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Simultaneous X-Ray and Infrared Observations of Sagittarius A''s Variability

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Abstract

Emission from Sagittarius A* is highly variable at both X-ray and infrared (IR) wavelengths. Observations over the last ~20 yr have revealed X-ray flares that rise above a quiescent thermal background about once per day, while faint X-ray flares from Sgr A* are undetectable below the constant thermal emission. In contrast, the IR emission of Sgr A* is observed to be continuously variable. Recently, simultaneous observations have indicated a rise in IR flux density around the same time as every distinct X-ray flare, while the opposite is not always true (peaks in the IR emission may not be coincident with an X-ray flare). Characterizing the behavior of these simultaneous X-ray/IR events and measuring any time lag between them can constrain models of Sgr A*'s accretion flow and the flare emission mechanism. Using 100+ hours of data from a coordinated campaign between the Spitzer Space Telescope and the Chandra X-ray Observatory, we present results of the longest simultaneous IR and X-ray observations of Sgr A* taken to date. The cross-correlation between the IR and X-ray light curves in this unprecedented data set, which includes four modest X-ray/IR flares, indicates that flaring in the X-ray may lead the IR by approximately 10–20 min with 68% confidence. However, the 99.7% confidence interval on the time-lag also includes zero, i.e., the flaring remains statistically consistent with simultaneity. Long-duration and simultaneous multi-wavelength observations of additional bright flares will improve our ability to constrain the flare timing characteristics and emission mechanisms, and must be a priority for Galactic Center observing campaigns.

Key words: accretion, accretion disks – black hole physics – Galaxy: center – radiation mechanisms: non-thermal

1. Introduction

Sagittarius A* (Sgr A*) sits at the center of the Milky Way and is the closest supermassive black hole (SMBH), at a distance of only ~8 kpc (e.g., Boehle et al. 2016; Gillessen et al. 2017). Monitored in the radio since its discovery, and more recently in the infrared (IR) and the X-ray, Sgr A* has a mass of ~4 × 10^6 M_☉ (Boehle et al. 2016; Gillessen et al. 2017), an extremely low bolometric-to-Eddington luminosity ratio (L/L_Edd ~ 10^-9; Genzel et al. 2010), and appears to be accreting material at a very low rate (< 10^-7 M_☉ yr^-1; Baganoff et al. 2003; Marrone et al. 2006, 2007; Shcherbakov et al. 2012; Yusef-Zadeh et al. 2015).

In X-rays, Sgr A* appears as a persistent source, with a flux of about 3 × 10^-3 erg s^-1 (Baganoff et al. 2001, 2003). This faint, steady emission arises from thermal bremsstrahlung radiation from a hot accretion flow dominated by regions near the Bondi radius (Quataert 2002; Baganoff et al. 2003; Yuan et al. 2003; Liu et al. 2004; Xu et al. 2006; Wang et al. 2013) and is interrupted about once per day by distinct flares of non-thermal emission coming from very close to the black hole (Neilsen et al. 2013, 2015). First detections of Sgr A* in the IR also revealed a highly variable source (Genzel et al. 2003; Ghez et al. 2004) with peaks in the IR emission detected more frequently than in X-rays. Since these first discoveries, the statistical behavior of both the X-ray (e.g., Li et al. 2015; Neilsen et al. 2015; Ponti et al. 2015) and the IR (e.g., Dodd-Eden et al. 2011; Witzel et al. 2012, 2018; Hora et al. 2014) activity have been well studied and the flux density distributions of these two wavelengths can both be described by a power law (Witzel et al. 2012, 2018; Neilsen et al. 2015). Though the X-ray and IR have no other clear statistical similarities, the coincidence of peaks in the variability hints that there may be a physical connection between them (e.g., Eckart et al. 2006a, 2008; Yusef-Zadeh et al. 2006, 2009; Dodd-Eden et al. 2009; Mossoux et al. 2016). The picture is even less clear at longer wavelengths, where some submm flares have been tentatively linked to IR flares with some delay (Marrone et al. 2008; Morris et al. 2012) or no delay (Fazio et al. 2018), and radio variability has not been observed to coincide with IR or X-ray activity (e.g., Capellupo et al. 2017). However, upcoming submm and coordinated multi-wavelength observations undertaken by the Event Horizon Telescope collaboration (Doeleman et al. 2008) may shed light on these connections in the near future.

Despite the intensive efforts that have been made to characterize Sgr A*’s variable emission, the physical mechanisms producing the variability are still unknown. Suggested physical models include particle acceleration due to magnetic...
reconnection events, violent disk instabilities, jets, other stochastic processes in the accretion flow (e.g., Markoff et al. 2001; Liu & Melia 2002; Yuan et al. 2003; Liu et al. 2004; Dexter et al. 2009; Maitra et al. 2009; Dodds-Eden et al. 2010; Ball et al. 2016; Li et al. 2017), and even tidal disruption of asteroids (Cadevè et al. 2008; Kostić et al. 2009; Zubovas et al. 2012). Additional models attempt to explain the variability in the context of expanding plasma blobs (e.g., van der Laan 1966; Marrone et al. 2008; Yousni & Wu 2015; Li et al. 2017), themselves launched by magnetic reconnection events or unsteady jet emission. Finally, gravitational lensing near the horizon of the SMBH might add an amplifying effect to the observed emission (Chan et al. 2015).

One possible picture for the IR flares describes a population of electrons undergoing continuous acceleration due to turbulent processes in the inner accretion flow and subsequently emitting synchrotron radiation. This is supported by observed timescales for the IR variability, with factors of $\gtrsim 10$ changes within $\sim 10$ min (e.g., Genzel et al. 2003; Ghez et al. 2004; Witzel et al. 2018), the spectral index at high flux densities ($\alpha \approx 0.6$; Hornstein et al. 2007; Bremer et al. 2011; Witzel et al. 2014), and the high linear polarization of the IR emission (Eckart et al. 2006b, 2008; Meyer et al. 2006, 2007; Tripp et al. 2007; Yusef-Zadeh et al. 2007; Witzel et al. 2011; Shahzamanian et al. 2015). The exact physical parameters of this turbulent acceleration of electrons (e.g., background magnetic field strength and $\gamma$, the Lorentz factor of the electrons) and the details of the radiative processes linking the IR flares to the X-ray remain unknown.

The radiation mechanisms often invoked to connect IR flares to simultaneous X-ray outbursts include pure synchrotron (e.g., Markoff et al. 2001; Dodds-Eden et al. 2009; Barrière et al. 2014; Ponti et al. 2017), synchrotron self-Compton (SSC) (Markoff et al. 2001; Eckart et al. 2008), and inverse Compton (IC; Yusef-Zadeh et al. 2012). Monitoring the black hole simultaneously in multiple wavelengths can constrain these variability models via the association of peaks at different energies and times. With this motivation, Sgr A* has recently been monitored simultaneously at wavelengths ranging from the radio, mm, submm, IR, X-rays, to gamma-rays (e.g., Dodds-Eden et al. 2009; Yusef-Zadeh et al. 2009, 2006; Trap et al. 2011; Eckart et al. 2012, 2006a; Fazio et al. 2018).

Sgr A*'s flaring radiation mechanisms can also be constrained by the statistical behavior of the flares. For example, Dibi et al. (2016) showed that the measured flux distributions of flares in the X-ray (Neilsen et al. 2013, 2015) and IR (Witzel et al. 2012, 2018) were consistent with both synchrotron and SSC models, though their different shapes are difficult to understand if driven by the same particle populations.

The first simultaneous IR and X-ray observations of Sgr A* were carried out by Eckart et al. (2004). Since then, studies with both X-ray and IR monitoring of Sgr A* have often reported that flares in the two wavelengths are simultaneous with each other or with the X-ray leading the IR by no more than 10 min (Eckart et al. 2006a, 2008; Yusef-Zadeh et al. 2006, 2009; Dodds-Eden et al. 2009; Mossoux et al. 2016). The exception is Yusef-Zadeh et al. (2012), who reported a short time lag between the maxima of the IR and X-ray flares, with the X-ray flares possibly lagging the IR. Such a behavior would point toward an interpretation of an IC scattering model, where a fraction of IR photons are up-scattered to X-ray energies and are seen as an X-ray “echo.”

These studies support the inference of a physical connection between the X-ray and IR flaring. However, the ground-based IR observations are often significantly shorter than the X-ray observations and there are frequently gaps in the data. When flares occur in one wavelength outside the observing window of another observatory, it is difficult to robustly associate two events (e.g., Marrone et al. 2008; Mossoux et al. 2016; Capellupo et al. 2017), leading to uncertainty in the cross-correlation.

The first observations of Sgr A* with the Spitzer Space Telescope (Hora et al. 2014) provided a continuous $>23$ hr light curve of Sgr A* at IR wavelengths that was more than a factor of two longer than the previous record holder (600 min; Meyer et al. 2008). Building on this study, we utilized two space telescopes (Spitzer and the Chandra X-ray Observatory) to obtain six simultaneous observations of Sgr A* at $4.5 \mu m$ and 2–8 keV. Four of these epochs have $\sim 24$ hr of overlapping coverage from the two observatories, maximizing the probability of catching the relatively rare X-ray flares with simultaneous IR monitoring. A detailed statistical analysis of the IR data was presented in Witzel et al. (2018), and our first multi-wavelength results were reported in Fazio et al. (2018).

In this work we investigated the temporal correlations between X-ray and IR variability using these six contemporaneous Chandra and Spitzer observations of Sgr A*, Section 2 describes these observations and the reduction of the data, while Section 3 details our characterization of the variability of Sgr A* by cross-correlating the light curves. Section 4 explores the results in the context of previous studies and discusses their implications for models of the variability.

## 2. Observations and Data Reduction

The IRAC instrument (Fazio et al. 2004) on Spitzer (Werner et al. 2004) was used to observe Sgr A* at $4.5 \mu m$ for eight $\sim 24$ hr long stretches between 2013 and 2017. Six of these observations had simultaneous monitoring from Chandra (Weisskopf et al. 2000), centered on Sgr A*’s radio position (R.A., decl. $= 17:45:40.0409, \sim -29:00:28.118$; Reid & Brunthaler 2004) and are listed in Table 1. The first two epochs had overlapping Chandra coverage for $\sim 5.5$ hr, while the other four had continuous $\sim 24$ hr coverage from both observatories. This results in a total of $\sim 107$ hr for which we collected simultaneous X-ray and IR data—the largest such data set to date. Figure 1 shows the resulting Sgr A* light curves.

### 2.1. Spitzer

Sgr A* and its immediate surroundings are unresolved with the Spitzer/IRAC detector. The measured flux of a single pixel is also highly sensitive to any changes in the telescope’s pointing, even on the subpixel level. This makes detecting the intrinsic variability of Sgr A* a significant challenge. Recovering the signal requires modeling the flux variations of the pixel containing the black hole as a property of both (1) the varying intra-pixel sensitivity of Spitzer/IRAC detector and (2) the telescope’s sub-pixel motion. Hora et al. (2014) presented the first detection of Sgr A* with Spitzer and demonstrated that a very precise relative flux measurement can be recovered for Sgr A* using this approach.
Table 1
Observing Log for Simultaneous Chandra and Spitzer Observations

<table>
<thead>
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<th>Obs Date</th>
<th>ObsID</th>
<th>Obs. Start (UT)</th>
<th>Obs. End (UT)</th>
<th>Chandra</th>
<th>AORKEY</th>
<th>AOR Start (UT)</th>
<th>AOR End (UT)</th>
<th>Mode</th>
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<td>22:56:01</td>
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</tr>
</tbody>
</table>

Notes.

a A (+1) in the start/end time columns indicates the offset of one day from the first date listed in the first column. For example, the first Chandra observation began at 02:59:23 on June 3. Times are UTC at the observatory. The time-lag analysis was based on corrected heliocentric times.

b Mode of operation (either Map or Stare modes). The “Map” mode was a short operation performed after the initial slew to Sgr A*, while the two “Stare” modes were each ∼12 hr long observations taken with a ∼4 min break between the two.

The Spitzer data presented in this work were obtained and reduced using the procedures described by Hora et al. (2014) and Witzel et al. (2018), the latter of which analyzes these and additional Spitzer light curves in the context of other IR data sets from Keck and the Very Large Telescope (VLT). As documented in Table 1, each observation was conducted as a set of three Astronomical Observation Requests (AORs). Each epoch followed the same observing sequence: an initial mapping operation performed after the slew to the Sgr A* field followed by two successive staring operations. Each staring operation began by using the “PCRS Peakup” mode to position Sgr A* on the center of pixel (16, 16) in the IRAC subarray. The subarray mode for Spitzer/IRAC reads out 64 consecutive images (a “frame set”) of a 32 × 32 pixel region on the IRAC detector. This frame set is known as one basic calibrated data product (BCD), which is the data format downloaded from the Spitzer Heritage Archive.11 Each frame in the frame set is a 0.1 s 32 × 32 image, so one frame set takes 6.4 s to complete. After converting the pixel intensity into mJy, each frame set was combined into a single 32 × 32 image referred to as a “6.4 s BCD coadd.” Since frame sets were typically separated by 2 s of telescope overheads, this resulted in an observation cadence of approximately 8.4 s.

The data reduction is described in Section 2.1 of Witzel et al. (2018), which is an improved version of the procedure in Appendix A1 of Hora et al. (2014). This procedure corrects for the effect of nearby sources on the measured flux of Sgr A* as the telescope jitters on a sub-pixel basis throughout the observations. The resulting light curves are displayed in Figure 1. The baseline flux density of these IR light curves is unknown, though the value has been inferred to be 1.9 mJy from the cumulative distributions of flux densities of Sgr A* in Witzel et al. (2018). Therefore, the Spitzer light curves in Figure 1 plot the excess flux density above 1.9 mJy from pixel (16,16), attributed to Sgr A*’s variability.

2.2. Chandra

All Chandra observations were acquired using the ACIS-S3 chip in the FAINT mode with a one-eighth subarray. The small subarray was chosen to avoid photon pileup in any bright flares from Sgr A* and in the nearby magnetar, SGR J1745-2900 (Coti Zelati et al. 2015, 2017). We performed Chandra data reduction and analysis with CIAO v4.9 tools12 and calibration database 4.7.3. The chandra_repro script was used to reprocess level 2 events files before the WCS coordinate system was updated (wcs_update). We used the CIAO tool axbary to apply barycentric corrections to the event times. Finally, we extracted a 2–8 keV light curve from a circular region of radius 1°25 centered on Sgr A*. The small extraction region and energy range isolate Sgr A*’s emission from the nearby magnetar (e.g., Coti Zelati et al. 2017) and the diffuse X-ray background (e.g., Bagannotto et al. 2003; Nowak et al. 2012; Wang et al. 2013). These X-ray light curves are plotted in Figure 1.

11 The Spitzer Heritage Archive (http://irsa.ipac.caltech.edu) is part of the NASA/IPAC Infrared Science Archive, which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

12 Chandra Interactive Analysis of Observations (CIAO) software is available at http://cxc.harvard.edu/ciao/.
3. Analysis

3.1. Flare Detection and Characteristics

To detect and characterize X-ray flares, we used the Bayesian Blocks algorithm as described by Scargle (1998) and Scargle et al. (2013) and provided by Peter K. G. Williams (bblocks; Williams et al. 2017). This algorithm takes an unbinned, filtered Chandra events file as input and models the light curve as a sequence of blocks with constant rates. A single block characterizes a region of the light curve for which there is no significant variability, and a light curve with a flare is made up of multiple blocks of differing count rates separated by change points. The code adopts a geometric prior on the number of blocks, preventing over-fitting the light curve by favoring fewer blocks when fewer events are present in the Chandra events file. We ran the algorithm requiring a detection significance of 95% for a single change point (a false positive rate of $p_0 = 0.05$), which implies that the overall detection significance of a flare (at least two change points) is $1 - p_0^2 \approx 99.8\%$ (see, e.g., Nowak et al. 2012; Neilsen et al. 2013).

We detected four X-ray flares during the total overlapping coverage of X-ray and IR coverage, one on 2016 July 12, one on 2016 July 18, and two on 2017 July 15. Increasing $p_0$ to 0.1 resulted in the detection of five flares (where the narrow peak around 910 min on 2017 July 15 is also identified as an individual flare). Both numbers are consistent with past measurements of the average number of X-ray flares from Sgr A* ($\sim$1.1/day; Neilsen et al. 2015; Ponti et al. 2015). The results of the Bayesian Blocks analysis are overplotted on the X-ray curves in orange. Labels 1–4 indicate the four IR flux peaks associated with significant X-ray activity.

Figure 1. Simultaneous IR and X-ray light curves of Sgr A*. Plotted in gray/red is the excess flux density (mJy) of the pixel containing Sgr A* from Spitzer 4.5 μm observations (see Section 2.1 of Witzel et al. 2018). Times for each epoch are relative to the beginning of the Spitzer observations, measured in Heliocentric Modified Julian Date. Gray dots are the flux densities of each 6.4 s BCD coadd, while the red line shows the data binned over 3.5 min. Chandra light curves of Sgr A* at 2–8 keV are plotted in purple with 300 s binning. The $p_0 = 0.05$ Bayesian Blocks results are overplotted on the X-ray curves in orange. Labels 1–4 indicate the four IR flux peaks associated with significant X-ray activity.
X-ray variability; Neilsen et al. 2013), the emission from Sgr A* at IR wavelengths is constantly varying. In the three epochs where we observe significant X-ray activity, the IR flux from Sgr A* rises above \(\sim -2\) mJy within tens of minutes of the peak in the X-ray (see the top row of Figure 2). The IR activity associated with significant X-ray flares appears to rise at the same time as the X-ray but lasts for a longer time (FWHM$_{IR} \gtrsim 2 \times$ FWHM$_{X-ray}$). These longer IR “flares” are labeled 1–4 in Figures 1 and 2, with flares 1–3 roughly associated with the first three X-ray flares and flare 4 associated with both the fourth and fifth X-ray flares identified with the Bayesian Blocks analysis. There are also multiple IR peaks at \(\sim -2\) mJy where we see no significant X-ray emission (e.g., around minute 300 on 2014 June 2; minute 450 on 2014 July 4; minute 500 and 1300 on 2016 July 18; and minute 900 on 2017 July 25).

### 3.2. Cross-correlation

To quantify lags between the peaks of potentially associated activity in the X-ray and IR, we used the z-transform discrete correlation function (ZDCF; Alexander 1997), a tool that estimates the cross-correlation function without penalty for having a sparse or unevenly sampled light curve. We used the FORTRAN 95 implementation\(^{13}\) of both the ZDCF and the maximum likelihood function used to estimate the location of the ZDCF peak described in Alexander (2013).

The ZDCF is not sensitive to the relative amplitudes of the input light curves. This allows us to run the ZDCF without renormalizing the Chandra and Spitzer data. The ZDCF is, however, sensitive to the shape of the light curves, meaning that two flares with similar rise times, envelopes, and decay times will result in the ZDCF having a stronger correlation with less uncertainty in the time lag. In our case, this means that if a flare in the X-ray light curve has a significantly shorter duration than a coincident flare in the IR light curve, the peak in the ZDCF will be weaker and flatter than if they had the same duration, limiting the precision in the measured time lag.

For our analysis we used the 3.5 min binned IR light curves as the first input and the 300 s binned X-ray light curves as the second input. A positive time lag corresponds to a feature in the IR leading the X-ray, and a negative time lag corresponds to a feature in the IR lagging the X-ray. Figure 2 shows the results of running the ZDCF on the three epochs for which we detect significant X-ray activity (2016 July 12, 2016 July 18, and 2017 July 15).

To identify significant correlation peaks, we generated 10,000 Monte Carlo (MC) realizations of the X-ray and IR light curves and ran them through the ZDCF. These results are visualized as the blue envelope in the bottom row of Figure 2. The X-ray MC realizations were generated by adding Poisson noise to a smooth model containing the Gaussians fit to the

\(^{13}\) www.weizmann.ac.il/weizsites/tal/research/software/
flares. The IR realizations were generated from the 3.5 min binned light curves by perturbing each point by a random amount drawn from a Gaussian distribution with standard deviation equal to that of the mean of the \( \sim 25 \) BCD coadds in the bin. We confirmed that these distributions in the Spitzer light curve bins are normal and dominated by white noise (see Witzel et al. 2018). For comparison, we cross-correlated the 10,000 IR MC realizations with a separate set of 10,000 X-ray MC light curves containing only Poisson noise at the level of the quiescent flux (no flares, plotted in gray in Figure 2). Significant correlation peaks are those that rise above the gray MC “no flare” envelope. There are correlation peaks near zero time lag in all three epochs. Peaks in the 2016 July 18 and 2017 July 15 epochs are considered significant and the small peak in the 2016 July 12 epoch is considered marginally significant. The highest points in the peak of these correlations all occur at negative time lags, implying that flares in the X-ray may lead activity in the IR.

To robustly measure the uncertainty on the time lags, we located the position of the peak in all 10,000 results from ZDCF and defined confidence intervals based on the distribution of these 10,000 time lags. This method for estimating the uncertainties takes the errors on the light curve data into account. Table 2 compiles the time lags and the confidence intervals found in the MC analysis. All three epochs indicate the X-ray emission peaks approximately 10–20 min before the IR flux density peak.

As an alternative analysis to the ZDCF, we also calculated the cross-power spectra of each epoch, but did not find it to show any strong trends or constrain any relevant timing between the two data sets. We also compared the ZDCF to the ccf function in R, which yielded almost identical results, but had the disadvantage of requiring the two time series be sampled concurrently and at equally spaced points in time.

### 4. Discussion

#### 4.1. Are the X-Ray and IR Flares Truly Correlated?

We detected IR activity nearly coincident with every X-ray flare, though not every IR peak with flux density level \( \gtrsim 2 \) mJy had corresponding X-ray emission (perhaps because the weaker X-ray flares are hidden beneath the blanket of constant thermal emission). To investigate whether or not the observed X-ray and IR variabilities are truly associated, we consider the alternative possibility that the apparently correlated peaks in X-ray and IR emission are a chance association of peaks in typical X-ray and IR variability.

We generated 1000 simulations of each of our six Spitzer IR light curves by randomly drawing from the posterior of case 3 in Witzel et al. (2018), and then producing a random light curve for the given parameter set. The simulated light curves have the measurement noise properties of Spitzer, and they are distributed according to the log-normal flux density distribution determined for the M-band (4.5 \( \mu m \)). We ran the 1000 simulated light curves for each epoch through the ZDCF with the corresponding Chandra X-ray light curves as the second input. This resulted in a total of 6000 cross-correlations between Chandra X-ray light curves and simulated Spitzer IR light curves.

To test the probability of detecting a time lag similar to the one we measured by chance, we counted the occurrences of significant time lags measured between \( -20 \) and 0 min in our 6000 cross-correlations. A detection of a time lag within this window occurred in 138 of our 6000 instances. In other words, given Chandra X-ray data and a random Spitzer light curve simulating the typical IR variability of Sgr A*, a detection of a time lag between 0 and \(-20 \) min arose by chance 2.3% of the time. Increasing the window of coincidence to include positive time lags (\(-20 \) min to \(+20 \) min) resulted in 279 instances of coincidence, or a 4.7% chance occurrence rate. Figure 3 shows a distribution of time lags measured in all 6000 simulations. Since we detected a time lag of negative 10–20 min in all three of the real Spitzer/Chandra epochs containing significant X-ray flares, we consider this strong evidence that the X-ray and IR flares from Sgr A* are indeed physically connected.

#### 4.2. Comparisons to Previous Work and Flaring Models

There are many models in the literature that attempt to connect Sgr A*’s flares across multiple wavelengths, and particular attention has been paid to the near-simultaneous X-ray and IR peaks. Most models assume the IR peaks are caused by a population of non-thermal electrons emitting synchrotron radiation. Models then evoke synchrotron and/or IC radiation to connect these IR peaks to their corresponding X-ray flares. Figure 4 displays the time lags found by the cross-correlation analysis in this work in the context of lags (or lack thereof) reported in the literature with flares in both X-ray and IR. Most previous observations with overlapping X-ray and IR coverage reported no time lag between flares seen in both wavelengths, using language like “simultaneous to within \( x \) minutes” but quoting no uncertainties.
Figure 4. Time lags between IR and X-ray flares as reported in this work and in the literature. Plotted in black are the time lags from the three epochs in this work with significant X-ray and IR activity and their 68% confidence intervals determined from the distribution of 10,000 time lags measured from our MC realizations of our light curves. Plotted in solid gray are the results of the cross-correlation of the isolated sections of the 2017 July 15 light curve containing IR and X-ray data. The upper limit from Hornstein et al. (2007) indicates an X-ray flare whose peak occurred 36 min before IR observations began. Yusef-Zadeh et al. (2012) is the only work to report any time lag between the X-ray and IR with error bars. We re-analyze the seven flares presented in their work and plot the results of our re-analysis here. Five of these flares come from previously reported data sets (color coded as green, blue, magenta, and orange for Eckart et al. 2004, 2006a, 2006b; Dodds-Eden et al. 2009; Ponti et al. 2017). The second column lists the date the simultaneous data were taken. Column three lists the facilities with which the IR and X-ray data were collected. The signif

<table>
<thead>
<tr>
<th>Flare</th>
<th>Date</th>
<th>Orig. Paper</th>
<th>IR Facility</th>
<th>X-Ray Facility</th>
<th>YZ Time Lag (min)</th>
<th>Our Time Lag (min)</th>
</tr>
</thead>
<tbody>
<tr>
<td>A</td>
<td>2004 Jul 6/7</td>
<td>Eck+2006b</td>
<td>VLT (K band)</td>
<td>Chandra</td>
<td>7.2±1.2</td>
<td>7.2±1.2</td>
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<tr>
<td>B</td>
<td>2005 Jul 30</td>
<td>Eck+2008</td>
<td>VLT (K band)</td>
<td>Chandra</td>
<td>8.5±1.6</td>
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<td>2007 Apr 4</td>
<td>DE+2009</td>
<td>VLT (K band)</td>
<td>Chandra</td>
<td>-0.5±2.5</td>
<td>-0.5±2.5</td>
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<tr>
<td>D</td>
<td>2007 Apr 4</td>
<td>YZ+2009</td>
<td>HST/NICMOS</td>
<td>XMM-Newton</td>
<td>5.1±1.4</td>
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<tr>
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<td>2007 Apr 4</td>
<td>YZ+2009</td>
<td>HST/NICMOS</td>
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<td>F</td>
<td>2008 May 5</td>
<td>YZ+2012</td>
<td>VLT (K band)</td>
<td>Chandra</td>
<td>19.0±4.8</td>
<td>19.0±4.8</td>
</tr>
<tr>
<td>G</td>
<td>2008 Jul 26/27</td>
<td>YZ+2012</td>
<td>VLT (K band)</td>
<td>Chandra</td>
<td>14.6±3.4</td>
<td>14.6±3.4</td>
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Table 3

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<thead>
<tr>
<th>Flare</th>
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Note. Negative values mean X-ray leads IR. The first column labels the flares. The second and third columns list the date the simultaneous data were taken. Column four lists the original paper where the data were reported (Eck+2006b = Eckart et al. 2006a, Eck+2008 = Eckart et al. 2008, DE+2009 = Dodds-Eden et al. 2009, YZ+2009 = Yusef-Zadeh et al. 2009, YZ+2012 = Yusef-Zadeh et al. 2012). Columns four and five list the facilities with which the IR and X-ray data were collected. Column six lists the time lag and 1σ errors reported from the ZDCF analysis in Yusef-Zadeh et al. (2012). The last column lists the time lags and 68% confidence intervals we find from our Monte Carlo analysis (10,000 realizations).

4.2.1. Synchrotron

In models where both the IR and X-ray emission are produced through direct synchrotron radiation, the population of electrons responsible for the emission would have very high energies ($\gamma \gtrsim 10^5$), and the cooling timescale would be much shorter than the duration of a typical X-ray flare. This requires that the energized electrons have a sustained source over the duration of the flare (Baganoff et al. 2001; Markoff et al. 2001; Yuan et al. 2004; Ponti et al. 2017). Such processes often predict a simultaneous rise in the X-ray and IR flares, with the higher-energy X-rays fading faster as the electrons lose energy through synchrotron cooling and/or adiabatic expansion (e.g., Dodds-Eden et al. 2010). These predictions from pure synchrotron models are consistent with our tentative measurement of a time lag between the peak in the X-rays and a peak in the IR (equivalently, the X-ray flare rising with the IR but falling faster). Alternatively, if the IR emission is characterized by a number of “subpeaks” then the X-ray flare could be coincident with an individual subpeak. This could lead to an offset between the X-ray peak and the centroid of the broader IR envelope.

Two studies with observations of simultaneous X-ray and IR flares (Dodds-Eden et al. 2009; Ponti et al. 2017) found the X-ray flares to be consistent with synchrotron radiation. They reported particularly bright X-ray flares coincident with IR peaks where the IR rise had already begun when the X-ray flux rose, and the X-ray flare fell before the IR peak ended. Both of
these studies used measurements of the spectral indices of the flares to argue that a synchrotron emission mechanism with a cooling break was responsible for both the IR and X-ray flares. The simultaneous X-ray and IR peaks of Dodds-Eden et al. (2009) were again interpreted by Dodds-Eden et al. (2010) in the context of synchrotron emission due to accelerated electrons from magnetic reconnection, and Li et al. (2017) re-interpreted the data in the context of magnetic reconnection accelerating electrons in the coronal region rather than the main body of the accretion flow.

During a low flux density phase of activity, Witzel et al. (2018) measured the spectral index of the Spitzer IR measurements to be significantly redder than typically observed at high flux densities. Using this measurement as a constraint, they determined that the majority of NIR variability data is consistent with a variable spectral index that linearly depends on the logarithm of flux density. Both this determination of the variable IR spectral index for a large sample of variability data and the individual multi-wavelength flare analyses in Dodds-Eden et al. (2010) and Ponti et al. (2017) suggest synchrotron mechanisms, which also dominate the submm to IR, and at times reach energies high enough to cause X-ray flares. Indeed, the unchanging spectral index of X-ray variability at high energies is also consistent with a synchrotron scenario (Zhang et al. 2017).

4.2.2. Synchrotron Self-Compton

In the scenario where the electrons responsible for IR synchrotron radiation are accelerated to energies of $\gamma \sim 100$–1000, they could scatter the synchrotron IR photons up to X-ray energies through SSC, predicting near simultaneity for the flares. Electrons of these energies in a magnetic field with a strength typically assumed for the accretion flow around Sgr A* would have synchrotron cooling timescales on the order of hours, which fits the duration of the strongest X-ray flares observed in other works (e.g., Eckart et al. 2008; Dodds-Eden et al. 2009; Ponti et al. 2017, D. Haggard et al. 2018 in preparation). Several authors have modeled the flares as a SSC process (e.g., Markoff et al. 2001; Eckart et al. 2004, 2006a, 2008; Trap et al. 2011; Dibi et al. 2014).

Two of these studies (Eckart et al. 2006a, 2008) reported an X-ray and IR flare to be simultaneous to within 10 min. Eckart et al. (2006a) used this simultaneity and the suggestion that the IR flare spectrum is relatively red with a variable spectral index (e.g., Eisenhauer et al. 2005; Witzel et al. 2018) to argue in favor of the SSC picture. Eckart et al. (2008) observed a polarized IR flare with an X-ray counterpart. They also found that the flares fit the SSC picture of submm synchrotron photons being up-scattered to IR and X-ray wavelengths and discussed this in the context of a model involving a temporary
curves with 1500 s

Orange displays the Bayesian Blocks results with

$p_{0.1} = 0.1$. Bottom: light
curves with 1500 s (25 min) binning. Gray dashed lines indicate the same
interval analyzed by Yusef-Zadeh et al. (2012).

Figure 6. Sgr A*’ X-ray light curves extracted from Chandra ObsID 9169 (light
curve F) and ObsID 9173 (light curve G). Top: light curves with 300 s binning.

Eckart et al. (2004) reported a possible time lag of \(\sim 10\) min
between the first flares detected in simultaneous observations,
with the X-ray leading the IR. Due to the large binning of their
X-ray light curves they quote \(\sim 15\) min as a conservative upper
limit for any time lag. They describe the flares with a SSC
model in which IR flares are due to both synchrotron and
the up-scattering of synchrotron submm photons, and X-ray flares
are produced through the IC scattering of submm and IR
photons.

Finally, reporting no constraints on the simultaneity of the X-ray
and IR variability in their work, Trap et al. (2011) interpreted their X-ray/IR/submm data in the context of a
model involving an expanding plasmoid releasing synchrotron
submm and IR radiation and up-scattering to X-ray energies
through SSC processes. These models are often invoked to
explain the tentative time lags between the X-ray/IR flares and
submm or radio activity. The authors found that the simplest
version of this expanding plasmoid model did not adequately
explain their data.

4.2.3. Inverse Compton

In IC models connecting the IR and X-ray variability, IR-
synchrotron emitting electrons with energies \(\gamma \sim 100–1000\)
scatter submm seed photons up to X-ray energies, predicting a
potential lag in the X-rays relative to the IR outburst. This is
inconsistent with our measurement of a negative X-ray time
lag, which illustrates (at least tentatively) that the X-rays lead
the IR or, at the very least, rise simultaneously and fall faster.

Yusef-Zadeh et al. (2006, 2009) detected several faint X-ray
flares with IR counterparts, but quoted no uncertainty on any
time lag. Yusef-Zadeh et al. (2006) predicted that in the IC
scenario, the spectral index of an IR flare must be the same as
the correlated X-ray flare (i.e., \(\alpha \sim 0.6\)). Unfortunately, the
X-ray flare they observed did not contain enough counts to
provide any spectral information. In their campaign to observe
Sgr A* across many wavelengths (cm, mm, submm, IR, X-ray),
Yusef-Zadeh et al. (2009) also argued in favor of an IC model,
but one in which synchrotron IR photons are up-scattered to
X-ray wavelengths by electrons responsible for the radio and
submm emission of Sgr A*.

As the only other work to quote uncertainties on the time
lags found in their cross-correlations, Yusef-Zadeh et al. (2012)
re-analyzed archival data (Eckart et al. 2006a, 2008; Yusef-
Zadeh et al. 2006, 2009; Dodds-Eden et al. 2009) along with
previously unreported observations. They reported evidence for
the peak of the X-ray flares lagging the IR peaks with a time
delay ranging from a few to tens of minutes. Assuming there is
not more than one population of flares, this is in tension with
our finding of the X-rays tentatively leading the IR. These
authors also employed the ZDCF, and the time lags they find
are quoted in Table 3.

4.3. Re-analyzing Light Curves from the Literature

Due to the very low signal-to-noise in the individual cross-
correlation results and the large binning of some of the X-ray
data in Yusef-Zadeh et al. (2012), we suspect their reported \(1\sigma\)
error bars to be underestimated. For these reasons, we elected
to run their light curves through our cross-correlation and MC
analysis to verify their results, provide us with a consistent
comparison, and determine the effect that the signal-to-noise
ratio of a flare has on the cross-correlation.

Using the seven light curves presented in Yusef-Zadeh et al.
(2012), we performed a cross-correlation analysis identical to
the analysis we applied to our own light curves in Section 3.2.
We generated 10,000 MC instances of the light curves scaled in
proportion to the errors on the data and ran the ZDCF on them.
The mean of the resulting distribution of time lags was adopted
as our measurement, with uncertainties determined by the
interval within which 68% of the time lags fall. Table 3 labels
the flares A to G and displays the time lags and \(1\sigma\) errors
reported by Yusef-Zadeh et al. (2012) as well as the time lags
and uncertainties that we measure from our MC analysis.
Figure 5 re-plots the light curves found in their work (with the
exception of Flare C, which is not plotted by Yusef-Zadeh et al.
2012 but still included in their analysis) along with the results
of our MC analysis with the ZDCF. Figure 4 displays our
measurements for these seven flares in the context of this work
and the literature.

The differences between our results and those reported by
Yusef-Zadeh et al. (2012) are due to our MC analysis
incorporating the signal-to-noise of the input light curves. In
a light curve with low signal-to-noise, larger uncertainties on the data points will produce MC simulations that span a larger flux range, resulting in a broader range of time lags found between features in the light curves. Both methods employ the ZDCF, but while Yusef-Zadeh et al. (2012) quote the time lag of a single ZDCF/PLIKE run, we measure the time lag and estimate the uncertainties from 10,000 runs of the ZDCF.

In comparison to Yusef-Zadeh et al. (2012), this method estimates larger and more realistic uncertainties on the time lags for flares A, D, E, F, and G, similar uncertainties for flare B, and a smaller uncertainty for the time lag of flare C, the brightest simultaneous X-ray and IR flare observed to date (Dodds-Eden et al. 2009). Our analysis of flares A, C, D, and E found the time lag to be closer to zero or even negative compared to the lag reported by Yusef-Zadeh et al. (2012). The low signal-to-noise in the light curves of flares E, F and G is reflected in the poor constraints on a time-lag.

In the case of flares F and G, we question whether cross-correlating these X-ray light curves binned at 25 min with IR light curves binned at 2 min is meaningful. To test the variability of the light curves containing flares F and G, we opted to download the raw Chandra data and run the flare detection algorithm we used for our X-ray light curves (Bayesian Blocks algorithm; see Section 3.1). The data are found within Chandra ObsIDs 9169 and 9173. Figure 6 plots the results of running the Bayesian Blocks flare detection algorithm on light curves F and G, the lowest signal-to-noise light curves in Figure 5. In our Bayesian Blocks analysis of these light curves (orange lines in Figure 6), we do not detect any statistically significant X-ray peaks near the IR flares discussed by Yusef-Zadeh et al. (2012). The cross-correlations for these observations should therefore not be considered measurements of meaningful lags between IR and X-ray variability.

5. Conclusion

We have presented results based on new observations which are the longest simultaneous IR and X-ray observations of Sgr A* to date. These overlapping light curves offer the best tests of the connection between these two wavelengths, and provide a crucial probe of Sgr A*’s variable emission. We detect four X-ray flares (∼4× quiescence) and no “strong” X-ray flares (>10× quiescence) during the combined >100 hr that Chandra observed Sgr A*. The IR emission of Sgr A* showed peaks coinciding with the weak X-ray flares and also occurring at times when no X-ray flares are detectable. A cross-correlation analysis of all our simultaneous light curves suggests that the X-ray flares may lead the IR by approximately 10–20 min, but the 99.7% confidence intervals are still consistent with zero time-lag. This is in agreement with models that describe both the X-ray and IR flares as synchrotron emission originating from particle acceleration events involving magnetic reconnection and shocks in the accretion flow (e.g., see 4.1 of Dodds-Eden et al. 2010) and consistent with models that predict simultaneity of the flares through SSC processes. Our results are inconsistent with models invoking external populations of electrons through IC processes as described by Yusef-Zadeh et al. (2012), though it is not obvious that all X-ray/IR flares are produced by the same process.

It remains difficult to distinguish between the suggested flaring mechanisms connecting the X-ray and IR. Despite having the longest uninterrupted and simultaneous X-ray/IR data set of Sgr A* to date, we observed no bright X-ray flares during the 4+ days of observations reported here. Though this prevented us from gaining spectral information from the faint X-ray flares, future coordinated observations may catch significantly brighter simultaneous outbursts as has happened in the past (Dodds-Eden et al. 2009) and will certainly add to the growing statistical strength of the catalog of multi-wavelength flares. Previous observations have revealed substructure in the outbursts of both wavelengths (e.g., Dodds-Eden et al. 2009, D. Haggard et al. 2018, in preparation), and future observations of bright simultaneous outbursts could allow for a more detailed cross-correlation of sub-components in the X-ray and IR peaks.

In the immediate future, upcoming Spitzer/Chandra observations approved for the summer of 201914 may detect multiple bright flares, which may be key to constraining the time lag between the X-ray and the IR. In the longer term, a better understanding of the time-dependent emission from Sgr A* will allow for the characterization of the accretion physics around the black hole and inform the next generation of GRMHD simulations. Not only will long epochs of observations at multiple wavelengths be ideal data sets for distinguishing between semi-analytical flaring models, but Sgr A*’s variability will provide a strict benchmark for testing whether or not state-of-the-art simulations are probing the real physical scales of the turbulent accretion flow. Additionally, the multi-wavelength efforts coordinated with the Event Horizon Campaigns in April 2017 and 2018 hold promise for narrowing in on the physical processes that drive Sgr A*’s variability.

The scientific results reported in this article are based on observations made by the Chandra X-ray Observatory and the Spitzer Space Telescope. We thank the Chandra and Spitzer scheduling, data processing, and archive teams for making these observations possible. Support for this work was provided by the National Aeronautics and Space Administration through Chandra Award Number GO7-18135B issued by the Chandra X-ray Center, which is operated by the Smithsonian Astrophysical Observatory for and on behalf of the National Aeronautics Space Administration under contract NAS8-03060. D.H. acknowledges support from a Natural Sciences and Engineering Research Council of Canada (NSERC) Discovery Grant and a Fonds de recherche du Québec–Nature et Technologies (FRQNT) Nouveaux Chercheurs Grant. G.W. acknowledges support from the NSF grants AST-0909218, AST-1412615. J.L.H. acknowledges support from NASA Grant 80NSSC18K0416. G.P. acknowledges financial support from BMWi/DLR grants FKZ 50 OR 1604, 50 OR 1715 and 50 OR 1812.

Facilities: Spitzer/IRAC, Chandra/ACIS.
Software: CIAO Fruscione et al. (2006), NumPy (Jones et al. 2015), AstroPy (The Astropy Collaboration et al. 2018), Matplotlib (Hunter 2007), Bayesian Blocks (Williams et al. 2017), ZDCF (Alexander 2013).

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