons. In case of a legitimate complaint, the Library will make the material inaccessible and/or remove it from the website. Please Ask the Library: <http://uba.uva.nl/en/contact>, or a letter to: Library of the University of Amsterdam, Secretariat, Singel 425, 1012 WP Amsterdam, The Netherlands. You will be contacted as soon as possible.

## THE *ISO*/SWS SPECTRUM OF IRC +10216: THE VIBRATIONAL BANDS OF C<sub>2</sub>H<sub>2</sub> AND HCN<sup>1</sup>

JOSÉ CERNICHARO,<sup>2</sup> ISSEI YAMAMURA,<sup>3</sup> EDUARDO GONZÁLEZ-ALFONSO,<sup>2</sup> TEIJE DE JONG,<sup>3,4</sup>

ANA HERAS,<sup>5</sup> RAFAEL ESCRIBANO,<sup>2</sup> AND JUAN ORTIGOSO<sup>2</sup>

Received 1999 July 20; accepted 1999 September 13; published 1999 October 19

### ABSTRACT

We present the *Infrared Space Observatory*/Short-Wavelength Spectrometer full grating resolution spectrum of IRC +10216, which is dominated by strong absorption/emission bands of C<sub>2</sub>H<sub>2</sub> and HCN. All C<sub>2</sub>H<sub>2</sub> bands and the strong near-infrared stretching bands of HCN are observed in absorption, whereas the fundamental, hot, and combination bands of HCN involving the  $\nu_2$  bending mode around 14  $\mu\text{m}$  are observed in emission. Particularly strong is the HCN  $\nu_2 = 2^0 \rightarrow \nu_2 = 1^1$  vibrational transition at 14.3  $\mu\text{m}$ . The most plausible mechanism for such emission is the radiative pumping of molecules from the ground to the  $\nu_2 = 2^0$  state (7.1  $\mu\text{m}$ ) followed by radiative decay:  $\nu_2 = 2^0 \rightarrow \nu_2 = 1^1$ . We present detailed models for HCN that verify the efficiency of the mentioned effect. The HCN abundance inferred from these models is  $(1.5\text{--}3) \times 10^{-5}$ .

*Subject headings:* circumstellar matter — infrared: stars — radiative transfer — stars: AGB and post-AGB — stars: atmospheres — stars: individual (IRC +10216)

### 1. INTRODUCTION

IRC +10216 is the brightest C-rich, evolved object in the sky. It has an extended circumstellar envelope (CSE) in which around 50 molecular species have been detected, and it is probably one of the best-studied stellar objects at infrared and radio wavelengths (Cernicharo, Guélin, & Kahane 1999). IRC +10216 has a particularly rich carbon chemistry, and most molecular species are carbon-chain radicals (Cernicharo & Guélin 1996) that are formed in the external layers of the CSE (Guélin, Lucas, & Cernicharo 1993). Other species of interest detected in IRC +10216 include silicon carbide (SiC; Cernicharo et al. 1989) and the metal halides NaCl, AlCl, AlF, and KCl (Cernicharo & Guélin 1987). Some metal-bearing species are also detected in the external shell of IRC +10216, e.g., MgNC (Guélin et al. 1993). While polar molecules can be detected at radio wavelengths, nonpolar molecules can only be detected in the infrared through their active ro-vibrational transitions.

We present in this Letter a full grating resolution spectrum of IRC +10216 taken with the *Infrared Space Observatory* (*ISO*)/Short-Wavelength Spectrometer (SWS; de Graauw et al. 1996). It covers the 2.38–45  $\mu\text{m}$  range with a spectral resolution of  $\approx 2000$ . In this work, attention will be confined to the molecular species that most strongly contribute to the emission/absorption spectrum of IRC +10216 in the near- and mid-infrared: C<sub>2</sub>H<sub>2</sub> and HCN (see Cernicharo 1998). A detailed scheme for the radiative pumping of the vibrational levels of HCN is proposed.

<sup>1</sup> Based on observations with *ISO*, an ESA project with instruments funded by ESA member states (especially the PI countries: France Germany, the Netherlands, and the United Kingdom) and with the participation of ISAS and NASA.

<sup>2</sup> Consejo Superior de Investigaciones Científicas, Instituto de Estructura de la Materia, Departamento Física Molecular, Serrano 121, E-28006 Madrid, Spain; cerni@astro.iem.csic.es.

<sup>3</sup> Astronomical Institute “Anton Pannekoek,” University of Amsterdam, Kruislaan 403, 1098 SJ, Amsterdam, Netherlands.

<sup>4</sup> SRON-Utrecht, Sorbonnelaan 2, 3584 CA Utrecht, Netherlands.

<sup>5</sup> Astrophysics Division, Space Science Department of ESA, ESTEC, P.O. Box 299, 2200 AG Noordwijk, Netherlands.

### 2. OBSERVATIONS AND RESULTS

The full range AOT06 spectrum of IRC +10216 was taken on 1996 June 2 ( $\phi = 0.3$ ). Another grating spectrum with lower spectral resolution (AOT01, speed 4) was taken 2 days before. The data have been reduced with the SWS Interactive Analysis system. The very strong signal of IRC +10216 causes significant fringes in band 3 (12–28  $\mu\text{m}$ ) and memory effects in bands 2 (4–12  $\mu\text{m}$ ) and 4 (29–45.2  $\mu\text{m}$ ). Nevertheless, the AOT06 and AOT01 spectra can be compared to ascertain the reality of the observed features. Fringes in band 3 (12–29  $\mu\text{m}$ ) are removed by shifting the relative spectral responsivity function slightly in wavelength. Fringes still remaining after this procedure were removed by a fast Fourier transform method. There are small discrepancies in flux between the different bands in the spectrum. Such inconsistencies have been corrected by scaling each band with respect to bands 1b (2.6–3.0  $\mu\text{m}$ ) and 3d (19.5–27.5  $\mu\text{m}$ ), which are adopted as references.

Figure 1 shows the spectrum of IRC +10216 between 2.4 and 197  $\mu\text{m}$ ; it shows the 11  $\mu\text{m}$  band due to SiC dust and the prominent “30  $\mu\text{m}$  feature” usually attributed to MgS (Yamamura et al. 1998). A crude fit to the dust emission provides an estimate for the mass-loss rate of a few times  $10^{-5} M_{\odot} \text{yr}^{-1}$ . Figure 2 shows the spectrum, normalized to the continuum flux, of IRC +10216 around 3  $\mu\text{m}$ . The broad absorption feature is due to the stretching modes of C<sub>2</sub>H<sub>2</sub> and HCN. In addition, the  $Q$ -branches of the  $\nu_3 + \nu_5$ ,  $\nu_3 + \nu_4$ , and  $\nu_2 + 2\nu_4 + \nu_5$  modes of C<sub>2</sub>H<sub>2</sub> are detected around 2.5  $\mu\text{m}$ . Many combination bands of the same molecule, like  $\nu_2 + \nu_5$  among others, are detected around 3.8  $\mu\text{m}$ . The  $Q$ -branch of the  $\nu_2 + \nu_3$  mode of HCN and those of the  $\nu_1 + \nu_2$  and  $\nu_1 + 2\nu_2 - \nu_2$  vibrational transitions of HCN are detected in absorption around 2.5 and 3.6  $\mu\text{m}$ , respectively. The  $\nu_3 - \nu_2$  mode of HCN and other combination bands are also seen in Figure 2. However, they seem to be in emission rather than in absorption (see below).

Figure 3 shows the 14  $\mu\text{m}$  spectrum of IRC +10216. The absorption features are due to the  $\nu_5$  bending mode of acetylene and its associated hot and combination bands. This figure shows that the HCN bands appear in emission with the  $Q$ -branch of the  $2\nu_2^0 - \nu_2^1$  transition being particularly strong. Many other weak bands are also present in the 2.4–9 and 15–25  $\mu\text{m}$  ranges. However, instrumental problems preclude at the present the definite assignation of these features to real spectral bands.

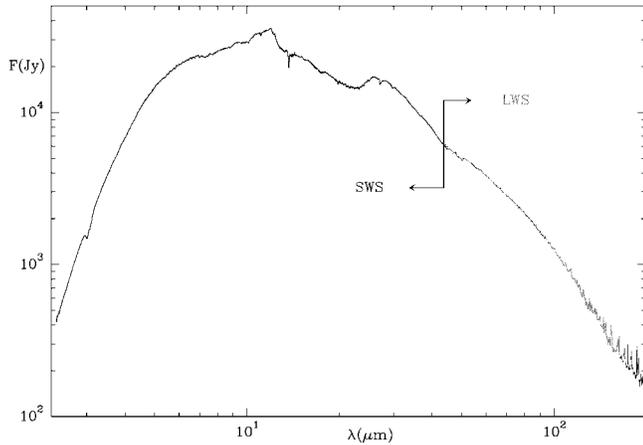


FIG. 1.—Spectrum of IRC +10216 between 2.4 and 197  $\mu\text{m}$  as observed with the *ISO/SWS* and *ISO/LWS* spectrometers. At long wavelengths, the pure rotational lines of CO and HCN are clearly superposed over the weak continuum (see Cernicharo et al. 1996).

Future data products from improved calibration are needed to unambiguously derive the nature of observed features and to assign them to new species or to abundant molecules already detected in IRC +10216 such as SiS,  $\text{NH}_3$ ,  $\text{CH}_4$ , and  $\text{SiH}_4$ , which have been observed in the infrared (see, e.g., Boyle et al. 1994; Keady & Ridgway 1993).

### 3. DISCUSSION

Observations of many  $\text{C}_2\text{H}_2$  ro-vibrational lines at 3 and 14  $\mu\text{m}$  with high spectral resolution have been reported by Ridgway, Carbon, & Hall (1978), Keady & Hinkle (1988), and Keady & Ridgway (1993). Although our data suffer from lower spectral resolution, we avail of the observations from all the  $\text{C}_2\text{H}_2$  bands (Figs. 1 and 2), thus allowing a better estimate of the temperature of the emitting/absorbing gas and the pumping mechanisms prevailing in the innermost region of the CSE. A first attempt to fit the absorption with the  $\nu_3$  and  $\nu_2+\nu_4+\nu_5$  bands of  $\text{C}_2\text{H}_2$  ( $\mu = 0.061$  and  $0.067$  D; Rinsland, Baldacci, & Rao 1982) and with the  $\nu_3$  band of HCN indicated that the contribution from several hot and combination bands of  $\text{C}_2\text{H}_2$ , together with bands from  $\text{H}^{13}\text{CCH}$  [ $x(\text{C}_2\text{H}_2)/x(\text{H}^{13}\text{CCH}) \approx 22$ ; Cernicharo et al. 1991], should be included. By assuming an LTE distribution of the  $\text{C}_2\text{H}_2$  and HCN levels' population, we have modeled the absorption produced by the  $\nu_3$  and  $\nu_2+\nu_4+\nu_5$  bands and all the hot bands arising from  $\nu_4 = 1, 2$ ,  $\nu_5 = 1, 2$ , and  $\nu_4+\nu_5$ . In total we have included in our models 21 bands of  $\text{C}_2\text{H}_2$  and two bands of  $\text{H}^{13}\text{CCH}$ , together with three bands of HCN and the  $\nu_3$  band of  $\text{H}^{13}\text{CN}$ .

A model with a single layer cannot explain the observed band profile (see Keady & Hinkle 1988). The adopted model, which is a simplification of the detailed CSE structure derived by Keady & Hinkle, includes several layers between 1 and 500 stellar radii. Figure 2 shows the best fit to the band shape at 3  $\mu\text{m}$ . We have adopted  $X(\text{C}_2\text{H}_2) = 5 \times 10^{-5}$  and  $X(\text{HCN}) = 3 \times 10^{-5}$  (Cernicharo et al. 1996; Wiedemann et al. 1991). In

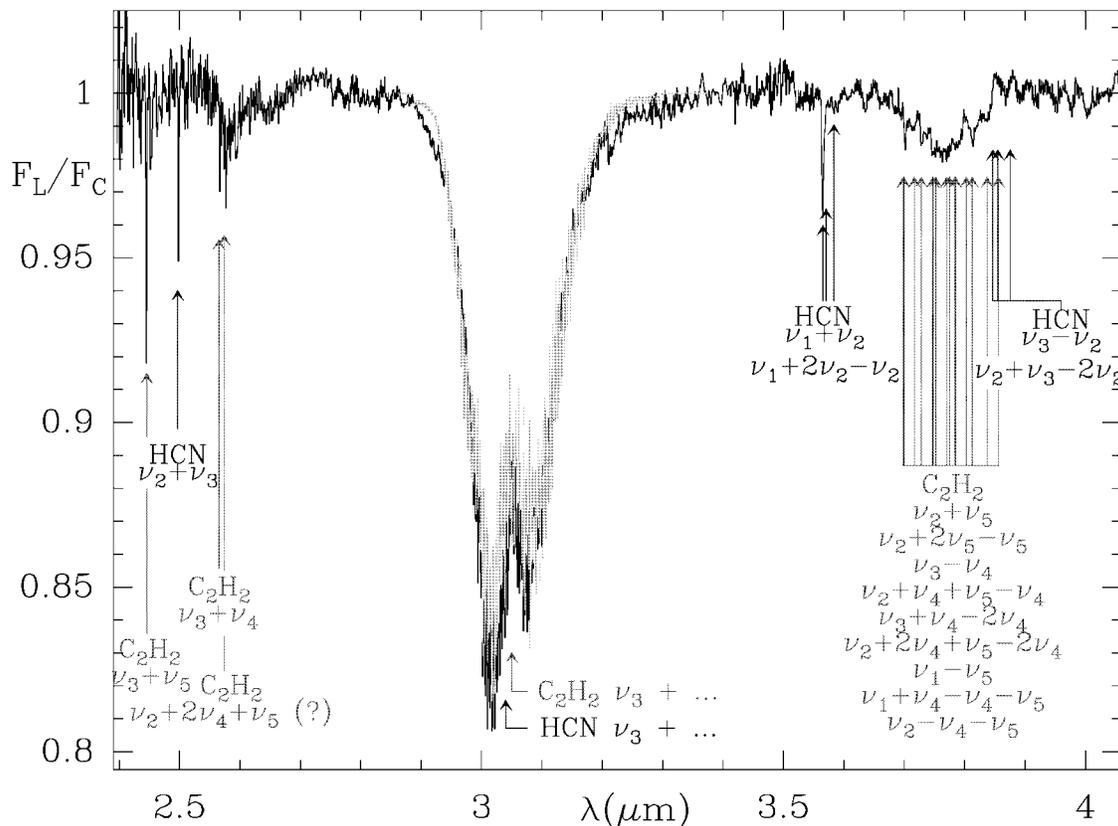


FIG. 2.—Observed spectrum between 2.4 and 4.0  $\mu\text{m}$  in IRC +10216. The spectral resolution is  $\approx 150 \text{ km s}^{-1}$ . The thin line shows the modeling of the  $\text{C}_2\text{H}_2$  and HCN stretching mode at 3  $\mu\text{m}$  and their hot and combination bands. The positions of the bands are indicated by arrows.

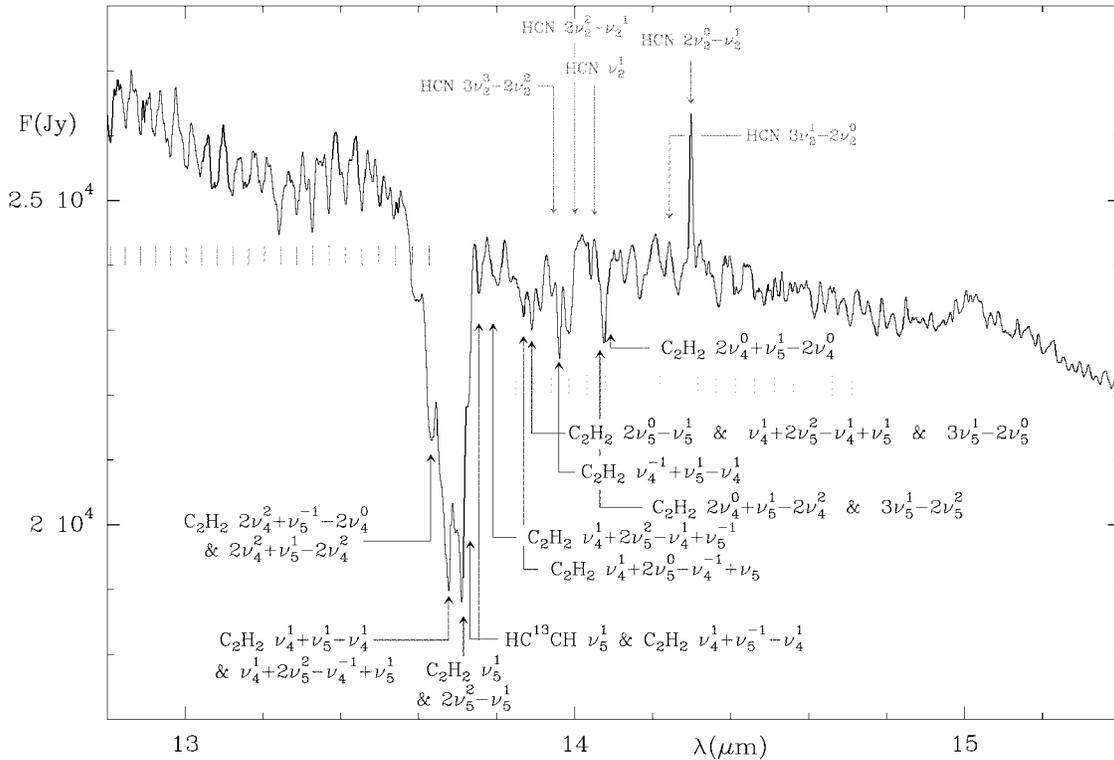


FIG. 3.—Observed spectrum around  $15 \mu\text{m}$ . The spectral resolution is  $\sim 200 \text{ km s}^{-1}$ . The position of the bands are indicated. Vertical lines indicate the positions of the individual  $R$  and  $P$  lines of the fundamental bending mode of  $\text{C}_2\text{H}_2$ . Note that the HCN bands are in emission as indicated by the shape of the  $P$ -branch of  $\text{C}_2\text{H}_2$  relative to the  $R$  one and by the strong  $Q$ -branch of the  $2\nu_2$   $l=0$ - $\nu_2$   $l=1$  transition.

order to fit the wings of the  $3 \mu\text{m}$  feature, we have included an additional contribution from a warm gas in the inner region ( $T_K \approx 1700 \text{ K}$ ;  $1 < r < 3R_*$ ). The  $\text{C}_2\text{H}_2$  and HCN abundances in this layer are  $2.5 \times 10^{-5}$  and  $1.5 \times 10^{-5}$ , respectively. This layer could be similar to the warm molecular envelope found by Tsuji et al. (1997) and Yamamura et al. (1999) in other CSEs, and its origin is probably related to the pulsation of the central star (Woitke et al. 1999). Observations of high- $J$  lines of  $\text{C}_2\text{H}_2$  and HCN with much better spectral and angular resolution are needed to infer the structure of the inner envelope and the abundance of both molecular species. Fuente, Cernicharo, & Omont (1998) have derived an abundance ratio  $\text{C}_2\text{H}_2/\text{HCN} \approx 1$  from observations of CCH and HCN. This value corresponds to the external layers of the CSE in which CCH is observed.

The HCN emission has to be modeled in a more complicated way. The  $\nu_3$  band of HCN at  $3 \mu\text{m}$ , although contaminated by  $\text{C}_2\text{H}_2$ , seems to be in absorption, whereas the bending states are detected in emission against the continuum. The  $Q$ -branch of the  $2\nu_2^0-\nu_2^1$  band is particularly strong, in part due to low-velocity separation between the individual lines ( $\approx 0.4 \text{ km s}^{-1}$ ). All the other  $Q$ -branches ( $\nu_2^1$ ,  $2\nu_2^2-\nu_2^1$ ,  $3\nu_2^1-2\nu_2^0$ ) also appear in emission, but are weaker. P Cygni profiles are in fact expected for HCN and  $\text{C}_2\text{H}_2$  (Wiedemann et al. 1991). The larger beam size of the SWS compared to the high angular resolution data of Wiedemann et al. could favor the observation of reemission from HCN through resonant scattering in the molecular envelope. Obviously the same could happen for acetylene. The different behavior of both molecules must be related to the larger abundance of  $\text{C}_2\text{H}_2$  through the whole envelope (which increases the importance of stellar blocking and absorption by

the dust due to the larger number of absorption and reemission processes; see González-Alfonso & Cernicharo 1999a, 1999b) and/or to the excitation mechanism that gives rise to the HCN band emission. An important difference between both molecules would be established if ro-vibrational excitation is mainly radiative instead of collisional: while the  $2\nu_2^0$  state of HCN is radiatively connected to the ground state, the equivalent transition in  $\text{C}_2\text{H}_2$  is forbidden, i.e., no photons can be absorbed from the ground to the  $2\nu_5^0$  state due to the fact that both levels have the same symmetry ( $\sigma_g^+$ ). In HCN, photons absorbed from the ground to the  $2\nu_2^0$  level will decay preferentially to the  $\nu_2^1$  level, because this transition has an Einstein coefficient for spontaneous emission of  $\sim 4 \text{ s}^{-1}$ , i.e., twice the value corresponding to the decay to the ground state. The equivalent process in  $\text{C}_2\text{H}_2$  is radiative excitation from the ground to the  $\nu_4+\nu_5$  state followed by radiative decay to the  $\nu_4$  level. However, this transition is very close in frequency to the  $Q$ -branch of the  $\nu_5$  transition, and any possible reemission could not compensate for the strong absorption from the ground to the  $\nu_5$  bending mode.

With the use of the nonlocal, non-LTE code described by González-Alfonso & Cernicharo (1997), we have modeled the HCN emission/absorption in IRC +10216. The rotational transitions of  $2\nu_2+\nu_3$  ( $l=0, 2$ ),  $\nu_2+\nu_3$ ,  $\nu_3$ ,  $\nu_3+\nu_2-\nu_2$ ,  $\nu_3-\nu_2$ ,  $2\nu_2$  ( $l=0, 2$ ),  $\nu_2$ ,  $2\nu_2-\nu_2$ , and  $3\nu_2-2\nu_2$  bands of HCN have been calculated using the rotational constants and transition dipole moments for these bands from Maki et al. (1996), Smith et al. (1989), Kim & King (1979), and Smith (1981). We have adopted a mass-loss rate of  $1.5 \times 10^{-5} M_\odot \text{ yr}^{-1}$  (Cernicharo et al. 1996).

The results (see Fig. 4) predict that the  $\nu_3$ ,  $\nu_3+\nu_2$ , and  $2\nu_2$

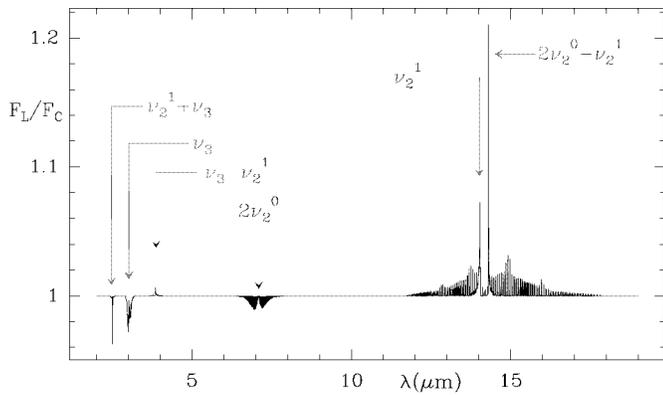


FIG. 4.—Modeled HCN spectrum in IRC +10216. The model is described in the text. Note the large intensity predicted for the  $Q$ -branch of the  $2\nu_2$   $l = 0 - \nu_2$   $l = 1$  transition, in good agreement with the observations. In addition, the  $\nu_3 - \nu_2^1$  band is also predicted in emission as observed in Fig. 1.

bands of HCN will be in absorption while the  $\nu_2$  and its overtone bands will be in emission. The model also predicts that the  $\nu_3 - \nu_2$  band will be in emission. In fact, a detailed analysis of Figure 2 shows that this band appears indeed in emission. With the adopted HCN abundance ( $3 \times 10^{-5}$  from Cernicharo et al. 1996), which matches the HCN rotational emission from a more simple large velocity gradient analysis, the  $Q$ -branch emission of the  $2\nu_2^0 - \nu_2^1$  transition is overestimated by a factor of  $\approx 2$ , so that a better estimate of the HCN abundance would be  $\sim (1-1.5) \times 10^{-5}$ . The pumping mechanisms are different for the different regions of the envelope. In the innermost regions, before the dust formation zone, absorption of photons at  $3 \mu\text{m}$  in the  $\nu_3$  mode will be very efficient. Radiative decay to the  $2\nu_2$   $l = 0$  and  $\nu_2$   $l = 1$  states will be an important mechanism

to populate both levels. Also in this region the  $\nu_3 - \nu_2$  band will appear in emission. Outside the dust formation region the flux at  $3 \mu\text{m}$  decreases, but it increases considerably at  $7.1 \mu\text{m}$ . The pumping from the ground state to the  $2\nu_2$   $l = 0$  vibrational level, followed by radiative decay to the  $\nu_2$   $l = 1$  level, will be the main mechanism in producing the observed emission at  $14.3 \mu\text{m}$ . Consequently, the intensity of the HCN bands, as observed by *ISO*, will be very different in stars with low, moderate, or high mass-loss rates (see, e.g., Aoki, Tsuji, & Ohnaka 1999). Our calculations show that resonant scattering through the envelope is very efficient for HCN and also probably for  $\text{C}_2\text{H}_2$  (see Cernicharo 1998; González-Alfonso et al. 1998 have modeled similar effects for water vapor). Resonant scattering in CSEs could also have an effect on the heating and cooling of the molecular gas (Cernicharo 1998). The behavior of the observed HCN emission bands is similar to that found in O-rich stars by Justtanont et al. (1998) for  $\text{CO}_2$ ; in fact, González-Alfonso & Cernicharo (1999a, 1999b) have shown that the  $\text{CO}_2$  emission around  $15 \mu\text{m}$  could also be explained by resonant scattering through the CSEs. The radiative pumping mechanism for HCN was already mentioned by Lucas & Cernicharo (1989) to explain the strong  $\nu_2 = 1$   $J = 2-1$  maser in IRC +10216 and other C-rich CSEs. Our calculations indicate that the low- $J$  lines of the  $\nu_2 = 1$  vibrational state are inverted and that the proposed radiative pumping could explain the maser emission. However, the ro-vibrational lines of the  $Q$ -branch are overlapped in frequency, and these effects have to be considered in order to explain the millimeter and submillimeter masers of HCN (see, e.g., González-Alfonso & Cernicharo 1997).

J. C. and E. G.-A. acknowledge Spanish DGES for this research under grants PB96-0883 and ESP98-1351E.

#### REFERENCES

- Aoki, W., Tsuji, T., & Ohnaka, K. 1999, *A&A*, 345, 945  
 Boyle, R. J., Keady, J. J., Jennings, D. E., Hirsch, K. L., & Wiedemann, G. R. 1994, *ApJ*, 420, 863  
 Cernicharo, J. 1998, *Ap&SS*, 255, 303  
 Cernicharo, J., et al. 1996, *A&A*, 315, L201  
 Cernicharo, J., Gottlieb, C. A., Guélin, M., Thaddeus, P., & Vrtilke, J. M. 1989, *ApJ*, 341, L25  
 Cernicharo, J., & Guélin, M. 1987, *A&A*, 183, L10  
 ———. 1996, *A&A*, 309, L27  
 Cernicharo, J., Guélin, M., Bogey, M., Demuynck, C., & Destombes, J. L. 1991, *A&A*, 246, 213  
 Cernicharo, J., Guélin, M., & Kahane, C. 1999, *A&A*, in press  
 de Graauw, T., et al. 1996, *A&A*, 315, L49  
 Fuente, A., Cernicharo, J., & Omont, A. 1998, *A&A*, 330, 232  
 González-Alfonso, E., & Cernicharo, J. 1997, *A&A*, 322, 938  
 ———. 1999a, in *The Universe as Seen by ISO*, ed. P. Cox & M. Kessler (ESA SP-427; Noordwijk: ESA), 325  
 ———. 1999b, *ApJ*, 525, 845  
 González-Alfonso, E., Cernicharo, J., van Dishoeck, E. F., Wright, C. M., & Heras, A. 1998, *ApJ*, 502, L169  
 Guélin, M., Lucas, R., & Cernicharo, J. 1993, *A&A*, 280, L19  
 Justtanont, K., Feuchtgruber, J., de Jong, T., Cami, J., Waters, L. B. F. M., Yamamura, I., & Onaka, T. 1998, *A&A*, 330, L17  
 Keady, J. J., & Hinkle, K. H. 1988, *ApJ*, 331, 539  
 Keady, J. J., & Ridway, S. T. 1993, *ApJ*, 406, 199  
 Kim, K., & King, W. T. 1979, *J. Chem. Phys.*, 71, 1967  
 Lucas, R., & Cernicharo, J. 1989, *A&A*, 218, L20  
 Maki, A., Quapp, W., Klee, S., Mellau, G. C., & Albert, S. 1996, *J. Mol. Spectrosc.*, 180, 323  
 Ridgway, S. T., Carbon, D. F., & Hall, D. N. B. 1978, *ApJ*, 225, 138  
 Rinsland, C. P., Baldacci, A., & Rao, K. N. 1982, *ApJS*, 49, 487  
 Smith, I. W. M. 1981, *J. Chem. Soc. Faraday Trans.*, 77, 2357  
 Smith, M. A. H., Coy, S. L., Klemperer, W., & Lehmann, K. K. 1989, *J. Mol. Spectrosc.*, 134, 134  
 Tsuji, T., Ohnaka, K., Aoki, W., & Yamamura, I. 1997, *A&A*, 320, L1  
 Wiedemann, G. R., Hinkle, K. H., Keady, J. J., Deming, D., & Jennings, D. E. 1991, *ApJ*, 382, 321  
 Woitke, P., Helling, Ch., Winters, J. M., & Jeong, K. S. 1999, *A&A*, 348, L17  
 Yamamura, I., de Jong, T., Justtanont, K., Cami, J., & Waters, L. B. F. M. 1998, *Ap&SS*, 255, 351  
 Yamamura, I., de Jong, T., Waters, L. B. F. M., Cami, J., & Justtanont, K. 1999, in *IAU Symp. 191, Asymptotic Giant Branch Stars*, ed. T. Le Bertre, A. Lébre, & C. Waelkens (San Francisco: ASP), 267