Discovery of a planetary-mass companion within the gap of the transition disk around PDS 70


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ABSTRACT

Context. Young circumstellar disks are the birthplaces of planets. Their study is of prime interest to understand the physical and chemical conditions under which planet formation takes place. Only very few detections of planet candidates within these disks exist, and most of them are currently suspected to be disk features.

Aims. In this context, the transition disk around the young star PDS 70 is of particular interest, due to its large gap identified in previous observations, indicative of ongoing planet formation. We aim to search for the presence of an embedded young planet and search for disk structures that may be the result of disk–planet interactions and other evolutionary processes.

Methods. We analyse new and archival near-infrared images of the transition disk PDS 70 obtained with the VLTI/SPHERE, VLT/NaCo, and Gemini/NICI instruments in polarimetric differential imaging and angular differential imaging modes.

Results. We detect a point source within the gap of the disk at about 195 mas (~22 au) projected separation. The detection is confirmed at five different epochs, in three filter bands and using different instruments. The astrometry results in an object of bound nature, and continuing astrometry will allow us to test theoretical predictions regarding planet–disk interactions, planetary atmospheres, and evolutionary models.

Key words. stars: individual: PDS 70 – techniques: high angular resolution – protoplanetary disks – scattering – radiative transfer – planets and satellites: detection

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** The reduced images and datacubes are only available at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via http://cdsarc.u-strasbg.fr/viz-bin/qcat?J/A+A/617/A44
1. Introduction

More than two decades after the first detection of an extrasolar planet (Mayor & Queloz 1995), we are facing an extraordinary diversity of planetary system architectures (Winn & Fabrycky 2015). Exploring which of the properties of these systems are imprinted by the initial conditions of the disks and which develop through a variety of dynamical interactions is crucial for understanding the planet population. It is therefore of high importance to study planets and their environments at the stage during which these objects are formed.

Transition disks (TDs) are of key interest in this context, as many of them are believed to bear direct witness to the process of planet formation. These objects were initially identified by a significantly reduced near-infrared (NIR) excess compared to the median spectral energy distribution (SED) of young stars with disks (Strom et al. 1989). In the meantime, recent high-resolution imaging observations of TDs at different wavelengths have revealed large gaps (e.g. Thalmann et al. 2010; Andrews et al. 2011; Avenhaus et al. 2014b; Perez et al. 2015; Pohl et al. 2017b; Kendler et al. 2018), azimuthal asymmetries (e.g. Casassus et al. 2013; van der Marel et al. 2013; Pérez et al. 2014), spirals (e.g. Muto et al. 2012; Grady et al. 2013; Tang et al. 2017), and multiple rings (e.g. de Boer et al. 2016; Ginski et al. 2016; Andrews et al. 2016; van Boekel et al. 2017; Pohl et al. 2017a; Bertrang et al. 2018), as well as shadowed regions and brightness dips, some of them varying with time (e.g. Pinilla et al. 2015a; Stolker et al. 2017; Debes et al. 2017). In many cases, cavities and substructures are present within these young, gas-rich disks, which have often been interpreted as tracers of ongoining planet formation and are suspected to originate from planet–disk interactions (see Espaillat et al. 2014 for a review).

The characterisation of TDs is therefore of prime interest to understand the physical and chemical conditions under which planet formation takes place. Observations of forming planets within those disks are extremely challenging as the disk often obstructs the planet, requiring observations at high contrast and angular resolution. Such detections have been reported for a few targets only, namely HD 100546 (Quanz et al. 2013a, 2015; Brittain et al. 2014; Currie et al. 2015), LkCa15 (Kraus & Ireland 2012; Sallum et al. 2015), HD 169142 (Quanz et al. 2013b; Biller et al. 2014; Reggiani et al. 2014), and MWC 758 (Reggiani et al. 2018), but most of the detections are presently being challenged (e.g. Rameau et al. 2017; Sissa et al. 2017; Ligi et al. 2018). The point sources identified as planet candidates in high-contrast imaging observations with the angular differential technique could be confused with brightness spots in asymmetric disks.

The aim of this paper is to study the pre-main sequence star PDS 70 (V* V1032 Cen) with high-contrast imaging. PDS 70 could be confused with brightness spots in asymmetric disks. Such detections have been reported for a few targets only, namely HD 100546 (Quanz et al. 2013a, 2015; Brittain et al. 2014; Currie et al. 2015), LkCa15 (Kraus & Ireland 2012; Sallum et al. 2015), HD 169142 (Quanz et al. 2013b; Biller et al. 2014; Reggiani et al. 2014), and MWC 758 (Reggiani et al. 2018), but most of the detections are presently being challenged (e.g. Rameau et al. 2017; Sissa et al. 2017; Ligi et al. 2018). The point sources identified as planet candidates in high-contrast imaging observations with the angular differential technique could be confused with brightness spots in asymmetric disks.

1 Following de Juan Ovelar et al. (2013), the NIR cavity wall is defined as the radial position where the flux equals half the value between the cavity wall. This deflection is an extraordinarily difficult radial position where the flux equals half the value between the minimum flux at the bottom of the gap and the flux maximum at the wall.
segregation of dust grains makes PDS 70 a prime candidate for hosting planets that are carving the gap.

In this paper, we present an extensive dataset on PDS 70 using the high-contrast imager Spectro-Polarimetric High-contrast Exoplanet REsearch (SPHERE; Beuzit et al. 2008), complemented with datasets obtained with VLT/NaCo and Gemini/NICI. Our observations include both angular differential imaging (ADI) and polarimetric observations at multiple wavelengths and epochs, from the optical to the NIR, covering a time period of more than four years. We report on the robust detection of a point source within the gap of the disk, which is interpreted as a planetary-mass companion. Furthermore, we detect for the first time scattered light emerging from the inner disk. We analyse the well-known outer disk with respect to its morphological appearance and investigate its structure with a radiative transfer (RT) model.

This paper is structured as follows: the observational setup and reduction strategy of the different datasets are described in Sect. 2. The analysis of the disk is presented in Sect. 3, and Sect. 4 describes our modelling efforts. Section 5 is dedicated to investigating any (physical) polarisation of the central unresolved source, consisting of the stellar photosphere plus thermal emission and/or scattered light from the inner disk, if the latter is present, is unpolarised. This implies that they cannot differentiate between intrinsic polarisation of the central source and the IP. In this case, correcting for the IP would mean subtracting the IP. In this case, correcting for the IP would mean subtracting the IP.


during the nights of March 25, 2016, and July 31, 2017, we obtained eight polarimetric cycles with two long exposures (64 s) per HWP position. Near the end of the sequence, the seeing degraded, and the telescope guiding was lost such that we had to discard the last polarimetric cycle. Before and after each sequence, we obtained a sequence of short, non-saturated images of the star outside the coronagraphic mask. In addition, to enable an accurate determination of the stellar position behind the coronagraph, we took calibration images in which four satellite spots in a square pattern centered on the star were generated by introducing a sinusoidal modulation onto the deformable mirror.

The second epoch (July 31, 2017) was obtained without a coronagraph in order to provide access to the innermost disk regions. We therefore chose very short exposures (2 s) to prevent the star from saturating the detector. The observations consisted of seven polarimetric cycles, with 20 exposures taken per HWP position. We aligned the images by fitting two-dimensional (2D) Gaussians to the target star in each frame.

The first step of data reduction consisted of dark subtraction, flatfielding, interpolation of bad pixels, and recentering of the frames. The rotation of the HWP to four different angles allowed us to determine a set of linear polarisation components \( Q, U, Q', U' \), from which we then obtained the clean Stokes \( Q \) and \( U \) frames using the double-difference method (Canovas et al. 2011), and the normalised double-difference method (van Holstein et al., in prep.) for the non-coronagraphic and coronagraphic dataset, respectively.

One of the most important steps in the reduction of polarimetric data is the subtraction of instrumental polarisation (IP). Standard techniques to correct for the IP, as described by Canovas et al. (2011), estimate the IP directly from the data itself. Here, the main assumption is that the central unresolved source, consisting of the stellar photosphere plus thermal emission and/or scattered light from the inner disk, if the latter is present, is unpolarised. This implies that they cannot differentiate between intrinsic polarisation of the central source and the IP. In this case, correcting for the IP would mean subtracting any (physical) polarisation of the central unresolved source. In contrast to that, to correct our measurements for the IP and cross-talk effects, we applied the detailed Mueller matrix model.

Table 2. Observing log of data used within this study.

<table>
<thead>
<tr>
<th>Date</th>
<th>ProgID</th>
<th>Instrument</th>
<th>Mode(^a)</th>
<th>Filter</th>
<th>R(^b)</th>
<th>Coronagraph</th>
<th>( \Delta \theta )</th>
<th>DIT(^c) (s)</th>
<th>( \Delta t ) (min)</th>
<th>( \epsilon ^{(*)} )</th>
</tr>
</thead>
<tbody>
<tr>
<td>2012-03-31</td>
<td>GS-2012A-C-3</td>
<td>NICI(^b)</td>
<td>ADI</td>
<td>L'</td>
<td>–</td>
<td>–</td>
<td>–</td>
<td>99.4(^{\circ})</td>
<td>0.76</td>
<td>118</td>
</tr>
<tr>
<td>2015-05-03</td>
<td>095.C-0298(A)</td>
<td>IRDIS</td>
<td>ADI</td>
<td>H2H3</td>
<td>–</td>
<td>–</td>
<td>–</td>
<td>52.0(^{\circ})</td>
<td>64.0</td>
<td>70</td>
</tr>
<tr>
<td>2015-05-03</td>
<td>095.C-0298(A)</td>
<td>IFS</td>
<td>ADI</td>
<td>YJ</td>
<td>54</td>
<td>ALC_YJH_S</td>
<td>–</td>
<td>52.0(^{\circ})</td>
<td>64.0</td>
<td>70</td>
</tr>
<tr>
<td>2015-05-31</td>
<td>095.C-0298(B)</td>
<td>IRDIS</td>
<td>ADI</td>
<td>H2H3</td>
<td>–</td>
<td>ALC_YJH_S</td>
<td>–</td>
<td>40.8(^{\circ})</td>
<td>64.0</td>
<td>70</td>
</tr>
<tr>
<td>2015-05-31</td>
<td>095.C-0298(B)</td>
<td>IFS</td>
<td>ADI</td>
<td>YJ</td>
<td>54</td>
<td>ALC_YJH_S</td>
<td>–</td>
<td>40.8(^{\circ})</td>
<td>64.0</td>
<td>70</td>
</tr>
<tr>
<td>2015-07-09</td>
<td>095.C-0404(A)</td>
<td>ZIMPOL</td>
<td>PDI</td>
<td>VBB</td>
<td>–</td>
<td>–</td>
<td>–</td>
<td>40.0</td>
<td>114</td>
<td>1.1</td>
</tr>
<tr>
<td>2016-03-25</td>
<td>096.C-0333(A)</td>
<td>IRDIS</td>
<td>PDI</td>
<td>–</td>
<td>ALC_Y1</td>
<td>–</td>
<td>–</td>
<td>64.0</td>
<td>94</td>
<td>1.9</td>
</tr>
<tr>
<td>2016-05-14</td>
<td>097.C-1001(A)</td>
<td>IRDIS</td>
<td>ADI</td>
<td>K1K2</td>
<td>30</td>
<td>–</td>
<td>16.9(^{\circ})</td>
<td>0.837</td>
<td>22</td>
<td>1.0</td>
</tr>
<tr>
<td>2016-05-14</td>
<td>097.C-1001(A)</td>
<td>IFS</td>
<td>ADI</td>
<td>YJH</td>
<td>–</td>
<td>–</td>
<td>16.9(^{\circ})</td>
<td>4.0</td>
<td>23</td>
<td>1.0</td>
</tr>
<tr>
<td>2016-06-01</td>
<td>097.C-0206(A)</td>
<td>NaCo</td>
<td>ADI</td>
<td>L'</td>
<td>–</td>
<td>–</td>
<td>83.7(^{\circ})</td>
<td>0.2</td>
<td>155</td>
<td>0.5</td>
</tr>
<tr>
<td>2017-07-31</td>
<td>099.C-0891(A)</td>
<td>IRDIS</td>
<td>PDI</td>
<td>J</td>
<td>–</td>
<td>–</td>
<td>2.0</td>
<td>36</td>
<td>0.7</td>
<td></td>
</tr>
</tbody>
</table>

Notes. \(^a\)Observing mode: angular differential imaging (ADI) or polarimetric differential imaging (PDI); \(^b\)spectral resolution; \(^c\)total field rotation, after frame selection; \(^d\)detector integration time; \(^e\)total time on target (including overheads); \(^f\)mean MASS/DIMM seeing; \(^g\)archival data, published in Hashimoto et al. (2012).
This method models the complete optical path that the beam traverses on its way from entering the telescope to the detector, and has already been applied in the analysis of several circumstellar disks observed with SPHERE (Pohl et al. 2017b; Canovas et al. 2018). The incident Stokes Q and U images are recovered by solving a set of equations describing every measurement of Q and U for each pixel individually. These Stokes Q and U images correspond to the images as they enter the telescope (star and disk convolved with telescope point spread function (PSF) and noise). With this method, we can therefore make a model prediction of the IP, and correct for the IP alone without subtracting the polarised signal from the central source. After the correction for the IP, any remaining polarised signal at the location of the central source, induced by, for example, unresolved material close to the star, such as an inner disk, would then become visible in the form of a central butterfly pattern in the Q or U image (see definition of Q and U below) and can affect the signal of the outer disk. This leftover signal can then be chosen to be subtracted following the method by Canovas et al. (2011). On the other hand, it can also be kept, if one desires to study the unresolved inner disk (see van Holstein et al., in prep.). We prepared two reductions, one for which we subtracted the central source polarisation (allowing us to study the morphology of the outer disk; see Sect. 3.1), and one where we did not subtract it (for analysing the inner disk; see Sect. 3.2). We corrected our data for true north and accounted for the instrument anamorphism (Maire et al. 2016). We measured a PSF full width at half maximum (FWHM) of ~52 and ~49 milli-arcseconds (mas) on the unsaturated flux frames of the coronagraphic observations, and of the total intensity frame for the non-coronagraphic data, respectively.

We transform our Stokes images into polar coordinates (Q, U), according to the following definition from Schmid et al. (2006):

\[
Q_\phi = +Q \cos(2\phi) + U \sin(2\phi), \\
U_\phi = -Q \sin(2\phi) + U \cos(2\phi),
\]

(1)

where \(\phi\) denotes the position angle measured east of north with respect to the position of the star. In this formulation, a positive signal in the \(Q_\phi\) image corresponds to a signal that is linearly polarised in azimuthal direction, whereas radially polarised light causes a negative signal in \(Q_\phi\). Any signal polarised in the direction \(\pm 45^\circ\) with respect to the radial direction is contained in \(U_\phi\). Therefore, in the case of low or mildly inclined disks, almost all scattered light is expected to be contained as positive signal in \(Q_\phi\). However, due to the non-negligible inclination of the disk around PDS 70 (49.7°; Hashimoto et al. 2012), and because the disk is optically thick such that multiple scattering processes cannot be neglected, we expect some physical signal in \(U_\phi\) (Canovas et al. 2015). \(U_\phi\) can therefore only be considered as an upper limit to the noise level.

### 2.2. ZIMPOL polarimetric observations (VBB band)

PDS 70 was observed during the night of July 9, 2015, with the SPHERE/ZIMPOL instrument (Thalmann et al. 2008). These non-coronagraphic observations were performed in the Slow-Polarmetry readout mode (P2), using the Very Broad Band (VBB, 590–881 nm) filter, which covers the wavelength range from the R- to the I-band. Especially in the second half of the sequence, the conditions were poor (seeing above 1″), resulting in a PSF FWHM of ~159 mas. Since the detailed Mueller matrix model for the correction of the instrumental polarisation by van Holstein et al. (in prep.) only applies to the IRDIS data, our correction for IP effects was performed by equalising the ordinary and extraordinary beams for each frame, as described by Avenhaus et al. (2014b). We interpolated two pixel columns in the image that were affected by readout problems.

#### 2.3. IRDISF angular and spectral differential imaging observations (Y–H-band)

During the guaranteed time observations (GTO) of the SPHERE consortium, PDS 70 was observed twice within the SHINE (SpHERE InfInfrared survey for Exoplanets; Chauvin et al. 2017a) program on the nights of May 3, 2015, and May 31, 2015. The data were taken in the IRDIS observing mode, with IRDIS working in the dual-band imaging mode (Vigan et al. 2010) making use of the H2H3 narrow-band filter pair (\(\lambda_{\text{H2}} = 1.593 \mu m\), \(\lambda_{\text{H3}} = 1.667 \mu m\)), and with IFS operating simultaneously in the wavelength range of the Y and J broadband filter (0.95–1.33 \(\mu m\)) with a spectral resolution of \(R \sim 50\) (Claudi et al. 2008). We made use of the N_ALC_YJH_S coronagraphic mask (apodized Lyot, diameter 185 mas). The observations were performed in pupil-tracking mode to allow for ADI (Marois et al. 2006). Before and after the sequences, we obtained calibration frames for measuring the location of the star behind the coronagraph and unsaturated images of the star without coronagraph for photometric calibration. Each sequence consisted of 64 exposures, from which we removed 10 and 14 bad quality frames for the May 3 and May 31 epochs, respectively.

After basic reduction steps applied to the IRDIS data (flat fielding, bad-pixel correction, sky subtraction, frame registration, frame selection to remove poor-quality frames based on the frame-to-frame photometric variability of the background object north from the star, correction of the instrument distortion, and correction of the flux calibration for the neutral density filter transmission), we used several different strategies to model and subtract the stellar speckle pattern. First of all, we applied the classical ADI (cADI) method (Marois et al. 2006). We then ran a sPCA (smart Principal Component Analysis) algorithm, adapted from Absil et al. (2013), which itself is based on the KLIP algorithm of Soummer et al. (2012). Further, we used the ANDROMEDA package (Cantalloube et al. 2015) which applies a statistical approach to search for point sources. Further, we applied the PCA and TLOCI (Template Locally Optimized Combination of Images; Marois et al. 2014) approach using the SpeCal implementation (Galicher et al. 2018). The main difference between our two PCA reductions is that the SpeCal PCA implementation does not select the frames for building the PCA library, and is therefore considered to be more aggressive than the former one. Finally, to obtain a non-ADI view of the disk morphology, we simply derotated and stacked the frames and applied a Laplacian filter, here referred to as ‘gradient reduction’, in order to enhance low spatial frequencies (i.e. fine disk structures) in the image. However, since this reduction is not flux conservative, we used it only for a qualitative analysis of the outer disk structures.

Concerning the IFS data, the basic data reduction was performed using the Data Reduction and Handling software (Pavlov et al. 2008) and custom IDL routines adapted from Vigan et al. (2015) and Mesa et al. (2015). We post-processed the data using the cADI, ANDROMEDA, and PCA-SpeCal algorithms. The IRDIS and IFS data were astrometrically calibrated following the methods in Maire et al. (2016).
right column

Fig. 1. SPHERE PDI observations. Top row: $Q_\phi$ images; bottom row: $U_\phi$ images. The left and middle columns correspond to the IRDIS J-band observations, taken with coronagraph (March 25, 2016) and without coronagraph (July 31, 2017), after correcting for the instrumental polarisation (IP) and subtraction of the central source polarisation. The right column presents the ZIMPOL observations (July 9, 2015). The colour scale was chosen arbitrarily but is the same for each pair of $Q_\phi$ and $U_\phi$. We note that negative values are saturated to enable a better visual contrast. North is up and east is to the left.

2.4. IRDIS_EXT angular and spectral differential imaging observations (Y-K-band)

The star was also observed with SPHERE in IRDIS_EXT mode during the night of May 14, 2016. In this mode, the IRDIS K1K2 narrow-band filter pair is used ($\lambda K_1 = 2.11 \mu m$, $\lambda K_2 = 2.25 \mu m$), whereas IFS is operating in the wavelength range of the YJH broad-band filter (0.97–1.64 $\mu m$) at a spectral resolution of $R \sim 30$. No coronagraph was used during the observations, and short detector integration times (DIT: 0.837 s for IRDIS, 4 s for IFS) were chosen to prevent any saturation of the detector. As the observing conditions were relatively stable during the sequence, no frame selection was performed. The data reduction strategy was identical to the one in Sect. 2.3, and we post-processed the data with sPCA, PCA-SpeCal, TLOCI, and ANDROMEDA.

2.5. NaCo angular differential imaging observations (L'-band)

We also made use of observations of PDS 70 carried out with VLT/NaCo within the ISPY (Imaging Survey for Planets around Young stars; Launhardt et al., in prep.) GTO program on the night of June 1, 2016. The sequence was obtained in pupil-stabilised mode, making use of the L'-band filter (3.8 $\mu m$) and the 27.1 mas pixel$^{-1}$ scale. No coronagraph was employed. A DIT of 0.2 s was used to prevent any saturation during the sequence. The seeing was rather stable during the observations (average seeing 0.5$''$), and seven frames were rejected. The data reduction and post-processing strategy was identical to the one in Sect. 2.3.

2.6. Archival NICI angular differential imaging observations (L'-band)

Finally, we used archival Gemini/NICI non-coronographic data taken on March 31, 2012 using the L' filter. The data together with the observing conditions and strategy were published in Hashimoto et al. (2012). As for the NaCo observations, no coronagraph was applied. Thirty-seven out of 144 frames were sorted out. We re-reduced and post-processed the data with the same approach as presented in Sect. 2.3.

3. Disk analysis

3.1. The outer disk in polarised scattered light

In total, we have three observations of the disk taken in PDI mode: two observations in IRDIS J-band (coronagraphic and non-coronagraphic), as well as one observation with the ZIMPOL VBB filter. Figure 1 shows the respective $Q_\phi$ and $U_\phi$ images. As expected from previous observations, we detect the disk in all three datasets as an elliptical ring. Because the setup of the non-coronographic IRDIS data was not optimal to detect the outer disk (very short DITs, hence lower S/N), we focus here on the IRDIS coronagraphic dataset, and use the reduction where we subtracted the central source polarisation (see Sect. 2.1). The $Q_\phi$ images show evidence of residual signal that is contained within a region smaller than $\sim$12 pixels (17 au). In both the IRDIS and ZIMPOL data, there is some signal in $U_\phi$, but this signal is mostly detected at radii inward of the outer disk and does therefore not impact our analysis of the outer disk.
In Fig. 2 we present the azimuthally averaged radial brightness profile after deprojecting the disk (using a position angle of 158.6° and an inclination of 49.7°, as determined by Hashimoto et al. 2012). To avoid effects from small-scale noise, we smoothed our images with a small Gaussian kernel with a FWHM of 50% of the measured image resolution. The uncertainties were computed from the standard deviation of $U_\phi$ in the corresponding radial bins, divided by the square-root of the number of resolution elements fitting in that bin. We remind the reader that especially towards the region close to the star, $U_\phi$ might contain physical signals, and therefore the error bars only indicate an upper limit for the noise. The mean radius of the disk brightness peak is determined to be \(\sim 54 \text{ au}\). The outer disk ring appears wider in the VBB profile than in the $J$-band profile. We note however that the PSF FWHM was about three times larger during the VBB observations than during the $J$-band observations. Within \(\sim 25 \text{ au}\), the profile rises towards the centre which is associated with emission from the inner disk (see Sect. 3.2). The slope is much stronger in the $J$-band than in the VBB-band. This can be explained by the fact that the regions close to the star in PDI observations are affected by PSF damping effects which are more strongly pronounced at shorter wavelengths where the Strehl ratio is significantly lower (Avenhaus et al. 2014a, 2018).

In Fig. 3, we plot the azimuthal profile of the deprojected $Q_\phi$ images. The profile was derived by averaging over azimuthal bins with a size of 8° between 35 and 70 au in radial direction. We note an azimuthal brightness modulation for both datasets. In each of them, the east side (corresponding to PA \(\leq 160°\)) appears on average brighter than the west side. By averaging the brightness in azimuthal bins of size \(\pm 20°\) around the PA of the semi-major axes, we derive a brightness contrast of \(\sim 1.8\) on the brightness maxima along the semi-major axis in the SE and NW in the VBB and \(\sim 0.8\) in the $J$-band images. We also found a disk brightness ratio of \(\sim 3.1\) and \(\sim 1.4\) along the minor axes, in the NE and SW for the VBB and $J$-band, respectively. Therefore, the brightness ratio shows the same trend along the minor axis in the VBB and $J$-band, with the NE side being brighter than the SW side in both the VBB band and the $J$-band. However, along the major axis, the trend is opposite between the two bands, with the SE side being brighter than the NW side in the VBB-band, but fainter than the NW side in the $J$-band. By fitting an ellipse to the disk, Hashimoto et al. (2012) showed that the center of such an ellipse is offset towards the east side with respect to the star. This is due to the flaring geometry of the disk and implies that the east side of the disk corresponds to the far side.

### 3.2. Detection of the inner disk in polarised light

The presence of an optically thick inner disk in the innermost few astronomical units was predicted from SED fitting of the NIR excess (Hashimoto et al. 2012; Dong et al. 2012). The ALMA observations by Long et al. (2018; beam size of \(0.19'' \times 0.15''\)) detected thermal emission from an inner disk component, which appears to be depleted of large grains. However, this inner disk component has not been directly detected in scattered light until now.

For the analysis of the inner region, we mainly rely on the IRDIS polarimetric non-coronographic dataset, as it allows us to probe regions closer to the star than does the coronagraphic dataset. As seen in Sect. 3.1, the outer disk is well recovered in the IRDIS PDI images by subtracting the remaining central source polarisation. To study the inner disk region, we therefore choose to focus on the dataset corrected for the IP, but without subtracting the central source polarisation. The corresponding $Q_\phi$ and $U_\phi$ images (coronagraphic and non-coronagraphic) are presented in Fig. 4 (left and middle columns).

We detect a strong butterfly pattern in both $Q_\phi$ and $U_\phi$. This pattern affects the outer disk. We determine a central source polarisation degree of \(0.75 \pm 0.2\%\) and an angle of linear polarisation of \(66 \pm 11°\). Hence, the polarisation direction in the inner region is approximately perpendicular to the disk semi-major axis. This butterfly pattern can be explained by the fact that we detect signal from the inner disk which is unresolved. If the inner disk is oriented in the same direction as the outer disk (PA of \(\sim 158.6°\); Hashimoto et al. 2012), the majority of this signal will be polarised in perpendicular direction to the semi-major axis, because the polarisation degree is highest along the semi-major axis (scattering angle of \(\sim 90°\)). This signal gets smeared out when convolved with the instrument PSF, causing
Fig. 4. IRDIS coronagraphic (left column) and non-coronagraphic (middle column) PDI images, corrected for the instrumental polarisation but without subtracting the central source polarisation. The right column shows the model image as a comparison, including an inner disk with an outer radius of 2 au. The first row corresponds to the $Q_0$ images, the second row to the $U_0$ images. North is up and east is to the left.

Fig. 5. SHINE IRDIS observations of May 31, 2015, showing the cADI (left panel), TLOCI (middle panel), and gradient reduction (right panel). North is up and east is to the left.

a large butterfly pattern in the resulting $Q_0$ and $U_0$ images. The fact that the unresolved signal is polarised perpendicular to the disk semi-major axis implies that the inner and outer disks are approximately aligned. We note that although we cannot determine the inclination of the inner disk, we can infer that it must be larger than zero, because had the inner disk been seen face-on, its signal would have cancelled out due to axial symmetry. Further, the absence of shadows on the outer disk indicates that the inclination of the inner disk should be similar to that of the outer disk.

We note that even after subtracting the central source polarisation from the non-coronagraphic (as well as the coronagraphic) data, a signal in $Q_0$ is detected inside about 17 au. The subtraction of the central source polarisation removes almost all signal from an unresolved source, and the leftover signal could originate from a partially resolved inner disk (larger than the resolution element). We therefore suspect that the disk is slightly larger than the resolution element, but not extending farther than 17 au, because otherwise we would have detected larger residuals after subtracting the central source polarisation.

We note that the polarisation of the central source is unlikely to be caused by interstellar dust due to the low extinction measured towards PDS 70 ($A_V = 0.05^{+0.03}_{-0.05}$ mag; Müller et al. 2018). We further study the inner disk characteristics by comparison with our RT models (Sect. 4).

3.3. The disk in IRDIS angular differential imaging

We considered the SHINE IRDIS ADI observations for the characterisation of the disk. In comparison to the observations presented in the previous section, they trace the total intensity and were taken at a longer wavelength ($H$-band). However, whereas the ADI technique is optimised for detecting point sources (see Sect. 5), when applied to extended sources as disks, the images suffer from self-subtraction effects. Figure 5 shows
the resulting cADI image, TLOCI, and gradient images, and the sPCA, PCA-SpeCal images are shown in Fig. C.1. In all reductions and epochs, the disk’s west side is clearly detected. As previously mentioned, this side corresponds to the near side of the disk. The disk’s extension along the semi-major axis is larger in the ADI images than in polarised light. When overlaying the two images, it appears that the signal we see in ADI in fact corresponds to the outer region of the disk as seen in polarised light (see Fig. C.1). Furthermore, in some reductions, the inner edge of the disk’s far side appears to be detected. This is especially true in the gradient image, but also in the cADI reduction.

The gradient image exhibits many circular artifacts while the disk signal deviates from this circular symmetry, and the inner edge of the outer disk is well detected at all position angles. The cADI, TLOCI, and both PCA reductions show a feature near the disk signal deviates from this circular symmetry, and the inner true in the gradient image, but also in the cADI reduction. Of the disk’s far side appears to be detected. This is especially responds to the in the ADI images than in polarised light. When overlaying the previously mentioned, this side corresponds to the near side of the disk’s far side of 60 au, inwards of which we multiplied the surface density by a Gaussian profile parametrized by a standard deviation of 8 au. The surface density of the inner disk is multiplied with a depletion factor $\delta_{\text{disk}}$, with respect to that of the outer disk. The general shape of the surface density is plotted in Fig. 6 (left panel). We consider two grain-size distributions (small and large ways, we assign a lower scale height (2 au at a radial distance of 100 au) to the disk of big grains, for both the inner and outer disk, we considered an outer disk radius of 60 au, inwards of which we multiplied the surface density with a Gaussian profile parametrized by a standard deviation of 8 au. The surface density of the inner disk is multiplied with a depletion factor $\delta_{\text{disk}}$ with respect to that of the outer disk. The general shape of the surface density is plotted in Fig. 6 (left panel). We consider two grain-size distributions (small and large grains), whose number density follows a power law as a function of the grain size $a$ with $n(a) \propto a^{-3.5}$. The population of small grains ranges from 0.001 to 0.15 $\mu$m, and the large grains from 0.15 to 1000 $\mu$m in size. To mimic dust settling in a simplified way, we assign a lower scale height (2 au at a radial distance of 100 au) to the disk of big grains, for both the inner and outer disk parts. The relative mass fraction of small to large grains is 1/31. We determine the optical properties for spherical, compact dust grains according to the Mie theory using the BHMIE code (Böhren & Huffman 1983). A total dust mass of $3.0 \times 10^{-5}$ $M_\odot$ was used. The dust mixture is composed of 70% astronomical silicates (Draine 2003), and 30% amorphous carbon grains (Zubko et al. 1996). The opacity mixture was generated according to the Bruggeman mixing rule. The scattering mode in RADMC-3D was set to anisotropic scattering with full treatment.

4. Radiative transfer modeling

4.1. Model setup

To compare our multiwavelength observations with a physical model, we built a RT model, where the basic parametric approach by Dong et al. (2012) is taken as a starting point. We used the RT Monte Carlo code RADMC-3D (Dullemond et al. 2012). Our aim is to find a plausible model which reproduces our observations and the SED. We note that we are not looking for a globally best fitting model. RADMC-3D computes the thermal structure of the disk and produces images in scattered polarised light and total intensity by ray-tracing at any desired wavelength. Our grid has an inner radius of 0.04 au and an outer radius of 120 au. The surface density is proportional to $r^{-1}$ and is truncated by a tapered edge with a characteristic radius $R_c$. We radially parametrize the dust surface density by:

$$\Sigma_{\text{disk}}(r) = \Sigma_0 R_c r \exp \left( -\frac{r}{R_c} \right),$$

where $\Sigma_0$ scales the amount of dust contained within the disk.

We assume a Gaussian distribution profile in the vertical direction and parametrize the density distribution in terms of the scale height $h$ as follows:

$$\rho(r, z) = \frac{\Sigma(r)}{\sqrt{2\pi}h} \exp(-z^2/2h^2).$$

The disk is assumed to be flared with a constant power law index $\beta$, such that the radial dependence of the scale height can be written as

$$h(r) = h_{100} \times \left( \frac{r}{100 \text{ au}} \right)^\beta,$$

where $h_{100}$ is the scale height at 100 au. To mimic the gap in the disk, we heavily deplete (by a factor of $10^{-15}$) the surface density between the outer radius of the inner disk and the inner radius of the outer disk. To ensure a smooth transition from the gap to the outer disk, we considered an outer disk radius of 60 au, inwards of which we multiplied the surface density with a Gaussian profile parametrized by a standard deviation of 8 au. The surface density of the inner disk is multiplied with a depletion factor $\delta_{\text{disk}}$ with respect to that of the outer disk. The general shape of the surface density is plotted in Fig. 6 (left panel). We consider two grain-size distributions (small and large grains), whose number density follows a power law as a function of the grain size $a$ with $n(a) \propto a^{-3.5}$. The population of small grains ranges from 0.001 to 0.15 $\mu$m, and the large grains from 0.15 to 1000 $\mu$m in size. To mimic dust settling in a simplified way, we assign a lower scale height (2 au at a radial distance of 100 au) to the disk of big grains, for both the inner and outer disk parts. The relative mass fraction of small to large grains is 1/31. We determine the optical properties for spherical, compact dust grains according to the Mie theory using the BHMIE code (Böhren & Huffman 1983). A total dust mass of $3.0 \times 10^{-5}$ $M_\odot$ was used. The dust mixture is composed of 70% astronomical silicates (Draine 2003), and 30% amorphous carbon grains (Zubko et al. 1996). The opacity mixture was generated according to the Bruggeman mixing rule. The scattering mode in RADMC-3D was set to anisotropic scattering with full treatment.
of polarisation. We computed the Stokes $Q$ and $U$ frames, as well as images in total intensity at the wavelengths of interest (0.7 and 1.25 $\mu$m), and convolved our images with the total intensity frames obtained during the corresponding observations. We then computed the $Q_{\rho}$ and $U_{\rho}$ frames according to Eq. (1) as well as a synthetic SED. Table 3 summarises the parameters used for our model.

![Fig. 7. Linear polarisation degrees for silicates of different (semi-mono-dispersive) grain sizes (left panel), and their phase function (right panel, normalised to $Z_{\perp}(0)$). The curves were computed for an observing wavelength of 0.7 $\mu$m.]

### Table 3. Parameters for our RT model.

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Inner disk</th>
<th>Outer disk</th>
</tr>
</thead>
<tbody>
<tr>
<td>$R_{\text{in}}$ [au]</td>
<td>0.04</td>
<td>60</td>
</tr>
<tr>
<td>$R_{\text{out}}$ [au]</td>
<td>[2.4,8,12,16,20]</td>
<td>120</td>
</tr>
<tr>
<td>$R_e$ [au]</td>
<td>40</td>
<td>40</td>
</tr>
<tr>
<td>$h_{100,\text{small}}$ [au]</td>
<td>13</td>
<td>13</td>
</tr>
<tr>
<td>$h_{100,\text{big}}$ [au]</td>
<td>2</td>
<td>2</td>
</tr>
<tr>
<td>$\beta$</td>
<td>1.25</td>
<td>1.25</td>
</tr>
<tr>
<td>$\delta_{\text{disk}}$</td>
<td>[0.05,0.1,0.25,1.0]</td>
<td>1</td>
</tr>
<tr>
<td>Disk inclination $i$ [°]</td>
<td>49.7</td>
<td>49.7</td>
</tr>
<tr>
<td>Disk pos. angle PA [°]</td>
<td>158.6</td>
<td>158.6</td>
</tr>
</tbody>
</table>

**Notes.** Parameters for our model. $R_{\text{in}}$ and $R_{\text{out}}$ denote the inner and outer radius of the inner disk, respectively, whereas $R_e$ is the characteristic radius (truncation radius of exponential cutoff). $\delta_{\text{disk}}$ corresponds to the depletion factor applied to the surface density of the inner and outer disks. $h_{100,\text{small}}$ quantifies the scale height of small and big grains at 100 au, and $\beta$ the flaring index, respectively.

### 4.2. Modelling the inner disk

We test different models containing an inner disk and compare the butterfly pattern in the model with the observations. For this purpose, we consider models with different inner disk configurations in terms of two free parameters: outer radius $R_{\text{out}}$ of the inner disk, and depletion factor $\delta_{\text{disk}}$ of the surface density of the inner disk. We compare cuts through the convolved $Q_{\rho}$ model images along the semi-major axis with the non-coronagraphic observations (without subtracting the central source polarisation; see Fig. 6, middle panel). The depletion factor was varied between 0.05 and 1.0, and the inner disk outer radius between 2 au (corresponding to the completely unresolved case) and 20 au, by keeping the total dust mass constant. For each radius value, we identified the best matching depletion factor. Our modelling appears to be degenerate since we could obtain, for both small and large radii values, model images reproducing reasonably well the observations. The range of solutions includes configurations with a small inner disk with a high surface density, as well as those with a larger inner disk with lower surface density. In any case, the inferred surface density outside 20 au in all configurations is very low, indicating that only a small amount of material is left at the location of the companion ($\sim$22 au; see Sect. 5).

### 4.3. Discussion of brightness asymmetries in the disk

In Sect. 3.1, we discussed the outer disk asymmetries and found that the east side is brightest in polarised light but is marginally detected in total intensity. Furthermore, along the north–south axis, the brightness maxima differ between the optical and NIR images taken a year apart.

The observed asymmetries between the east and west sides might be connected to the scattering properties of the dust. The polarised intensity is the product of the degree of polarisation and total intensity. The phase function of both depends on particle size and observing wavelength. In Fig. 7 (left), we plot the linear polarisation degree expressed by the Mueller matrix elements ($-Z_{\parallel}/Z_{\perp}$) as a function of scattering angle. For this test, we computed the scattering matrix using Mie theory for different grain sizes. To smooth out the resonances in the phase function appearing when considering perfect spheres of a single size, instead of using a mono-dispersive grain size distribution, we generated for each grain size a narrow Gaussian distribution centred around the considered grain size and with a FWHM of 10% of the grain size. The scattering matrix was evaluated at 0.7 $\mu$m which corresponds approximately to the central wavelength of the VBB filter. The plot shows that for very small grains (0.001–0.01 $\mu$m), the polarisation degree is symmetric around 90°. For larger grains, the polarisation degree has several minima, until at grain sizes much larger than the observing wavelength (≥2 $\mu$m), the maximum of the polarisation degree is shifted towards angles smaller than 90°. Regarding the scattering phase function (Fig. 7, right panel), particles smaller than ≤0.1 $\mu$m scatter symmetrically around 90°. Above that value, the phase function becomes asymmetric and forward scattering is increasingly pronounced.

The observed polarised intensity is now an interplay between the polarisation degree and total intensity, and the disk’s geometry. Due to the disk’s flared geometry, the scattering angles at the far side are closer to values of 90° than at the near side (the scattering angles are symmetric around $\theta = 90° - \psi - i$ and $\theta = 90° - \psi + i$ on the near and far sides, respectively, when $i$ denotes the inclination and $\psi$ the opening angle of the disk;
see Min et al. 2012). For smaller grains that are in the Rayleigh scattering regime and therefore not strongly forward scattering, this could therefore make the far side appear brighter than the near side (e.g. Murakawa 2010; Min et al. 2012). On the other hand, the larger the grains that are considered, the more forward scattering they are, and the more the polarised intensity dominated by the phase function of the total intensity. The near side then becomes the brighter side, the same in total intensity as in polarised intensity.

To test this hypothesis, we computed ray-traced $Q_\phi$ images at 0.7 $\mu$m using different grain sizes. We subtracted the central source polarisation after convolving the Stokes $Q$ and $U$ images from the $Q_\phi$ frame using a $U_\phi$ minimisation. In this procedure, scaled versions of the total intensity $I$ frame are subtracted from the Stokes $Q$ and $U$ model images, such that the absolute value of $U$ in a defined region is minimised. Figure 8 compares the resulting images. As expected, the disk model with larger grains ($\geq 0.18$ $\mu$m) shows strong forward scattering even in polarised light, whereas for the disk containing small grains, the far side appears brighter. As a mono-dispersive grain size distribution would be unrealistic, we also tested a grain size distribution of small grains ($0.001-0.15$ $\mu$m), and still find that the far side appears brighter than the near side (Fig. 8, lower middle panel). We conclude that if the brightness asymmetry between the east and west sides is real, and not dominated by effects from poor seeing conditions and reduction artifacts, we need predominantly small sub micrometer sized grains ($\leq 0.15$ $\mu$m) to reproduce the observations in a qualitative way. Although we are able to reproduce the qualitative behaviour of the brightness asymmetries, we are not able to reproduce contrast ratios as large as in the observations between the two sides. Furthermore, our brightness contrast is very similar in the VBB-band and $J$-band (contrary to the observations).

In total intensity, on the other hand, the near side is always expected to be brighter than the far side, as even in the Rayleigh scattering regime (where the scattering phase function is symmetric around a minimum at 90° and where grains are not forward scattering), the scattering angles at the near side are farther from 90° than at the far side, corresponding to a higher scattering efficiency.

We note that our model is based on Mie scattering. It is certainly worth testing the impact of particles that are not spherical and homogeneous, or that are of a somewhat different chemical composition; for example, including water ice mantles that are neglected in our modelling approach. However, this is beyond the scope of this study.

Summarising, by retaining only small ($\leq 0.15$ $\mu$m) grains in the disk surface layer, we are able to reproduce the brightness contrasts between the east and west sides in a qualitative (although not quantitative) way. The existence of the north/south brightness asymmetry and its different behaviour in the VBB band and $J$-band on the other hand cannot be explained solely with grain scattering properties, as the scattering angles are expected to be symmetric with respect to the semi-minor axis. One could speculate that the grain properties are different in the north and south regions, but more complex modelling would be needed to explain this behaviour.

It should be mentioned that the strong butterfly patterns detected in $Q_\phi$ and $U_\phi$ after correcting for IP effects affect the outer disk. In the $Q_\phi$ image, this adds positive signal along the disk semi-major axis, and subtracts signal along the semi-minor axis. The subtraction of this central component is likely imperfect and some residuals may be present in the images of the outer disk in Fig. 1. In addition, all our polarimetric observations suffered from rather poor seeing conditions, which might further influence the apparent brightness distribution in the disk. Therefore, we cannot rule out that artifacts from the data reduction and/or weather conditions affect the azimuthal brightness distribution of the polarimetric datasets.

![Fig. 8. Convolved ray-traced $Q_\phi$ images evaluated at 0.7 $\mu$m using different grain size distributions. Successively, the central source polarisation emerging from the unresolved inner disk (radius 2 au) was subtracted using a $U_\phi$-minimisation. The lower right panel shows the VBB observation for comparison. North is up and east is to the left.](image-url)
Fig. 9. Images of the point source detection as retrieved with the sPCA reduction (from left to right): NICI L′-band (2012-03-31), IRDIS H2H3-band (2015-05-03), IRDIS H2H3-band (2015-05-31), IRDIS K1K2-band (2016-05-14), NaCo L′-band (2016-06-01). North is up and east is to the left. The images were smoothed with a Gaussian kernel of size $0.5 \times \text{FWHM}$.

Table 4. Astrometric calibrations used for the ADI datasets.

<table>
<thead>
<tr>
<th>Date</th>
<th>Instrument</th>
<th>Filter</th>
<th>True north correction angle (°)</th>
<th>Rotator offset position (°)</th>
<th>Pixel scale (mas px$^{-1}$)</th>
</tr>
</thead>
<tbody>
<tr>
<td>2012-03-31</td>
<td>NICI</td>
<td>L′</td>
<td>0.0 ± 0.1$^b$</td>
<td>180.0 ± 0.1$^f$</td>
<td>17.95 ± 0.01$^d$</td>
</tr>
<tr>
<td>2015-05-03</td>
<td>IRDIS</td>
<td>H2</td>
<td>−1.700 ± 0.076$^a$</td>
<td>−135.99 ± 0.11</td>
<td>12.255 ± 0.021$^e$</td>
</tr>
<tr>
<td>2015-05-03</td>
<td>IRDIS</td>
<td>H3</td>
<td>−1.700 ± 0.076$^a$</td>
<td>−135.99 ± 0.11</td>
<td>12.250 ± 0.021$^e$</td>
</tr>
<tr>
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<td>IRDIS</td>
<td>H2</td>
<td>−1.700 ± 0.076$^a$</td>
<td>−135.99 ± 0.11</td>
<td>12.255 ± 0.021$^e$</td>
</tr>
<tr>
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<td>IRDIS</td>
<td>H3</td>
<td>−1.700 ± 0.076$^a$</td>
<td>−135.99 ± 0.11</td>
<td>12.250 ± 0.021$^e$</td>
</tr>
<tr>
<td>2016-05-14</td>
<td>IRDIS</td>
<td>K1</td>
<td>−1.675 ± 0.080$^e$</td>
<td>−135.99 ± 0.11</td>
<td>12.243 ± 0.021$^e$</td>
</tr>
<tr>
<td>2016-05-14</td>
<td>IRDIS</td>
<td>K2</td>
<td>−1.675 ± 0.080$^e$</td>
<td>−135.99 ± 0.11</td>
<td>12.238 ± 0.021$^e$</td>
</tr>
<tr>
<td>2016-06-01</td>
<td>NaCo</td>
<td>L′</td>
<td>0.518 ± 0.120$^c$</td>
<td>89.5 ± 0.11$^c$</td>
<td>27.195 ± 0.063$^e$</td>
</tr>
</tbody>
</table>

Notes. ($^a$Maire et al. (2016); $^b$)assumed value (no astrometric measurement around the NICI epoch available). ($^c$)Adopted from Chauvin et al. (2012). ($^d$)Adopted from Hayward et al. (2014) and Wahhaj et al. (2014). ($^e$)Launhardt et al. (in prep.). ($^f$)Cassegrain rotator position angle. There is no information on the uncertainty available; we therefore adopt an uncertainty of 0.1°.

5. Detection of a planetary mass companion

In the two SHINE IRDIS epochs, we identified a point source at a separation of about 195 mas and a position angle of about 155° from PDS 70. Upon analysis of the IRDIS K12 Open Time data, as well as of the NaCo data, the point source was detected around the same location as in the IRDIS data. In addition, we reanalysed archival data taken with Gemini/NICI in L′-band on March 31, 2012, published by Hashimoto et al. (2012). The authors reported a non-detection of companion candidates (apart from the previously mentioned background source to the north), but their analysis considered only regions outside of ~200 mas. Upon our re-reduction of this dataset, we detected the point source around the expected location. The final detection images obtained with sPCA are shown in Fig. 9. The corresponding images retrieved using ANDROMEDA, PCA-SpeCal, and TLOCI are presented in Fig. A.1. We detected the point source in all our available ADI epochs, spanning a total time range of four years.

We also noted another structure present in some of the PCA and TLOCI processed images, located at a similar separation and a position angle of about 280°. To check the point-like nature of this structure, we processed the data with the ANDROMEDA algorithm, which is optimised for the retrieval of point sources (Cantalloube et al. 2015). Figure A.2 shows the corresponding S/N maps, which are a result of forward modelling under the assumption of the presence of a point-like source. It can be seen that this structure is not consistently recovered by ANDROMEDA in the different epochs at significant S/N. Especially towards longer wavelengths, no other source apart from the above-described source at about 155° is detected with any significance. This implies that the structure found at about 280° is not point-like and, if physical, we can associate this structure with a disk feature. The latter interpretation is supported by the projected proximity to the outer disk ring.

5.1. Astrometry

For the characterisation of the point source, we extracted the astrometry and the photometry for all epochs. We detected it with all algorithms (sPCA, ANDROMEDA, PCA-SpeCal, and TLOCI), and in the following focus on the analysis of the sPCA reduction. In this sPCA reduction, we divided the image in concentric annuli with a width of $1 \times \text{FWHM}$. For each annulus we adjusted the number of modes in such a way that the protection angle was maximised. A maximum number of 20 modes was applied and we set the maximum protection angle to 0.75 × FWHM. We extracted the astrometry and photometry by injection of a PSF taken from the unsaturated frames out of the coronagraph with negative flux, as proposed by Lagrange et al. (2010). Our approach to find the location and flux of the point source consisted of varying the parameters of this negative signal using a predefined grid to minimise the residuals in the resulting sPCA-processed dataset. We therefore computed for each parameter set of position and flux the effective signal using a predefined grid to minimise the residuals in the resulting sPCA-processed dataset. We therefore computed for each parameter set of position and flux the $\chi^2$ value within a segment having a radial and azimuthal extension of $2 \times \text{FWHM}$ and $4 \times \text{FWHM}$ around the point source, respectively. To derive uncertainties in the astrometric and photometric values, posterior probability distributions for each parameter were computed following the method described by Olofsson et al. (2016). The astrometric uncertainties related to the calibration error take into
The point source is detected at a mean projected separation of 12.5 mas within the 1-σ uncertainties. The relative position as measured in the NICI data taken with the expected trajectory for a background star relative to PDS 70, we compared its measured position relative to the rotator offset and the pixel scale. The corresponding values are reported in Table 4. We derived the final astrometric uncertainties at each epoch by quadratically summing the errors from these individual contributions. Our astrometric measurements obtained at the different epochs are presented in Table 5. As a cross-check, the results of the ANDROMEDA, PCA-SpeCal, and TLOCI reductions are listed in Table A.1.

To test whether the point source is part of a physical system with PDS 70, we compared its measured position relative to the star at the different epochs. Due to the proper motion ($\mu_r \cos \delta = -29.7$ mas yr$^{-1}$, $\mu_\delta = -23.8$ mas yr$^{-1}$; Gaia Collaboration 2016, 2018), a stationary background star would have moved by $\sim 160$ mas within the given time span. As the relative motion ($\sim 40$ mas during the $\sim 4$ years observational time span) differs significantly from the prediction for a stationary background object, the astrometric results strongly imply that the point source is comoving with PDS 70. The measurements, together with the expected trajectory for a background star relative to PDS 70, are displayed in Fig. 10. Further, the probability of detecting at least one background contaminant of similar brightness or brighter within the mean separation of the companion is less than 0.033% according to the Besançon galactic population model (Robin et al. 2003).

The relative position as measured in the NICI data taken in 2012 does not coincide with the positions derived from the SPHERE and NaCo observations performed in 2015 and 2016 within the 1-σ uncertainties. This difference in measured position between the epochs is possibly due to orbital motion. The point source is detected at a mean projected separation of $\sim 195$ mas, corresponding to $\sim 22$ au. The orbital period of such a bound object, assuming a stellar mass of 0.76 $M_\odot$, would be $\sim 119$ years. For a face-on circular orbit, this implies a displacement of $\sim 3^\circ$ per year, resulting in a total change of position angle of 12.5$^\circ$ within the time covered by our observations, which is in good agreement with the observations. Further, the observed change in position angle is in clockwise direction, which corresponds to the sense of rotation of the disk (Hashimoto et al. 2015). Therefore, this displacement is consistent with an object on a circular face-on orbit rotating in the same sense as the disk. However, regarding the relatively large uncertainties on the astrometry and the short time span covered by our data, detailed orbital fitting exploring the possibility of an inclined and/or eccentric orbit will be performed in a follow-up study on this source (Müller et al. 2018). Although the possibility of the point source being a background star with almost the same proper motion as PDS 70 is very small, only the detection of orbital motion over a significant part of the orbit will allow excluding fully the background star scenario.

### 5.2. Photometry

Our current information on the physical properties of the companion candidate relies on the $H$, $K$, and $L'$ photometry as derived from our SPHERE/IRDIS, NaCo, and NICI images. It is marginally detected in the IFS data, when the channels corresponding to $J$-band and $H$-band are collapsed. Due to the large uncertainties, these data are not considered here. The low S/N detection of the companion candidate in the IFS data can be explained by its faintness and red colours, the larger IFS thermal background (IFS is not cooled contrary to IRDIS), and the smaller IFS pixel scale (7.46 mas pixel$^{-1}$ vs. $\sim 12.25$ mas pixel$^{-1}$). HD 95086 b offers a similar case of a faint companion with red colours for which a detection with IRDIS is achieved in the $K$-band in individual observation sequences.
whereas the detection with IFS in the J- and H-bands required the combination of several epochs (Chauvin et al. 2018).

The companion has very red colours, with a magnitude difference of $H_2 - K_1 = 1.82 \pm 0.36$ mag and $1.59 \pm 0.39$ mag, considering the first and second SHINE H-band epochs, respectively. Accordingly, we measured a magnitude difference of $H_2 - L' = 3.67 \pm 0.46$ mag and $3.44 \pm 0.48$ mag (considering the NICI L'-band photometry). The properties of the companion are further discussed in Sect. 5.4.

5.3. The nature of the point-like source

Due to the detection of the point source at multiple epochs and using several different instruments, filter bands, and image post-processing algorithms at about the same location, we can exclude that the source is due to an instrumental or atmospheric artifact (speckle). Furthermore, the discrepancy of the relative position of the point source with respect to the predicted trajectory of a stationary background star strongly implies a gravitationally bound object. Several of the proposed companion candidates within protoplanetary disks are currently under debate, because they are not detected consistently at all available wavelengths. A reason for possible confusion is that the ADI process acts as a spatial frequency filter and enhances sharp and asymmetric disk features, as shown by Milli et al. (2012), applied on extended disk structures, it can cause distortions or even create artifacts that are not related to physical structures. Disk structures such as rings, spiral arms, or clumps, when processed with ADI, could therefore lead to a misinterpretation of point-like sources. One of the point sources detected around HD 169142 (Biller et al. 2014; Reggiani et al. 2014) was shown to be related to an inhomogeneous ring structure in the inner region of the disk by Ligi et al. (2018). In addition, these authors found that an additional compact structure in that system detected at about 100 µas projected separation could possibly be related to a further ring structure at the given separation. Two companion candidates were also found around HD 100546 (Quanz et al. 2013a, 2015; Brittain et al. 2014; Currie et al. 2015), but remain debated under consideration of recent GPI and SPHERE observations, as they might not appear point-like at all wavelengths (Rameau et al. 2017; Follette et al. 2017; Currie et al. 2017; Sissa et al. 2017). Regarding LkCa15, three companion candidates have been reported in the literature (Kraus & Ireland 2012; Sallum et al. 2015). Thalmann et al. (2016) found that the location of the planet candidates coincides with the bright, near side of the inner component of the LkCa15 disk and conclude that this inner disk might account for at least some of the signal attributed to the protoplanets detected by Sallum et al. (2015). Due to the detection of $H_2$ emission at the location of LkCa15b, only this candidate makes a convincing case for a protoplanet in this system.

All these debates illustrate that a careful analysis of companion candidates located in protoplanetary disks with respect to a possible link to disk features is required. We therefore address the hypothesis that some asymmetric dust structure at or close to the given location of the point source is responsible for our detection. As stated above, we are not able to resolve the detailed structure of the inner disk with our PDI observations, but suspect the inner disk to be smaller than 17 au in radius. To test the impact of the inner disk signal on the structures seen in ADI, we follow the approach by Ligi et al. (2018), and simulate a cADI observation using the IRDIS PDI J-band $Q_8$ images, after subtracting the central source polarisation. For this purpose, we created a datacube whose 50 frames correspond to identical copies of the PDI $Q_8$ image, rotated by the respective parallactic angles encountered during the ADI epoch of May 31, 2015. We then subtracted the median of this datacube from each single frame, before de-rotating them and computing their median. In addition, we applied the same procedure to the $Q_8$ model image, computed at 1.25 µm. We convolved our image with the total intensity frame acquired during the non-coronagraphic J-band observations and subtracted the central source polarisation using a $U_8$ minimisation before applying the cADI algorithm. The inner disk in the model configuration used for this test extends out to 2 au. The result is shown in Fig. 12. There is no prominent structure appearing at the distance of interest (∼200 mas). While this test does not allow us to completely rule out a disk structure as the cause of this feature, there is at least no obvious polarised inner disk structure that would create this kind of artifact. One further argument against the hypothesis of the companion being a disk feature is the fact that we do not detect strong polarised signal at the location of the companion in the PDI data, which would be consistent with the signal detected in ADI being of thermal origin. We therefore conclude that, given the present data, a physically bound companion is the most plausible explanation for the detected point source, and refer to it as PDS 70b hereafter.

5.4. Companion properties

Figure 11 shows the location of PDS 70b in SPHERE H-band and K-band-based colour-magnitude diagrams (CMD). The diagrams are complemented with the synthetic photometry of M, L, and T dwarfs, as well as with the measurements of young companions and red dwarfs of the Sco-Cen association and other regions. We refer the reader to Chauvin et al. (2018) and Bonnefoy et al. (2018) for details regarding these diagrams. The diagrams show that the absolute $H_2$ magnitude of the companion is consistent with those of L-type dwarfs. PDS 70b is located between the ∼5–11 Myr old, ∼8–14 $M_{\text{Jup}}$ planet 1RXS1609b (Lafrenière et al. 2008, 2010; Lachapelle et al. 2015) and the 30 Myr-old ∼7 $M_{\text{Jup}}$ planets HR 8799 c,d,e (see Bonnefoy et al. 2016). The location of PDS 70b in the H2-H3 CMD is remarkably close to the recently discovered dusty giant planet HIP 65426b (Chauvin et al. 2017b). Indeed, HIP 65426b’s mass (6–12 $M_{\text{Jup}}$) derived from evolutionary models is similar to the one of PDS 70b (see below), although significantly older (14 ± 4 Myr; Chauvin et al. 2017b). In addition, the K1-K2 diagram reveals a similar photometry to the 11 ± 2 $M_{\text{Jup}}$ massive and 13 ± 2 Myr old companion HD 106906b (Bailey et al. 2014).

The colours of PDS 70b are very red. Its H2-K1 colour is redder than most L dwarfs in the field, and is consistent with the very red companions to CD-35 2722 (Wahhaj et al. 2011) and 2M1207 (Chauvin et al. 2004), as well as with HIP 65426 (Chauvin et al. 2017b) within the uncertainties. If due to a photosphere, the red colour is only compatible with an L-type object or with a reddened background object, but this latter possibility is very unlikely due to the proper motion test. The absolute $L'$-band magnitude is brighter than most of the detected companions and consistent with those of late M- to early L-type objects, but again significantly redder than these sources. The $H_2 - L'$ colour is as red as the >50 Myr old, very dusty companion to PDS 70b has also recently been observed by MagAO, revealing a potential detection of $H_2$ emission at the expected location of PDS 70b (Wagner et al. 2018). If this detection is due to the accretion of gas on the planetary object, this would provide further evidence that the object is neither a disk feature nor a background star.
HD 206893 ($H2 - L' = 3.36 \pm 0.18$ mag), which is one of the reddest brown dwarf companions known (Milli et al. 2017; Delorme et al. 2017). Therefore, the location of PDS 70b on the CMD is quite unusual. However, it should be kept in mind that only very few of these objects are of similarly young age as PDS 70, and none of the above objects are detected within the TD of its host. PDS 70b might therefore be the only of these objects that is directly observed during its formation process.

In order to estimate the mass of the companion, we compared the photometry of PDS 70b to the Bern Exoplanet (BEX) evolution models. These tracks are obtained from the synthetic planetary populations of Mordasini et al. (2017), which predict the post-formation planetary luminosity as a function of time, considering different efficiencies of the accretional heating during the formation process and including the effect of deuterium burning. The planets formed in the planetary population synthesis are classified into four different populations ("hottest", "hot", "warm", "coldest"), according to their luminosity as a function of mass at the moment when the disk disappears (see Mordasini et al. 2017 for details). The planets according to the "hottest" and "coldest" populations have the highest and lowest luminosities, and correspond to the traditional hot and cold start models (e.g. Chabrier et al. 2000; Baraffe et al. 2003; Marley et al. 2007). They describe the two extreme cases where the entire gas accretion shock luminosity is either deposited in the planet’s interior or radiated away during the formation process. These two populations are superseded by the more realistic cases of the "hot" and "warm" populations, which are representative for cases with intermediate initial entropies between the extreme "hot" and "cold" start models. For our comparison, we made use of the "hottest", "hot", and "warm" populations, but discarded the "coldest" population. For this scenario, a planet mass larger than 10 $M_{\text{Jup}}$ would be required to reproduce the observed magnitudes of PDS 70b. However, to be classified into the "coldest" population requires small core masses which do not develop into planets with such high masses in the planetary population synthesis models (see Mordasini et al. 2017, Fig. 13). In addition, observations suggest that the pure cold start formation is in reality not the preferred formation mechanism and the truth most probably lies somewhere between the two extrema of purely hot and cold start (Bonnefoy et al. 2014; Samland et al. 2017). From the theoretical side, recent detailed simulations of
the accretion shock suggest that hot starts are preferred (e.g. Marleau et al. 2017, and prep.). To follow the post-formation cooling, the outcome of the population synthesis was combined with the boundary conditions for the atmospheric structure from the COND models (Baraffe et al. 2003). The synthetic SPHERE magnitudes were then computed using the DUSTY atmospheric model (Chabrier et al. 2000). These results were linearly interpolated in time to the stellar age (5.4 Myr). Figure 13 compares the photometry of PDS 70b with the synthetic colours of the “hottest”, “hot”, and “warm” tracks from Mordasini et al. (2017), as well as with the original DUSTY model of Chabrier et al. (2000). We find a mass between 5 and 9 \( M_{\text{Jup}} \) for the hot start models (“hottest”, “hot”, “DUSTY”), and a mass between 12 and 14 \( M_{\text{Jup}} \) for the coldest (“warm”) population considered, implying that in the case of lowest entropy, deuterium burning might play a role. It is important to note that the evolutionary tracks do not take into account the time needed for the planet to form. Since this may take up to several million years, the stellar age is only an upper limit on the age of the planet, and consequently, the estimated masses should be considered as conservative upper limits.

For completeness, Fig. 13 (lower panel) shows the H2-L’ colour of PDS 70b in comparison with the DUSTY tracks, which implies a similar mass estimate (5–8 \( M_{\text{Jup}} \)). We also plotted the evolutionary tracks corresponding to a planetary age of 1 and 3 Myr to illustrate the mass range in which the planet would be found in the case that it has formed considerably more recently than the star (down to 2–4 \( M_{\text{Jup}} \) for a planetary age of 1 Myr). We also compared the \( H \)-band photometry to the warm-start models from Spiegel & Burrows (2012), as well as the hot-start COND models from Baraffe et al. (2003), resulting in a similar finding (5–10 \( M_{\text{Jup}} \) for the former, and 4–5 \( M_{\text{Jup}} \) for the latter). However, the colours from the COND model do not match those observed for PDS 70b since they are significantly redder than predicted by the COND models, which suggests the presence of a dusty or cloudy atmosphere. We emphasise that none of these models considers the presence of circumplanetary material, which could affect the observed SED and the corresponding mass estimate. The presence of a circumplanetary disk could also cause an IR excess in the object’s SED pushing the photometry towards redder colours. Future ALMA observations will allow us to search for the presence of such material around PDS 70b.

Finally, we used the Exoplanet Radiative-convective Equilibrium Model (Exo-REM) to analyse the SED of PDS 70b (Baudino et al. 2015, 2017). We performed a grid search to determine the parameters that best minimise the \( \chi^2 \) taking into account all photometric points from the \( H \) to the \( L \)-band (Baudino et al. 2015). All the determined radii were larger than or equal to 2 \( R_{\text{Jup}} \), which is a large value compared to the evolutionary models. Therefore, instead of simply minimising the \( \chi^2 \) to find the radius, we determined the minimal radius that gives a spectrum similar to the observation at 1, 3, and 5-H2 (when applicable). We obtained at 5-H2 a surface gravity of \( \log g = 3.9 \pm 0.9 \) dex, and a temperature of \( T_{\text{eff}} = 1200 \pm 200 \) K. Our grid takes into account solar metallicity and high cloud absorption (\( \tau_{\text{ref}} = 3 \)), suggested by the fact that the object is redder than usual on the CMD. We note that the photometry is also in good agreement with a simple blackbody with a temperature range between 1150 and 1350 K, further indicating a very dusty or cloudy atmosphere with few spectral features. Also in the blackbody case, large effective radii of several \( R_{\text{Jup}} \) are needed to fit the absolute flux density, which, again, may be explained by the possible existence of spatially unresolved circumplanetary material contributing to the measured flux. Our atmospheric models are in this respect oversimplified and the possible presence of circumplanetary material would require substantial modifications of the underlying models, which is beyond the scope of our paper, but will be discussed in a forthcoming paper (Muller et al. 2018).

6. Summary and conclusions

PDS 70 is a young T Tauri star hosting a known transition disk with a large gap. Transition disks are thought to host gap-carving planets, and are therefore prime targets to observe ongoing planet formation and planet–disk interactions.

We have presented VLT/SPHERE optical and NIR observations in PDI mode, carried out with SPHERE/ZIMPOL in the VBB-band and SPHERE/IRDIS in the J-band. In addition, we have obtained total intensity images with SPHERE/IRDIS.

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3 We note that the BEX tracks are currently not available for the NaCo magnitudes, and are therefore not considered for the comparison in the \( L' \)-H2L’ diagram.
in H2H3 and K1K2 dual-band imaging, with simultaneous spectro-imaging using IFS working in YJ-band and YJH-band, respectively. Our observations of the PDS 70 system obtained within this work have been complemented with data taken with VLT/NaCo and archival Gemini/NICI observations in the L’-band using ADI. The presented data comprise eight different epochs spanning a time range of five years, leading to the following results.

The disk is clearly detected in all data sets presented in this work and resolved in scattered light with high angular resolution. We confirm the previously reported gap with a size of \(\sim 54\) au. We detect for the first time scattered light from the inner disk. By comparison with our RT model, we derive that the position angle of the inner disk is approximately the same as the outer disk. We also infer that the inner disk is not seen pole-on and has a maximum outer radius of \(<17\) au. The disk’s far side is brighter than the near side in PDI (VBB-band, J-band), whereas the disk’s near side is brighter in ADI (H-band, K’, L’-band). We suggest that this can be explained by the flared geometry of the disk in connection with Rayleigh scattering from small, sub micron-meter-sized grains.

We detect a point source at approximately 195 mas separation and 155° position angle. The detection is achieved at five different epochs, including the SPHERE/IRDIS, Gemini/NICI, and VLT/NaCo instruments in the \(H, K, L’\) band filter. The astrometry of the point source implies that the confusion with a reddened background object is unlikely, and that the object is bound. Due to the astrometric coverage of 4 years, we might see first hints of orbital motion. Astrometric follow-up observations will be performed to confirm the orbital motion and to constrain the orbital parameters.

The photometry of the companion shows evidence of very red colours. Comparison with evolutionary models suggests that the photometry is most compatible with a young planetary-mass body with a dusty or cloudy atmosphere.

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## Appendix A: PDS 70b astrometric and photometric results

### Table A.1. Comparison of photometry and astrometry of PDS 70b, as derived from the sPCA, ANDROMEDA, PCA-SpeCal, and TLOCI reductions.

<table>
<thead>
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<th>Date</th>
<th>Instrument</th>
<th>Filter</th>
<th>Sep (mas)</th>
<th>PA (deg)</th>
<th>Δ mag</th>
<th>S/N</th>
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</thead>
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<tr>
<td>Results from sPCA</td>
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<td>2012-03-31</td>
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<td>$L'$</td>
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<td>162.2 ± 3.7</td>
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<td>H2</td>
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<td>154.5 ± 1.2</td>
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<td>$L'$</td>
<td>181.7 ± 54.4</td>
<td>147.8 ± 17.2</td>
<td>7.1 ± 1.3</td>
<td>0.9</td>
</tr>
</tbody>
</table>

Notes. The current implementation of the astrometric error estimation for our TLOCI reduction is unreliable for the $L'$ datasets. We therefore used a conservative uncertainty of 2 pixels for the angular separation and the position angle for these datasets.
Fig. A.1. Contrast maps of the point source detection as retrieved with the sPCA reduction (first row), ANDROMEDA (second row), PCA SpeCal (third row), and TLOCI (fourth row). From left to right: NICI $L'$-band (2012-03-31), IRDIS H2H3-band (2015-05-03), IRDIS H2H3-band (2015-05-31), IRDIS K1K2-band (2016-05-14), NaCo $L'$-band (2016-06-01). The sPCA images were smoothed with a Gaussian kernel of size $0.5 \times$ FWHM. North is up and east is to the left. The brightness levels were adapted individually for visibility purposes.

Fig. A.2. ANDROMEDA S/N maps of the ADI epochs.
Appendix B: Stellar parameters of PDS 70

Fig. B.1. The location of PDS 70 on a HR diagram in comparison with the PMS evolutionary tracks and isochrones from Tognelli et al. (2011).

Appendix C: IRDIS ADI view of the disk

Fig. C.1. SHINE IRDIS observations of May 31, 2015: sPCA reduction (left panel), PCA-SpeCal reduction (middle panel), and cADI with an contour overlay of the PDI coronagraphic J-band image (right panel). The contours are drawn with respect to the peak value of the PDI image. For visibility purposes, the images are shown on individual colourscales. North is up and east is to the left.

PDS 70 is located at a distance of $113.43 \pm 0.52$ pc (Gaia Collaboration 2018). Using low-resolution optical spectra, Pecaut & Mamajek (2016) derived a stellar temperature of 3972 K, corresponding to a K7 spectral type. Assuming a distance of 98.9 pc, these authors derived a luminosity of 0.27 $L_\odot$, which, scaled to a distance of 113.4 pc, corresponds to a luminosity of 0.35 $L_\odot$.

We compare the position of PDS 70 with the PMS evolutionary tracks and isochrones from Tognelli et al. (2011) in Fig. B.1 (assuming a metallicity $Z = 0.02$, an initial helium abundance of 0.27, a mixing length of 1.68 and a deuterium abundance of $2 \times 10^{-5}$), which implies a stellar mass of 0.7–0.85 $M_\odot$, and an age of 6 ± 2 Myr. As an additional approach, we fitted the stellar evolutionary models from the MIST project (Paxton et al. 2011, 2013, 2015; Dotter 2016; Choi et al. 2016) using a Markov Chain Monte Carlo approach to determine the stellar parameters. This method is described in detail in Müller et al. (2018). This approach resulted in a stellar age of 5.4 ± 1.0 Myr and a mass of 0.76 ± 0.02 $M_\odot$, which is consistent with the above given values. It is worth noting that our mass estimates are in good agreement with the dynamical mass of PDS 70 (0.6–0.8 $M_\odot$; Hashimoto et al. 2015; Long et al. 2018).