Measurement of the atmospheric $\nu_e$ and $\nu_\mu$ energy spectra with the ANTARES neutrino telescope


DOI
10.1016/j.physletb.2021.136228

Publication date
2021

Document Version
Final published version

Published in
Physics Letters B

License
CC BY

Citation for published version (APA):
https://doi.org/10.1016/j.physletb.2021.136228

General rights
It is not permitted to download or to forward/distribute the text or part of it without the consent of the author(s) and/or copyright holder(s), other than for strictly personal, individual use, unless the work is under an open content license (like Creative Commons).

Disclaimer/Complaints regulations
If you believe that digital publication of certain material infringes any of your rights or (privacy) interests, please let the Library know, stating your reasons. In case of a legitimate complaint, the Library will make the material inaccessible and/or remove it from the website. Please Ask the Library: https://uba.uva.nl/en/contact, or a letter to: Library of the University of Amsterdam, Secretariat, Singel 425, 1012 WP Amsterdam, The Netherlands. You will be contacted as soon as possible.
Measurement of the atmospheric $\nu_e$ and $\nu_\mu$ energy spectra with the ANTARES neutrino telescope

The ANTARES Collaboration

A. Albert$^{a,b}$, S. Alves$^c$, M. André$^d$, M. Anghinolfi$^d$, G. Anton$^e$, M. Ardid$^f$, J.-J. Aubert$^g$, J. Aubin$^h$, B. Baret$^h$, S. Basa$^i$, B. Belhorma$^j$, J. Bendahan$^{k,h}$, V. Bertin$^g$, S. Biagi$^i$, M. Bissinger$^e$, J. Boumaaza$^k$, M. Bouta$^m$, M.C. Bouwhuis$^n$, H. Brânzaș$^o$, R. Brujin$^{n,p}$, J. Brunner$^g$, J. Busto$^g$, A. Capone$^{q,r}$, L. Caramete$^o$, J. Carr$^g$, V. Carretero$^o$, S. Celli$^{q,r}$, M. Chabab$^l$, T.N. Chau$^h$, R. Cherkaoût El Moursli$^k$, T. Chiarusi$^u$, M. Circella$^v$, A. Coleiro$^h$, M. Colomer-Molla$^{h,s}$, R. Coniglione$^i$, P. Coyle$^h$, A. Creusot$^h$, A.F. Díaz$^w$, G. de Wasseige$^h$, A. Deschamps$^c$, D. Distefano$^l$, I. Di Palma$^{q,r}$, A. Domínguez$^{d,y}$, C. Donzaud$^{h,z}$, D. Dornic$^g$, D. Drouhin$^{a,b}$, T. Eber$^e$, N. El Khayati$^k$, A. Enzenhöfer$^g$, F. Fermi$^{q,r}$, G. Ferrara$^l$, F. Filippini$^{u,aa}$, L. Fusco$^{h,g}$, R. García$^h$, Y. Gatelet$^h$, P. Gay$^{ab,h}$, H. Glotin$^ac$, R. Gozzini$^{s,e}$, K. Graf$^e$, C. Guidi$^{d,y}$, S. Hallmann$^e$, H. van Haren$^{ad}$, A.J. Heijboer$^o$, Y. Hello$^x$, J.J. Hernández-Rey$^s$, J. Höbl$^e$, J. Hofestädtere, F. Huang$^a$, G. Illuminati$^{u,aa,h}$, C.W. James$^{ae}$, B. Jisse-Jung$^m$, M. de Jong$^{n,af}$, P. de Jong$^h$, M. Jongen$^n$, M. Kadler$^{ag}$, O. Kalezin$^e$, U. Katz$^e$, N.R. Khan-Chowdhury$^s$, A. Kouchner$^h$, I. Kreykenbohm$^{ai}$, V. Kulikovskiy$^{d,aj}$, R. Lahmann$^e$, R. Le Breton$^h$, D. Lefèvre$^{ak,al}$, E. Leonora$^{am}$, G. Levi$^{u,aa}$, M. Lencetto$^g$, D. Lopez-Coto$^{an}$, S. Loucatos$^{ao,h}$, L. Maderer$^h$, J. Manczak$^l$, M. Marcelin$^l$, A. Margiotta$^{u,aa}$, A. Marinelli$^{ap}$, J.A. Martínez-Mora$^l$, K. Melis$^{h,p}$, P. Migliozzi$^{am}$, M. Moser$^e$, A. Moussa$^m$, R. Muller$^n$, L. Nauta$^n$, S. Navas$^{an}$, E. Nezri$^a$, A. Nuñez-Castiñeayra$^{g,l}$, B. O’Fearraigh$^n$, M. Organokov$^a$, G.E. Pávallas$^c$, C. Pellegrino$^{u,aq,ar}$, M. Perrin-Terrin$^g$, P. Piattelli$^l$, C. Pieterse$^s$, C. Poirè$^i$, V. Popa$^e$, T. Pradier$^a$, N. Randazzo$^{am}$, S. Reck$^g$, G. Riccobene$^k$, F. Salesa Greus$^s$, D.F.E. Samtleben$^{h,af}$, A. Sánchez-Losa$^v$, M. Sanguineti$^{d,y}$, P. Sapienza$^l$, J. Schnabel$^e$, J. Schumann$^f$, F. Schüssler$^{ao}$, M. Spurio$^{u,aa,*}$, T. Stolarczyk$^{ab}$, M. Taituti$^{d,y}$, Y. Tayalati$^k$, T. Thakore$^s$, S.J. Tingay$^{ae}$, B. Vallage$^{ao,h}$, V. Van Elewyck$^{nh,ah}$, F. Versari$^{u,aa,h,*}$, S. Viola$^l$, D. Vivolo$^{ap,as}$, J. Wilms$^{ai}$, A. Zegarelli$^{q,r}$, J.D. Zornozas$^k$, J. Zúñiga$^s$

$^a$ Université de Strasbourg, CNRS, IPHC UMR 7178, F-67000 Strasbourg, France
$^b$ Université de Haute Alsace, F-68200 Mulhouse, France
$^c$ Technical University of Catalonia, Laboratory for Applied Biocoustics, Rambla Exposició, 08800 Vilanova i la Geltrú, Barcelona, Spain
$^d$ INFN - Sezione di Genova, Via Dodecaneso 33, 16146 Genova, Italy
$^e$ Friedrich-Alexander-Universität Erlangen-Nürnberg, Erlangen Centre for Astroparticle Physics, Erwin-Rommel-Str. 1, 91058 Erlangen, Germany
$^f$ Institut d’Investigació per a la Gestió Integra de les Zones Costaneres (IGIC) - Universitat Politècnica de València. C/ Paranimf 1, 46730 Gandia, Spain
$^g$ Institut d’Astrophysique de Paris, CNRS/INSU/CNRS, 98bis boulevard Arago, 75014 Paris, France
$^h$ Universität Mainz, Mainz, Germany
$^i$ University Mohammed V in Rabat, Morocco
$^j$ University Mohammed I in Salé, Morocco
$^k$ INFN - Laboratori Nazionali del Sud (LNS), Via S. Sofia 62, 95123 Catania, Italy
$^l$ University Mohammed I Laboratories of Physics and Radiations, B.P. 717, Oujda 6000, Morocco
$^m$ Nikhef, Science Park, Amsterdam, the Netherlands
$^n$ Institute of Space Science, RO-077125 Bucharest, Romania
$^o$ Universiteit van Amsterdam, Instituut voor Hoge-Energie Fysica, Science Park 105, 1098 XG Amsterdam, the Netherlands

$^{*}$ Corresponding authors.
E-mail addresses: spuro@bo.infn.it (M. Spurio), versari@bo.infn.it (F. Versari).

https://doi.org/10.1016/j.physletb.2021.136228

0370-2693/© 2021 The Author(s). Published by Elsevier B.V. This is an open access article under the CC BY license (http://creativecommons.org/licenses/by/4.0/). Funded by SCOAP3.
This letter presents a combined measurement of the energy spectra of atmospheric $\nu_\tau$ and $\nu_\mu$ in the energy range between $\sim100$ GeV and $\sim50$ TeV with the ANTARES neutrino telescope. The analysis uses 3012 days of detector livetime in the period 2007–2017, and selects 1016 neutrinos interacting in (close to) the instrumented volume of the detector, yielding shower-like events (mainly from $\nu_\tau + \pi_\tau$ charged current plus all neutrino neutral current interactions) and starting track events (mainly from $\nu_\mu + \pi_\mu$ charged current interactions). The contamination by atmospheric muons in the final sample is suppressed at the level of a few per mill by different steps in the selection analysis, including a Boosted Decision Tree classifier. The distribution of reconstructed events is unfolded in terms of electron and muon neutrino fluxes. The derived energy spectra are compared with previous measurements that, above 100 GeV, are limited to experiments in polar ice and, for $\nu_\mu$, to Super-Kamiokande.

© 2021 The Author(s). Published by Elsevier B.V. This is an open access article under the CC BY license (http://creativecommons.org/licenses/by/4.0/). Funded by SCOAP3.
with those obtained by other experiments [§6]. The ANTARES telescope is not able to distinguish between neutrino and antineutrino events. Hence, the unfolded spectra are the sum of $\nu_e + \bar{\nu}_e$ and of $\nu_\mu + \bar{\nu}_\mu$, averaged over the zenith region $90^\circ$–$180^\circ$.

2. The ANTARES detector and neutrino reconstruction algorithms

The ANTARES telescope [18] is a deep-sea Cherenkov neutrino detector, located 40 km offshore Toulon, France, in the Mediterranean Sea. The detector comprises a three dimensional array of 885 optical modules [19], each one housing a 10-inch photomultiplier tube (PMT). The optical modules are distributed along 12 vertical strings anchored to the sea floor at distances of about 70 m from each other and at a depth of about 2500 m. The detection of light from upward going charged particles is optimised with the PMTs facing $45^\circ$ downward. Particles above the Cherenkov threshold induce a coherent radiation emitted in a cone with a characteristic angle $\theta_C \simeq 42^\circ$ in water. For high-energy muons ($E_\mu > 1$ TeV), the contribution of the energy losses due to radiative processes increases linearly with $E_\mu$, and the resulting electromagnetic showers produce additional light. Completed in 2008, the telescope aims primarily at the detection of neutrino-induced through-going muons.

The signals induced in the PMT by detected photons are referred to as hits [20]. The position, time, and collected charge of the hits are used to reconstruct the direction and energy of events induced by neutrino interactions and atmospheric muons. Trigger conditions based on combinations of local coincidences are applied to identify signals due to physics events over the environmental light background due to $^{40}\text{K}$ decays and bioluminescence [21]. For astronomy studies, atmospheric muons and atmospheric neutrinos constitute the main source of background.

This analysis focuses on events induced by neutrinos whose interaction vertices are contained inside (or close to) the instrumented detector volume. These events include:

- $\nu_e$ charged current (CC) interactions producing electromagnetic and hadronic cascades, and neutral current (NC) interactions of neutrinos of all flavours inducing hadronic cascades. Due to the radiation and nuclear interaction lengths in water, the cascades extend up to a maximum distance of $\sim 10$ m from the interaction vertex, much shorter than the distance between detector strings. These events are thus almost point-like at the scale of the detector and are referred to as shower-like events in the following.
- $\nu_\mu$, CC interactions, with a hadronic cascade near the vertex and a starting muon. Most of these muons are minimum ionising particles, and can travel in water about 4 m per GeV of energy, inducing Cherenkov light over large distances with respect to the interaction vertex position. These events with a cascade and a track are referred to as starting track events in the following.

All neutrino candidates used in this letter are selected with an algorithm (denoted in the following as TANTRA) devoted to reconstruct events with interaction vertex inside, or close to, the instrumented volume. The initial reconstructed data sample is dominated by downward going atmospheric muons, exceeding by a factor $10^4$ the expected signal of shower-like and starting track events. As detailed in [22], the TANTRA likelihood-based method allows reconstructing the vertex coordinates, the neutrino direction, and the neutrino energy, yielding a parameter (a modified $\chi^2$-like quantity) associated to the quality of the fit that is denoted as $M_{\text{reco}}$. When the background of atmospheric muons is suppressed, and in dependence of the reconstruction quality $M_{\text{reco}}$, the neutrino vertex position is determined with a precision up to $\sim 1$ m; the neutrino direction is estimated with a median angular resolution of $\sim 3^\circ$ for a $E^{-2}_{\nu}$ energy spectrum in the range $1$–1000 TeV. Under these assumptions, the uncertainty on the reconstructed neutrino energy can be as low as $\sim 10\%$ for $\nu_e$.

If arriving in the detector an atmospheric muon, or a muon produced in a neutrino CC interaction with vertex far from the instrumented volume, induces a long sequence of hits characteristic of a long track. The track reconstruction algorithm used in off-line ANTARES analyses is called AAFit [23] and it is based on a likelihood fit that exploits a detailed parametrisation of the probability density function for the time of the hits. The algorithm provides the track direction with its estimated angular uncertainty and a proxy for the muon energy loss, which can be used to estimate the parent neutrino energy. The reconstruction quality is determined by a parameter, referred to as $\Lambda$, which is based on the maximum value of the likelihood fit and the number of degrees of freedom of the fit. The AAFit method is described in [23] and in this analysis it is mainly used to remove the largest fraction of atmospheric muons in the data sample. These events are downward going and can be significantly suppressed by a combination of cuts based on the reconstructed track direction and the $\Lambda$ quality parameter, as described in [24].

Finally, to improve the rejection of downward going atmospheric muon events, this analysis uses an auxiliary algorithm denoted as GridFit [25]. GridFit searches for tracks in 500 different directions covering the full solid angle. The number of hits compatible with a muon track coming from each direction is evaluated and a likelihood fit is performed. The outcomes of the AAFit and GridFit methods are used in the event selection in order to define cuts allowing to remove (within the available statistic of simulated events) the contamination of atmospheric muons, as detailed in the following section.

All ANTARES analyses follow a blinding policy to avoid possible biases. The cuts and the selection criteria are studied and optimised on a sample of Monte Carlo (MC) simulated events and only at the end of the full selection chain, these cuts are applied to data. A small sample containing 10% of the real data uniformly distributed over livetime is used to verify the agreement with MC events along the selection.

The simulation chain [26] starts with the generation of the event and comprises the generation of Cherenkov light, the inclusion of the environmental optical background extracted from real data, and the digitisation of the PMT signals following a run-by-run strategy. This strategy accounts for seasonal variations related to biological activities and for inefficiencies due to the ageing of the PMTs and to biofouling [27].

At the end of the full simulation chain, a set of MC files is available for each run of real data, stored in the same format. Simulated files are processed with the same reconstruction algorithms and analysis procedures used for the corresponding data. Monte Carlo neutrino events have been generated in the energy range $10 \leq E_\nu \leq 10^8$ GeV, separately for $\nu_e$, $\nu_\mu$, and their antineutrinos, and for CC and NC processes. Details on the simulation chain, hadronic model for cross sections, interaction kinematics and particle distribution functions are given in [26]. The same MC sample can be differently weighted to reproduce the conventional atmospheric neutrinos, the prompt neutrinos and theoretical astrophysical signals. In the present letter, the atmospheric $\nu_e$ and $\bar{\nu}_e$ fluxes are represented with the same models used in [28], namely, the conventional component follows the spectrum described in [12], extrapolated at higher energy as in [5], and the prompt contribution as calculated in [15]. The MC statistics for the atmospheric neutrino sample corresponds to more than two orders of magnitude than for real data.

Finally, for each data run, a file with simulated atmospheric muons (CR$\mu$) is produced with the MUPAGE package [29,30]: in this case, the equivalent MC livetime corresponds to 1/3 of the real run livetime.
Table 1
Number of events in different Monte Carlo samples surviving the preselection and the cut on track quality parameter $\Lambda$ (second column) and the final BDT cut (third column). The last row shows the number of data events in 3012 days of livetime. Cosmic neutrinos are the searched signal in neutrino telescopes; for this analysis they represent an additional background (see §4). The MC Cosmic $\nu$ expectation is computed assuming the flux estimated in [28].

<table>
<thead>
<tr>
<th>Preselection</th>
<th>$\Lambda &gt; -5.7$</th>
<th>+BDT &gt; 0.33</th>
</tr>
</thead>
<tbody>
<tr>
<td>MC CR$\mu$</td>
<td>136700</td>
<td>~3</td>
</tr>
<tr>
<td>MC Atmospheric $\nu_e$ CC</td>
<td>242</td>
<td>96</td>
</tr>
<tr>
<td>MC Atmospheric $\nu_e$ NC</td>
<td>22</td>
<td>9</td>
</tr>
<tr>
<td>MC Atmospheric $\nu_\mu$ CC</td>
<td>3780</td>
<td>620</td>
</tr>
<tr>
<td>MC Atmospheric $\nu_\mu$ NC</td>
<td>400</td>
<td>180</td>
</tr>
<tr>
<td>MC Cosmic $\nu$</td>
<td>30.4</td>
<td>9.2</td>
</tr>
<tr>
<td>MC total</td>
<td>141200</td>
<td>917</td>
</tr>
<tr>
<td>Data (3012 days)</td>
<td>133676</td>
<td>1016</td>
</tr>
</tbody>
</table>

3. Event selection: signal and background

Data collected from 2007 until the end of 2017 have been used. Only runs without high bioluminescence level have been selected. The total livetime corresponds to 3012 days. The background is almost entirely due to CR$\mu$’s: after trigger and reconstruction, the expected signal-to-background rate is $\sim 10^{-4}$. The background suppression is organised in three different steps.

An initial preselection of shower-like and starting track events combines information from both TANTRA and AAFit reconstruction algorithms, according to the following four requirements: i) the direction of the event as reconstructed by AAFit must be upward going (i.e., zenith angle $> 90^\circ$), to reject the largest fraction of CR$\mu$’s; ii) the TANTRA’s reconstructed event interaction vertex must be contained in a cylindrical volume of axial radius of 300 m and height of 500 m, centred at the centre-of-gravity of the detector modules; iii) the TANTRA estimated angular uncertainty on the event direction must be $< 30^\circ$ and the quality parameter $M_{\text{est}} < 1000$, to remove poorly reconstructed events; iv) upward going tracks with AAFit quality parameter $\Lambda > -5.2$ are discarded. This removes muons generated from neutrino interactions outside the detector volume that could have survived previous cuts. These through-going events have already been used in the previous measurement of the $\nu_\mu$ spectrum [10].

After this preselection, the MC signal is reduced by a factor of two with respect to the trigger and reconstruction level, with $\sim 350$ survived CR$\mu$’s for each atmospheric neutrino candidate.

The second step, following [10,24,31], uses the AAFit quality parameter, $\Lambda$. The best compromise to suppress the largest percentage of background while keeping a large enough fraction of signal events is obtained by removing events with $\Lambda > -5.7$. After this cut, 25% of the signal survives, with about 30 remaining background events for each atmospheric neutrino. Table 1 summarises the number of events passing the preselection and the $\Lambda$ cut for each MC sample. The last row shows the events in the data sample, after unblinding.

The final classification of events as signal or background is performed with a Boosted Decision Tree (BDT), defined on a multidimensional parameter space. A BDT is an algorithm that belongs to the family of supervised machine learning techniques. To build the classification function, training samples are necessary. CR$\mu$ events generated with MUPAGE constitute the background sample; CC+NC interactions of atmospheric $\nu_e$ are used as signal. The CC+NC interactions of $\nu_\mu$ are not used for training the algorithm to reject background. This choice is motivated by the fact that the $\nu_\mu$ flavour produces the cleanest case of shower-like events and it is the most difficult channel to measure in neutrino telescopes.

For each CR$\mu$ or $\nu_e$ event, the classifier is trained using the following 15 quantities. From the TANTRA algorithm, the reconstructed 1) zenith angle and 2) azimuth angle in the local reference frame; 3) interaction vertex coordinates; 4) quality parameter estimator, $M_{\text{est}}$; 5) number of detector lines with at least one hit; 6) total number of hits used to reconstruct the event; 7) angular resolution associated to a shower-like event. From the AAFit track-like algorithm, the reconstructed 8) zenith angle and 9) azimuth angle in the local reference frame; 10) track length inside the detector volume; 11) quality parameter estimator, $\Lambda$; 12) angular resolution associated to a track-like event. From the GridFit track-like algorithm, 13) the quality parameter; 14) the CR$\mu$ veto parameter, a likelihood variable based on time sequence and charge of the hits in different storeys of the detector, causally-connected under the assumption of a downward going, minimum ionising particle; 15) the number of on-time hits, which assumes that the photons are produced at the Cherenkov angle and arrive at the PMT unscattered.

A ranking of the BDT input variables is derived by counting how often each variable is used to split decision tree nodes, and by weighting each split occurrence by its squared separation gain and by the number of events in the node [32]. None of the variables is found to be significantly dominant; the variable with the highest ranking is the TANTRA zenith angle (1) with score 0.12, followed with the GridFit quality (13) with score 0.10; in the last two positions, the estimators of the angular resolution from TANTRA (8), with score 0.04, and that from AAFit (12), with score 0.023.

As shown in Fig. 1, the BDT output is an excellent discriminator between events in the atmospheric $\nu_e$ and background CR$\mu$ samples. The BDT distribution obtained from events induced by atmospheric $\nu_e$, CC+NC interactions is also included in the plot. As expected, this distribution resembles that of the $\nu_e$ signal. The BDT condition $-0.33 >$BDT $-0.2$ is due to the uncertainties in the modelling of the bulk of atmospheric muons in a parameter-space (i.e., direction, energy, muon bundle multiplicity) far from that of neutrino-induced events. The excess of events at large BDT
values is consistent with an overall normalisation of the MC atmospheric neutrino flux.

The last column of Table 1 shows the number of events in the final sample, after BDT cut. The neutrino signal is reduced to ~20% of that present in the previous step, while the background of atmospheric muons is removed from the generated MC sample. Only a maximum residual contamination of 3 events is estimated to survive in the data sample, figure obtained with an extrapolation of the BDT distribution. As expected, the neutrino signal is still dominated by atmospheric νμ producing starting tracks: only ~10% of the selected events originate from νe. At 1 TeV, the expected flux ratio is Φνμ/Φνe ~ 20.

4. Unfolding procedure and detector acceptance

In order to derive from data the νe and νμ energy spectra, an unfolding method is used. The two true distributions are deconvolved from the experimentally measured one, based on the best knowledge of the detector and on assumptions made on the interaction rates of the different neutrino flavours.

In counting experiments, events are grouped into certain regions of phase-space, called bins. The main observable quantities in neutrino telescopes are the neutrino direction and energy, which are measured only with finite precision due to inevitable detector effects. Consequently, an event may be assigned to a wrong bin.

The outcome of the unfolding procedure, folded with the detector acceptance and livetime, results in a spectrum that allows a direct comparison with other experiments. Two major classes of unfolding methods exist: algorithms based on matrix inversion or singular value decomposition, such as the TUnfold [33] algorithm used in this analysis; algorithms based on iterative methods or on the use of Bayes’ theorem [34]. A Bayesian approach has been used, e.g., by the Super-Kamiokande experiment [8] and in our previous measurement of the νμ energy spectrum using through-going muons [10]. For an overview of the commonly used unfolding algorithms, see also [35,36].

The TUnfold algorithm [33] is a widely tested and validated algorithm in the context of high-energy physics and it can handle one or more background sources. The algorithm allows to estimate the number of events in m bins of a true distribution xj, given an observed distribution of yi in n bins:

\[ y_i = \sum_{j=1}^{m} A_{ij} \cdot x_j + b_i, \quad 1 \leq i \leq n, \tag{1} \]

where each bin has a background contribution bi. Aij is a matrix of probabilities describing the migrations from bin j to any of the n bins. The method, interfaced to the ROOT analysis package [37], uses a least square method with Tikhonov regularisation [38] and a constraint to fix the total number of events. The least square minimisation requires a number of degrees of freedom such that \( n - m > 0 \), meaning that the data yi have to be measured in finer bins than extracted by the unfolding procedure.

The energy estimated by the TANTRA reconstruction algorithm, \( E_{\text{reco}} \), is used to construct the distribution of \( y_j \). Events in Fig. 1 with BDT> 0.33 are atmospheric νμ or νe, with a contamination of less than a few per mil from CRμ and a ~1% fraction of cosmic neutrinos; both samples are considered as background. The unfolding method requires a \((n \times m)\) matrix for the νμ and νe energies, with m bins of \( E_{\text{reco}} \) and m bins of true energy \( E_{\nu} \). Monte Carlo samples allow the construction of:

- \( A_{ij}^{\nu e} \), a \((6 \times 3)\) response matrix obtained with the simulated samples of νe CC+NC interactions;
- \( A_{ij}^{\nu \mu} \), a \((15 \times 5)\) response matrix obtained with the simulated samples of νμ CC+NC interactions.

The chosen number of bins in \((E_{\text{reco}}, E_{\nu})\) for the two samples provides the highest stability in terms of unfolding results applied on MC samples with the same number of events as real data. In the unfolding procedure, the use of \( E_{\text{reco}} \) is limited to energies between ~100 GeV and ~50 TeV. The lower bound is determined by the fact that our reconstruction algorithm cannot reliably reconstruct neutrino energies below 100 GeV. Above 50 TeV, the event statistics are significantly reduced by the requirement of the containment of the interaction vertex within, or near to, the instrumented volume. In addition, cosmic neutrinos, whose flux suffers large uncertainties, start to be the dominant “background”.

Fig. 2 shows the distribution of \( E_{\text{reco}} \) for the νμ sample (blue) with the bin size used for the construction of the \( A_{ij}^{\nu \mu} \) matrix. For completeness, the distribution of \( E_{\text{reco}} \) for the νe sample (red) is superimposed, although the distribution used for the construction of the \( A_{ij}^{\nu e} \) matrix has a different binning. The expected contribution in the sample of cosmic νs’ of all flavours is also shown.

Concerning the background terms, it includes an extrapolated contribution of 3 track-like CRμ events. Based on the behaviour of atmospheric muons before the BDT cut, the \( b_i^{CR\mu} \) terms are assumed to affect only the νμ sample, and uniformly in the \( E_{\text{reco}} \) range. The background from the cosmic neutrino flux (terms \( b_i^{\nu} \)) assuming equipartition (νμ: νμ: νe) = (1:1:1), contributes about equivalently to the νμ and νe samples, following the \( E_{\text{reco}} \) distribution shown in Fig. 2.

The unfolding procedure assumes that in the νμ distribution there are three background components corresponding to CR’s (\( b_i^{\nu CR} \)), 50% of cosmic neutrinos (\( b_i^{\nu} \)), and the νe fraction (\( b_i^{\nu} \)). In the νe distribution, there are two background components corresponding to 50% of cosmic neutrinos (\( b_i^{\nu} \)) and the νμ fraction (\( b_i^{\nu} \)). The algorithm assumes that \( \sum (b_i^{\nu \mu} + b_i^{\nu} + b_i^{\nu CR \mu} + b_i^{\nu}) \) is equal to the total number of events. The events \( b_i^{\nu \mu} \) and \( b_i^{\nu} \) are assumed to be produced with the fluxes given in [12], as is the default Monte Carlo simulation, with free normalisation. Possible variations in their spectral indexes are accounted for in the treatment of systematic effects (see §5).

Table 2 presents the information on the unfolded energy in 5 (3) bins for the νμ (νe) sample. The first two columns contain the energy range of the corresponding bin and the weighted central value of the neutrino energy bin, calculated taking into account the steep decrease of the energy spectrum and the detector response. The third column shows the unfolded number of data events as obtained by the algorithm.
5. The unfolded energy spectrum

To transform the unfolded number of events, \( N^{\text{evt}} \), given in Table 2 into a differential energy flux in the proper units (GeV\(^{-1}\) cm\(^{-2}\) s\(^{-1}\) sr\(^{-1}\)), the following steps are required: i) divide each bin by the livetime of 3012 days, obtaining the event rate integrated in the \( \log_{10} \) of the neutrino energy over the bin; ii) divide by the width of the bin (0.54 for \( \nu_\mu \) and 0.9 for \( \nu_e \)); then, transform the \( \frac{dN^{\text{evt}}}{d\log_{10}E_\nu} \) distribution into the \( \frac{d\phi}{dE_\nu} \) one; iii) divide by the integrated value of the observation solid angle, i.e., 2\( \pi \) sr; iv) divide by the detector effective area, \( A_{\text{eff}}(E_\nu) \), averaged over the distribution of zenith angles, as reconstructed by the TANTRA algorithm.

The effective area is the figure of merit for a neutrino telescope, representing the size of a 100% efficient hypothetical target that the detector offers to a certain simulated neutrino flux. It is calculated as

\[
A_{\text{eff}}(E_\nu) = \frac{N_{\text{sel}}(E_\nu)}{N_{\text{gen}}(E_\nu)} \cdot V_{\text{gen}} \cdot \rho N_A \cdot \sigma(E_\nu) \cdot P_{\text{Earth}}(E_\nu),
\]

where \( N_{\text{sel}}(E_\nu) \) and \( N_{\text{gen}}(E_\nu) \) are, respectively, the number of selected and generated events of a given neutrino energy \( E_\nu \), in the generation volume \( V_{\text{gen}} \); \( \rho N_A \) are the matter density and the Avogadro’s number; \( \sigma(E_\nu) \) is the neutrino cross section; \( P_{\text{Earth}}(E_\nu) \) is the probability of the neutrino to traverse the Earth without being absorbed. Above 100 GeV, there are no corrections needed for oscillation effects. Fig. 3 shows the effective area obtained from the selection of events described in this work.

The fourth column of Table 2 presents the differential flux obtained with the overall procedure. The reported statistical error is determined by the TUnfold method.

The result of the unfolding process depends on the MC simulation via the construction of the response matrix. In turn, the simulation depends on a number of parameters with associated uncertainties. The effects inducing systematic uncertainties on the measurement of the \( \nu_e \) flux through-going events have been extensively described in [10]. The same systematics affect both the \( \nu_\mu \) and \( \nu_e \) samples in this analysis and the effects are estimated in dedicated MC simulation datasets (either for \( \nu_\mu \) and \( \nu_e \)), by varying each time only one of the following parameters:

- Overall sensitivity of the optical modules, changed by +10% and -10%. This includes the uncertainty on the conversion of a photon into a photoelectron as well as the angular dependence of the light collection efficiency of each optical module.
- The uncertainties on water properties, by scaling up and down by 10% the absorption length of light in water with respect to the nominal value.
- The uncertainties related to the neutrino fluxes used in the default response matrix of the unfolding procedures, including a slope change of \( \pm 0.1 \) in the spectral index, independently for \( \nu_e \) and \( \nu_\mu \).

Each modified MC sample was then used as pseudo-data to construct a new response matrix, used for unfolding. The deviation in the content of each \( E_\nu \) bin from the spectrum obtained with the default response matrix, \( A_{\text{eff}}^i \), or \( A_{\text{eff}}^i \), corresponds to the systematic uncertainty associated with the parameter variation. For each energy bin, the total uncertainty is computed as the quadratic sum of each contribution, and the resulting value is reported in the last column of Table 2.

6. Results and conclusions

Fig. 4 shows the \( (\nu_e + \bar{\nu}_e) \) and \( (\nu_\mu + \bar{\nu}_\mu) \) fluxes measured in this work, together with the results from previous experiments. Our
unfolded atmospheric neutrino spectra, whose statistical errors are largely dominant over the systematic ones, are 20%-25% below the most recent computations using the SIBYLL-2.3c hadronic interaction model [17,39].

The measurement of the electron neutrino flux at high energy is challenging, because very large detectors are needed to collect sufficient statistics, and due to large systematic uncertainties. Each measurement of IceCube-DeepCore [4] and IceCube [5] relies on about 200 interacting $\nu_e$ in the polar ice medium. The present measurement is performed in seawater, under completely different environmental conditions and systematic uncertainties, yielding consistent results with the ones obtained in polar ice. During a lifetime of 3012 days, $\sim$130 $\nu_e$ interactions have been reconstructed within the instrumented ANTARES volume. The statistics of the $\nu_e$ sample is not sufficient to test models above a few tens of TeV, where a significant cosmic flux is present and the transition from the conventional to the prompt flux is expected. Below 100 GeV, the PMT density of the ANTARES detector is insufficient to reconstruct a significant number of events.

Concerning the unfolded $\nu_e$ flux, our previous measurement [10] with a sample of $\sim$650 through-going events collected in 855 days of livetime and generated by neutrino interactions external to the instrumented volume almost superimposed the SIBYLL-2.3c model. The present analysis relies on a totally independent data sample, provided by neutrinos whose reconstructed interaction vertex is inside (or nearby) the instrumented volume of the detector. The $\nu_e$ data sample, summing the N$^{\text{ext}}$ events in Table 2, corresponds to $\sim$850 events. By scaling the relative livetimes, the number of events in this sample would be $\sim 1/3$ of the through-going muons reconstructed as in [10]. However, due to the superior energy estimate of these (semi)-contained events with respect to the through-going sample, the overall uncertainty on the measured flux is smaller with respect to our previous measurement in the central bins, reaching a precision equivalent to that obtained by the Super-Kamiokande Collaboration in [8]. The present unfolded $\nu_e$ flux is more close to that of IceCube with 40 strings (2011) [6] and is 20%-25% below both the flux reported in our previous measurement and the one reported by IceCube using 59 strings (2015) [7], although consistent within errors.

Declaration of competing interest

The authors declare that they have no known competing financial interests or personal relationships that could have appeared to influence the work reported in this paper.

Acknowledgements

The authors acknowledge the financial support of the funding agencies: Centre National de la Recherche Scientifique (CNRS), Commissariat à l’Energie Atomique et aux Énergies Alternatives (CEA), Commission Européenne (FEDER fund and Marie Curie Program), Institut Universitaire de France (IUF), Labex UnivEarthS (ANR-10-LABX-0023 and ANR-18-IDEX-0001), Région Île-de-France (DIM-ACAV), Région Alsace (contrat CPER), Région Provence-Alpes-Côte d’Azur, Département du Var and Ville de La Seyne-sur-Mer, France; Bundesministerium für Bildung und Forschung (BMBF), Germany; Instituto Nazionale di Fisica Nucleare (INFN), Italy; Nederlandse Organisatie voor Wetenschappelijk Onderzoek (NWO), the Netherlands; Council of the President of the Russian Federation for young scientists and leading scientific schools supporting grants, Russia; Executive Unit for Financing Higher Education, Research, Development and Innovation (UEFISCDI), Romania; Ministerio de Ciencia e Innovación (MCI) and Agencia Estatal de Investigación: Programa Estatal de Generación de Conocimiento (refs. PGC2018-096663-B-C41, -A-C42, -B-C43, -B-C44) (MCI/FEDER), Severo Ochoa Centre of Excellence and MultiDark Consolider, Junta de Andalucía (ref. SOMM17/6104/UGR and A-FQM-053-UGR18), Generalitat Valenciana: Grisolia (ref. GRISOLIA/2018/119) and GenT (ref. CIDEGENT/2018/034) programs, Spain; Ministry of Higher Education, Scientific Research and Professional Training, Morocco. We also acknowledge the technical support of Irrometer, AIM and Foseley Marine for the sea operation and the CC-IN2P3 for the computing facilities.

References