Imaging Jupiter's radiation belts down to 127 MHz with LOFAR


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Imaging Jupiter’s radiation belts down to 127 MHz with LOFAR


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ABSTRACT

Context. With the limited amount of in-situ particle data available for the innermost region of Jupiter’s magnetosphere, Earth-based observations of the giant planets synchrotron emission remain today the sole method to scrutinize the distribution and dynamical behavior of the ultra energetic electrons magnetically trapped around the planet. Radio observations ultimately provide key information addressing the origin and control parameters of the harsh radiation environment known as of today.

Aims. Here we perform the first resolved and low-frequency imaging of the synchrotron emission with LOFAR. At a frequency as low as of 127 MHz, the radiation from electrons with energies of ~1–30 MeV are expected, for the first time, to be measured and mapped over a broad region of Jupiter’s inner magnetosphere.

Methods. Measurements consist of interferometric visibilities taken during a single 10 hour rotation of the Jovian system. These visibilities were processed in a custom pipeline developed for planetary observations, combining flagging, calibration, wide-field imaging, direction-dependent calibration and specific visibility correction for planetary targets. We produced spectral image cubes of Jupiter’s radiation belts at various angular, temporal and spectral resolutions from which flux densities were measured.

Results. The first resolved images of Jupiter’s radiation belts at 127–172 MHz are obtained, with a noise level ~20–25 mJy/beam, along with total integrated flux densities. They are compared with previous observations at higher frequencies. A larger extent of the synchrotron emission source (~4 R_J) is measured in the LOFAR range, that is the signature – as at higher frequencies – of the superposition of a “pancake” and an isotropic electron distribution. Asymmetry of east–west emission peaks is measured, as well as the longitudinal dependence of the radial distance of the belts, and the presence of a hot spot at ~230 ± 25°. Spectral flux density measurements are on the low side of previous (unresolved) ones, suggesting a low-frequency turnover and/or time variations of the Jovian synchrotron spectrum.

Conclusions. LOFAR proves to be a powerful and flexible planetary imager. In the case of Jupiter, observations at 127 MHz depict the distribution of ~1–30 MeV energy electrons up to ~4–5 planetary radii. The similarities of the observations at 127 MHz with those at higher frequencies reinforce the conclusion that the magnetic field morphology primarily shapes the brightness distribution features of Jupiter’s synchrotron emission, as well as how the radiating electrons are likely radially and latitudinally distributed inside about 2 planetary radii. Nonetheless, the detection of an emission region that extends to larger distances than at higher frequencies, combined with the overall lower flux density, yields new information on Jupiter’s electron distribution, information that ultimately may shed light on the origin and mode of transport of these particles.

Key words. Jupiter – radiation belts – synchrotron emission – radio interferometry – LOFAR

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1. Introduction

Jupiter is among the most intense radio emitters in our Solar System. It has a strong magnetic field dominated by a dipole component of moment ~4.3 R_J^3 G (1 G = 10^-4 T, 1 R_J ≈ 71 492 km), much larger than that of the Earth (Baganu et al. 2014). This dipole is tilted by ~9.6° relative to the rotation axis, toward a longitude of ~200° (Fig. 1). The rotation of this magnetic field with a period of 9h55m29.71s defines a coordinate system of reference called “System III” (1965) as described in Dessler (1983). Its interaction with the Solar wind creates a large magnetosphere, in which charged particles are accelerated to keV – MeV energies. Three main radio components are produced by Jupiter and its magnetosphere (Fig. 1): (1) the thermal emission coming from the planetary disk dominates the spectrum above ~4 GHz with a brightness temperature of ~150 K (Kloosterman et al. 2008; Hafez et al. 2008); (2) auroral emission is produced below 40 MHz by electrons accelerated to keV energies in the magnetosphere at ~20–50 R_J from the planet, and then precipitated along magnetic field lines toward high latitudes where they...
produce aurorae and associated cyclotron radio emission; (3) synchrotron emission is produced between ~30 MHz and ~30 GHz by electrons accelerated to MeV energies and trapped in the so-called radiation belts of the inner magnetosphere, within a few radii of the surface, mostly at low latitudes. In the present paper, we are interested in this synchrotron emission.

Fig. 1. a) Sketch of the location of Jovian radio sources in Jupiter’s inner magnetospheric field lines (in orange): (1) thermal, (2) auroral, (3) radiation belts. Ω, Ω, and M are respectively the rotation vector, the magnetic field and the magnetic moment. The angle β is the tilt between the rotation axis and the magnetic moment. b) Corresponding typical spectra (in Jansky normalized to 4.04 AU, adapted from Zarka [2004], with indication of LOFAR’s low (LBA = low-band antennas) and high (HBA = high band antennas) spectral bands. DAM and DIM are the official denomination of the decameter and the decimeter emissions.

Since its discovery in the mid-fifties, the synchrotron emission has been imaged between 330 MHz and 22 GHz using various instruments (VLA, WSRT, ATCA, . . . ), and a few un-resolved measurements have also been performed down to 74 MHz (VLA, CLFST) (see de Pater et al. [2003]). Ground-based synchrotron measurements provide valuable information about the angular and frequency distribution of high-energy electrons trapped in Jupiter’s inner radiation belt (<6 R_J). Relying on the well understood physics of synchrotron emission, they are used to test physical models of the radiation belts, incorporating various physical processes such as radial diffusion of the electrons, interaction with the magnetospheric plasma, satellites, rings and plasma waves, and synchrotron losses (see e.g., de Pater [1981]; de Pater et al. [1997]; Santos-Costa [2001]; Bolton et al. [2004]; de Pater [2004]; Santos-Costa & Bolton [2008].

Energetic electrons are in fast gyration around Jupiter’s magnetic field lines. This gyration motion is associated to an invariant that is the magnetic moment of the electron $E_{\text{L}}/B$, that causes bouncing of the electrons between magnetic mirror points where the parallel velocity reverses and the pitch angle is ~90°. The synchrotron emission taps the perpendicular energy of the electrons and is beamed in the direction of electron motion. As a consequence, the bulk of emission comes from electrons having a velocity near-perpendicular to magnetic field lines. Because the Earth always lies within ~13° of the Jovian magnetic equator, the field lines are themselves near-perpendicular to the line of sight (Fig. 1). An accumulation of such particles exists around Jupiter’s magnetic equator (for the electron’s with a large equatorial pitch angle, trapped between magnetic mirror points at low latitudes) and at high magnetic latitudes (where the mirror points of energetic electrons with small pitch angles lie; due to their small parallel velocity there, electrons reside a long time near these mirror points, leading to enhanced synchrotron emission). Since the emitted power is proportional to $E^2 \times B^2$, the peak frequency proportional to $E^2 \times B^2$, with $E$ the electron’s energy and $B$ the magnetic field strength at the source, synchrotron spectra as well as images at different frequencies allow us to probe the distribution of electrons at various energies in Jupiter’s inner magnetosphere. Lower radio frequencies are associated with lower energy electrons (typically several MeV) in a strong B field and higher energy electrons at greater distances from the planet (i.e., in a weaker magnetic field). Hence, it is difficult to disentangle the energy distribution of the electron in observations without any spatial resolution, since this information is entangled with information about the pitch angle distribution and the line-of-sight integration through a complicated magnetic field configuration. High resolution imaging is thus crucial to derive sound constraints. No resolved image has been obtained yet below 330 MHz (de Pater 2004). It is in particular interesting to map Jupiter at frequencies in the 70–300 MHz range since the disk-integrated spectral measurements are suggestive of a turnover in the spectrum at these lower frequencies (de Pater & Butler [2003]). Moreover, at LOFAR frequencies, resolved imaging is valuable as it enables “scanning” for the first time the 1–30 MeV electron population through their contribution to the synchrotron emission located further away from the planet. In the equatorial plane, assuming a dipolar magnetic field in the region from 1 R_J to 4 R_J, the detectable synchrotron emission at 1.4 GHz is associated to electrons with energy ranging from 7.9 MeV up to 67 MeV. With a rule of thumb, at 127 MHz, in the same region, we can probe electrons populations from to 2 MeV up to 20 MeV. Therefore, the study of the resolved synchrotron emission with LOFAR at low frequencies and in distant regions of the belts contribute to constraint the electron populations originating from the middle magnetosphere and undergoing inward diffusion and acceleration processes. In this paper, we present the first resolved images of Jupiter’s synchrotron emission obtained (with LOFAR) at a frequency as low as 127 MHz, as well as disk-integrated spectral measurements, and we derive preliminary constraints on the morphology and variability of the emission. Observations and the custom pipeline that we developed for analyzing LOFAR planetary observations are described in Section 2. The resulting images and spectra are presented in Section 3, and quantitatively analyzed in Section 4. Section 5 discusses these first low-frequency observations and perspectives for further studies.

2. Observations and planetary imaging pipeline

2.1. Observational requirements for planetary imaging

Planetary imaging requires a special observing strategy and calibration as compared to other radio observations. For Jupiter, two main effects have to be taken into account: i) the proper motion of Jupiter on the sky background, ii) the intrinsic motion of the radiation belts in the reference frame of the planet.

First, as we observe from the ground, Earth’s (and Jupiter’s) orbital motion induces an additional apparent motion of planetary targets with respect the rest of the sky, the apparent motion of which is due to Earth’s rotation. This cause the planetary source to travel over the course of the year between radio sources with the consequence of impacting the calibration of long integrated observations. This motion is relatively fast for Jupiter, causing a shift of 3.16 arcminutes – i.e., nearly 4 times its diameter – during one 10h planetary rotation, relative to the “fixed” RA/DEC sources (e.g., NVSS source J020457+114145). Although this is large compared to our ~7″ synthesized beam width, it is much smaller than the ~5° primary beam of the telescope).
Second, Jupiter’s radiation belts are fixed – at zero order – relative to the Jovian magnetic field. But as Jupiter’s magnetic dipole axis is tilted by 9.6° with respect to its rotation axis, the former precesses around the latter with Jupiter’s rotation (see Fig. 1). As a consequence, the magnetic equator and the whole image of the radiation belts wobbles or rocks on the plane of the sky at the System III period (Fig. 2). This rocking is discernable between panel a) and d) of Fig. 2 where the main axis of the image of Jupiter has changed in orientation.

As a third and minor effect, the varying distance between the observer and Jupiter has to be taken into consideration. Therefore, all measured flux densities must be scaled to the common reference distance of 4.04 AU to enable comparison between epochs.

Jupiter’s synchrotron emission is a few jansky (1 Jy = 10⁻²⁶ Wm⁻²Hz⁻¹) radio source that is resolved by LOFAR; therefore a long time-integration is necessary to obtain an image with a good signal-to-noise ratio (SNR, defined as the ratio of the peak flux to the background r.m.s. noise). If no precaution is taken while producing long time and frequency integrated images, the displacement motion of Jupiter and the rocking motion of its radiation belts will lead to a large smearing of their image (in addition to time and frequency smearing effects which also slightly distorts the shape of the sources located at the edge of the beam). The former motion must be compensated in the Fourier domain via a time-dependent translation of the phase center including antenna delay correction, and the latter by a rotation of the visibility reference frame, i.e., the (u,v) axes, prior to imaging. However, these corrections should only be applied to Jupiter visibility data, otherwise they will cause a systematic smearing of other fixed coordinate radio sources in the field, increasing the difficulty of imaging sources that are no longer point sources.

Therefore, to enable posterior correction of these effects, an observation should be carried toward an arbitrary pointing direction with a constant RA/DEC coordinate. As illustrated in Fig. 2, the target beam was pointed to Jupiter’s mean RA/DEC position during the full observing period. We can see the motion and rocking effects in the preliminary LOFAR images derived for five consecutive 2 hours intervals.

2.2. Observational setup

The observations analyzed here were recorded during the commissioning period of LOFAR. They consist of visibilities recorded within a 10 hour window from 18:24 UT on 2011/11/10 to 04:24 UT on 2011/11/11. At the observation epoch, Jupiter was at 3.99 AU from Earth and subtended an angle of ~49" in the sky. The Sun–Earth–Jupiter angle was 165° and the Earth was at a Jovigraphic latitude (DJ) of ~3.29° (www.imcce.fr).

Observations were carried out using 49 High Band Antenna (HBA) fields (2 per station for 20 Core Stations + 1 per station for 9 Dutch Remote Stations (see [van Haarlem et al. 2013]). Two station beams were synthesized by phasing the antennas at station level: a “target” beam centered on Jupiter, and a “calibration” beam centered on the radio source 4C15.05 (for phase calibration) four degrees away from Jupiter. The approximate half power beam width (HPBW) of the beams is ~5° at 150 MHz. The same ~23 MHz of total bandwidth were recorded from each beam, in the form of 121 sub-bands of 195 kHz, in twelve groups of ten contiguous subbands (each group is therefore 2 MHz wide), regularly distributed between 127 and 172 MHz (hence a spectral coverage of 50%). The raw data consist of complex visibilities produced at ~1 s time resolution and in 3 kHz-wide frequency channels for all available baselines. Base-line lengths were distributed between ~15 λ and ~30 kλ (with λ the wavelength). The (u,v) radial density peaks at ~500 λ (corresponding to Core Station baselines) and is then approximately flat up to 30 kλ, providing a maximum theoretical angular resolution of ~6.5°. The two co-polarization (XX, YY) and two cross-polarization (XY, YX) terms were recorded, but only total intensity I measurements were reliable at this early stage of LOFAR exploitation, thus we limit our present analysis to those measurements and we did not exploit the Q, U and V data.

2.3. Flagging and direction-independent calibration

A classical data pre-processing was applied (van Haarlem et al. 2013): flagging of radio frequency interference (RFI) using the AOflagger (Offringa et al. 2012) followed by time integration on 3 s steps and by frequency integration on 195 kHz (1 LOFAR sub-band) by the LOFAR NDPPP pipeline (Pizzo 2015), calibration using the phase calibrator field using BBS (Pandey et al. 2009), then derivation of complex gain solutions for all antennas in 9 s bins (i.e., 1 gain solution every 3 time bins). Gain amplitudes and phases were then visually inspected and bad data were flagged. The gain solutions were significantly more noisy during approximately the first and last hour of the observation (due to the low elevation of the source and probably the ionosphere turbulence state). Strong radio sources such as the “A”-team (e.g., Cas A, Cyg A, Vir A, Tau A, . . .) can contaminate LOFAR data if they are present in the station side lobes. As the HBA band is less affected than the LBA by the A-team, and as visual inspection of visibilities did not reveal the contribution of any A-team source in the data, we did not apply any specific treatment to these strong radio sources.

2.4. Direction-dependent calibration, background subtraction and proper motion correction

To alleviate the spatial smearing caused by the planetary corrections in the visibility plane, we need to detect and subtract all other radio sources in the data to improve the dynamic range of the image of the target. Because the field of view (FoV) of the LOFAR stations is large (~5° HPBW at ~150 MHz), wide-field imaging within the full FoV is required.

Thus, the planetary imaging pipeline that we developed includes the following steps: (i) make a wide-field image of the target field from the calibrated visibilities; (ii) detect in the image the sources other than Jupiter above a given threshold and identify them using a radio source catalog; (iii) subtract these sources with direction-dependent (DD) calibration solutions; (iv) apply the above motion corrections to the peeled visibilities and (u,v,w) coordinates; (v) build final Jupiter images integrated over selected intervals of time and frequency.

For building the wide-field image (i) we used the AWMager (Tasse et al. 2013) that does beam correction (A-projection) and wide-field imaging corrections (W-projection, Cormell et al. 2008). Automatic source detection (ii) was performed using the Duchamp source finder (Whiting 2012) and a sky model creator buildsky (Yatavwatta et al. 2013 and references therein). Most of the detected sources could be associated with the GSM (Global Sky Model, Pizzo 2015) that contains radio sources from the VLSS, NVSS and WENSS surveys. The GSM provides a realistic model of the sky with reliable flux densities and spectral indices. However at LOFAR wavelengths, the spectral index of some radio sources decreases, which introduces a systematic bias when their flux densities are extrapolated from high fre-

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Fig. 2. Calibrated images of the Jupiter in a 35′ × 35′ field, integrated over 20 sub-bands of 195 kHz within the range 166–172 MHz, for five consecutive 2 hours intervals (indicated on each panel). The selected subset of baselines provides an angular resolution of 35″. Pixel size is 5″. All panels are centered around the target phase direction (in RA/Dec). The belts are resolved near the center of the image and the beam is displayed in the bottom-left corner in yellow. The point source, east of Jupiter, is J020457+114145. Observing time and CML (CML = central meridian longitude in System III) are indicated on each panel. The motion of Jupiter from image to image is clearly visible. The rocking of the main axis of the image of Jupiter is also discernable. The last image is more distorted, due to lower quality data in the last time window.

Fig. 3. Wobbling of the Jovian magnetic equator (red circle) on the plane of the sky during one planetary rotation. The rings (in black) lie in the rotational equator. The blue line represents the main axis of the projection on the sky of the magnetic equator, which should also be the main axis of the radiation belts image as a function of time or longitude. Observer’s longitude (CML) is indicated on each panel as well as the reference meridian (in red).

The rocking of the radiation belts have indeed been corrected. We also note that the detailed shape of the radiation belts varies from panel to panel, which we attribute to the limited SNR of each image. The last image is very distorted, due to the noisy character of the last portion of the data and the poor (a.u.) coverage due to the low elevation of the source at the end of the observation (≈10°). A more detailed analysis of intermediate images shows that the interval with highest quality data is the 7 hour interval from ~19:00 to ~02:00 UT, that we used for building the 12 rotation-averaged images (not displayed). Finally, from these 7 hours and the entire 23 MHz bandwidth of observation, we built the time-and-frequency-averaged image of Fig. 6, which is the first resolved image of Jupiter obtained in the 127–172 MHz band. The residual noise in this image is 4.7 mJy beam⁻¹. At the extremity of the extended emission (around the 30 % of the peak flux), the “local” SNR ratio is ≈14.

In order to calibrate the flux density in the images, and to derive total integrated flux densities over the entire radiation belts that can be compared to previous measurements, we have performed source-integrated flux measurements similarly on Jupiter and on 3 nearby sources (Fig. 7a), before the DD subtraction step (iii), in each of the twelve 2 MHz bands. We have compared the measured total flux at each frequency with the spectra deduced from the catalogued fluxes and spectral indices for the 3 nearby sources. Measured values lie within 30 % of values deduced from catalogues for all 3 sources surrounding Jupiter.

Step (v) consist of deconvolution and image cube creation. At step (i), we used AWImager [Tasse et al. (2013)] to perform wide field imaging of all sources with beam corrections and W-projection. After steps (iii) and (iv), the visibilities are mainly dominated by the synchrotron emission from the radiation belts in a small region near the center of the field. Therefore, we used the Cotton-Schwab CLEAN algorithm implemented in CASA [NRAO (2013)] to produce final images of the radiation itself.

2.5. Image and spectrum processing of Jupiter’s radiation belts

We have built a 12 × 5 image cube (one image per 2 MHz band and per 2 hours of integration), 5 frequency-averaged images (one image per 2 hours, integrated over the 23 MHz of bandwidth between 127 and 172 MHz), and 12 rotation-averaged images (one image per 2 MHz band, integrated over ~7 hours – from ~19:00 to ~02:00 UT).

The five frequency-averaged images are displayed in Fig. 5 centered on the position of Jupiter. We can see that the shift and the rocking of the radiation belts have indeed been corrected. We also note that the detailed shape of the radiation belts varies from panel to panel, which we attribute to the limited SNR of each image. The last image is very distorted, due to the noisy character of the last portion of the data and the poor (a.u.) coverage due to the low elevation of the source at the end of the observation (≈10°). A more detailed analysis of intermediate images shows that the interval with highest quality data is the 7 hour interval from ~19:00 to ~02:00 UT, that we used for building the 12 rotation-averaged images (not displayed). Finally, from these 7 hours and the entire 23 MHz bandwidth of observation, we built the time-and-frequency-averaged image of Fig. 6 which is the first resolved image of Jupiter obtained in the 127–172 MHz band. The residual noise in this image is 4.7 mJy beam⁻¹, giving a peak SNR of 37. At the extremity of the extended emission (around the 30 % of the peak flux), the “local” SNR ratio is ≈14.

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We have carried out a first analysis of the LOFAR images and spectrum. In Fig. 5 after planetary motion and wobble correction, we note that the brightness maximum peak is first located on the west side of the planet and is located in the east side approximately half a rotation later. This effect is relatively well known and associated to the “beaming” curve highlighting the variation of the peak maxima with CML over a rotation and depending of the observing geometry (controlled by the observer latitude $D_{\phi}$, (Dulk et al. 1999a). To backup this assertion, modeling of the electron population is required as well as a synchrotron model. We reproduced the situations of panel b) and d) of Fig. 5 with simulated synchrotron images (Fig. 9) derived from Salammbô-3D particle code coupled with a synchrotron imaging model taking into account LOFAR observation parameters (time/CML, coverage, frequency band and angular resolution). Salammbô-3D was originally developed for Earth radiation belts computation, but was later adapted to Jupiter’s belt system and used to study the dynamics of inner belts (Bourdarie et al. 1996; Santos-Costa et al. 2001; Santos-Costa 2001; Sicard et al. 2004; Sicard & Bourdarie 2004). At present, the code uses the [O6 + Khurana] coupled magnetic model (Khurana 1997) and models the dynamics of electrons from 0.025 to 712 MeV, each contributing to the synchrotron emission at different frequencies and in different regions of the inner belts. While being refined for Earth and Jupiter, it was also adapted to other magnetic bodies such as Saturn (Lorenzato et al. 2012). The output of the simulation (assuming an infinite angular resolution) was convolved by the median beam over the whole frequency range in the two CML ranges. These simulations show that the maximum brightness peak effectively changes sides from west to east over few hours, consistent with the LOFAR observation in Fig. 5. The radial position of the brightness peaks and the extent of the belts are also consistent with the observation. Such preliminary comparisons suggest that a further quantitative investigation at all CMLs with detailed fitting of the physical model of the radiation belts (and of the electron populations) can lead to an accurate understanding of their morphology at low frequencies, an area that is fairly unexplored in a resolved regime of time, frequency and angular resolution. Exploitation of new wide-band data (using LOFAR & WSRT) and a complete modeling using this kind of particle code is currently ongoing and will be subject to a future publication.

In Fig. 6 (right), the mean synchrotron emission appears to extend above the noise level up to a distance of $\geq 4 R_J$ of Jupiter’s center, farther out than images at higher frequency (represented as contours derived from C-band VLA data taken in 1997 by Santos-Costa et al. 2009 and convolled down to the LOFAR angular resolution). Especially, at a distance of $\sim 3.5 R_J$, the brightness is 10% of the peak flux at high frequency whereas at low frequencies, still 30%–40% of the brightness is present at the same location, suggesting a larger extent of the radiation belts at low frequencies. This concurs with the samples of in-situ particle data collected by Pioneers 10/11 and the Galileo probe/orbiter in the early 1970s and late 1990s and early 2000s, which have showed dense population of electrons with energies of $\sim 1$–30 MeV in Jupiter’s inner magnetosphere. This is compatible with VLA observations of Santos-Costa et al. (2014) where the radiation zone of Jupiter at P band is observed to be slightly more extended than at L and C bands (quiet state or while varying). As in images at higher frequencies, the intensity distribution in the image reveals a near-equatorial “pancake” distribution.
of electrons (with equatorial pitch angles close to 90°) plus high-latitude lobes which require a component with a more isotropic distribution of pitch angles near L = 2.

We have measured the position of the “east” and “west” emission peaks as a function of frequency and time in (respectively) time-integrated and frequency-integrated images. The results are displayed in Fig. 10: in panels a) and b) an offset is measured between the average radial distance of the east maximum (1.51 R_J) and that of the west maximum (1.36 R_J). The accuracy of this measurement is limited to the size of the synthesized PSF for each of the reconstructed images. Although the determination of the peak flux is precise to the pixel level, we estimate the global uncertainty on the true position of the peaks to be ~0.5 R_J (as depicted by the error bars). Even with the lack of precision in our measurements, such east–west asymmetry was already observed (e.g. Dulk et al. [1997] and Santos-Costa et al. [2009] at 5 GHz). It reveals the local time (dawn–dusk) asymmetry of the inner Jovian magnetosphere, also visible in the radial distance of the Io plasma torus and attributed to the presence of an east–west electric field (see Brice & Mcdonough [1977], Smyth & Marconi [1998], Kita et al. [2013], and references therein). The time variations displayed in panels c) and d) of Fig. 10 measured at a few time steps, are consistent with radial excursions measured at higher frequencies (~0.25 R_J from 1.45 to 1.7 R_J in Dulk et al. [1997]). Those are due to the longitudinal asymmetries of Jupiter’s internal magnetic field that cause the average distance of the radiation belts to be dependent on the longitude, combined with projection effects on the sky at various phases of the planetary rotation. More accurate measurements are required to investigate this effect at low frequencies.

Panel a) of Fig. 12 displays the peak intensity (in Jy/beam) measured on the east and west sides of Jupiter in each of the 5 frequency-averaged images, as a function of the CML at the middle of the 2 hour interval corresponding to each image. Following Dulk et al. [1999a,b] and as illustrated in Fig. 11 although the observed emission from any point of the image results from integration along the line of sight through the optically thin radiation belts, the main contribution to the intensity observed at a given CML from the east side originates from a “source” at System III longitude λ_{III} = CML+90°. Conversely, intensity observed on the west side originates from a “source” at λ_{III} = CML−90°. Assuming that source characteristics (i.e., synchrotron emissivity at any point of the radiation belts) do not vary at timescales shorter than Jupiter’s rotation period and
that asymmetries in the magnetic field between the east and west sides can be ignored, it is possible by shifting by ±90° the observed points on Fig. 12a to deduce a profile of the peak emissivity as a function of longitude, displayed on Fig. 12b as a solid line. To test the consistency of the above transformations from CML to Δλ, we adjusted separately the east and west measurements by a spline function and inferred intermediate values at each longitude where a measurement exists on the other side of Jupiter (open diamonds), resulting in pairs of values (derived from east and west peaks) at each longitude where one actual measurement exists. The two (dashed) profiles deduced from east and west peak intensities display similar overall variations.

A broad hot spot is observed around Δλ = 230° ± 25°, that was already noted in previous observations at higher frequencies (Branson 1968, Conway & Stannard 1972), and was suggested to be caused by the geometry of Jupiter’s magnetic field configuration (de Pater 1980, 1981), east-to-west peak intensity ratios derived from panels a) and b) of Fig. 12 are plotted in panels c) and d). The east/west ratio as a function of CML is reminiscent – albeit with a lower amplitude – of that measured at higher frequencies (e.g. de Pater et al. 1997, Kloosterman et al. 2005, Santos-Costa et al. 2009).

The amplitude of the emission in our work is much lower, probably due to a combination of the long integration time (2 hr), the lower angular resolution in our images, and perhaps the lower frequency content of the source. As shown in Kloosterman et al. (2005), the detailed curves of the east/west ratio as a function of CML depend on the declination of the Earth relative to Jupiter (D_E). D_E was different in each case (∼3.3° for Leblanc et al. 1997, 0.07°–0.34° for Santos-Costa et al. 2009, and +3.29° in our observation), so that these comparisons are necessarily preliminary.

Finally, our spectral measurements of Fig. 8 are ∼35% lower than earlier measurements from 1998 at the same frequency (de Pater et al. 2003), i.e., marginally compatible with them taking into account our rather high estimated error bar (∼30%) on the LOFAR flux density measurements. But they are significantly lower than the model fit to the VLA measurement from 1994 (de Pater et al. 2003). This suggests a possible turnover of the spectrum below ∼300 MHz and/or time variations of the spectral flux density overall (such as shown by the 1998 vs. 1994 data in Fig. 7), or just at low frequencies.
Synchrotron emission is a well-understood process. Observations allow us to probe the energetic electron population in the inner magnetosphere. At low frequencies, the thermal component is negligible, so that the emitted power is proportional to the peak position with frequency is \(1.43 \pm 1.67\ R_J\) for the east peak and \(1.30 \pm 1.78\ R_J\) for the west peak. With time, it is \(1.13 \pm 1.70\ R_J\) for the east peak and \(1.22 \pm 1.54\ R_J\) for the west peak. Shaded surfaces represent an uncertainty of \(\pm 1\ R_J\).

**4. Discussion and conclusion**

LOFAR proves to be a powerful and flexible planetary imager, providing images complementary to the VLA at a spatial resolution only 4 times lower than the typical resolution at the VLA, but at frequencies >10 times lower, due to its long baselines. We have obtained here the first resolved images well below 300 MHz. Although still not perfect (LOFAR was still at commissioning stage) the image of Fig. 6 roughly agrees with maps at higher frequencies; the shape of the emission confirms that two electrons populations coexist: a pancake and an isotropic one. The latter produces emission at high latitudes (near electron mirror points). We have characterized east–west or longitudinal asymmetries. Although the uncertainty of the LOFAR flux density is high (~30%), the disk-integrated data points are stay marginally compatible with previous observations, suggesting the possible existence of a spectral turnover below 300 MHz and/or time variations of the spectrum.

LOFAR is now fully running. Further observations can be done with 24 core and 14 (possibly 16) remote stations (compared to 20 core and 9 remote stations in the paper) which will improve significantly the sensitivity and the angular resolution by the increase of long baselines (216 in this paper compared to 427 with 14 remote stations). Along with advanced hardware and software applied to the data, the set of 12 international (European) stations brings us the maximal baseline to 1500 km (instead of ~100 km in the paper). The Low Band of LOFAR will permit imaging of the Jovian synchrotron emission down to 40 MHz (upper limit of the decameter emission) and even less taking into account predictable absences of DAM emission (Cecconi et al. 2012), bringing the first very low frequency images of Jupiter’s radiation belts, diagnosing very low energy electrons and weak magnetic fields. Along with new LOFAR observations, joint synchrotron emission modeling is necessary.

Another campaign was conducted on 19–20 Feb. 2013 with LOFAR LBA & HBA and simultaneous Westerbork Synthesis
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Ratio Telescope observations at higher frequencies, ensuring together a frequency coverage from 50 MHz to 5 GHz ($\lambda = 6$ m to 6 cm). It will allow us to address spectral variations and the search for a low-frequency turnover. Further studies will also rely on the analysis of the polarization of the emission (long known to be dominantly linear, e.g., [Radhakrishnan & Roberts 1960] de Pater 1980). Advanced imaging methods such as sparse image reconstruction (Garsden et al. 2015) Girard et al. 2015 of the extended emission may improve the quality of snapshot images to better constrain the shape of the belts in smaller CML integration windows. Finally, synchrotron observations in the context of the JUNO mission around Jupiter will also be of high interest, as JUNO will provide in situ particle measurements and a very accurate model of the Jovian magnetic field (Bagenal et al. 2014).

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Appendix A: Coordinate transformations and Jupiter tracking

Appendix A.1: Phase center correction

We defined the rotation $R_u$ (resp. $R_v$), the rotation of angle $-\alpha_0$ (resp. $\delta_0$) around the axis $w$ (resp. $u$) in the $(u,v,w)$ space. The $(\alpha_0,\delta_0)$ defines the equatorial coordinates of the phase center, which was maintained constant during the observation. We want to apply the angular transformation from $(\alpha_0,\delta_0)$ to $(\alpha_t,\delta_t)$ where $\alpha_t$ and $\delta_t$ are the time-dependent center coordinates of the Jupiter disk during the observation. We used the ephemeris from the Institut de Mécanique Céleste et de Calcul des Éphémérides (IMCCE) to locate the center of the disk in equatorial coordinates. The correction was performed at a 5 minute rate. Given the orientation of the declination and right ascension axes, we can define two rotation matrices around the axis $w$ and the axis $u$ as follows:

$$R_u(-\alpha_t) = \begin{pmatrix} \cos \alpha_t & \sin \alpha_t & 0 \\ -\sin \alpha_t & \cos \alpha_t & 0 \\ 0 & 0 & 1 \end{pmatrix}$$

(A.1)

$$R_u(\delta_t) = \begin{pmatrix} 1 & 0 & 0 \\ 0 & \cos \delta_t & -\sin \delta_t \\ 0 & \sin \delta_t & \cos \delta_t \end{pmatrix}$$

(A.2)

The operator $T$ to transform the frame toward the direction of Jupiter at time $t$ is therefore:

$$T_t = R_u(\delta_t)R_u(-\alpha_t) = \begin{pmatrix} \cos \alpha_t & \sin \alpha_t & \cos \delta_t & -\sin \alpha_t & \sin \delta_t & 0 \\ -\sin \alpha_t & \cos \alpha_t & \cos \delta_t & -\sin \alpha_t & \sin \delta_t & 0 \\ 0 & 0 & \cos \delta_t & 0 & \cos \delta_t & \sin \delta_t \end{pmatrix}$$

(A.3)

In addition, it is required to apply a phase correction to the complex visibility data as the plane wave coming from direction $u_0$ should now come in phase from direction $u_t$. This factor is expressed as a function of the transformation:

$$\phi_{\text{cor}}(l,t) = \exp \left( \frac{2\pi}{\lambda} (w_0 - w_t)^T T_t [u_t, v_t, w_t]^T \right)$$

(4.4)

where $w_0$ and $w_t$ are the third column of matrix Eq. A.3 with the corresponding indices.

Appendix A.2: Intrinsic rotation correction

Once the previous phase and axis corrections have been performed, we need to apply a correction on the $(u,v)$ axes to follow the intrinsic oscillation of the radiation belts. By applying a time-dependent rotation of angle $\beta_j(t)$ (with $\beta_j(t) = -111.6^\circ \pm 9.6^\circ$, counting from the increasing declination axis), the mean direction of the apparent magnetic equator on the sky, around the axis defined by $w_j$ by the following transformation:

$$R_u(\beta_{\text{rot}}) = \begin{pmatrix} \cos \beta_{u}(t) & -\sin \beta_{u}(t) & 0 \\ \sin \beta_{u}(t) & \cos \beta_{u}(t) & 0 \\ 0 & 0 & 1 \end{pmatrix}$$

(A.5)

In first approximation, no phase correction is necessary after applying the rotation of Eq. A.2 on the $(u,v,w)$ coordinates.